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Herschel

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OUTLINES

OF

ASTRONOMY.

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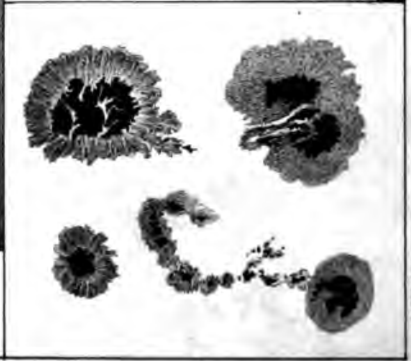

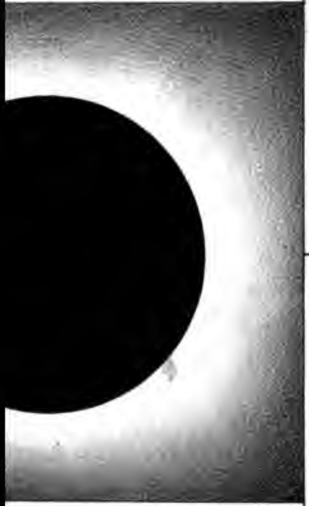

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OUTLINES

OF

ASTRONOMY:

BY

SIR JOHN F. W. HERSCHEL, BART. K.H.

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&c. &c. &c.

A NEW EDITION.

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NOTE

TO

THE FOURTH EDITION.

Several alterations and additions are made in this Edition, besides what have been introduced into the Third, to bring it up to the actual state of astronomical discovery. The elements of four new planets (Parthenope, Egeria, Victoria, and Irene) have been added, and improved elements of Iris, Metis, Hebe, and Hygeia, substituted for the provisional elements before given. The remarkable discovery of an additional ring of Saturn, and the curious researches of M. Peters on the proper motion of Sirius, with several minor features, are noticed. Where such additions are introduced in the text, they are indicated by being enclosed in brackets [ ].

J. F. W. Herschel.

London, Aug. 5, 1851.

PREFACE.

The work here offered to the Public is based upon and may be considered as an extension, and, it is hoped, an improvement of a treatise on the same subject, forming Part 43, of the Cabinet Cyclopædia, published in the year 1833. Its object and general character are sufficiently stated in the introductory chapter of that volume, here reprinted with little alteration; but an opportunity having been afforded me by the Proprietors, preparatory to its re-appearance in a form of more pretension, I have gladly availed myself of it, not only to correct some errors which, to my regret, subsisted in the former volume, but to remodel it altogether (though in complete accordance with its original design as *a work of explanation*); to introduce much new matter in the earlier portions of it; to re-write, upon a far more matured and comprehensive plan, the part relating to the lunar and planetary perturbations, and to bring the subjects of sidereal and nebular astronomy to the level of the present state of our knowledge in those departments.

The chief novelty in the volume, as it now stands, will be found in the manner in which the subject of Perturbations is treated. It is not — it cannot be made *elementary*, in the sense in which that word is understood in these days of light reading. The chapters devoted to it must, therefore, be considered as addressed to a class of readers in possession of somewhat more mathematical knowledge than those who will find the rest of the work readily

r accessible; to readers desirous of preparing them-
the possession of a sort of *carte du pays*, for a cam-
he most difficult, but at the same time the most attract-
e most remunerative of all the applications of modern
More especially they may be considered as addressed
s in that university, where the "Principia" of Newton
r ever will be, put aside as an obsolete book, behind the
where the grand, though rude outlines of the lunar
s delivered in the eleventh section of that immortal
studied less for the sake of the theory itself than for
of far-reaching thought, superior to and disencumbered
al aids, which distinguishes that beyond any other pro-
f the human intellect.

vering a rational as distinguished from a technical expo-
this subject, however, the course pursued by Newton in
n of the Principia alluded to, has by no means been
followed. As regards the perturbations of the nodes
nations, indeed, nothing equally luminous can ever be
d for his explanation. But as respects the other dis-

steps (for the subject) of the most elementary character, to the general formulæ for these perturbations, delivered by Laplace in the Mécanique Céleste.[1]

The reader will find one class of the lunar and planetary inequalities handled in a very different manner from that in which their explanation is usually presented. It comprehends those which are characterized as incident on the epoch, the principal among them being the annual and secular equations of the moon, and that very delicate and obscure part of the perturbational theory (so little satisfactory in the manner in which it emerges from the analytical treatment of the subject), the constant or permanent effect of the disturbing force in altering the disturbed orbit. I will venture to hope that what is here stated will tend to remove some rather generally diffused misapprehensions as to the true bearings of Newton's explanation of the annual equation.[2]

If proof were wanted of the inexhaustible fertility of astronomical science in points of novelty and interest, it would suffice to adduce the addition to the list of members of our system of no less than eight new planets and satellites during the preparation of these sheets for the press. Among them is one whose discovery must ever be regarded as one of the noblest triumphs of theory. In the account here given of this discovery, I trust to have expressed myself with complete impartiality; and in the exposition of the perturbative action on Uranus, by which the existence and situation of the disturbing planet became revealed to us, I have endeavoured, in pursuance of the general plan of this work, rather to exhibit a rational view of the dynamical action, than to convey the slightest idea of the conduct of those masterpieces of analytical skill which the researches of Messrs. Leverrier and Adams exhibit.

To the latter of these eminent geometers, as well as to my excellent and esteemed friend the Astronomer Royal, I have to

[1] Livre ii. chap. viii. art. 67.

[2] Principia, lib. i. prop. 66, cor. 6.

y best thanks for communications which would have
relieved some doubts I at one period entertained, had
eeded in the interim in getting clear of them, as to the
lity of my views on the subject of the annual equation
luded to, with the tenor of Newton's account of it. To
friend, Professor De Morgan, I am indebted for some
nious suggestions on the subject of the mistakes com-
the early working of the Julian reformation of the
of which I should have availed myself, had it not ap-
eferable, on mature consideration, to present the sub-
s simplest form, avoiding altogether entering into mi-
chronological discussion.

J. F. W. Herschel.

ood, April 12, 1849.

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OUTLINES

OF

ASTRONOMY.

INTRODUCTION.

(1.) EVERY student who enters upon a scientific pursuit, especially if at a somewhat advanced period of life, will find not only that he has much to learn, but much also to unlearn. Familiar objects and events are far from presenting themselves to our senses in that aspect and with those connections under which science requires them to be viewed, and which constitute their rational explanation. There is, therefore, every reason to expect that those objects and relations which, taken together, constitute the subject he is about to enter upon will have been previously apprehended by him, at least imperfectly, because much has hitherto escaped his notice which is essential to its right understanding: and not only so, but too often also erroneously, owing to mistaken analogies, and the general prevalence of vulgar errors. As a first preparation, therefore, for the course he is about to commence, he must loosen his hold on all crude and hastily adopted notions, and must strengthen himself, by something of an effort and a resolve, for the unprejudiced admission of any conclusion which shall appear to be supported by careful observation and logical argument, even should it prove of a nature adverse to notions he may have previously formed for himself, or taken up, without examination, on the credit of others. Such an effort is, in fact, a commencement of that intellectual discipline which forms one of the most important ends of all science. It is the first movement of approach towards that state of mental purity which alone can fit us for a full and steady perception of moral beauty as well as physical adaptation. It is the "euphrasy and rue" with which we must "purge our sight" before we can receive and contemplate as they are the lineaments of truth and nature.

(2.) There is no science which, more than astronomy, stands in need of such a preparation, or draws more largely on that intellectual liberality which is ready to adopt whatever is demonstrated, or concede whatever is rendered highly probable, however new and uncommon the points of view may be in which objects the most familiar may thereby become placed. Almost all its conclusions stand in open and striking contradiction with those of superficial and vulgar observation, and with what appears to every one, until he has understood and weighed the proofs to the contrary, the most positive evidence of his senses. Thus, the earth on which he stands, and which has served for ages as the unshaken foundation of the firmest structures, either of art or nature, is divested by the astronomer of its attribute of fixity, and conceived by him as turning swiftly on its centre, and at the same time moving onwards through space with great rapidity. The sun and the moon, which appear to untaught eyes round bodies of no very considerable size, become enlarged in his imagination into vast globes, — the one approaching in magnitude to the earth itself, the other immensely surpassing it. The planets, which appear only as stars somewhat brighter than the rest, are to him spacious, elaborate, and habitable worlds; several of them much greater and far more curiously furnished than the earth he inhabits, as there are also others less so; and the stars themselves, properly so called, which to ordinary apprehension present only lucid sparks or brilliant atoms, are to him suns of various and transcendent glory — effulgent centres of life and light to myriads of unseen worlds. So that when, after dilating his thoughts to comprehend the grandeur of those ideas his calculations have called up, and exhausting his imagination and the powers of his language to devise similes and metaphors illustrative of the immensity of the scale on which his universe is constructed, he shrinks back to his native sphere; he finds it, in comparison, a mere point; so lost — even in the minute system to which it belongs — as to be invisible and unsuspected from some of its principal and remoter members.

(3.) There is hardly any thing which sets in a stronger light the inherent power of truth over the mind of man, when opposed by no motives of interest or passion, than the perfect readiness with which all these conclusions are assented to as soon as their evidence is clearly apprehended, and the tenacious hold they acquire over our belief when once admitted. In the conduct, therefore, of this volume, I shall take it for granted that the reader is more desirous to learn the system which it is its object to teach as it now stands, than to raise or revive objections against it; and that, in short, he comes to the task with a willing mind; an assumption which will not only save the trouble of piling argument on argument to convince the sceptical, but will greatly facilitate his actual progress; inasmuch as he will find it at once easier and more satisfactory to pursue from the outset a straight and definite path, than to be constantly stepping aside, involving himself in perplexities and circuits, which, after all, can only terminate in finding himself compelled to adopt the same road.

(4.) The method, therefore, we propose to follow in this work is neither strictly the analytic nor the synthetic, but rather such a combination of both, with a leaning to the latter, as may best suit with a *didactic* composition. Its object is not to convince or refute opponents, nor to inquire, under the semblance of an assumed ignorance, for principles of which we are all the time in full possession — but simply to *teach* what is *known*. The moderate limit of a single volume, to which it will be confined, and the necessity of being on every point, within that limit, rather diffuse and copious in explanation, as well as the eminently matured and ascertained character of the science itself, render this course both practicable and eligible. Practicable, because there is now no danger of any revolution in astronomy, like those which are daily changing the features of the less advanced sciences, supervening, to destroy all our hypotheses, and throw our statements into confusion. Eligible, because the space to be bestowed, either in combating refuted systems, or in leading the reader forward by slow and measured steps from the known to the unknown, may be more advantageously devoted to such explanatory illustrations as will impress on him a familiar and, as it were, a practical sense of the sequence of phenomena, and the manner in which they are produced. We shall not, then, reject the analytic course where it leads more easily and directly to our objects, or in any way fetter ourselves by a rigid adherence to method. Writing only to be understood, and to communicate as much information in as little space as possible, consistently with its *distinct* and *effectual* communication, no sacrifice can be afforded to system, to form, or to affectation.

(5.) We shall take for granted, from the outset, the Copernican system of the world; relying on the easy, obvious, and natural explanation it affords of all the phenomena as they come to be described, to impress the student with a sense of its truth, without either the formality of demonstration or the superfluous tedium of eulogy, calling to mind that important remark of Bacon: — "Theoriarum vires, arcta et quasi se mutuo sustinente partium adaptatione, quâ quasi in orbem cohærent, firmantur[1];"

---

[1] "The confirmation of theories relies on the compact adaptation of their parts, by which, like those of an arch or dome, they mutually sustain each other, and form a coherent whole." This is what Dr. Whewell expressively terms the *consiliences* of inductions.

not failing, however, to point out to the reader, as occasion offers, the contrast which its superior simplicity offers to the complication of other hypotheses.

(6.) The preliminary knowledge which it is desirable that the student should possess, in order for the more advantageous perusal of the following pages, consists in the familiar practice of decimal and sexagesimal arithmetic; some moderate acquaintance with geometry and trigonometry, both plane and spherical; the elementary principles of mechanics; and enough of optics to understand the construction and use of the telescope, and some other of the simpler instruments. Of course, the more of such knowledge he brings to the perusal, the easier will be his progress, and the more complete the information gained; but we shall endeavour in every case, as far as it can be done without a sacrifice of clearness, and of that useful brevity which consists in the absence of prolixity and episode, to render what we have to say as independent of other books as possible.

(7.) After all, I must distinctly caution such of my readers as may commence and terminate their astronomical studies with the present work (though of such, — at least in the latter predicament, — I trust the number will be few), that its utmost pretension is to place them on the threshold of this particular wing of the temple of Science, or rather on an eminence exterior to it, whence they may obtain something like a general notion of its structure; or, at most, to give those who may wish to enter a ground-plan of its accesses, and put them in possession of the pass-word. Admission to its sanctuary, and to the privileges and feelings of a votary, is only to be gained by one means, — *sound and sufficient knowledge of mathematics, the great instrument of all exact inquiry, without which no man can ever make such advances in this or any other of the higher departments of science as can entitle him to form an independent opinion on any subject of discussion within their range.* It is not without an effort that those who possess this knowledge can communicate on such subjects with those who do not, and adapt their language and their illustrations to the necessities of such an intercourse. Propositions which to the one are almost identical, are theorems of import and difficulty to the other; nor is their evidence presented in the same way to the mind of each. In teaching such propositions, under such circumstances, the appeal has to be made, not to the pure and abstract reason, but to the sense of analogy — to practice and experience: principles and modes of action have to be established not by direct argument from acknowledged axioms, but by continually recurring to the sources from which the axioms themselves have been drawn; viz. examples; that is to say, by bringing forward and dwelling on simple and familiar instances in which the same principles and the same or similar modes of action take place: thus erecting, as it were, in each particular case, a separate induction, and constructing at each step a little body of science to meet its exigencies. The difference is that of pioneering a road through an untraversed country and advancing at ease along a broad and beaten highway; that is to say, if we are determined to make ourselves distinctly understood, and will appeal to reason at all. As for the method of *assertion*, or a direct demand on the *faith* of the student (though in some complex cases indispensable, where illustrative explanation would defeat its own end by becoming tedious and burdensome to both parties), it is one which I shall neither willingly adopt nor would recommend to others.

(8.) On the other hand, although it is something new to abandon the road of mathematical demonstration in the treatment of subjects susceptible of it, and to teach any considerable branch of science entirely or chiefly by the way of illustration and familiar parallels, it is yet not impossible that those who are already well acquainted with our subject, and whose knowledge has been acquired by that confessedly higher practice which is incompatible with the avowed objects of the present work, may yet find their account in its perusal, — for this reason, that it is always of advantage to present any given body of knowledge to the mind in as great a variety of different lights as possible. It is a property of illustrations of this kind to strike no two minds in the same manner, or with the same force; because no two minds are stored with the same images, or have acquired their notions of them by similar habits. Accordingly, it may very well happen, that a proposition, even to one best acquainted with it, may be placed not merely in a new and uncommon, but in a more impressive and satisfactory light by such a course — some obscurity may be dissipated, some inward misgivings cleared up, or even some links supplied which may lead to the perception of connections and deductions altogether unknown before. And the probability of this is increased when, as in the present instance, the illustrations chosen have not been studiously selected from books, but are such as have presented themselves freely to the author's mind as being most in harmony with his own views; by which, of course, he means to lay no claim to originality in all or any of them beyond what they may really possess.

(9.) Besides, there are cases in the application of mechanical principles with which the mathematical student is but too familiar, where, when the data are before him, and the numerical and geometrical relations of his problems all clear to his conception,—when his forces are estimated and his lines measured, — nay, when even he has followed up the application of his technical processes, and fairly arrived at his conclusion, — there is still something wanting in his mind — not in the evidence, for he has examined each link, and finds the chain complete — not in the principles, for those he well knows are too firmly established to be shaken — but precisely in the *mode of action*. He has followed out a train of reasoning by logical and technical rules, but the signs he has employed are not pictures of nature, or have lost their original meaning as such to his mind: he has not seen, as it were, the process of nature passing under his eye in an instant of time, and presented as a consecutive whole to his imagination. A familiar parallel, or an illustration drawn from some artificial or natural process, of which he has that direct and individual impression which gives it a reality and associates it with a name, will, in almost every such case, supply in a moment this deficient feature, will convert all his symbols into real pictures, and infuse an animated meaning into what was before a lifeless succession of words and signs. I cannot, indeed, always promise myself to attain this degree of vividness of illustration, nor are the points to be elucidated themselves always capable of being so *paraphrased* (if I may use the expression) by any single instance adducible in the ordinary course of experience; but the object will at least be kept in view; and, as I am very conscious of having, in making such attempts, gained for myself much clearer views of several of the more concealed effects of planetary perturbation than I had acquired by their mathematical investigation in detail, it may reasonably be hoped that the endeavour will not always be unattended with a similar success in others.

(10.) From what has been said, it will be evident that our aim is not to offer to the public a technical treatise, in which the student of practical or theoretical astronomy shall find consigned the minute description of methods of observation, or the formulæ he requires prepared to his hand, or their demonstrations drawn out in detail. In all these the present work will be found meagre, and quite inadequate to his wants. Its aim is entirely different; being to present in each case the mere ultimate *rationale* of facts, arguments, and processes; and, in all cases of mathematical application, avoiding whatever would tend to encumber its pages with algebraic or geometrical symbols, to place under his inspection that central thread of common sense on which the pearls of analytical research are invariably strung; but which, by the attention the latter claim for themselves, is often concealed from the eye of the gazer, and not always disposed in the straightest and most convenient form to follow by those who string them. This is no fault of those who have conducted the inquiries to which we allude. The contention of mind for which they call is enormous; and it may, perhaps, be owing to their experience of *how little* can be accomplished in carrying such processes on to their conclusion, by mere ordinary *clearness of head;* and how necessary it often is to pay more attention to the purely mathematical conditions which ensure success, — the hooks-and-eyes of their equations and series, — than to those which enchain causes with their effects, and both with the human reason, — that we must attribute something of that indistinctness of view which is often complained of as a grievance by the earnest student, and still more commonly ascribed ironically to the native cloudiness of an atmosphere too sublime for vulgar comprehension. We think we shall render good service to both classes of readers, by dissipating, so far as lies in our power, that accidental obscurity, and by showing ordinary untutored comprehension clearly what it *can*, and what it *cannot*, hope to attain.

CHAPTER I.

GENERAL NOTIONS. — APPARENT AND REAL MOTIONS. — SHAPE AND SIZE OF THE EARTH. — THE HORIZON AND ITS DIP. — THE ATMOSPHERE. — REFRACTION. — TWILIGHT. — APPEARANCES RESULTING FROM DIURNAL MOTION. — FROM CHANGE OF STATION IN GENERAL. — PARALLACTIC MOTIONS. — TERRESTRIAL PARALLAX. — THAT OF THE STARS INSENSIBLE. — FIRST STEP TOWARDS FORMING AN IDEA OF THE DISTANCE OF THE STARS. — COPERNICAN VIEW OF THE EARTH'S MOTION. — RELATIVE MOTION. — MOTIONS PARTLY REAL, PARTLY APPARENT.—GEOCENTRIC ASTRONOMY, OR IDEAL REFERENCE OF PHÆNOMENA TO THE EARTH'S CENTRE AS A COMMON CONVENTIONAL STATION.

(11.) The magnitudes, distances, arrangement, and motions of the great bodies which make up the visible universe, their constitution and physical condition, so far as they can be known to us, with their mutual influences and actions on each other, so far as they can be traced by the effects produced, and established by legitimate reasoning, form the assemblage of objects to which the intention of the astronomer is directed. The term astronomy[1] itself, which denotes the *law* or rule of the *astra* (by which the ancients understood not only the stars properly so called, but the sun, the moon, and all the visible constituents of the heavens), sufficiently indicates this; and, although the term astrology, which denotes the *reason*, *theory*, or *interpretation* of the stars,[2] has become degraded in its application, and confined to superstitious and delusive attempts to divine future events by their dependence on pretended planetary influences, the same meaning originally attached itself to that epithet.

(12.) But, besides the stars and other celestial bodies, the earth itself, regarded as an individual body, is one principal object of the astronomer's consideration, and indeed, the chief of all. It derives its importance, in a practical as well as theoretical sense, not only from its proximity, and its relation to us as animated beings, who draw from it the supply of all our wants, but as the station from which we see all the rest, and as the only one among them to which we can, in the first instance, refer for any determinate marks and measures by which to recognize their changes of situation, or with which to compare their distances.

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[1] Αστηρ, *a star*; νομος, a *law*; or νεμειν, to tend, as a shepherd, his flock; so that αστρονομος means "shepherd of the stars." The two two etymologies are, however, coincident.

[2] Λογος, *reason*, or *a word*, the vehicle of reason; the interpreter of thought.

(13.) To the reader who now for the first time takes up a book on astronomy, it will no doubt seem strange to class the earth with the heavenly bodies, and to assume any community of nature among things apparently so different. For what, in fact, can be more apparently different than the vast and seemingly immeasurable extent of the earth, and the stars? The earth is a dark and opaque, while the celestial bodies are brilliant. We perceive in it no motion, while in them we observe a continual change of place, as we view them at different hours of the day or night, or at different seasons of the year. The ancients, accordingly, one or two of the more enlightened of them only excepted, admitted no such community of nature; and, by thus placing the heavenly bodies and their movements without the pale of analogy and experience, effectually intercepted the progress of all reasoning from what passes here below, to what is going on in the regions where they exist and move. Under such conventions, astronomy, as a science of cause and effect, could not exist, but must be limited to a mere registry of appearances, unconnected with any attempt to account for them on reasonable principles, however successful to a certain extent might be the attempt to follow out their order of sequence, and to establish empirical laws expressive of this order. To get rid of this prejudice, therefore, is the first step towards acquiring a knowledge of what is really the case; and the student has made his first effort towards the acquisition of sound knowledge, when he has learnt to familiarize himself with the idea that the earth, after all, *may* be nothing but a great star. How correct such an idea may be, and with what limitations and modifications it is to be admitted, we shall see presently.

(14.) It is evident, that, to form any just notions of the arrangement, in space, of a number of objects which we cannot approach and examine, but of which all the information we can gain is by sitting still and watching their evolutions, it must be very important for us to know, in the first instance, whether what we call sitting still is *really* such: whether the station from which we view them, with ourselves, and all objects which immediately surround us, be not itself in motion, unperceived by us; and if so, of what nature that motion is. The apparent places of a number of objects, and their apparent arrangement with respect to each other, will of course be materially dependent on the situation of the spectator among them; and if this situation be liable to change, unknown to the spectator himself, an appearance of change in the respective situations of the objects will arise, without the reality. If, then, such be actually the case, it will follow that *all* the movements we think we perceive among the stars will not be real movements, but that some part, at least, of whatever changes of relative place we perceive among them must be merely apparent, the results of the shifting of our own point of view; and that, if we would ever arrive at a knowledge of their real motions, it can only be by first investigating our own, and making due allowance for its effects. Thus, the question whether the earth is in motion or at rest, and if in motion, what that motion is, is no idle inquiry, but one on which depends our only chance of arriving at true conclusions respecting the constitution of the universe.

(15.) Nor let it be thought strange that we should speak of a motion existing in the earth, unperceived by its inhabitants; we must remember that it is of the earth *as a whole,* with all that it holds within its substance or sustains on its surface, that we are speaking; of a motion common to the solid mass beneath, to the ocean which flows around it, the air that rests upon it, and the clouds which float above it in the air. Such a motion, which should displace no terrestrial object from its relative situation among others, interfere with no natural processes, and produce no sensations of shocks or jerks, might, it is very evident, subsist undetected by us. There is no peculiar sensation which advertises us that we are *in motion.* We perceive *jerks,* or *shocks,* it is true, because these are sudden *changes* of motion, produced, as the laws of mechanics teach us, by sudden and powerful forces acting during short times; and these forces, applied to our bodies, are what we *feel.* When, for example, we are carried along in a carriage with the blinds down, or with our eyes closed (to keep us from seeing external objects), we perceive a tremor arising from inequalities in the road, over which the carriage is successively lifted and let fall, but we have no sense of *progress.* As the road is smoother, our sense of motion is diminished, though our rate of travelling is accelerated. Railway travelling, especially by night, or in a tunnel, has familiarized every one with this remark. Those who have made aeronautic voyages testify that with closed eyes, and under the influence of a steady breeze communicating no oscillatory or revolving motion to the car, the *sensation* is that of perfect rest, however rapid the transfer from place to place.

(16.) But it is on shipboard, where a great system is maintained in motion, and where we are surrounded with a multitude of objects which participate with ourselves and each other in the common progress of the whole mass, that we feel most satisfactorily the identity of sensation between a state of motion and one of rest. In the cabin of a large and heavy vessel, going smoothly before the wind in still water, or drawn along a canal, not the smallest indication acquaints us with the way it is making. We read, sit, walk, and perform every customary action as if we were on land. If we throw a ball into the air, it falls back into our hand; or if we drop it, it alights at our feet. Insects buzz around us as in the free air; and smoke ascends in the same manner as it would do in an apartment on shore. If, indeed, we come on deck, the case is, in some respects, different; the air, not being carried along with us, drifts away smoke and other light bodies — such as feathers abandoned to it — apparently, in the opposite direction to that of the ship's progress; but, in reality, *they* remain at rest, and we leave them behind in the air. Still, the illusion, so far as massive objects and our own movements are concerned, remains complete; and when we look at the shore, we then perceive the effect of our own motion transferred, in a contrary direction, to external objects — *external*, that is, *to the system of which we form a part.*

"Provehimur, portu, terraque urbesque recedunt.

(17.) In order, however, to conceive the earth as in motion, we must form to ourselves a conception of its shape and size. Now, an object cannot have shape and size unless it is *limited* on all sides by some definite outline, so as to admit of our imagining it, at least, disconnected from other bodies, and existing insulated in space. The first rude notion we form of the earth is that of a flat surface, of indefinite extent in all directions from the spot where we stand, *above* which are *the air* and *sky;* below, to an indefinite profundity, solid matter. This is a prejudice to be got rid of, like that of the earth's immobility; — but it is one much easier to rid ourselves of, inasmuch as it originates only in our own mental inactivity, in not questioning ourselves *where* we will place a limit to a thing we have been accustomed from infancy to regard as immensely large; and does not, like that, originate in the testimony of our senses unduly interpreted. On the contrary, the direct testimony of our senses lies the other way. When we see the sun set in the evening in the west, and rise again in the east, as we cannot doubt that it is the *same* sun we see after a temporary absence, we must do violence to all our notions of solid matter, to suppose it to have made its way *through* the substance of the earth. It must, therefore, have gone *under* it, and that not by a mere subterraneous *channel;* for if we notice the points where it sets and rises for many successive days, or for a whole year, we shall find them constantly shifting, round a very large extent of the horizon; and, besides, the moon and stars also set and rise again in *all* points of the visible horizon. The conclusion is plain: the earth cannot extend indefinitely in depth downwards, nor indefinitely in surface laterally; it must have not only bounds in a horizontal direction, but also an *under side* round which the sun, moon, and stars can pass; and that side must, at least, be so far like what we see, that it must have a sky and sunshine, and a day when it is night to us, and *vice versâ;* where, in short,

> —"redit à nobis Aurora, diemque reducit.
> Nosque ubi primus equis oriens afflavit anhelis,
> Illic sera rubens accendit lumina Vesper. *Georg.*

(18.) As soon as we have familiarized ourselves with the conception of an earth without *foundations* or fixed supports—existing insulated in space from contact of every thing external, it becomes easy to imagine it in motion—or, rather, difficult to imagine it otherwise: for, since there is nothing to *retain* it in one place, should any causes of motion exist, or any *forces* act upon it, it must obey their impulse. Let us next see what obvious circumstances there are to help us to a knowledge of the *shape* of the earth.

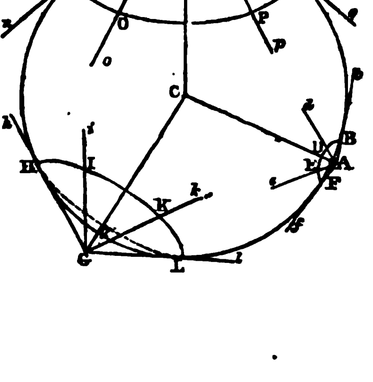
(19.) Let us first examine what we can actually *see* of its shape. Now, it is not on land (unless, indeed, on uncommonly level and extensive plains), that we can see any thing of the *general* figure of the earth;—the hills, trees, and other objects which roughen its surface, and break and elevate the line of the horizon, though obviously bearing a most minute proportion to the *whole* earth, are yet too considerable with respect to ourselves and to that small portion of it which we can see at a single view, to allow of our forming any judgment of the form of the whole, from that of a part so disfigured. But with the surface of the sea or any vastly extended level plain, the case is otherwise. If we sail out of sight of land, whether we stand on the deck of the ship or climb the mast, we see the surface of the sea—not losing itself in distance and mist, but terminated by a sharp, clear, well-defined line or *offing* as it is called, which runs all round us in a circle, having our station for its centre. That this line is really a circle, we conclude, first, from the perfect apparent similarity of all its parts; and, secondly, from the fact of all its parts appearing at the same distance from us, and that, evidently, a moderate one; and thirdly, from this, that its apparent *diameter*, measured with an instrument called the *dip sector*, is the same (except under some singular atmospheric circumstances, which produce a temporary distortion of the outline), in whatever direction the measure is taken,—properties which belong only to the circle among geometrical figures. If we ascend a high eminence on a plain (for instance, one of the Egyptian pyramids,) the same holds good.

(20.) Masts of ships, however, and the edifices erected by man, are trifling eminences compared to what nature itself affords; Ætna, Teneriffe, Mowna Roa, are eminences from which no contemptible *aliquot* part of the whole earth's surface can be seen; but from these again — in those few and rare occasions when the transparency of the air will permit the real boundary of the horizon, the true sea-line, to be seen — the very same appearances are witnessed, but with this remarkable addition, viz., that the angular *diameter* of the visible area, as measured by the dip sector, is materially *less* than at a lower level; or, in other words, that the *apparent size* of the earth has sensibly diminished as we have receded from its surface, while yet the *absolute quantity* of it seen at once has been increased.

(21.) The same appearances are observed universally, in every part of the earth's surface visited by man. Now, the figure of a body which, however seen, appears always *circular*, can be no other than a sphere or globe.

(22.) A diagram will elucidate this. Suppose the earth to be represented by the sphere LHNQ, whose centre is C, and let A, G, M be stations at different elevations above various points of its surface, represented

Fig. 1.



by *a*, *g*, *m* respectively. From each of them (as from M) let a line be drawn, as MN*n*, a tangent to the surface at N, then will this line represent the visual ray along which the spectator at M will see the visible horizon; and as this tangent sweeps round M, and comes successively into the positions MO*o*, MP*p*, MQ*q*, the point of contact N will mark out on the surface the circle NOPQ. The area of the spherical surface comprehended within this circle is the portion of the earth's surface visible to a spectator at M, and the angle NMQ included between the two extreme visual rays is the measure of its apparent angular diameter. Leaving, at present, out of consideration the effect of refraction in the air below M, of which more hereafter, and which always tends, in some degree, to *increase* that angle, or render it more *obtuse*, this is the angle measured by the dip sector. Now, it is evident, 1st, that as the point M is more elevated above *m*, the point immediately below it on the sphere, the visible area, *i. e.* the spherical segment or slice NOPQ, increases; 2dly, that the distance of the visible *horizon*[1] or boundary of our view from the eye, viz. the line MN, increases; and, 3dly, that the angle NMQ becomes *less obtuse*, or, in other words, the apparent angular diameter of the earth diminishes, being nowhere so great as 180°, or two right angles, but falling short of it by some sensible quantity, and that more and more the higher we ascend. The figure exhibits three states or stages of elevation, with the horizon, &c., corresponding to each, a glance at which will explain our meaning; or, limiting ourselves to the larger and more distinct, MNOPQ, let the reader imagine *n*NM, MQ*q* to be the two legs of a ruler jointed at M, and kept extended by the globe N*m*Q between them. It is clear, that as the joint M is urged home towards the surface, the legs will open, and the ruler will become more nearly *straight*, but will not attain *perfect* straightness till M is brought fairly up to contact with the surface at *m*, in which case its whole length will become a *tangent* to the sphere at *m*, as is the line *xy*.

(23.) This explains what is meant by the *dip of the horizon*. M*m*, which is perpendicular to the general surface of the sphere at *m*, is also the direction in which a *plumb-line*[2] would hang; for it is an observed fact, that in all situations, in every part of the earth, the direction of a plumb-line is exactly perpendicular to the surface of still water; and, moreover, that it is also exactly perpendicular to a line or surface truly adjusted by a *spirit-level*.[2] Suppose, then, that at our station M we were to adjust a line (a wooden ruler for instance) by a spirit-level, with perfect exactness; then, if we suppose the direction of this line indefinitely prolonged both ways, as XMY, the line so drawn will be at right angles to M*m*, and therefore parallel to *xmy*, the tangent to the sphere at *m*. A spectator placed at M will therefore see not only all the vault of the sky *above* this line, as XZY, but also that portion or zone of it which lies between XN and YQ; in other words, his sky will be more than a hemisphere by the zone YQXN. It is the angular breadth of this redundant zone — the angle YMQ, by which the *visible* horizon appears depressed below the direction of a spirit-level — that is called the *dip of the horizon*. It is a correction of constant use in nautical astronomy.

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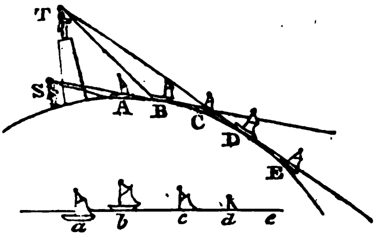
[1] ‘Οριζω, *to terminate.*

[2] See these instruments described in Chap. III.

(24.) From the foregoing explanations it appears, 1st, That the general figure of the earth (so far as it can be gathered from this kind of observation) is that of a sphere or globe. In this we also include that of the sea, which, wherever it extends, covers and fills in those inequalities and local irregularities which exist on land, but which can of course only be regarded as trifling deviations from the general outline of the whole mass, as we consider an orange not the less round for the roughness on its rind. 2dly, That the appearance of a *visible* horizon, or sea-offing, is a consequence of the curvature of the surface, and does not arise from the inability of the eye to follow objects to a greater distance, or from atmospheric indistinctness. It will be worth while to pursue the general notion thus acquired into some of its consequences, by which its consistency with observations of a different kind, and on a larger scale, will be put to the test, and a clear conception be formed of the manner in which the parts of the earth are related to each other, and held together as a whole.

(25.) In the first place, then, every one who has passed a little while at the sea-side is aware that objects may be seen perfectly well beyond the *offing* or visible horizon — but not the *whole* of them. We only see their upper parts. Their bases where they rest on, or rise out of the water, are hid from view by the spherical surface of the sea, which protrudes between them and ourselves. Suppose a ship, for instance, to sail directly away from our station; — at first, when the distance of the ship is small, a spectator, S, situated at some certain height above the sea, sees the whole of the ship, even to the *water line* where it rests on the sea, as at A. As it recedes it diminishes, it is true, in apparent size, but still the *whole* is seen down to the water line, till it reaches the *visible* horizon at B. But as soon as it has passed this distance, not only does the visible portion still continue to diminish in apparent *size*, but the hull begins to disappear bodily, as if sunk below the surface. When it has reached a certain distance, as at C, its hull has entirely vanished, but the masts and sails remain, presenting the appearance *c*. But if, in this state of things, the spectator quickly ascends to a higher station, T, whose visible horizon is at D, the hull comes again in sight; and, when he descends again he loses it. The ship still receding, the lower sails seem to sink below the water, as at *d*, and at length the whole disappears: while yet the distinctness with which the last portion of the sail *d* is seen is such as to satisfy us that were it not for the interposed segment of the sea, ABCDE, the distance TE is not so great as to have prevented an equally perfect view of the whole.

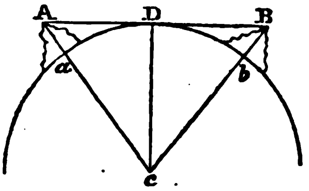
Fig. 2.



(26.) The history of aëronautic adventure affords a curious illustration of the same principle. The late Mr. Sadler, the celebrated aëronaut, ascended on one occasion in a balloon from Dublin, and was wafted across the Irish Channel, when, on his approach to the Welsh coast, the balloon descended nearly to the surface of the sea. By this time the sun was set, and the shades of evening began to close in. He threw out nearly all his ballast, and suddenly sprang upwards to a great height, and by so doing brought his horizon to *dip* below the sun, producing the whole phenomenon of a western sunrise. Subsequently descending in Wales, he of course witnessed a second sunset on the same evening.

(27.) If we could measure the heights and exact distance of two stations which could barely be discerned from each other over the edge of the horizon, we could ascertain the actual size of the earth itself: and, in fact, were it not for the effect of refraction, by which we are enabled to see in some small degree *round* the interposed segment (as will be hereafter explained), this would be a tolerably good method of ascertaining it. Suppose A and B to be two eminences, whose perpendicular heights A *a* and B *b* (which, for simplicity, we will suppose to be exactly equal) are known, as well as their exact horizontal interval *a* D *b*, by measurement; then it is clear that D, the visible horizon of both, will lie just half-way between them, and if we suppose *a* D *b* to be the sphere of the earth, and C its centre in the figure C D *b* B, we know D *b*, the length of the arch of the circle between D and *b*,—viz. half the measured interval, and b B, the excess of its secant above its radius — which is the height of B, — data which, by the solution of an easy geometrical problem, enable

Fig. 3.



us to find the length of the radius D C. If, as is really the case, we suppose both the heights and distance of the stations inconsiderable in comparison with the size of the earth, the solution alluded to is contained in the following proposition: —

The earth's diameter bears the same proportion to the distance of the visible horizon from the eye as that distance does to the height of the eye above the sea level.

When the stations are unequal in height, the problem is a little more complicated.

(28.) Although, as we have observed, the effect of refraction prevents this from being an exact method of ascertaining the dimensions of the earth, yet it will suffice to afford such an approximation to it as shall be of use in the present stage of the reader's knowledge, and help him to many just conceptions, on which account we shall exemplify its application in numbers. Now, it appears by observation, that two points, each ten feet above the surface, cease to be visible from each other over still water, and in average atmospheric circumstances, at a distance of about 8 miles. But 10 feet is the 528th part of a mile, so that half their distance, or 4 miles, is to the height of each as 4 × 528 or 2112 : 1, and therefore in the same proportion to 4 miles is the length of the earth's diameter. It must, therefore, be equal to 4 × 2112 = 8448, or, in round numbers, about 8000 miles, which is not very far from the truth.

(29.) Such is the first rough result of an attempt to ascertain the earth's magnitude; and it will not be amiss, if we take advantage of it to compare it with objects we have been accustomed to consider as of vast size, so as to interpose a few steps between it and our ordinary ideas of dimension. We have before likened the inequalities on the earth's surface, arising from mountains, valleys, buildings, &c. to the roughnesses on the rind of an orange, compared with its general mass. The comparison is quite free from exaggeration. The highest mountain known hardly exceeds five miles in perpendicular elevation: this is only one 1600th part of the earth's diameter; consequently, on a globe of sixteen inches in diameter, such a mountain would be represented by a protuberance of no more than one hundredth part of an inch, which is about the thickness of ordinary drawing-paper. Now, as there is no entire continent, or even any very extensive tract of land, known, whose general elevation above the sea is any thing like half this quantity, it follows, that if we would construct a correct model of our earth, with its seas, continents, and mountains, on a globe sixteen inches in diameter, the whole of the land, with the exception of a few prominent points and ridges, must be comprised on it within the thickness of thin writing-paper; and the highest hills would be represented by the smallest visible grains of sand.

(30.) The deepest mine existing does not penetrate half a mile below the surface: a scratch, or pin-hole, duly representing it, on the surface of such a globe as our model, would be imperceptible without a magnifier.

(31.) The greatest depth of sea, probably, does not very much exceed the greatest elevation of the continents; and would, of course, be represented by an excavation, in about the same proportion, into the substance of the globe: so that the ocean comes to be conceived as a mere film of liquid, such as, on our model, would be left by a brush dipped in colour, and drawn over those parts intended to represent the sea: only, in so conceiving it, we must bear in mind that the resemblance extends no farther than to proportion in point of quantity. The mechanical laws which would regulate the distribution and movements of such a film, and its adhesion to the surface, are altogether different from those which govern the phenomena of the sea.

(32.) Lastly, the greatest extent of the earth's surface which has ever been seen at once by man, was that exposed to the view of MM. Biot and Gay-Lussac, in their celebrated aëronautic expedition to the enormous height of 25,000 feet, or rather less than five miles. To estimate the proportion of the area visible from this elevation to the whole earth's surface, we must have recourse to the geometry of the sphere, which informs us that the convex surface of a spherical segment is to the whole surface of the sphere to which it belongs as the versed sine or thickness of the segment is to the diameter of the sphere; and further, that this thickness, in the case we are considering, is almost exactly equal to the perpendicular elevation of the point of sight above the surface. The proportion, therefore, of the visible area, in this case, to the whole earth's surface, is that of five miles to 8000, or 1 to 1600. The portion visible from Ætna, the Peak of Teneriffe, or Mowna Roa, is about one 4000th.

(33.) When we ascend to any very considerable elevation above the surface of the earth, either in a balloon, or on mountains, we are made aware, by many uneasy sensations, of an insufficient supply of *air*. The barometer, an instrument which informs us of the weight of air incumbent on a given horizontal surface, confirms this impression, and affords a direct measure of the rate of diminution of the quantity of air which a given space includes as we recede from the surface. From its indications we learn, that when we have ascended to the height of 1000 feet, we have left below us about one-thirtieth of the whole mass of the atmosphere:—that at 10,600 feet of perpendicular elevation (which is rather less than that of the summit of Ætna[1]) we have ascended through about one-third; and at 18,000 feet (which is nearly that of Cotopaxi) through one-half the material, or, at least, the ponderable body of air incumbent on the earth's surface. From the progression of these numbers, as well as *à priori*, from the nature of the air itself, which is *compressible*, i. e. capable of being condensed or crowded into a smaller space in proportion to the incumbent pressure, it is easy to see that, although by rising still higher, we should continually get above more and more of the air, and so relieve ourselves more and more from the pressure with which it weighs upon us, yet the amount of this additional relief, or the *ponderable quantity* of air surmounted, would be by no means in proportion to the additional height ascended, but in a constantly decreasing ratio. An easy calculation, however, founded on our experimental knowledge of the properties of air, and the mechanical laws which regulate its dilatation and compression, is sufficient to show that, at an altitude above the surface of the earth not exceeding the hundredth part of its diameter, the tenuity, or rarefaction, of the air must be so excessive, that not only animal life could not subsist, or combustion be maintained in it, but that the most delicate means we possess of ascertaining the existence of *any air at all* would fail to afford the slightest perceptible indications of its presence.

(34.) Laying out of consideration, therefore, at present, all nice questions as to the probable existence of a definite limit to the atmosphere, beyond which there is, absolutely and rigorously speaking, *no* air, it is clear, that, for all practical purposes, we may speak of those regions which are more distant above the earth's surface than the hundredth part of its diameter as void of air, and of course of clouds (which are nothing but visible vapours, diffused and *floating* in the air, sustained by it, and rendering it *turbid* as mud does water). It seems probable, from many indications, that the greatest height at which visible clouds *ever exist* does not exceed ten miles; at which height the density of the air is about an eighth part of what it is at the level of the sea.

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[1] The height of Ætna above the Mediterranean (as it results from a barometrical measurement of my own, made in July, 1824, under very favourable circumstances) is 10,872 English feet.—*Author.*

(35.) We are thus led to regard the atmosphere of air, with the clouds it supports, as constituting a coating of equable or nearly equable thickness, enveloping our globe on all sides; or rather as an aërial ocean, of which the surface of the sea and land constitutes the bed, and whose inferior portions or strata, within a few miles of the earth, contain by far the greater part of the whole mass, the density diminishing with extreme rapidity as we recede upwards, till, within a very moderate distance (such as would be represented by the sixth of an inch on the model we have before spoken of, and which is not more in proportion to the globe on which it rests, than the downy skin of a peach in comparison with the fruit within it), all sensible trace of the existence of air disappears.

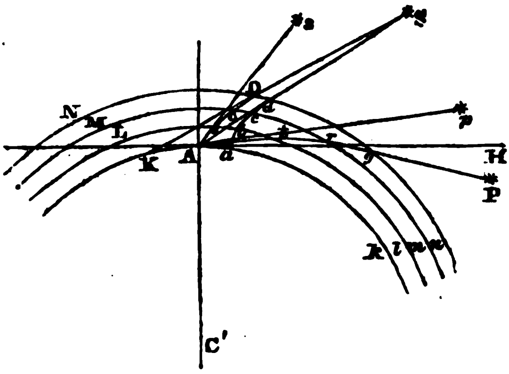
(36.) Arguments, however, are not wanting to render it, if not absolutely certain, at least in the highest degree probable, that the surface of the aërial, like that of the aqueous ocean, has a real and definite limit, as above hinted at; beyond which there is positively *no* air, and above which a fresh quantity of air, could it be added from without, or carried aloft from below, instead of dilating itself indefinitely upwards, would, after a certain very enormous but still finite enlargement of volume, sink and merge, as water poured into the sea, and distribute itself among the mass beneath. With the truth of this conclusion, however, astronomy has little concern; all the effects of the atmosphere in modifying astronomical phenomena being the same, whether it be supposed of definite extent or not.

(37.) Moreover, whichever idea we adopt, within those limits in which it possesses any appreciable density its constitution is the same over all points of the earth's surface; that is to say, on the great scale, and leaving out of consideration temporary and local causes of derangement, such as winds, and great fluctuations, of the nature of waves, which prevail in it to an immense extent. In other words, the law of diminution of the air's density as we recede upwards *from the level of the sea* is the same in every column into which we may conceive it divided, or from whatever point of the surface we may set out. It may therefore be considered as consisting of successively superposed strata or layers, each of the form of a spherical shell, concentric with the general surface of the sea and land, and each of which is *rarer*, or specifically lighter, than that immediately beneath it; and *denser*, or specifically heavier, than that immediately above it. This, at least, is the kind of distribution which alone would be consistent with the laws of the equilibrium of fluids. Inasmuch, however, as the atmosphere is not in perfect equilibrium, being always kept in a state of circulation, owing to the excess of heat in its equatorial regions over that at the poles, some slight deviation from the rigorous expression of this law takes place, and in peculiar localities there is reason to believe that even considerable permanent depressions of the contours of these strata, below their general or spherical level, subsist. But these are points of consideration rather for the meteorologist than the astronomer. It must be observed, moreover, that with this distribution of its strata the inequalities of mountains and valleys have little concern. These exercise hardly more influence in modifying their general spherical figure than the inequalities at the bottom of the sea interfere with the general sphericity of its surface. They would exercise absolutely none were it not for their effect in giving another than horizontal direction to the currents of air constituting winds, as shoals in the ocean throw up the currents which sweep over them towards the surface, and so in some small degree tend to disturb the perfect level of that surface.

(38.) It is the power which air possesses, in common with all transparent media, of *refracting* the rays of light, or bending them out of their straight course, which renders a knowledge of the constitution of the atmosphere important to the astronomer. Owing to this property, objects seen obliquely through it appear otherwise situated than they would to the same spectator, had the atmosphere no existence. It thus produces a false impression respecting their places, which must be rectified by ascertaining the amount and direction of the displacement so apparently produced on each, before we can come at a knowledge of the true directions in which they are situated from us at any assigned moment.

(39.) Suppose a spectator placed at A, any point of the earth's surface K A *k*; and let L *l*, M *m*, N *n*, represent the successive strata or layers, of decreasing density, into which we may conceive the atmosphere to be divided, and which are spherical surfaces concentric with K *k*, the earth's surface. Let S represent a star, or other heavenly body, beyond the utmost limit of the atmosphere. Then, if the air were away, the spectator would see it in the direction of the straight line A S. But, in reality, when the ray of light S A reaches the atmosphere, suppose at *d*, it will, by the laws of optics, begin to bend *downwards*, and take a more inclined direction, as *d c*. This bending will at first be imperceptible, owing to the extreme tenuity of the uppermost strata; but as it advances downwards, the strata continually increasing in density, it will continually undergo greater and greater *refraction* in the same direction; and thus, instead of pursuing the straight line S *d* A, it will describe a curve S *d c b a*, continually more and more concave downwards, and will reach the earth, not at A, but at a certain point *a*, nearer to S. *This* ray, consequently, will not reach the spectator's eye. The ray by which he will see the star is, therefore, not S *d* A, but another ray which, had there been no atmosphere, would have struck the earth at K, a point *behind* the spectator; but which, being bent by the air into the curve S D C B A, actually strikes on A. Now, it is a law of optics, that an object is seen in the direction which the visual ray has at the instant of *arriving at the eye*, without regard to what may have been otherwise its course between the object and the eye. Hence the star S will be seen, not in the direction A S, but in that of A *s*, a *tangent* to the curve S D C B A, at A.

Fig. 4.



But because the curve described by the refracted ray is concave downwards, the tangent A *s* will lie *above* A S, the unrefracted ray: consequently the object S will appear more elevated above the horizon A H, when seen through the refracting atmosphere, than it would appear were there no such atmosphere. Since, however, the disposition of the strata is the same in all directions around A, the visual ray will not be made to deviate *laterally*, but will remain constantly in the same vertical plane, S A C′, passing through the eye, the object, and the earth's centre.

(40.) The effect of the air's refraction, then, is to *raise* all the heavenly bodies higher above the horizon in appearance than they are in reality. Any such body, situated actually *in* the true horizon, will appear *above* it, or will have some certain apparent *altitude* (as it is called). Nay, even some of those actually below the horizon, and which would therefore be invisible but for the effect of refraction, are, by that effect, raised above it and brought into sight. Thus, the sun, when situated at P below the true horizon, A H, of the spectator, becomes visible to him, as if it stood at p, by the refracted ray P q r t A, to which A p is a tangent.

(41.) The exact estimation of the amount of atmospheric refraction, or the strict determination of the angle S A s, by which a celestial object at any assigned altitude, H A S, is raised in appearance above its true place, is, unfortunately, a very difficult subject of physical inquiry, and one on which geometers (from whom alone we can look for any information on the subject) are not yet entirely agreed. The difficulty arises from this, that the *density* of any stratum of air (on which its refracting power depends) is affected not *merely* by the superincumbent pressure, but also by its *temperature* or degree of heat. Now, although we know that as we recede from the earth's surface the temperature of the air is constantly diminishing, yet the *law*, or amount of this diminution at different heights, is not yet fully ascertained. Moreover, the refracting power of air is perceptibly affected by its *moisture;* and this, too, is not the same in every part of an aërial column; neither are we acquainted with the laws of its distribution. The consequence of our ignorance on these points is to introduce a corresponding degree of uncertainty into the determination of the amount of refraction, which affects, to a certain appreciable extent, our knowledge of several of the most important *data* of astronomy. The uncertainty thus induced, is, however, confined within such very narrow limits as to be no cause of embarrassment, except in the most delicate inquiries, and to call for no further allusion in a treatise like the present.

(42.) A "Table of Refraction," as it is called, or a statement of the amount of apparent displacement arising from this cause, at all altitudes, or in every situation of a heavenly body, from the horizon to the *zenith*,[1] or point of the sky vertically above the spectator, and, under all the circumstances in which astronomical observations are usually performed which may influence the result, is one of the most important and indispensable of all astronomical tables, since it is only by the use of such a table we are enabled to get rid of an illusion which must otherwise pervert all our notions respecting the celestial motions. Such have been, accordingly, constructed with great care, and are to be found in every collection of astronomical tables. Our design, in the present treatise, will not admit of the introduction of tables; and we must, therefore, content ourselves here, and in similar cases, with referring the reader to works especially destined to furnish these useful aids to calculation. It is, how-

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[1] From an Arabic word of this signification. See this term technically defined in Chap. II.

le that he should bear in mind the following general notions
nt, and law of variations.

In the *zenith* there is no refraction. A celestial object,
ically over-head, is seen in its true direction, as if there were
re, at least if the air be tranquil.

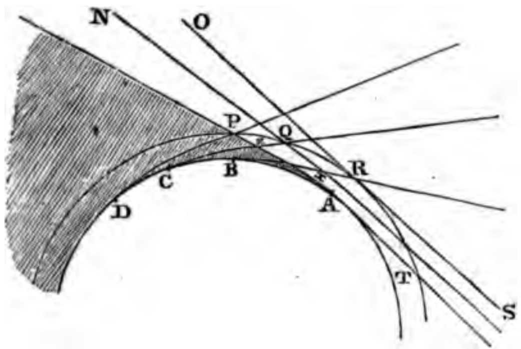
descending from the *zenith* to the horizon, the refraction con-
eases. Objects near the horizon appear more elevated by it
true directions than those at a high altitude.

e *rate* of its increase is nearly in proportion to the tangent of
t angular distance of the object from the zenith. But this
is not far from the truth, at moderate *zenith distances*, ceases
rrect results in the vicinity of the horizon, where the law
ch more complicated in its expression.

e average amount of refraction, for an object half-way between
nd the horizon, or at an apparent altitude of 45°, is about 1′
ly 57″), a quantity hardly sensible to the naked eye; but at
orizon it amounts to no less a quantity than 33′, which is
than the greatest apparent diameter of either the sun or the
nce it follows, that when we see the lower edge of the sun or
pparently resting on the horizon, its whole disk is in reality

would not otherwise reach our eyes is intercepted in its course, and thrown back, or laterally, upon us, becomes to us a means of illumination. Such reflective obstacles always exist floating in the air. The whole course of a sun-beam penetrating through the chink of a window-shutter into a dark room is *visible* as a bright line in the air; and even if it be stifled, or *let out* through an opposite crevice, the light scattered through the apartment from this source is sufficient to prevent entire darkness in the room. The luminous lines occasionally seen in the air, in a sky full of partially broken clouds, which the vulgar term "the sun drawing water," are similarly caused. They are sunbeams, through apertures in the clouds, partially intercepted and reflected on the dust and vapours of the air below. Thus it is with those solar rays which, after the sun is itself concealed by the convexity of the earth, continue to traverse the higher regions of the atmosphere above our heads, and pass through and out of it, without directly striking on the earth at all. Some portion of them is intercepted and reflected by the floating particles above mentioned, and thrown back, or laterally, so as to reach us, and afford us that secondary illumination,

Fig. 5.



which is twilight. The course of such rays will be immediately understood from the annexed figure, in which A B C D is the earth; A a point on its surface, where the sun S is in the act of setting; its last lower ray S A M just grazing the surface at A, while its superior rays S N, S O, traverse the atmosphere above A without striking the earth, leaving it finally at the points P Q R, after being more or less bent in passing through it, the lower most, the higher less, and that which, like S R O, merely grazes the exterior limit of the atmosphere, not at all. Let us consider several points, A, B, C, D, each more remote than the last from A, and each more

lved in the *earth's shadow*, which occupies the whole space
eath the line A M. Now, A just receives the sun's last direct
esides, is illuminated by the whole reflective atmosphere PQ
erefore receives twilight from the whole sky. The point B,
e sun has set, receives no direct solar light, nor any, direct or
om all that part of *its* visible atmosphere which is below A P M;
e lenticular portion P R x, which is traversed by the sun's rays,
lies above the visible horizon B R of B, it receives a twilight,
ongest at R, the point immediately below which the sun is,
way gradually towards P, as the luminous part of the atmo-
s off. At C, only the last or thinnest portion, P Q z of the
gment, thus illuminated, lies above the horizon, C Q, of that
, then, the twilight is feeble, and confined to a small space in
e horizon, which the sun has quitted, while at D the twilight
ltogether.

hen the sun is above the horizon, it illuminates the atmosphere
and these again disperse and scatter a portion of its light in
s, so as to send some of its rays to every exposed point, from
of the sky. The generally diffused light, therefore, which we
e daytime, is a phenomenon originating in the very same causes

mits of mountains seen from the plains, or other terrestrial stations at different levels seen from each other, immersed in it. Every difference of level, accompanied, as it must be, with a difference of density in the aërial strata, must also have, corresponding to it, a certain amount of refraction; less, indeed, than what would be produced by the *whole* atmosphere, but still often of very appreciable, and even considerable, amount. This refraction between terrestrial stations is termed *terrestrial refraction*, to distinguish it from that total effect which is only produced on celestial objects, or such as are beyond the atmosphere, and which is called celestial or astronomical refraction.

(47.) Another effect of refraction is to distort the visible forms and proportions of objects seen near the horizon. The sun, for instance, which at a considerable altitude always appears round, assumes, as it approaches the horizon, a flattened or oval outline; its horizontal diameter being visibly greater than that in a vertical direction. When very near the horizon, this flattening is evidently more considerable on the lower side than on the upper; so that the apparent form is neither circular nor elliptic, but a species of oval, which deviates more from a circle below than above. This singular effect, which any one may notice in a fine sunset, arises from the rapid rate at which the refraction increases in approaching the horizon. Were every visible point in the sun's circumference equally raised by refraction, it would still appear circular, though displaced; but the lower portions being *more* raised than the upper, the vertical diameter is thereby shortened, while the two extremities of its horizontal diameter are equally raised, and in parallel directions, so that its apparent length remains the same. The dilated size (generally) of the sun or moon, when seen near the horizon, beyond what they appear to have when high up in the sky, has nothing to do with refraction. It is an illusion of the judgment, arising from the terrestrial objects interposed, or placed in close comparison with them. In that situation we view and judge of them as we do of terrestrial objects — in detail, and with an acquired habit of attention to parts. Aloft we have no associations to guide us, and their insulation in the expanse of sky leads us rather to undervalue than to over-rate their apparent magnitudes. Actual measurement with a proper instrument corrects our error, without, however, dispelling our illusion. By this we learn, that the sun, when just on the horizon, subtends at our eyes almost exactly the same, and the moon a materially *less* angle, than when seen at a great altitude in the sky, owing to its greater distance from us in the former situation as compared with the latter, as will be explained farther on.

(48.) After what has been said of the small extent of the atmosphere

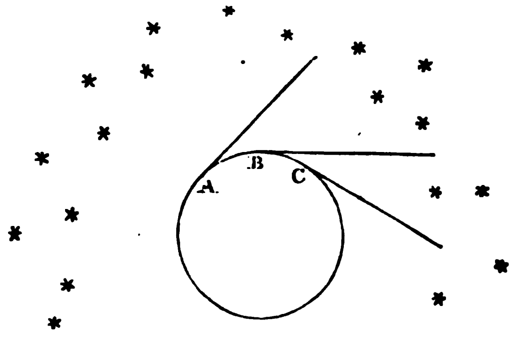
rison with the mass of the earth, we shall have little hesitation
ing those luminaries which people and adorn the sky, and which,
y obviously form no part of the earth, and receive no support
re yet not borne along at random like clouds upon the air, or
the winds, to be external to our atmosphere. As such we have
l them while speaking of their refractions — as existing in the
y of space beyond, and situated, perhaps, for any thing we can
to the contrary, at enormous distances from us and from each

Could a spectator exist unsustained by the earth, or any solid
he would see around him at one view the whole contents of space
sible constituents of the universe: and, in the absence of any
judging of their distances from him, would refer them, in the
in which they were seen from his station, to the concave sur-
imaginary sphere, having his eye for a centre, and its surface
vast indeterminate distance. Perhaps he might judge those
pear to him large and bright, to be nearer to him than the
nd less brilliant; but, independent of other means of judging
have no warrant for this opinion, any more than for the idea
were equidistant from him, and *really* arranged on such a

of the horizon will bring within the scope of vision a little more than a hemisphere, and refraction, wherever he may be situated, will enable him to look, as it were, a little round the corner; but the zone thus added to his visual range can hardly ever, unless in very extraordinary circumstances, exceed a couple of degrees in breadth, and is always ill seen on account of the vapours near the horizon. Unless, then, by a change of his geographical situation, he should shift his horizon (which is always a plane passing through his eye, and touching the spherical convexity of the earth); or unless, by some movements proper to the heavenly bodies, they should of themselves come above his horizon; or, lastly, unless, by some rotation of the earth itself on its centre, the point of its surface which he occupies should be carried round, and presented towards a different region of space; he would never obtain a sight of almost one half the objects external to our atmosphere. But if any of these cases be supposed, more, or all, may come into view according to the circumstances.

(51.) A traveller, for example, shifting his locality on our globe, will obtain a view of celestial objects invisible from his original station, in a way which may be not inaptly illustrated by comparing him to a person standing in a park close to a large tree. The massive obstacle presented by its trunk cuts off his view of all those parts of the landscape which it occupies as an object; but by walking round it a complete successive view of the whole panorama may be obtained. Just in the same way, if we set off from any station, as London, and travel southwards, we shall not fail to notice that many celestial objects which are never seen from London come successively into view, as if rising up above the horizon, night after night, from the south, although it is in reality our horizon,

Fig. 6.



elling with us southwards round the sphere, sinks in succes-
h them. The novelty and splendour of fresh constellations
ally brought into view in the clear calm nights of tropical
long voyages to the south, is dwelt upon by all who have
is spectacle, and never fails to impress itself on the recollection
most delightful and interesting of the associations connected
sive travel. A glance at the accompanying figure, exhibiting
ssive stations of a traveller, A, B, C, with the horizon corre-
o each, will place this process in clearer evidence than any

gain: suppose the earth itself to have a motion of rotation on
It is evident that a spectator at rest (as it appears to him)
t of it will, unperceived by himself, be carried round with it:
d, we say, because his horizon will constantly contain, and he
, the same terrestrial objects. He will have the same land-
tantly before his eyes, in which all the familiar objects in it,
e him for landmarks and directions, retain, with respect to
to each other, the same invariable situations. The perfect
s and equality of the motion of so vast a mass, in which every
sees around him participates alike, will (art. 15) prevent his

admitting into view the former, and successively hiding the latter. As the horizon of every such spectator, however, appears *to him* motionless, all such changes will be referred by him to a motion in the objects themselves so successively disclosed and concealed. In place of his horizon approaching the stars, therefore, he will judge the stars to approach his horizon; and when it passes over and hides any of them, he will consider them as having sunk below it, or *set;* while those it has just disclosed, and from which it is receding, will seem to be rising above it.

(54.) If we suppose this rotation of the earth to continue in one and the same direction, — that is to say, to be performed round one and the same *axis*, till it has completed an entire revolution, and come back to the position from which it set out when the spectator began his observations, — it is manifest that every thing will then be in precisely the same relative position as at the outset: all the heavenly bodies will appear to occupy the same places in the concave of the sky which they did at that instant, except such as may have actually moved in the interim; and if the rotation still continue, the same phenomena of their successive rising and setting, and return to the same places, will continue to be repeated in the same order, and (if the velocity of rotation be uniform) in equal intervals of time, *ad infinitum.*

(55.) Now, in this we have a lively picture of that grand phenomenon, the most important beyond all comparison which nature presents, the daily rising and setting of the sun and stars, their progress through the vault of the heavens, and their return to the same apparent places at the same hours of the day and night. The accomplishment of this restoration in the regular interval of twenty-four hours is the first instance we encounter of that great law of *periodicity*,[1] which, as we shall see, pervades all astronomy; by which expression we understand the continual reproduction of the same phenomena, in the same order, at equal intervals of time.

(56.) A free rotation of the earth round its centre, if it exist and be performed in consonance with the same mechanical laws which obtain in the motions of masses of matter under our immediate control, and within our ordinary experience, must be such as to satisfy two essential conditions. It must be invariable in its direction *with respect to the sphere itself*, and uniform in its velocity. The rotation must be performed *round an axis* or diameter of the sphere, whose *poles* or extremities, where it meets the surface, correspond always to the same points on the sphere. Modes of rotation of a solid body under the influence of external agency are conceivable, in which the poles of the imaginary line or axis about which it is at any moment revolving shall hold no fixed places on the surface, but shift upon it every moment. Such changes, however, are inconsistent with the idea of a rotation of a body of regular figure about its axis of symmetry, performed in free space, and without resistance or obstruction from any surrounding medium, or disturbing influences. The complete absence of such obstructions draws with it, of necessity, the strict fulfilment of the two conditions above mentioned.

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[1] Περίοδος, a *going round*, a circulation or revolution.

(57.) Now, these conditions are in perfect accordance with what we observe, and what recorded observation teaches us, in respect of the diurnal motions of the heavenly bodies. We have no reason to believe, from history, that any sensible change has taken place since the earliest ages in the interval of time elapsing between two successive returns of the same star to the same point of the sky; or, rather, it is demonstrable from astronomical records that no such change *has* taken place. And with respect to the other condition,—*the permanence of the axis of rotation*,—the appearances which any alteration in that respect must produce, would be marked, as we shall presently show, by a corresponding change of a very obvious kind in the apparent motions of the stars; which, again, history decidedly declares them *not* to have undergone.

(58.) But, before we proceed to examine more in detail how the hypothesis of the rotation of the earth about an axis accords with the phenomena which the diurnal motion of the heavenly bodies offers to our notice, it will be proper to describe, with precision, in what that diurnal motion consists, and how far it is participated in by them all; or whether any of them form exceptions, wholly or partially, to the common analogy of the rest. We will, therefore, suppose the reader to station himself, on a clear evening, just after sunset, when the first stars begin to appear, in some open situation whence a good general view of the heavens can be obtained. He will then perceive, above and around him, as it were, a vast concave hemispherical vault, beset with stars of various magnitudes, of which the brightest only will first catch his attention in the twilight; and more and more will appear as the darkness increases, till the whole sky is over-spangled with them. When he has awhile admired the calm magnificence of this glorious spectacle, the theme of so much song, and of so much thought,—a spectacle which no one can view without emotion, and without a longing desire to know something of its nature and purport,—let him fix his attention more particularly on a few of the most brilliant stars, such as he cannot fail to recognize again without mistake after looking away from them for some time, and let him refer their apparent situations to some surrounding objects, as buildings, trees, &c., selecting purposely such as are in different quarters of his horizon. On comparing them again with their respective points of reference, after a moderate interval, as the night advances, he will not fail to perceive that they have changed their places, and advanced, as by a general movement, in a westward direction; those towards the eastern quarter appearing to rise or recede from the horizon, while those which lie towards the west will be seen to approach it; and, if watched long enough, will, for the most part, finally sink beneath it, and disappear; while others, in the eastern quarter, will be seen to rise as if out of the earth, and, joining in the general procession, will take their course with the rest towards the opposite quarter.

(59.) If he persist for a considerable time in watching their motions, on the same or on several successive nights, he will perceive that each star appears to describe, as far as its course lies above the horizon, a circle in the sky; that the circles so described are not of the same magnitude for all the stars; and that those described by different stars differ greatly in respect of the parts of them which lie above the horizon. Some, which lie towards the quarter of the horizon which is denominated the SOUTH,[1] only remain for a short time above it, and disappear, after describing in sight only the small upper segment of their diurnal circle; others, which rise between the south and east, describe larger segments of their circles above the horizon, remain proportionally longer in sight, and set precisely as far to the westward of south as they rose to the eastward; while such as rise exactly in the east remain just twelve hours visible, describe a semicircle, and set exactly in the west. With those, again, which rise between the east and north, the same law obtains; at least, as far as regards the time of their remaining above the horizon, and the proportion of the visible segment of their diurnal circles to their whole circumferences. Both go on increasing; they remain in view more than twelve hours, and their visible diurnal arcs are more than semicircles. But the magnitudes of the circles themselves diminish, as we go from the east, northward; the greatest of all the circles being described by those which rise exactly in the east point. Carrying his eye farther northwards, he will notice, at length, stars which, in their diurnal motion, just graze the horizon at its north point, or only dip below it for a moment; while others never reach it at all, but continue always above it, revolving in entire circles round ONE POINT called the POLE, which appears to be the common centre of all their motions, and which alone, in the whole heavens, may be considered immoveable. Not that this point is marked by any star. It is a purely imaginary centre; but there is near it one considerably

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[1] We suppose our observer to be stationed in some northern latitude; some where in Europe, for example.

called the Pole Star, which is easily recognized by the very
it describes; so small, indeed, that, without paying particular
d referring its position very nicely to some fixed mark, it may
pposed at rest, and be, itself, mistaken for the common centre
all the others in that region describe their circles; or it may
y its configuration with a very splendid and remarkable *con-*
group of stars, called by astronomers the GREAT BEAR.
will further observe, that the apparent relative situations of
among one another, is not changed by their diurnal motion.
parts of their circles they are observed, or at whatever hour
, they form with each other the same identical groups or con-
to which the name of CONSTELLATIONS has been given. It
, in different parts of their course, these groups stand dif-
h respect to the horizon; and those towards the north, when
e of their diurnal movement they pass alternately above and
ommon centre of motion described in the last article, become
erted with respect to the horizon, while, on the other hand,
turn the same points towards the pole. In short, he will
t the whole assemblage of stars visible at once, or in succes-
heavens, may be regarded as one great constellation, which
volve with a uniform motion, as if it formed one coherent

(See art. 49.) There is always, therefore (he would justly argue), a star-bespangled canopy over his head, by day as well as by night, only that the glare of daylight (which he perceives gradually to efface the stars as the morning twilight comes on) prevents them from being seen. And such is really the case. The stars actually continue visible through telescopes in the day-time; and, in proportion to the power of the instrument, not only the largest and brightest of them, but even those of inferior lustre, such as scarcely strike the eye at night as at all conspicuous, are readily found and followed even at noonday,—unless in that part of the sky which is very near the sun,—by those who possess the means of pointing a telescope accurately to the proper places. Indeed, from the bottoms of deep narrow pits, such as a well, or the shaft of a mine, such bright stars as pass the zenith may even be discerned by the naked eye; and we have ourselves heard it stated by a celebrated optician, that the earliest circumstance which drew his attention to astronomy was the regular appearance, at a certain hour, for several successive days, of a considerable star, through the shaft of a chimney. Venus in our climate, and even Jupiter in the clearer skies of tropical countries, are often visible, without any artificial aid, to the naked eye of one who knows nearly where to look for them. During total eclipses of the sun, the larger stars also appear in their proper situations.

(62.) But to return to our incipient astronomer, whom we left contemplating the sphere of the heavens, as completed in imagination beneath his feet, and as rising up from thence in its diurnal course. There is one portion or segment of this sphere of which he will not thus obtain a view. As there is a segment towards the north, adjacent to the pole above his horizon, in which the stars *never set*, so there is a corresponding segment, about which the smaller circles of the more southern stars are described, in which they *never rise*. The stars which border upon the extreme circumference of this segment just graze the southern point of his horizon, and show themselves for a few moments above it, precisely as those near the circumference of the northern segment graze his northern horizon, and dip for a moment below it, to re-appear immediately. Every point in a spherical surface has, of course, another diametrically opposite to it; and as the spectator's horizon divides his sphere into two hemispheres — a superior and inferior — there must of necessity exist a depressed pole to the south, corresponding to the elevated one to the north, and a portion surrounding it, perpetually beneath, as there is another surrounding the north pole, perpetually above it.

"Hic vertex nobis semper sublimis; at illum
Sub pedibus nox atra videt, manesque profundi."—Virgil.

One pole rides high, one, plunged beneath the main,
Seeks the deep night, and Pluto's dusky reign.

get sight of this segment, he must travel southwards. In so
w set of phenomena come forward. In proportion as he
the south, some of those constellations which, at his original
ly grazed the northern horizon, will be observed to sink below
t first remaining hid only for a very short time, but gradually
part of the twenty-four hours. They will continue, however,
about the same point — that is, holding the same invariable
respect to them in the concave of the heavens among the
his point itself will become gradually depressed with respect
ator's horizon. The axis, in short, about which the diurnal
rformed, will appear to have become continually less and less
the horizon; and by the same degrees as the northern pole is
e southern will rise, and constellations surrounding it will
ew; at first momentarily, but by degrees for longer and longer
ch diurnal revolution — realizing, in short, what we have
d in art. 51.

he travel continually southwards, he will at length reach a
earth's surface, called *the equator*, at any point of which,
, if he take up his station and recommence his observations,

constantly invisible to him; and *vice versâ*, those stars which at his former station he never saw, he will now never cease to see.

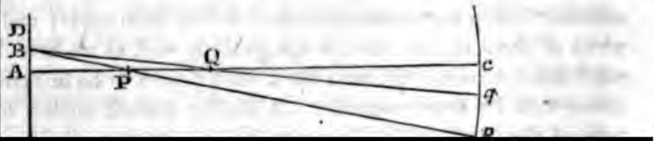
(66.) Finally, if, instead of advancing southwards from his first station, he travel northwards, he will observe the northern pole of the heavens to become more elevated above his horizon, and the southern more depressed below it. In consequence, his hemisphere will present a less variety of stars, because a greater proportion of the whole surface of the heavens remains constantly visible or constantly invisible: the circle described by each star, too, becomes more nearly parallel to the horizon; and, in short, every appearance leads to suppose that could he travel far enough to the north, he would at length attain a point *vertically under* the northern pole of the heavens, at which none of the stars would either rise or set, but each would circulate round the horizon in circles parallel to it. Many endeavours have been made to reach this point, which is called the north pole of the earth, but hitherto without success; a barrier of almost insurmountable difficulty being presented by the increasing rigour of the climate: but a very near approach to it has been made; and the phenomena of those regions, though not precisely such as we have described as what must subsist *at* the pole itself, have proved to be in exact correspondence with its near proximity. A similar remark applies to the south pole of the earth, which, however, is more unapproachable, or, at least, has been less nearly approached, than the north.

(67.) The above is an account of the phenomena of the diurnal motion of the stars, as modified by different geographical situations, not grounded on any speculation, but actually observed and recorded by travellers and voyagers. It is, however, in complete accordance with the hypothesis of a rotation of the earth round a fixed axis. In order to show this, however, it will be necessary to premise a few observations on *parallactic motion* in general, and on the appearances presented by an assemblage of remote objects, when viewed from different parts of a small and circumscribed station.

(68.) It has been shown (art. 16) that a spectator in smooth motion, and surrounded by, and forming part of, a great system partaking of the same motion, is unconscious of his own movement, and transfers it in idea to objects external and unconnected, in a contrary direction; those which he leaves behind appearing to recede from, and those which he advances towards to approach, him. Not only, however, do external objects at rest appear in motion generally, with respect to ourselves when we are in motion among them, but they appear to move one among the other — they shift their *relative* apparent places. Let any one travelling rapidly along a high road fix his eye steadily on any object, but at the

not entirely withdraw his attention from the general landscape,
ee, or think he sees, the whole landscape thrown into *rotation*,
g round that object as a centre; all objects between it and
pearing to move *backwards*, or the contrary way to his own
nd all beyond it, forwards, or in the direction in which he
t let him withdraw his eye from that object, and fix it on
a nearer one, for instance,—immediately the appearance of
ifts also, and the apparent centre about which this illusive
is performed is transferred to the new object, which, for the
ppears to rest. This apparent change of situation of objects
ct to one another, arising from a motion of the spectator, is
rallactic motion. To see the reason of it we must consider
sition of every object is referred by us to the surface of an
sphere of an indefinite radius, having our eye for its centre:

Fig. 7.



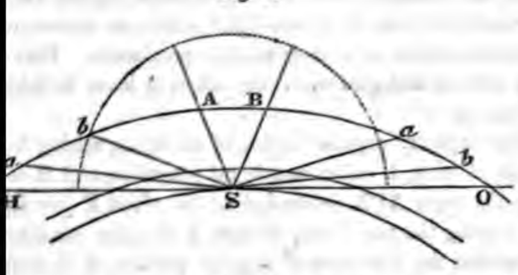
dark night, for instance, when all intermediate objects are unseen, the apparent relative movement of two lights which we are assured are themselves fixed, will decide as to their relative proximities. That which seems to advance with us and gain upon the other, or leave it behind it, is the farthest from us.

(70.) The apparent angular motion of an object, arising from a change of our point of view, is called in general *parallax*, and it is always expressed by *the angle* APB *subtended at the object* P (see fig. of art. 68) by a line joining the two points of view A B under consideration. For it is evident that the difference of angular position of P, with respect to the invariable direction ABD, when viewed from A and from B, is the difference of the two angles DBP and DAP; now, DBP being the exterior angle of the triangle ABP, is equal to the sum of the interior and opposite, $DBP = DAP + APB$, whence $DBP - DAP = APB$.

(71.) It follows from what has been said that the amount of parallactic motion arising from any given change of our point of view is, *cæteris paribus*, less, as the distance of an object viewed is greater; and when that distance is extremely great in comparison with the change in our point of view, the parallax becomes insensible; or, in other words, objects do not appear to vary in situation at all. It is on this principle, that in alpine regions visited for the first time we are surprised and confounded at the little progress we appear to make by a considerable change of place. An hour's walk, for instance, produces but a small parallactic change in the relative situations of the vast and distant masses which surround us. Whether we walk round a circle of a hundred yards in diameter, or merely turn ourselves round in its centre, the distant panorama presents almost exactly the same aspect, — we hardly seem to have changed our point of view.

(72.) Whatever notion, in other respects, we may form of the stars, it is quite clear they must be immensely distant. Were it not so, the apparent angular interval between any two of them seen over-head would be much greater than when seen near the horizon, and the constellations, instead of preserving the same appearances and dimensions during their whole diurnal course, would appear to enlarge as they rise higher in the sky, as we see a small cloud in the horizon swell into a great overshadowing canopy when drifted by the wind across our zenith, or as may be seen in the annexed figure, where ab, A B, ab, are three different positions of the same stars, as they would, if near the earth, be seen from a spectator S, under the visual angles aSb, ASB. No such change of apparent dimension, however, is observed. The nicest measurements of the apparent angular distance of any two *stars inter se*, taken in any

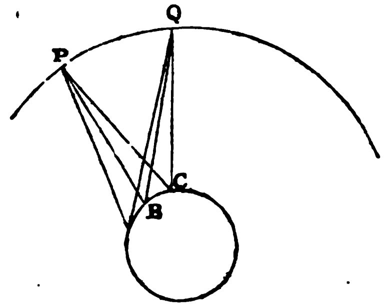
Fig. 8.



ir diurnal course, (after allowing for the unequal effects of
r when taken at such times that this cause of distortion shall
on both,) manifest *not the slightest* perceptible variation.
is, but at whatever point of the earth's surface the measure-
ormed, the results are *absolutely identical.* No instruments
ented by man are delicate enough to indicate, by an increase
n of the angle subtended, that one point of the earth is nearer
r from the *stars* than another.

e necessary conclusion from this is, that the dimensions of
rge as it is, are comparatively *nothing*, absolutely impercep-
compared with the interval which separates the stars from the

Fig. 9.



distance of the stars from the earth cannot be *so small* as 100,000 of the earth's diameters. It is, indeed, incomparably greater; for we shall hereafter find it fully demonstrated that the distance just named, immense as it may appear, is yet much underrated.

(74.) From such a distance, to a spectator with our faculties, and furnished with our instruments, the earth would be imperceptible; and, reciprocally, an object of the earth's size, placed at the distance of the stars, would be equally undiscernible. If, therefore, at the point on which a spectator stands, we draw a plane touching the globe, and prolong it in imagination till it attain the region of the stars, and through the centre of the earth conceive another plane parallel to the former, and co-extensive with it, to pass; these, although separated throughout their whole extent by the same interval, viz., a semi-diameter of the earth, will yet, on account of the vast distance at which that interval is seen, be confounded together, and undistinguishable from each other in the region of the stars, when viewed by a spectator on the earth. The zone they there include will be of evanescent breadth to his eye, and will only mark out a great circle in the heavens, one and the same for both the stations. This great circle, when spoken of as a circle of the sphere, is called the *celestial horizon* or simply *the* horizon, and the two planes just described are also spoken of as the *sensible* and the *rational* horizon of the observer's station.

(75.) From what has been said (art. 73) of the distance of the stars, it follows, that if we suppose a spectator at the centre of the earth to have his view bounded by the *rational* horizon, in exactly the same manner as that of a corresponding spectator on the surface is by his *sensible* horizon, the two observers will see the same stars in the same relative situations, each beholding that entire hemisphere of the heavens which is above the celestial horizon, corresponding to their common zenith. Now, so far as

go, it is clearly the same thing whether the heavens, that is,
h its contents, revolve round a spectator at rest in the earth's
ether that spectator simply turn round in the opposite direc-
lace, and view them in succession. The aspect of the heavens,
nt, as referred to his horizon (which must be supposed to
n), will be the same in both suppositions. And since, as has
appearances are also, so far as the stars are concerned, the
ectator on the surface as to one at the centre, it follows that,
suppose the heavens to revolve without the earth, or the earth
eavens, *in the opposite direction*, the diurnal phenomenon, to
tants, will be no way different.

Copernican astronomy adopts the latter as the true explana-
phenomena, avoiding thereby the necessity of otherwise re-
he cumbrous mechanism of a solid but invisible sphere, to
tars must be supposed attached, in order that they may be
d the earth without derangement of their relative situations
ich a contrivance would, indeed, suffice to explain the diurnal
the stars, so as to "save appearances;" but the movements
nd moon, as well as those of the planets, are incompatible with
sition, as will appear when we come to treat of these bodies.
r hand, that a spherical mass of moderate dimensions (or,

al account of the appearances they successively present,— that is to say, an account of which the several parts, postulates, propositions, deductions, intelligibly cohere, without contradicting each other or the nature of things as concluded from experience. In this view of the Copernican doctrine it is rather a geometrical conception than a physical theory, inasmuch as it simply assumes the requisite motions, without attempting to explain their mechanical origin, or assign them any dependence on physical causes. The Newtonian theory of gravitation supplies this deficiency, and, by showing that all the motions required by the Copernican conception *must*, and that no others can, result from a single, intelligible, and very simple dynamical law, has given a degree of certainty to this conception, as a matter of fact, which attaches to no other creation of the human mind.

(78.) To understand this conception in its further developments, the reader must bear steadily in mind the distinction between *relative* and *absolute* motion. Nothing is easier to perceive than that, if a spectator at rest view a certain number of moving objects, they will group and arrange themselves *to his eye*, at each successive moment, in a very different way from what they would do were he in active motion among them,— if he formed one of them, for instance, and joined in their dance. This is evident from what has been said before of parallactic motion; but it will be asked, How is such a spectator to disentangle from each other the two parts of the apparent motions of these external objects, — that which arises from the effect of his own change of place, and which is therefore only apparent (or, as a German metaphysician would say, *subjective* — having reference only to him as perceiving it), — and that which is real (or *objective* — having a positive existence, whether perceived by him or not)? By what rule is he to ascertain, from the appearances presented to him while himself in motion, what *would be* the appearances were he at rest? It by no means follows, indeed, that he would even then at once obtain a clear conception of all the motions of all the objects. The appearances so presented to him would have still something *subjective* about them. They would be still *appearances*, not geometrical realities. They would still have a reference to the point of view, which might be very unfavourably situated (as, indeed, is the case in our system) for affording a clear notion of the real movement of each object. No geometrical figure, or curve, is seen by the eye as it is conceived by the mind to exist in reality. The laws of perspective interfere and alter the apparent directions and foreshorten the dimensions of its several parts. If the spectator be unfavourably situated, as, for instance, nearly in the plane of the figure (which is the case we have to deal with), they may do so to such an extent, as to make

e effort of imagination necessary to pass from the sensible to
.

l, preparatory to this ultimate step, it is first necessary that
should free or clear the appearances from the disturbing
his own change of place. And this he can always do by the
eral rule or proposition : —

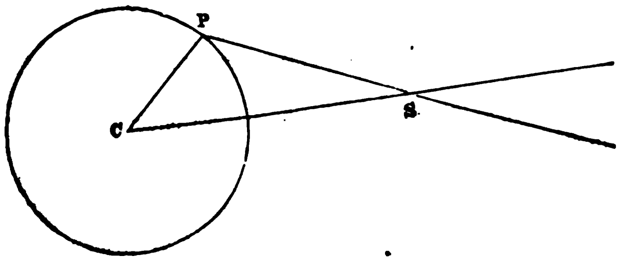
ive motion of two bodies is the same as if either of them
and all its motion communicated to the other in an opposite

two bodies move alike, they will, when seen from each other
erence to other near bodies, but only to the starry sphere),
st. Hence, also, if the absolute motions of two bodies be
rectilinear, their relative motion is so also.

stars are so distant, that as we have seen it is absolutely
om what point of the earth's surface we view them. Their
s *inter se* are identically the same. It is otherwise with the
nd planets, which are near enough (especially the moon) to
ically displaced by change of station from place to place on
In order that astronomers residing on different points of the
ce should be able to compare their observations with effect, it
that they should clearly understand and take account of this

upon a point equally related to all, as the common point of reference; and this can be no other than the centre of the globe itself. The parallactic change of apparent place which would arise in an object, could any observer suddenly transport himself to the centre of the earth, is evidently the angle C S P, subtended on the object S by that radius C P of the earth which joins the centre and the place P of observation.

Fig. 10.



CHAPTER II.

GY AND ELEMENTARY GEOMETRICAL CONCEPTIONS AND NS. — TERMINOLOGY RELATING TO THE GLOBE OF THE – TO THE CELESTIAL SPHERE. — CELESTIAL PERSPECTIVE.

VERAL of the terms in use among astronomers have been ex- he preceding chapter, and others used anticipatively. But the nguage of every subject requires to be formally stated, both ncy of usage and definiteness of conception. We shall there- d, in the first place, to define a number of terms in perpetual relation to the globe of the earth and the celestial sphere.

ION 1. The *axis* of the earth is that diameter about which with a uniform motion, *from west to east;* performing one in the interval which elapses between any star leaving a cer-

(88.) **Def.** 7. The *latitude* of a place on the earth's surface is its angular distance from the equator, measured on its own terrestrial meridian: it is reckoned in degrees, minutes, and seconds, from 0 up to 90°, and northwards or southwards according to the hemisphere the place lies in. Thus, the observatory at Greenwich is situated in 51° 28′ 40″ north latitude. This definition of latitude, it will be observed, is to be considered as only temporary. A more exact knowledge of the physical structure and figure of the earth, and a better acquaintance with the niceties of astronomy, will render some modification of its terms, or a different manner of considering it, necessary.

(89.) **Def.** 8. *Parallels of latitude* are small circles on the earth's surface parallel to the equator. Every point in such a circle has the same latitude. Thus, Greenwich is said to be situated *in the parallel of* 51° 28′ 40″.

(90.) **Def.** 9. The *longitude* of a place on the earth's surface is the inclination of its meridian to that of some fixed station referred to as a point to reckon from. English astronomers and geographers use the observatory at Greenwich for this station; foreigners, the principal observatories of their respective nations. Some geographers have adopted the island of Ferro. Hereafter, when we speak of longitude, we reckon from Greenwich. The longitude of a place is, therefore, measured by the arc of the equator intercepted between the meridian of the place and that of Greenwich; or, which is the same thing, by the spherical angle at the pole included between these meridians.

(91.) As *latitude* is reckoned north or south, so *longitude* is usually said to be reckoned west or east. It would add greatly, however, to systematic regularity, and tend much to avoid confusion and ambiguity in computations, were this mode of expression abandoned, and longitudes reckoned invariably *westward* from their origin round the whole circle from 0 to 360°. Thus, the longitude of Paris is, in common parlance, either 2° 20′ 22″ east, or 357° 39′ 38″ west of Greenwich. But, in the sense in which we shall henceforth use and recommend others to use the term, the latter is its proper designation. Longitude is also reckoned in time at the rate of 24 h. for 360°, or 15° per hour. In this system the longitude of Paris is 23 h. 50 m. 38½ s.[1]

(92.) Knowing the longitude and latitude of a place, it may be laid down on an artificial globe; and thus a map of the earth may be con-

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[1] To distinguish minutes and seconds of time from those of angular measure we shall invariably adhere to the distinct system of notation here adopted (° ′ ″, and h. m. s.) Great confusion sometimes arises from the practice of using the same marks for both.

Maps of particular countries are detached portions of this
), extended into planes; or rather, they are representations
f such portions, executed according to certain conventional
rules, called *projections*, the object of which is either to
tle as possible the outlines of countries from what they are
—or to establish easy means of ascertaining, by inspection or
easurement, the latitudes and longitudes of places which
m, without referring to the globe or to books — or for other
s. See Chap. IV.

EF. 10. The *Tropics* are two parallels of latitude, one on the
the other on the south side of the equator, over every point
espectively, the sun in its diurnal course passes vertically on
June ~~March~~ and the 21st of December ~~September~~ in every year. Their
about 23° 28′ respectively, north and south.

EF. 11. The Arctic and Antarctic circles are two small circles
of latitude as distant from the north and south poles as the
rom the equator, that is to say, about 23° 28′; their latitudes,
re about 66° 32′. We say *about*, for the places of these
f the tropics are continually shifting on the earth's surface,
extreme slowness, as will be explained in its proper place.

F. 12. The sphere of the heavens or of the stars is an imatronomers, the *equinoctial,* is a great circle of the celestial sphere, marked out by the indefinite extension of the plane of the terrestrial equator.

(98.) Def. 15. The celestial horizon of any place is a great circle of the sphere marked out by the indefinite extension of the plane of any spectator's *sensible* or (which comes to the same thing as will presently be shown,) his *rational* horizon, as in the case of the equator.

(99.) Def. 16. The *zenith* and *nadir*[1] of a spectator are the two points of the sphere of the heavens, vertically over his head, and vertically under his feet, or the *poles* of the celestial horizon; that is to say, points 90° distant from every point in it.

(100.) Def. 17. *Vertical circles* of the sphere are great circles passing through the zenith and nadir, or great circles perpendicular to the horizon. On these are measured the *altitudes* of objects above the horizon — the complements to which are their *zenith distances.*

(101.) Def. 18. The *celestial meridian* of a spectator is the great circle marked out on the sphere by the prolongation of the plane of his terrestrial meridian. If the earth be supposed at rest, this is a fixed circle, and all the stars are carried across it in their diurnal courses from east to west. If the stars rest and the earth rotate, the spectator's meridian, like his horizon (art. 52), sweeps daily across the stars from west to east. Whenever in future we speak of the meridian of a spectator or observer, we intend the celestial meridian, which being a circle passing through the poles of the heavens and the zenith of the observer, is necessarily a vertical circle, and passes through the north and south points of the horizon.

(102.) Def. 19. The *prime vertical* is a vertical circle perpendicular to the meridian, and which therefore passes through the east and west points of the horizon.

[illegible]

(103.) Def. 20. *Azimuth* is the angular distance of a celestial object from the north or south point of the horizon (according as it is the north or south pole which is *elevated*), when the object is referred to the horizon by a vertical circle; or it is the angle comprised between two vertical planes — one passing through the elevated pole, the other through the object. Azimuth *may* be reckoned eastwards or westwards, from the north or south point, and is usually so reckoned only to 180° either way. But to avoid confusion, and to preserve continuity of interpretation when algebraic symbols are used (a point of essential importance, hitherto too little insisted on), we shall always reckon azimuth *from* the points of the horizon *most remote* from the *elevated* pole, *westward* (so as to agree in general directions with the apparent diurnal motion of the stars), and

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[1] From Arabic words. Nadir corresponds evidently to the German *nieder,* (down,) whence our *nether.*

ckoning from 0° to 360° if always reckoned positive, consider-
tward reckoning as negative.

)EF. 21. The *altitude* of a heavenly body is its apparent angular
bove the horizon. It is the complement to 90°, therefore, of
istance. The altitude and azimuth of an object being known,
the visible heavens is determined.

)EF. 22. The *declination* of a heavenly body is its angular
m the equinoctial or celestial equator, or the complement to
angular distance from the nearest pole, which latter distance is
Polar distance. Declinations are reckoned *plus* or *minus*,
s the object is situated in the northern or southern celestial
. *Polar distances* are always reckoned from the North Pole,
to 180°, by which all doubt or ambiguity of expression with
ign is avoided.

)EF. 23. Hour circles of the sphere, or circles of declination,
rcles passing through the poles, and of course perpendicular to
tial. The hour circle, passing through any particular heavenly
s to refer it to a point in the equinoctial, as a vertical circle
oint in the horizon.

)EF. 24. The hour angle of a heavenly body is the angle at
luded between the hour circle passing through the body, and

(110.) *Sidereal time* is reckoned by the diurnal motion of the stars, or rather of that point in the equinoctial from which right ascensions are reckoned. This point may be considered as a star, though no star is, in fact, there; and, moreover, the point itself is liable to a certain slow variation, — so slow, however, as not to affect, perceptibly, the interval, of any two of its successive returns to the meridian. This interval is called a sidereal day, and is divided into 24 sidereal hours, and these again into minutes and seconds. A clock which marks sidereal time, *i. e.* which goes at such a rate as always to show 0h. 0m. 0s. when the equinox comes on the meridian, is called a sidereal clock, and is an indispensable piece of furniture in every observatory. Hence the hour angle of an object reduced to time at the rate of 15° per hour, expresses the interval of sidereal time by which (if its reckoning be positive) it has past the meridian; or, if negative, the time it wants of arriving at the meridian of the place of observation. So also the right ascension of an object, if converted into time at the same rate (since 360° being described uniformly in 24 hours, 15° must be so described in 1 hour), will express the interval of sidereal time which elapses from the passage of the vernal equinox across the meridian to that of the object next subsequent.

(111.) As a globe or maps may be made of the whole or particular regions of the surface of the earth, so also a globe, or general map of the heavens, as well as charts of particular parts, may be constructed, and the stars laid down in their proper situations relative to each other, and to the poles of the heavens and the celestial equator. Such a representation, once made, will exhibit a true appearance of the stars as they present themselves in succession to every spectator on the surface, or as they may be conceived to be seen at once by one at the centre of the globe. It is, therefore, independent of all *geographical* localities. There will occur in such a representation neither zenith, nadir, nor horizon — neither east nor west points; and although great circles may be drawn on it from pole to pole, corresponding to terrestrial meridians, they can no longer, in this point of view, be regarded as the celestial meridians of fixed points on the earth's surface, since, in the course of one diurnal revolution, every point in it passes beneath each of them. It is on account of this change of conception, and with a view to establish a complete distinction between the two branches of *Geography* and *Uranography*,[1] that astronomers have adopted different terms, (viz. *declination* and *right ascension*) to represent those arcs in the heavens which correspond to *latitudes* and *longitudes* on the earth. It is for this reason that they quator of the heavens the *equinoctial;* that what are meridians h are called *hour circles* in the heavens, and the angles they tween them at the poles are called *hour angles.* All this is and intelligible; and had they been content with this nomen- confusion could ever have arisen. Unluckily, the early s have employed *also* the words latitude and longitude in their y, in speaking of arcs of circles not corresponding to those he same words on the earth, but having reference to the motion and planets among the stars. It is now too late to remedy ion, which is ingrafted into every existing work on astronomy: y regret, and warn the reader of it, that he may be on his n, at a more advanced period of our work, we shall have occa- ne and use the terms in their *celestial sense,* at the same time commending to future writers the adoption of others in their

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[1] Γη, the earth; γραφειν, to describe or represent; ουρανος, the heaven.

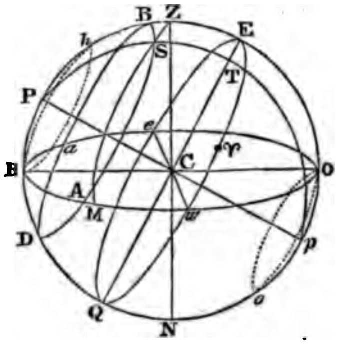
It remains to illustrate these descriptions by reference to a et C be the centre of the earth, N C S its axis; then are N

Fig. 11.

be his sensible horizon: *n* A *s* marked on that plane by its intersection with his meridian will be his meridian line, and *n* and *s* the north and south points of his horizon.

(113.) Again, neglecting the size of the earth, or conceiving him stationed at its centre, and referring every thing to his *rational* horizon; let the annexed figure represent the sphere of the *heavens;* C the spectator; Z his zenith; and N his nadir: then will H A O, a great circle of the sphere, whose poles are Z N, be his *celestial horizon;* P *p* the *elevated* and *depressed* POLES of the heavens; H P the *altitude of the pole*, and H P Z E O his *meridian;* E T Q, a great circle perpendicular to P *p*, will be the *equinoctial;* and if ♈ represent the equinox, ♈ T will be the *right ascension*, T S the *declination*, and P S the *polar distance* of any star or object S, referred to the equinoctial by the *hour circle* P S T *p;* and B S D will be the diurnal circle it will appear to describe about the pole. Again, if we refer it to the horizon by the *vertical circle* Z S M, O M will be its azimuth, M S its altitude, and Z S its zenith distance. H and O are the north and south, *e w* the east and west points of his horizon, or of the heavens. Moreover, if H *h*, O *o*, be small circles, or *parallels of declination*, touching the horizon in its north and south points, H *h* will be the circle of *perpetual apparition*, between which and the elevated pole the stars never set; O *o* that of *perpetual occultation*, between which and the depressed pole they *never rise*. In all the zone of the heavens between H *h* and O *o*, they rise and set; any one of them, as S, remaining above the horizon in that part of its diurnal circle represented by *a* B A, and below it throughout all the part represented by A D *a*. It will exercise the reader to construct this figure for

Fig. 12.



ɜrent *elevations of the pole*, and for a variety of positions of n each.

Celestial perspective is that branch of the general science of which teaches us to conclude, from a knowledge of the real id forms of objects, lines, angles, motions, &c. with respect to or, their apparent aspects, as seen by him projected on the concave of the heavens; and, *vice versâ*, from the apparent ns and movements of objects so seen projected, to conclude, ey can be thence concluded, their real geometrical relations to and to the spectator. It agrees with ordinary perspective when ll visual area is contemplated, because the concave ground of l sphere, for a small extent, may be regarded as a plane surhich objects are seen projected or depicted as in common per- But when large amplitudes of the visual area are considered, e whole contents of space are regarded as projected on the ior surface of the sphere, it becomes necessary to use a different y, and to resort to a different form of conception. In common there is a single "point of sight," or "centre of the picture," ine from the eye to which is perpendicular to the "plane of ," and all straight lines are represented by straight lines. In erspective, every point to which the view is for the moment

These lines which, marking the course of rays emanating from a point almost infinitely distant, are to be considered as parallel straight lines, are thrown into great circles of the sphere, having two apexes or points of common intersection — one in the place where the sun itself (if not obscured) would be seen. The other diametrically opposite. The first only is most commonly suggested when the spectator's view is towards the sun. But in mountainous countries, the phenomenon of sunbeams converging towards a point diametrically opposite to the sun, and as much depressed below the horizon as the sun is elevated above it, is not unfrequently noticed, the back of the spectator being turned to the sun's place. Occasionally, but much more rarely, the whole course of such a system of sunbeams, stretching in semicircles across the hemisphere from horizon to horizon (the sun being near setting), may be seen.[1] Thus again, the streamers of the Aurora Borealis, which are doubtless electrical rays, parallel, or nearly parallel to each other, and to the dipping needle, usually appear to diverge from the point towards which the needle, freely suspended, would dip northwards (*i. e.* about 70° below the horizon and 23° west of north from London), and in their upward progress pursue the course of great circles till they again converge (in appearance) towards the point diametrically opposite (*i. e.* 70° above the horizon, and 23° to the eastward of south), forming a sort of canopy over-head, having that point for its centre. So also in the phenomenon of shooting stars, the lines of direction which they appear to take on certain remarkable occasions of periodical recurrence, are observed, if prolonged backwards, apparently to meet nearly in one point of the sphere; a certain indication of a general near approach to parallelism in the real directions of their motions on those occasions. On which subject more hereafter.

(116.) In relation to this idea of celestial perspective, we may conceive the north and south poles of the sphere as the two vanishing points of a system of lines parallel to the axis of the earth; and the zenith and nadir of those of a system of perpendiculars to its surface at the place of observation, &c. It will be shown that the direction of a *plumb-line*, at every place is perpendicular to the surface of still water at that place which is the true horizon, and though mathematically speaking no two plumb-lines are exactly parallel (since they converge to the earth's centre), yet over very small tracts, such as the area of a building — in one and the same town, &c., the difference from exact parallelism is so small that it may be practically disregarded.[1] To a spectator looking upwards such a system of plumb-lines will appear to converge to his zenith; downwards, to his nadir.

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[1] It is in such cases only that we conceive them as circles, the ordinary conventions of plane perspective becoming untenable. The author had the good fortune to witness on one occasion the phenomenon described in the text under circumstances of more than usual grandeur. Approaching Lyons from the south on Sept. 30, 1826, about 5½ h. P. M., the sun was seen nearly setting behind broken masses of stormy cloud, from whose apertures streamed forth beams of rose-coloured light, traceable all across the hemisphere almost to their opposite point of convergence behind the snowy precipices of Mont Blanc, conspicuously visible at nearly 100 miles to the eastward. The impression produced was that of another but feebler sun about to rise from behind the mountain, and darting forth precursory beams to meet those of the real one opposite.

(117.) So also the celestial equator, or the equinoctial, must be conceived as the vanishing circle of a system of planes parallel to the earth's equator, or perpendicular to its axis. The celestial horizon of any spectator is in like manner the vanishing circle of all planes parallel to his true horizon, of which planes his *rational* horizon (passing through the earth's centre) is one, and his *sensible* horizon (the tangent plane of his station) another.

(118.) Owing, however, to the absence of all the ordinary indications of distance which influence our judgment in respect of terrestrial objects, owing to the want of determinate figure and magnitude in the stars and planets as commonly seen — the projection of the celestial bodies on the ground of the heavenly concave is not usually regarded in this its true light, of *a perspective representation or picture*, and it even requires an effort of imagination to conceive them in their true relations, as at vastly different distances, one behind the other, and forming with one another lines of junction violently foreshortened, and including angles altogether differing from those which their projected representations appear to make. To do so at all with effect presupposes a knowledge of their actual situations in space, which it is the business of astronomy to arrive at by appropriate considerations. But the connections which subsist among the several parts *of the picture*, the purely geometrical relations among the angles and sides of the spherical triangles of which it consists, constitute, under the name of Uranometry,[2] a preliminary and subordinate branch of the general science, with which it is necessary to be familiar before any further progress can be made. Some of the most elementary and frequently occurring of these relations we proceed to explain. And first, as immediate consequences of the above definitions, the following propositions will be borne in mind.

(119.) *The altitude of the elevated pole is equal to the latitude of the spectator's geographical station.*

For it appears, see *fig.* art. 112, that the angle P A Z between the pole and the zenith is equal to N C A, and the angles Z A *n* and N C E being right angles, we have P A *n* = A C E. Now the former of these is the elevation of the pole as seen from E, the latter is the angle at the earth's centre subtended by the arc E A, or the latitude of the place.

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[1] An interval of a mile corresponds to a convergence of plumb-lines amounting to somewhat less space than a minute.

[2] Ουρανος, the heavens; μετρειν, to measure: the measurement of the heavens.

(120.) Hence to a spectator at the north pole of the earth, the north pole of the heavens is in his zenith. As he travels southward it becomes less and less elevated till he reaches the equator, when both poles are in his horizon — south of the equator the north pole becomes depressed below, while the south rises above his horizon, and continues to do so till the south pole of the globe is reached, when that of the heavens will be in the zenith.

(121.) The same stars, in their diurnal revolution, come to the meridian, *successively*, of every place on the globe once in twenty-four sidereal hours. And, since the diurnal rotation is uniform, the interval, in sidereal time, which elapses between the same star coming upon the meridians of two different places is measured by the difference of longitudes of the places.

(122.) *Vice versâ* — the interval elapsing between two *different stars* coming on the meridian of *one and the same place*, expressed in sidereal time, is the measure of the difference of right ascensions of the stars.

(123.) The equinoctial intersects the horizon in the east and west points, and the meridian is a point whose altitude is equal to the co-latitude of the place. Thus, at Greenwich, of which the latitude is 51° 28′ 40″, the altitude of the intersection of the equinoctial and meridian is 38° 31′ 20″. The north and south poles of the heavens are the poles of the equinoctial. The east and west points of the horizon of a spectator are the poles of his celestial meridian. The north and south points of his horizon are the poles of his prime vertical, and his zenith and nadir are the poles of his horizon.

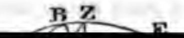
(124.) All the heavenly bodies *culminate* (*i. e.* come to their greatest altitudes) on the meridian; which is, therefore, the best situation to observe them, being least confused by the inequalities and vapours of the atmosphere, as well as least displaced by refraction.

(125.) All celestial objects within the circle of perpetual apparition come twice on the meridian, above the horizon, in every diurnal revolution; once *above* and once *below* the pole. These are called their *upper* and *lower culminations*.

(126.) The problems of uranometry, as we have described it, consist in the solution of a variety of spherical triangles, both right and oblique angled, according to the rules, and by the formulæ of spherical trigonometry, which we suppose known to the reader, or for which he will consult appropriate treatises. We shall only here observe generally, that in all which spherical geometry is concerned, the student will find ractical maxim rather to consider the poles of the great circles uestion before him refers to than the circles themselves. To mple, in the relations he has to consider, polar distances rather ations, zenith distances rather than altitudes, &c. Bearing l, there are few problems in uranometry which will offer any The following are the combinations which most commonly lution *when the place of one celestial object only* on the sphere l.

n the triangle Z P S, Z is the zenith, P the elevated pole, and sun, or other celestial object. In this triangle occur, 1st, P Z, g the complement of P H (the altitude of the pole), is ob- complement of the latitude (or the *co-latitude*, as it is called) ; 2d, P S, the *polar distance*, or the complement of the decli- *eclination*) of the star; 3d, Z S, the *zenith distance* or *co-alti-* star. If P S be greater than 90°, the object is situated on the equinoctial opposite to that of the elevated pole. If Z S bject is below the horizon.

Fig. 13.



The following five astronomical magnitudes, then, occur among the sides of this most useful triangle: viz., 1st, The co-latitude of the place of observation; 2d, the polar distance; 3d, the zenith distance; 4th, the hour angle; and 5th, the sub-azimuth (supplement of azimuth) of a given celestial object; and by its solution therefore may all problems be resolved, in which three of these magnitudes are directly or indirectly given, and the other two required to be found.

(128.) For example, suppose the time of rising or setting of the sun or of a star were required, having given its right ascension and polar distance. The star rises when *apparently* on the horizon, or *really* about 34′ below it (owing to refraction), so that, at the moment of its apparent rising, its zenith distance is 90° 34′ = Z S. Its polar distance PS being also given, and the co-latitude Z P of the place, we have given the three sides of the triangle, to find the hour angle Z P S, which, being known, is to be added to or subtracted from the star's right ascension, to give the sidereal time of setting or rising, which, if we please, may be converted into solar time by the proper rules and tables.

(129.) As another example of the use of the same triangle, we may propose to find the local sidereal time, and the latitude of the place of observation, by observing equal altitudes of the same star east and west of the meridian, and noting the interval of the observations in sidereal time.

The hour angles corresponding to equal altitudes of a fixed star being equal, the hour angle east or west will be measured by half the observed interval of the observations. In our triangle, then, we have given this hour angle Z P S, the polar distance P S of the star, and Z S, its co-altitude at the moment of observation. Hence we may find P Z, the co-latitude of the place. Moreover, the hour angle of the star being known, and also its right ascension, the point of the equinoctial is known, which is on the meridian at the moment of observation; and, therefore, the local sidereal time at that moment. This is a very useful observation for determining the latitude and time at an unknown station.

CHAPTER III.[1]

OF THE NATURE OF ASTRONOMICAL INSTRUMENTS AND OBSERVATIONS IN GENERAL.—OF SIDEREAL AND SOLAR TIME.—OF THE MEASUREMENTS OF TIME. — CLOCKS, CHRONOMETERS. — OF ASTRONOMICAL MEASUREMENTS.— PRINCIPLE OF TELESCOPIC SIGHTS TO INCREASE THE ACCURACY OF POINTING. — SIMPLEST APPLICATION OF THIS PRINCIPLE.— THE TRANSIT INSTRUMENT.— OF THE MEASUREMENT OF ANGULAR INTERVALS. — METHODS OF INCREASING THE ACCURACY OF READING.— THE VERNIER.— THE MICROSCOPE.— OF THE MURAL CIRCLE. — THE MERIDIAN CIRCLE. — FIXATION OF POLAR AND HORIZONTAL POINTS. — THE LEVEL, PLUMB-LINE, ARTIFICIAL HORIZON.—PRINCIPLE OF COLLIMATION.—COLLIMATORS OF RITTENHOUSE, KATER, AND BENZENBERG. — OF COMPOUND INSTRUMENTS WITH CO-ORDINATE CIRCLES. — THE EQUATORIAL, ALTITUDE, AND AZIMUTH INSTRUMENTS. — THEODOLITE. — OF THE SEXTANT AND REFLECTING CIRCLE. — PRINCIPLE OF REPETITION. — OF MICROMETERS. — PARALLEL WIRE MICROMETER. — PRINCIPLE OF THE DUPLICATION OF IMAGES.— THE HELIOMETER.— DOUBLE REFRACTING EYE-PIECE. — VARIABLE PRISM MICROMETER. — OF THE POSITION MICROMETER.

(130.) OUR first chapters have been devoted to the acquisition chiefly of preliminary notions respecting the globe we inhabit, its relation to the celestial objects which surround it, and the physical circumstances under which all astronomical observations must be made, as well as to provide ourselves with a stock of *technical words* and elementary ideas of most frequent and familiar use in the sequel. We might now proceed to a more exact and detailed statement of the facts and theories of astronomy; but, in order to do this with full effect, it will be desirable that the reader be made acquainted with the principal means which astronomers possess, of determining, with the degree of nicety their theories require, the data on which they ground their conclusions; in other words, of ascertaining by measurement the apparent and real magnitudes with which they are conversant. It is only when in possession of this knowledge that he can fully appreciate either the truth of the theories themselves, or the degree of reliance to be placed on any of their conclusions antecedent to trial: since it is only by knowing what amount of error can certainly be perceived and distinctly measured, that he can satisfy himself whether any theory offers so close an approximation, in its numerical results, to actual phenomena, as will justify him in receiving it as a true representation of nature.

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[1] The student who is anxious to become acquainted with the chief subject matter of this work, may defer the reading of that part of this chapter which is devoted to the description of particular instruments, or content himself with a cursory perusal of it, until farther advanced, when it will be necessary to return to it.

(181.) Astronomical instrument-making may be justly regarded as the most refined of the mechanical arts, and that in which the nearest approach to geometrical precision is required, and has been attained. It may be thought an easy thing, by one unacquainted with the niceties required, to turn a circle in metal, to divide its circumference into 360 equal parts, and these again into smaller subdivisions, — to place it accurately on its centre, and to adjust it in a given position; but practically it is found to be one of the most difficult. Nor will this appear extraordinary, when it is considered that, owing to the application of telescopes to the purposes of angular measurement, every imperfection of structure of division becomes magnified by the whole optical power of that instrument; and that thus, not only direct errors of workmanship, arising from unsteadiness of hand or imperfection of tools, but those inaccuracies which originate in far more uncontrollable causes, such as the unequal expansion and contraction of metallic masses, by a change of temperature, and their unavoidable flexure or bending by their own weight, become perceptible and measurable. An angle of one minute occupies, on the circumference of a circle of 10 inches in radius, only about $\frac{1}{350}$th part of an inch, a quantity too small to be *certainly* dealt with without the use of magnifying glasses; yet one minute is a gross quantity in the astronomical measurement of an angle. With the instruments now employed in observatories, a single second, or the 60th part of a minute, is rendered a distinctly visible and appreciable quantity. Now, the arc of a circle, subtended by one second, is less than the 200,000th part of the radius, so that on a circle of 6 feet in diameter it would occupy no greater linear extent than $\frac{1}{5700}$th part of an inch; a quantity requiring a powerful microscope to be *discerned* at all. Let any one figure to himself, therefore, the difficulty of placing on the circumference of a metallic circle of such dimensions (supposing the difficulty of its construction surmounted), 360 marks, dots, or cognizable divisions, which shall all be true to their places within such narrow limits; to say nothing of the subdivision of the degrees so marked off into minutes, and of these again into seconds. Such a work has probably baffled, and will probably for ever continue to baffle, the utmost stretch of human skill and industry; nor, if executed, could it endure. The ever-varying fluctuations of heat and cold have a tendency to produce not merely temporary and transient, but permanent, uncompensated changes of form in all considerable masses of those metals which alone are applicable to such uses; and their own weight, however symmetrically formed, must always be unequally sustained, since it is impossible to apply the sustaining power to *every part* separately: even could this be done, at all events force must be used to move and to fix them; which can never be done without producing temporary and risking permanent change of form. It is true, by dividing them on their centres, and in the identical places they are destined to occupy, and by a thousand ingenious and delicate contrivances, wonders have been accomplished in this department of art, and a degree of perfection has been given, not merely to *chefs d'œuvre*, but to instruments of moderate prices and dimensions, and in ordinary use, which, on due consideration, must appear very surprising. But though we are entitled to look for *wonders* at the hands of scientific artists, we are not to expect *miracles*. The demands of the astronomer will always surpass the power of the artist; and it must, therefore, be constantly the aim of the former to make himself, as far as possible, independent of the imperfections incident to every work the latter can place in his hands. He must, therefore, endeavour so to combine his observations, so to choose his opportunities, and so to familiarize himself with all the causes which may produce instrumental derangement, and with all the peculiarities of structure and material of each instrument he possesses, as not to allow himself to be misled by their errors, but to extract from their indications, as far as possible, all that is *true*, and reject all that is erroneous. It is in this that the art of the practical astronomer consists, — an art of itself of a curious and intricate nature, and of which we can here only notice some of the leading and general features.

(132.) The great aim of the practical astronomer being numerical correctness in the results of instrumental measurement, his constant care and vigilance must be directed to the detection and compensation of errors, either by annihilating, or by taking account of, and allowing for them. Now, if we examine the sources from which errors may arise in any instrumental determination, we shall find them chiefly reducible to three principal heads: —

(133.) 1st, External or incidental causes of error; comprehending such as depend on external, uncontrollable circumstances: such as, fluctuations of weather, which disturb the amount of refraction from its tabulated value, and, being reducible to no fixed law, induce uncertainty to the extent of their own possible magnitude; such as, by varying the temperature of the air, vary also the form and position of the instruments used, by altering the relative magnitudes and the tension of their parts; and others of the like nature.

(134.) 2dly, *Errors of observation:* such as arise, for example, from inexpertness, defective vision, slowness in seizing the exact *instant* of occurrence of a phenomenon, or precipitancy in anticipating it, &c.; from atmospheric indistinctness; insufficient optical power in the instrument, and the like. Under this head may also be classed all errors arising from momentary instrumental derangement,—slips in clamping, looseness of screws, &c.

(135.) 3dly, The third, and by far the most numerous class of errors to which astronomical measurements are liable, arise from causes which may be deemed instrumental, and which may be subdivided into two principal classes. The *first* comprehends those which arise from an instrument not *being* what it professes to be, which is *error* of *workmanship.* Thus, if a pivot or axis, instead of being, as it ought, exactly cylindrical, be slightly flattened, or elliptical,—if it be not exactly (as it is intended it should) concentric with the circle it carries;—if this circle (so called) be in reality *not* exactly circular, or not in one plane;—if its divisions, intended to be precisely equidistant, should be placed in reality at unequal intervals,—and a hundred other things of the same sort. These are not mere speculative sources of error, but practical annoyances, which every observer has to contend with.

(136.) The *other* subdivision of instrumental errors comprehends such as arise from an instrument not being placed in the *position* it ought to have; and from those of its parts, which are made purposely moveable, not being properly disposed *inter se.* These are *errors of adjustment.* Some are unavoidable, as they arise from a general unsteadiness of the soil or building in which the instruments are placed; which, though too minute to be noticed in any other way, become appreciable in delicate astronomical observations; others, again, are consequences of imperfect workmanship, as where an instrument once well adjusted will not remain so, but keeps deviating and shifting. But the most important of this class of errors arise from the non-existence of natural indications, *other* than those afforded by astronomical observations themselves, whether an instrument has or has not the exact position, with respect to the horizon and its cardinal points, the axis of the earth, or to other principal astronomical lines and circles, which it ought to have to fulfil properly its objects.

(137.) Now, with respect to the first two classes of error, it must be observed, that, in so far as they cannot be reduced to known laws, and thereby become subjects of calculation and due allowance, they actually vitiate, to their full extent, the results of any observations in which they subsist. Being, however, in their nature casual and accidental, their effects necessarily lie sometimes one way, sometimes the other; sometimes diminishing, sometimes tending to increase the results. Hence, by greatly multiplying observations, under varied circumstances, by avoiding unfavourable, and taking advantage of favourable circumstances of weather, or otherwise using opportunity to advantage — and finally, by taking the *mean or average* of the results obtained, this class of errors may be so far *subdued*, by setting them to destroy one another, as no longer sensibly to vitiate any theoretical or practical conclusion. This is the great and indeed only resource against such errors, not merely to the astronomer, but to the investigator of numerical results in every department of physical research.

(138.) With regard to errors of adjustment and workmanship, not only the *possibility*, but the *certainty* of their existence, in every imaginable form, in all instruments, must be contemplated. Human hands or machines never formed a circle, drew a straight line, or erected a perpendicular, nor ever placed an instrument in *perfect* adjustment, unless accidentally; and then only during an instant of time. This does not prevent, however, that a great approximation to all these desiderata should be attained. But it is the peculiarity of astronomical observation to be the *ultimate means of detection* of all mechanical defects which elude by their minuteness every other mode of detection. What the eye cannot discern nor the touch perceive, a course of astronomical observations will make distinctly evident. The imperfect products of man's hands are here tested by being brought into comparison under very great magnifying powers (corresponding in effect to a great increase in acuteness of perception) with the perfect workmanship of nature; and there is none which will bear the trial. Now, it may seem like arguing in a vicious circle, to deduce theoretical conclusions and laws from observation, and then to turn round upon the instruments with which those observations were made, accuse them of imperfection, and attempt to detect and rectify their errors by means of the very laws and theories which they have helped us to a knowledge of. A little consideration, however, will suffice to show that such a course of proceeding is perfectly legitimate.

(139.) The steps by which we arrive at the laws of natural phenomena, and especially those which depend for their verification on numerical determinations, are necessarily successive. Gross results and palpable laws are arrived at by rude observation with coarse instruments, or without any instruments at all, and are expressed in language which is not to be considered as absolute, but is to be interpreted with a degree of latitude commensurate to the imperfection of the observations themselves. These results are corrected and refined by nicer scrutiny, and with more delicate means. The first rude expressions of the laws which embody them are perceived to be inexact. The language used in their expression is corrected, its terms more rigidly defined, or fresh terms introduced, until the new state of language and terminology is brought to fit the improved state of knowledge of facts. In the progress of this scrutiny subordinate laws are brought into view which still further modify both the verbal statement and numerical results of those which first offered themselves to our notice; and when these are traced out and reduced to certainty, others, again, subordinate to them, make their appearance, and become subjects of further inquiry. Now, it invariably happens (and the reason is evident) that the first glimpse we catch of such subordinate laws — the first form in which they are dimly shadowed out to our minds —is that of *errors*. We perceive a discordance between what we *expect*, and what we *find*. The first occurrence of such a discordance we attribute to accident. It happens again and again; and we begin to suspect our instruments. We then inquire, to what amount of error their determinations can, *by possibility*, be liable. If their *limit of possible error* exceed the observed deviation, we at once condemn the instrument, and set about improving its construction or adjustments. Still the same deviations occur, and, so far from being palliated, are more marked and better defined than before. We are now sure that we are on the traces of a law of nature, and we pursue it till we have reduced it to a definite statement, and verified it by repeated observation, under every variety of circumstances.

(140.) Now, in the course of this inquiry, it will not fail to happen that other discordances will strike us. Taught by experience, we suspect the existence of some natural law, before unknown; we tabulate (*i. e.* draw out in order) the results of our observations; and we perceive, in this synoptic statement of them, distinct indications of a regular progression. Again we improve or vary our instruments, and we now lose sight of this supposed new law of nature altogether, or find it replaced by some other, of a totally different character. Thus we are led to suspect an instrumental cause for what we have noticed. We examine, therefore, the *theory* of our instrument; we suppose defects in its structure, and, by the aid of geometry, we trace their influence in introducing *actual errors* into its indications. These errors have *their laws*, which, so long as we have no knowledge of causes to guide us, may be confounded with laws of nature, as they are mixed up with them in their effects. They are not fortuitous, like errors of observation, but, as they arise from sources inherent in the instrument, and unchangeable while it and its adjustments remain unchanged, they are reducible to fixed and ascertainable forms; each particular defect, whether of structure or adjustment, producing its own appropriate *form* of error. When these are thoroughly investigated, we recognize among them one which coincides in its nature and progression with that of our observed discordances. The mystery is at once solved. We have detected, by direct observation, an instrumental defect.

(141.) It is, therefore, a chief requisite for the practical astronomer to make himself completely familiar with the *theory* of his instruments. By this alone is he enabled at once to decide what *effect* on his observations any given imperfection of structure or adjustment will produce in any given circumstances under which an observation can be made. This alone also can place him in a condition to derive available and practical means of destroying and eliminating altogether the influence of such imperfections, by so arranging his observations, that it shall affect their results in opposite ways, and that its influence shall thus disappear from their mean, which is one of the chief modes by which precision is attained in practical astronomy. Suppose, for example, the principle of an instrument required that a circle should be concentric with the axis on which it is made to turn. As this is a condition which no workmanship can *exactly* fulfil, it becomes necessary to inquire what errors will be produced in observations made and registered on the faith of such an instrument, by any assigned deviation in this respect; that is to say, what would be the disagreement between observations made with it and with one absolutely perfect, could such be obtained. Now, simple geometrical considerations suffice to show —1st. that if the axis be excentric by a given fraction (say one thousandth part) of the radius of the circle, all angles read off on that part of the circle *towards* which the excentricity lies, will appear by that fractional amount too small, and all on the opposite side too large. And, 2dly, that *whatever* be the amount of the excentricity, and on *whatever* part of the circle any proposed angle is measured, the effect of the error in question on the result of observations depending on the graduation of its circumference (or limb, as it is technically called) will be completely annihilated by the very easy method of always reading off the divisions on two diametrically opposite points of the circle, and taking a mean; for the effect of excentricity is always to increase the arc representing the angle in question on one side of the circle, by just the same quantity by which it diminishes that on the other. Again, suppose that the proper use of the instrument required that this axis should be exactly parallel to that of the earth. As it never can be *placed* or remain so, it becomes a question, what amount of error will arise, in its use, from any assigned deviation, whether in a horizontal or vertical plane, from this precise position. Such inquiries constitute the theory of instrumental errors; a theory of the utmost importance to practice, and one of which a complete knowledge will enable an observer, with moderate instrumental means, often to attain a degree of precision which might seem to belong only to the most refined and costly. This theory, as will readily be apprehended, turns almost entirely on considerations of pure geometry, and those for the most part not difficult. In the present work, however, we have no further concern with it. The Astronomical instruments we propose briefly to describe in this chapter will be considered as perfect both in construction and adjustment.[1]

(142.) As the above remarks are very essential to a right understanding of the philosophy of our subject and the spirit of astronomical methods, we shall elucidate them by taking one or two special cases. Observant persons, before the invention of astronomical instruments, had already concluded the apparent diurnal motions of the stars to be performed in circles about fixed poles in the heavens, as shown in the foregoing chapter. In drawing this conclusion, however, refraction was entirely overlooked, or, if forced on their notice by its great magnitude in the immediate neighbourhood of the horizon, was regarded as a local irregularity, and, as such, neglected, or slurred over. As soon, however, as the diurnal paths of the stars were attempted to be traced by instruments, even of the coarsest kind, it became evident that the notion of exact circles described about one and the same pole would not represent the phenomena correctly, but that, owing to some cause or other, the apparent diurnal orbit of every star is distorted from a circular into an oval form, its lower segment being *flatter* than its upper; and the deviation being greater the nearer the star approached the horizon, the effect being the same as if the circle had been squeezed upwards from below, and the lower parts more than the higher. For such an effect, as it was soon found to arise from no casual or instrumental cause, it became necessary to seek a natural one; and refraction readily occurred, to solve the difficulty. In fact, it is a case precisely analogous to what we have already noticed (art. 47), of the apparent distortion of the sun near the horizon, only on a larger scale, and traced up to greater altitudes. This new law once established, it became necessary to modify the expression of that anciently received, by inserting in it a *salvo* for the effect of refraction, or by making a distinction between the *apparent* diurnal orbits, as affected by refraction, and the *true* ones cleared of that effect. This distinction between the *apparent* and the *true*—between the *uncorrected* and *corrected*—between the *rough* and *obvious*, and the *refined* and *ultimate*—is of perpetual occurrence in every part of astronomy.

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[1] The principle on which the chief adjustments of two or three of the most useful and common instruments, such as the transit, the equatorial, and the sextant, are performed, are, however, noticed, for the convenience of readers who may use such instruments without going farther into the arcana of practical astronomy.

(143.) Again. The first impression produced by a view of the diurnal movement of the heavens is that *all* the heavenly bodies perform this revolution in one common period, viz. *a day*, or 24 hours. But no sooner do we come to examine the matter *instrumentally*, i. e. by noting, by time-keepers, their successive arrivals on the meridian, than we find differences which cannot be accounted for by any error of observation. All the *stars*, it is true, occupy the same interval of time between their successive appulses to the meridian, or to any vertical circle; but *this* is a very different one from that occupied by the sun. It is palpably shorter; being, in fact, only 23$^{\text{h}}$ 56′ 4·09″, instead of 24 hours, such hours as our common clocks mark. Here, then, we have already *two different days*, a *sidereal* and a *solar;* and if, instead of the sun, we observe the moon, we find a third, much longer than either,—a *lunar* day, whose average duration is 24$^{\text{h}}$ 54$^{\text{m}}$ of our ordinary time, which last is *solar* time, being of necessity conformable to the *sun's* successive re-appearances, on which all the business of life depends.

(144.) Now, all the *stars* are found to be unanimous in giving the same exact duration of 23$^{\text{h}}$ 56′ 4·09″, for the *sidereal day;* which, therefore, we cannot hesitate to receive as the period in which the earth makes one revolution on its axis. We are, therefore, compelled to look on the sun and moon as exceptions to the general law; as having a different nature, or at least a different relation to us, from the stars; and as having motions, real or apparent, of their own, independent of the rotation of the earth on its axis. Thus a great and most important distinction is disclosed to us.

(145.) To establish these facts, almost no apparatus is required. An observer need only station himself to the north of some well-defined vertical object, as the angle of a building, and, placing his eye exactly at a certain fixed point (such as a small hole in a plate of metal nailed to some immoveable support), notice the successive disappearances of any star behind the building, by a watch.[1] When he observes the sun, he must shade his eye with a dark-coloured or smoked glass, and notice the moments when its western and eastern edges successively come up to the wall, from which, by taking half the interval, he will ascertain (what he cannot directly *observe*) the moment of disappearance of its centre.

(146.) When, in pursuing and establishing this general fact, we are led to attend more nicely to the times of the daily arrival of the sun on the meridian, irregularities (such they first seem to be) begin to make their appearance. The intervals between two successive arrivals are not the same at all times of the year. They are sometimes greater, sometimes less, than 24 hours, as shown by the clock; that is to say, the *solar day* is not always of the same length. About the 21st of December, for example, it is half a minute *longer*, and about the same day of September nearly as much *shorter*, than its *average duration*. And thus a distinction is again pressed upon our notice between the *actual* solar day, which is never two days in succession alike, and the *mean solar day* of 24 hours, which is an average of all the solar days throughout the year. Here, then, a new source of inquiry opens to us. The sun's apparent motion is not only not the same with the stars, but it is not (as the latter is) uniform. It is subject to fluctuations, whose laws become matter of investigation. But to pursue these laws, we require nicer means of observation than what we have described, and are obliged to call in to our aid an instrument called the *transit instrument*, especially destined for such observations, and to attend minutely to all the causes of irregularity in the going of clocks and watches which may affect our reckoning of time. Thus we become involved by degrees in more and more delicate instrumental inquiries; and we speedily find that, in proportion as we ascertain the amount and law of one great or leading fluctuation, or inequality, as it is called, of the sun's diurnal motion, we bring into view others continually smaller and smaller, which were before obscured, or mixed up with errors of observation and instrumental imperfections. In short, we may not inaptly compare the *mean* length of the solar day to the mean or average height of water in a harbour, or the general level of the sea unagitated by tide or waves. The great annual fluctuation above noticed may be compared to the daily variations of level produced by the tides, which are nothing but enormous waves extending over the whole ocean, while the smaller subordinate inequalities may be assimilated to waves ordinarily so called, on which, when large, we perceive lesser undulations to ride, and on these, again, minuter ripplings, to the series of whose subordination we can perceive no end.

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[1] This is an excellent practical method of ascertaining the rate of a clock or watch, being exceeding accurate if a few precautions are attended to; the chief of which is, to take care that that part of the edge behind which the star (a bright one, *not a planet*) disappears shall be quite smooth; as otherwise variable refraction may transfer the point of disappearance from a protuberance to a notch, and thus vary the moment of observation unduly. This is easily secured, by nailing up a smooth-edged board. The verticality of its edge should be insured by the use of a plumb-line.

(147.) With the causes of these irregularities in the solar motion we have no concern at present; their explanation belongs to a more advanced part of our subject; but the distinction between the solar and sidereal days, as it pervades every part of astronomy, requires to be early introduced, and never lost sight of. It is, as already observed, the *mean* or *average* length of the solar day, which is used in the civil reckoning of time. It commences at midnight, but astronomers, even when they use mean solar time, depart from the civil reckoning, commencing their day at noon, and reckoning the hours from 0 round to 24. Thus, 11 o'clock in the forenoon of the second of January, in the civil reckoning of time, corresponds to January 1 day 23 hours in the astronomical reckoning; and one o'clock in the afternoon of the former, to January 2 days 1 hour of the latter reckoning. This usage has its advantages and disadvantages, but the latter seem to preponderate; and it would be well if, in consequence, it could be broken through, and the civil reckoning substituted. *Uniformity in nomenclature and modes of reckoning in all matters relating to time, space, weight, measure, &c., is of such vast and paramount importance in every relation of life as to outweigh every consideration of technical convenience or custom.*[1]

(148.) Both astronomers and civilians, however, who inhabit different points of the earth's surface, differ from each other in their reckoning of time; as it is obvious they must, if we consider that, when it is noon at one place, it is midnight at a place diametrically opposite; sunrise at another; and sunset, again, at a fourth. Hence arises considerable inconvenience, especially as respects places differing very widely in situation, and which may even in some critical cases involve the mistake of a whole day. To obviate this inconvenience, there has lately been introduced a system of reckoning time by mean solar days and parts of a day counted from a fixed instant, common to all the world, and determined by no *local* circumstance, such as noon or midnight, but by the motion of the sun among the stars. Time, so reckoned, is called equinoctial time; and is numerically the same, at the same instant, in every part of the globe. Its origin will be explained more fully at a more advanced stage of our work.

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[1] The only disadvantage to astronomers of using the civil reckoning is this — that their observations being chiefly carried on during the night, the day of their date will, in this reckoning, always have to be changed at midnight, and the former and latter portion of every night's observations will belong to two differently numbered civil days of the month. There is no denying this to be an inconvenience. Habit, however, would alleviate it; and *some* inconveniences must be cheerfully submitted to by all who resolve to act on general principles. All other classes of men, whose occupation extends to the night as well as day, submit to it, and find their advantage in doing so.

(149.) Time is an essential element in astronomical observation, in a twofold point of view: — 1st, As the representative of angular motion. The earth's diurnal motion being uniform, every star describes its diurnal circle uniformly; and the time elapsing between the passage of the stars in succession across the meridian of any observer becomes, therefore, a direct measure of their differences of right ascension. 2dly, As the fundamental element (or natural *independent variable*, to use the language of geometers) in all dynamical theories. The great object of astronomy is the determination of the laws of the celestial motions, and their reference to their proximate or remote causes. Now, the statement of the *law* of any observed motion in a celestial object can be no other than a proposition declaring what has been, is, and will be, the real or apparent situation of that object *at any time*, past, present, or future. To compare such laws, therefore, with observation, we must possess a register of the observed situations of the object in question, and of the *times when* they were observed.

(150.) The measurement of time is performed by clocks, chronometers, clepsydras, and hour-glasses. The two former are alone used in modern astronomy. The hour-glass is a coarse and rude contrivance for measuring, or rather counting out, fixed portions of time, and is entirely disused. The clepsydra, which measured time by the gradual emptying of a large vessel of water through a determinate orifice, is susceptible of considerable exactness, and was the only dependence of astronomers before the invention of clocks and watches. At present it is abandoned, owing to the greater convenience and exactness of the latter instruments. In one case only has the revival of its use been proposed; viz. for the accurate measurement of very small portions of time, by the flowing out of mercury from a small orifice in the bottom of a vessel, kept constantly full to a fixed height. The stream is intercepted at the moment of noting any event, and directed aside into a receiver, into which it continues to run, till the moment of noting any other event, when the intercepting cause is suddenly removed, the stream flows in its original course, and ceases to run into the receiver. The *weight* of mercury received, compared with the weight received in an interval of time observed by the clock, gives the interval between the events observed. This ingenious and simple method of resolving, with all possible precision, a problem of much importance in many physical inquiries, is due to the late Captain Kater.

(151.) The pendulum clock, however, and the balance watch, with those improvements and refinements in its structure which constitute it emphatically a *chronometer*,[1] are the instruments on which the astronomer depends for his knowledge of the lapse of time. These instruments are now brought to such perfection, that an habitual irregularity in the *rate* of going, to the extent of a single second in twenty-four hours in two consecutive days, is not tolerated in one of good character; so that any interval of time less than twenty-four hours may be certainly ascertained within a few tenths of a second, by their use. In proportion as intervals are longer, the risk of error, as well as the amount of error risked, becomes greater, because the accidental errors of many days may accumulate; and causes producing a slow progressive change in the rate of going may subsist unperceived. It is not safe, therefore, to trust the determination of time to clocks, or watches, for many days in succession, without checking them, and ascertaining their errors by reference to natural events which we know to happen, day after day, at equal intervals. But if this be done, the longest intervals may be fixed with the same precision as the shortest; since, in fact, it is then only the times intervening between the first and the last moments of such long intervals, and such of those periodically recurring events adopted for our points of reckoning, as occur within twenty-four hours respectively of either, that we measure by artificial means. The whole days are counted out for us by nature; the fractional parts only, at either end, are measured by our clocks. To keep the reckoning of the integer days correct, so that none shall be lost or counted twice, is the object of the calendar. Chronology marks out the order of succession of events, and refers them to their proper years and days; while chronometry, grounding its determinations on the precise observation of such regularly periodical events as can be conveniently and exactly subdivided, enables us to fix the moments in which phenomena occur, with the last degree of precision.

(152.) In the *culmination* or *transit* (*i. e.* the passage across the meridian of an observer,) of every star in the heavens, he is furnished with such a regularly periodical natural event as we allude to. Accordingly, it is to the *transits* of the brightest and most conveniently situated fixed stars that astronomers resort to ascertain their exact time, or, which comes to the same thing, to determine the exact amount of error of their clocks.

(153.) Before we describe the instrument destined for the purpose of observing such culminations, however, or those intended for the measurement of angular intervals in the sphere, it is requisite to place clearly before the reader the principle on which the telescope is applied in astronomy to the precise determination of a direction in space, — that, namely, of the visual ray by which we see a star or any other distant object.

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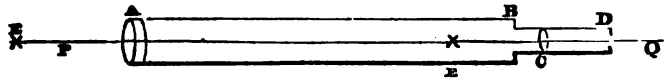
[1] Χρονος, time; *μετρειν*, to measure.

(154.) The telescope most commonly used in astronomy for these purposes is the refracting telescope, which consists of an object-glass (either single, or as is now almost universal, double, forming what is called in optics, an acromatic combination) A; a tube A B, into which the brass cell of the object-glass is firmly screwed, and an eye-lens C, for which is often substituted a combination of glasses designed to increase the magnifying power of the telescope, or otherwise give more distinctness of vision according to optical principles which we have no occasion here to refer to. This also is fitted into a cell, which is screwed firmly into the end B of the tube, so that object-glass, tube, and eye-glass may be considered as forming one piece, invariable in the relative position of its parts.

Fig. 14.



(155) The line P Q joining the centres of the object and eye-glasses and produced, is called the *axis*, or *line of collimation* of the telescope. And it is evident, that the situation of this line holds a fixed relation to the tube and its appendages, so long as the object and eye-glasses maintain *their* fixity in this respect.

(156.) Whatever distant object E, this line is directed to, an inverted picture or image of that object F is formed (according to the principles of optics), in the focus of the object-glass, and may there be viewed as if it *were a real object*, through the eye-lens C, which (if of short focus) enables us to *magnify* it just as such a lens would magnify a material object in the same place.

(157.) Now as this image is formed and viewed in the air, being itself immaterial and impalpable — nothing prevents our placing in that very place F in the axis of the telescope, a real, substantial object of very definite form and delicate make, such as a fine metallic point, as of a needle[*] — or better still, a cross formed by two very fine threads (spider-lines), thin metallic wires, or lines drawn on glass intersecting each other at right angles — and whose intersection is all but a mathematical point. If such a point, wire, or cross be carefully placed and firmly fixed in the exact focus F, both of the object and eye-glass, it will be seen through the latter *at the same time*, and occupying *the same precise place* as the image of the distant star E. The magnifying power of the lens renders perceptible the smallest deviation from perfect coincidence, which, should it exist, is a proof, that the axis Q P is not directed rigorously towards E. In that case, a fine motion (by means of a screw duly applied), communicated to the telescope, will be necessary to vary the direction of the axis till the coincidence is rendered perfect. So precise is this mode of pointing found in practice, that the axis of a telescope may be directed towards a star or other definite celestial object without an error of more than a few tenths of a second of angular measure.

(158.) This application of the telescope may be considered as completely annihilating that part of the error of observation which might otherwise arise from an erroneous estimation of the direction in which an object lies from the observer's eye, or from the centre of the instrument. It is, in fact, the grand source of all the precision of modern astronomy, without which all other refinements in instrumental workmanship would be thrown away; the errors capable of being committed in pointing to an object, without such assistance, being far greater than what could arise from any but the very coarsest graduation.[1] In fact, the telescope thus applied becomes, with respect to angular, what the microscope is with respect to linear dimension. By concentrating attention on its smallest parts, and magnifying into palpable intervals the minutest differences, it enables us not only to scrutinise the form and structure of the objects to which it is pointed, but to refer their apparent places, with all but geometrical precision, to the parts of any scale with which we propose to compare them.

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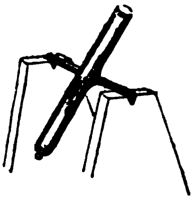
[1] The honour of this capital improvement has been successfully vindicated by Derham (Phil. Trans. xxx. 603) to our young, talented, and unfortunate countryman Gascoigne, from his correspondence with Crabtree and Horrockes, in his (Derham's) possession. The passages cited by Derham from these letters leave no doubt that, so early as 1640, Gascoigne had applied telescopes to his quadrants and sextants, *with threads in the common focus of the glasses*; and had even carried the invention so far as to illuminate the field of view by artificial light, which he found "*very helpful when the moon appeareth not, or it is not otherwise light enough.*" These inventions were freely communicated by him to Crabtree, and through him to his friend Horrockes, the pride and boast of British astronomy; both of whom expressed their unbounded admiration of this and many other of his delicate and admirable improvements in the art of observation. Gascoigne, however, perished, at the age of twenty-three, at the battle of Marston Moor; and the premature and sudden death of Horrockes, at a yet earlier age, will account for the temporary oblivion of the invention. It was revived, or re-invented, in 1667, by Picard and Auzout (Lalande, Astron. 2310), after which its use became universal. Morin, even earlier than Gascoigne (in 1635), had proposed to substitute the telescope for plain sights; but it is the thread or wire stretched in the focus with which the image of a star can be brought to exact coincidence, which gives the telescope its advantage in practice; and the idea of this does not seem to have occurred to Morin. See Lalande, *ubi suprà*.

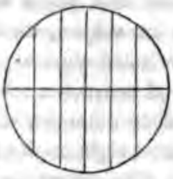
(159.) We now return to our subject, the determination of time by the transits or culminations of celestial objects. The instrument with which such culminations are observed is called a *transit instrument*. It consists of a telescope firmly fastened on a horizontal axis directed to the east and west points of the horizon, or at right angles to the plane of the meridian of the place of observation. The extremities of the axis are formed into cylindrical pivots of exactly equal diameters, which rest in notches formed in metallic supports, bedded (in the case of large instruments) on strong pieces of stone, and susceptible of nice adjustment by screws, both in a vertical and horizontal direction. By the former adjustment, the axis can be rendered precisely horizontal, by *levelling* it with a *level* made to rest on the pivots. By the latter adjustment the axis is brought precisely into the east and west directions, the criterion of which is furnished by the observations themselves made with the instrument, in a manner presently to be explained, or by a well-defined object, called a *meridian-mark*, originally determined by such observations, and then, for convenience of ready reference, permanently established, at a great distance, exactly in a *meridian line* passing through the central point of the whole instrument. It is evident, from this description, that, if the axis, or line of collimation of the telescope be once well adjusted at right angles to the axis of the transit, it will never quit the plane of the meridian, when the instrument is turned round on its axis of rotation.

Fig. 15.



(160.) In the focus of the eye-piece, and at right angles to the length of the telescope, is placed, not a single cross, as in our general explanation in art. 157, but a system of one horizontal and several equidistant vertical threads or wires, (five or seven are more usually employed,) as represented in the annexed figure, which always appear in *the field of view*, when properly illuminated, by day by the light of the sky, by night by that of a lamp introduced by a contrivance not necessary here to explain. The

Fig. 16.

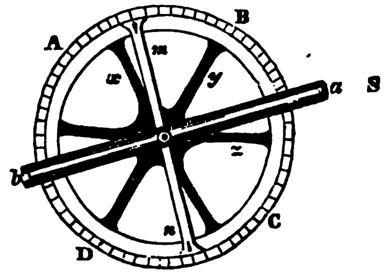


is system of wires may be altered by adjusting screws, giving l (horizontal) motion; and it is by this means brought to such that the middle one of the vertical wires shall intersect *the line tion* of the telescope, where it is arrested and permanently In this situation it is evident that the middle thread will be a esentation of that portion of the celestial meridian to which pe is pointed; and when a star is seen to cross this wire in pe, it is in the act of culminating, or passing the celestial meri- instant of this event is noted by the clock or chronometer, ns an indispensable accompaniment of the transit instrument. precision, the moment of its crossing all the vertical threads nd a mean taken, which (since the threads are equidistant) exactly the same result, were all the observations perfect, and other words, a clock may be compared with the earth's diurnal motion by a single observation, without risk of greater error. By multiplying observations, of course, a yet greater degree of precision may be obtained.

(162.) The plane described by the line of collimation of a transit ought to be that of the meridian of the place of observation. To ascertain whether it is so or not, celestial observation must be resorted to. Now, as the meridian is a great circle passing through the pole, it necessarily bisects the diurnal circles described by all the stars, all which describe the two semicircles so arising in equal intervals of 12 sidereal hours each. Hence, if we choose a star whose whole diurnal circle is above the horizon, or which never sets, and observe the moments of its upper and lower transits across the middle wire of the telescope, if we find the two semidiurnal portions east and west of the plane described by the telescope to be described in *precisely* equal times, we may be sure that plane is the meridian.

(163.) The angular intervals measured by means of the transit instrument and clock are arcs of the equinoctial, intercepted between circles of declination passing through the objects observed; and their measurement, in this case, is performed by no artificial graduation of circles, but by the help of the earth's diurnal motion, which carries equal arcs of the equinoctial across the meridian, in equal times, at the rate of 15° per sidereal hour. In all other cases, when we would measure angular intervals, it is

Fig. 17.



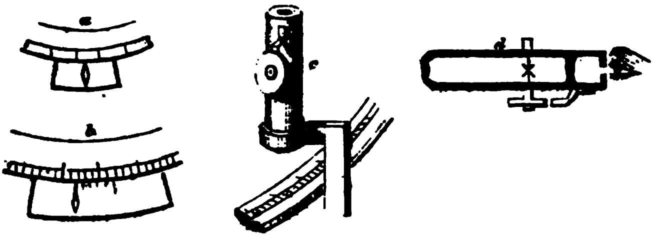
necessary to have recourse to circles, or portions of circles, constructed of metal or other firm and durable material, and mechanically subdivided into equal parts, such as degrees, minutes, &c. The simplest and most obvious mode in which the measurement of the angular interval between two directions in space can be performed is as follows. Let ABCD be a circle, divided into 360 degrees, (numbered in order from any point 0° in the circumference, round to the same point again,) and connected with its centre by spokes or rays, x, y, z, firmly united to its circumference or

the centre let a circular hole be pierced, in which shall move
actly fitting it, carrying a tube, whose axis, ab, is exactly
the plane of the circle, or perpendicular to the pivot; and also
m, n, at right angles to it, and forming one piece with the tube
s; so that the motion of the axis on the centre shall carry the
rms smoothly round the circle, to be arrested and fixed at any
please, by a contrivance called a *clamp*. Suppose, now, we
sure the angular interval between two fixed objects, S, T. The
e circle must first be adjusted so as to pass through them both,
eably fixed and maintained in that position. This done, let the
the tube be directed to one of them, S, and *clamped*. Then
k on the arm m point either exactly to some one of the divisions
b, or between two of them adjacent. In the former case, the
ust be noted as *the reading* of the arm m. In the latter, the
part of one whole interval between the consecutive divisions by
mark on m *surpasses* the last inferior division must be estimated
d by some mechanical or optical means. (See art. 165.) The
d fractional part thus noted, and reduced into degrees, minutes,
ds, is to be set down as the *reading of the limb* corresponding
sition of the *tube* ab, where it points to the object S. The
then be done for the object T; the tube pointed to it, and the

originates an arm, F, carrying at its extremity an index, or other proper mark, to point out and read off the exact division of the circle at B, the point close to it. It is evident that, as the telescope and circle revolve through any angle, the part of the limb of the latter, which by such revolution is carried past the index F, will measure the angle described. This is the most usual mode of applying divided circles in astronomy.

(165.) The index F may either be a simple pointer, like a clock hand (*fig. a*); or a vernier (*fig. b*); or, lastly, a compound microscope (*fig.* c), represented in section in *fig. d*, and furnished with a cross in the common focus of its object and eye-glass, moveable by a fine-threaded screw, by which the intersection of the cross may be brought to exact coincidence with the image of the nearest of the divisions of the circle formed in the focus of the object lens upon the very same principle with that explained, art. 157 for the pointing of the telescope, only that here the fiducial cross is made moveable; and by the turns and parts of a turn of the screw required for this purpose the distance of that division from the original or zero point of the microscope may be estimated. This simple but delicate contrivance gives to the reading off of a circle a degree of accuracy only limited by the power of the microscope, and the perfection with which a screw can be executed, and places the subdivision of angles on the same footing of optical certainty which is introduced into their measurement by the use of the telescope.

Fig. 19.



(166.) The exactness of the result thus obtained must depend, 1st, on the precision with which the tube *a b* can be pointed to the objects; 2dly, on the accuracy of graduation of the limb; 3dly, on the accuracy with which the subdivision of the intervals between any two consecutive graduations can be performed. The mode of accomplishing the latter object with any required exactness has been explained in the last article. With regard to the graduation of the limb, being merely of a mechanical nature,

ass it without remark, further than this, that, in the present
trument-making, the amount of error from this source of inac-
educed within very narrow limits indeed.[1] With regard to the
st be obvious that, if the sights *a b* be nothing more than
ses, or pin-holes at the ends of a hollow tube, or an eye-hole
and a cross at the other, no greater nicety in pointing can be
an what simple vision with the naked eye can command. But
of these simple but coarse contrivances, the tube itself be con-
a telescope, having an object-glass at *b*, an eye-piece at *a*, and
ross in their common focus, as explained in art. 157; and if
of the tube on the limb of the circle be arrested when the
rought just into coincidence with the intersectional point of
it is evident that a greater degree of exactness may be attained
ting of the tube than by the unassisted eye, in proportion to
ying power and distinctness of the telescope used.

he simplest mode in which the measurement of an angular in-
be executed, is what we have just described; but, in strictness,
s applicable only to terrestrial angles, such as those occupied
ible horizon by the objects which surround our station, — be-
only remain stationary during the interval while the telescope
on the limb from one object to the other. But the diurnal

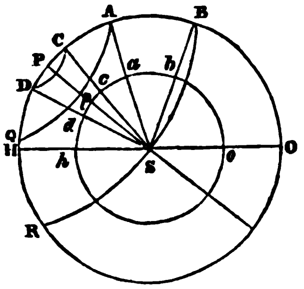
telescope is well clamped, we equally well secure the position of its diurnal circle as if we continued to *see* it ever so long. The reading off of the limb may then be performed at leisure; and when another star comes round *into the plane* of the circle, we may unclamp the telescope, and a similar observation will enable us to assign the place of *its* diurnal circle on the limb: and the observations may be repeated alternately, every day, as the stars pass, till we are satisfied with their result.

(168.) This is the principle of the mural circle, which is nothing more than such a circle as we have described in art. 163, firmly supported, in the plane of the meridian, on a long and powerful horizontal axis. This axis is let into a massive pier, or wall, of stone (whence the name of the instrument), and so secured by screws as to be capable of adjustment both in a vertical and horizontal direction; so that, like the axis of the transit, it can be maintained in the exact direction of the east and west points of the horizon, the plane of the circle being consequently truly meridional.

(169.) The meridian, being at right angles to all the diurnal circles described by the stars, its arc intercepted between any two of them will measure the least distance between these circles, and will be equal to the difference of the declinations, as also to the difference of the *meridian altitudes* of the objects—at least when corrected for refraction. These differences, then, are the angular intervals *directly* measured by the mural circle. But from these, supposing the law and amount of refraction known, it is easy to conclude, not their differences only, but the quantities themselves, as we shall now explain.

(170.) The declination of a heavenly body is the complement of its distance from the pole. The pole, being a point in the meridian, might be directly observed on the limb of the circle, if any star stood *exactly* therein; and thence the *polar distances*, and, of course, the declinations of all the rest, might be at once determined. But this not being the case, a bright star as near the pole as can be found is selected, and observed in its *upper* and *lower* culminations; that is, when it passes the meridian *above* and *below* the pole. Now, as its distance from the pole remains the same, the difference of reading off the circle in the two cases is, of course (when corrected for refraction), equal to twice the polar distance of the star; the arc intercepted on the limb of the circle being, in this case, equal to the angular diameter of the star's diurnal circle. In the annexed diagram, H P O represents the celestial meridian, P the pole, B R, A Q, C D the diurnal circles of stars which arrive on the meridian at B, A, and C in their upper and at R, Q, D in their lower culminations, of which D and Q happen above the horizon H O. P is the pole; and if

Fig. 20.



we suppose *hpo* to be the mural circle, having S for its centre, *bacpd* will be the points on its circumference corresponding to B A C P D in the heavens. Now the arcs *ba*, *bc*, *bd*, and *cd* are given immediately by observation; and since C P = P D, we have also *cp* = *pd*, and each of them = ½ *cd*, consequently the place of *the polar point*, as it is called, upon the limb of the circle becomes known, and the arcs *p b*, *p a*, *p c*, which represent on the circle the *polar-distances* required, become also known.

(171.) The situation of the pole star, which is a very brilliant one, is eminently favourable for this purpose, being only about a degree and an half from the pole; it is, therefore, the star usually and almost solely chosen for this important purpose; the more especially because, both its culminations taking place at great and not very different altitudes, the refractions by which they are affected are of small amount, and differ but slightly from each other, so that their correction is easily and safely applied. The brightness of the pole star, too, allows it to be easily observed in the daytime. In consequence of these peculiarities, this star is one of constant resort with astronomers for the adjustment and verification of instruments of almost every description. In the case of the transit, for instance, it furnishes an excellent object for the application of the method of testing the meridional situation of the instrument described in art. 162, in fact, the most advantageous of any for that purpose, owing to its being the most remote from the zenith, at its upper culmination, of all bright stars observable both above and below the pole.

(172.) The place of the *polar point* on the limb of the mural circle once determined, becomes an origin, or zero point, from which the polar distances of all objects, referred to other points on the same limb, reckon. It matters not whether the actual commencement 0° of the graduations stand there, or not; since it is only by the *differences* of the readings that the arcs on the limb are determined; and hence a great advantage is obtained in the power of commencing anew a fresh series of observations, in which a different part of the circumference of the circle shall be employed, and different graduations brought into use, by which inequalities of division may be detected and neutralized. This is accomplished practically by detaching the telescope from its old bearings on the circle, and fixing it afresh, by screws or clamps, on a different part of the circumference.

(173.) A point on the limb of the mural circle, not less important than the *polar point*, is the *horizontal point*, which, being once known, becomes in like manner an origin, or zero point, from which altitudes are reckoned. The principle of its determination is ultimately nearly the same with that of the polar point. As no star exists in the celestial horizon, the observer must seek to determine two points on the limb, the one of which shall be precisely as far *below* the horizontal point as the other is above it. For this purpose, a star is observed at its culmination on one night, by pointing the telescope directly to it, and the next, by pointing to *the image of the same star reflected* in the still, unruffled surface of a fluid at perfect rest. Mercury, as the most reflective fluid known, is generally chosen for that use. As the surface of a fluid at rest is necessarily horizontal, and as the angle of reflection, by the laws of optics, is equal to that of incidence, this image will be just as much depressed below the horizon as a star itself is above it (allowing for the difference of refraction at the moment of observation). The arc intercepted on the limb of the circle between the star and its reflected image thus consecutively observed, when corrected for refraction, is the double altitude of the star, and its point of bisection the horizontal point. The reflecting surface of a fluid so used for the determination of the altitudes of objects is called an *artificial horizon*.[1]

(174.) The mural circle is, in fact, at the same time, a transit instrument; and, if furnished with a proper system of vertical wires in the focus of its telescope, may be used as such. As the axis, however, is only supported at one end, it has not the strength and permanence necessary for the more delicate purposes of a transit; nor can it be verified, as a transit may, by the *reversal* of the two ends of its axis, east or west.

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[1] By a peculiar and delicate manipulation and management of the setting bisection and reading off of the circle, aided by the use of a moveable horizontal micrometic wire in the focus of the object-glass, it is found practicable to observe a slow moving star (as the pole star) *on one and the same night*, both by reflection and direct vision, sufficiently near to either culmination to give the horizontal point, without risking the change of refraction in twenty-four hours; so that this source of error is thus completely eliminated.

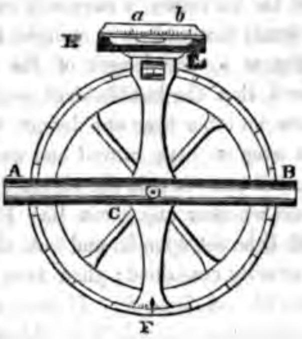
owever, prevents a divided circle being permanently fastened of a transit instrument, either near to one of its extremities, the telescope, so as to revolve with it, the reading off being by one or more microscopes fixed on one of its piers. Such an is called a TRANSIT CIRCLE, or a MERIDIAN CIRCLE, and serves ultaneous determination of the right ascensions and polar disbjects observed with it; the time of transit being noted by the the circle being read off by the lateral microscopes. There is ntage, when extensive catalogues of small stars have to be this simultaneous determination of both their celestial co-ordihich may be added the facility of applying to the meridian escope of any length and optical power. The construction of ircle renders this highly inconvenient, and indeed impracticable y moderate limits.

The determination of the horizontal point on the limb of an is of such essential importance in astronomy, that the student made acquainted with every means employed for this purpose. the artificial horizon, the plumb-line, the level, and the collihe artificial horizon has been already explained. The plumbe thread or wire, to which is suspended a weight, whose oscilimpeded and quickly reduced to rest by plunging it in water.

impossible to execute, and every tube having some slight curvature; if the convex side be placed upwards the bubble will occupy the higher part, as in the figure (where the curvature is purposely exaggerated). Suppose such a tube, as A B, firmly fastened on a straight bar, C D, and marked at *a b*, two points distant by the length of the bubble; then, if the instrument be so placed that the bubble shall occupy this interval, it is clear that C D can have no other than one definite inclination to the horizon; because, were it ever so little moved one way or other, the bubble would shift its place, and run towards the elevated side. Suppose, now, that we would ascertain whether any given line P Q be horizontal; let the base of the level C D be set upon it, and note the points *a b*, between which the bubble is exactly contained; then turn the level end for end, so that C shall rest on Q, and D on P. If then the bubble continue to occupy the same place between *a* and *b*, it is evident that P Q can be no otherwise than horizontal. If not, the side towards which the bubble runs is highest, and must be lowered. Astronomical levels are furnished with a divided scale, by which the places of the ends of the bubble can be nicely marked; and it is said that they can be executed with such delicacy, as to indicate a single second of angular deviation from exact horizontality. In such levels accident is not trusted to to give the requisite curvature. They are ground and polished internally by peculiar mechanical processes of great delicacy.

(177.) The mode in which a level may be applied to find the horizontal point on the limb of a vertical divided circle may be thus explained; let A B be a telescope firmly fixed to such a circle, D E F, and moveable in one with it on a horizontal axis C, which must be like that of a transit, susceptible of reversal (see art. 161) and with which the circle is inseparably connected. Direct the telescope on some distant well-defined object S, and bisect it by its horizontal wire, and in this position clamp it fast. Let L be a level fastened at right angles to an arm, L E F, furnished with a microscope, or vernier at F, and, if we please, another at E. Let this arm be fitted by grinding on the axis C, but capable of moving smoothly on it without carrying it round, and also of being clamped fast on it, so as to prevent it from moving until required. While the telescope is kept fixed on the object S, let the level be set so as to bring its bubble to the marks *a b*, and clamp it there. Then will the arm L C F have some certain determinate inclination (no matter what) to the horizon. In this position let the circle be read off at F, and then let the whole apparatus be *reversed* by turning its horizontal axis end for end, *without unclamping the level arm* from the axis. This done, by the motion of the whole instrument (level and all) on its axis, restore the *level* to its horizontal posi-

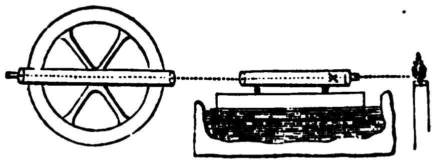
Fig. 22.



he bubble at $a\,b$. Then we are sure that the telescope has
me inclination to the horizon *the other way*, that it had when
S, and the reading off at F will not have been changed. Now
e level, and keeping it nearly horizontal, turn round the circle
, so as to carry back the telescope *through the zenith* to S, and
ition clamp the circle and telescope fast. Then it is evident

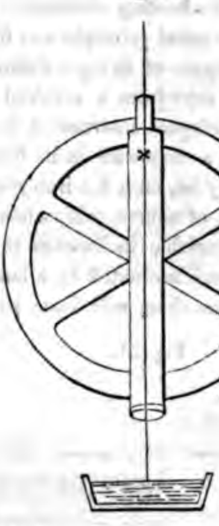
innumerable cases, the most convenient means of ascertaining the *horizontal point*, is that afforded by the floating collimator, an invention of Captain Kater, but of which the optical principle was first employed by Rittenhouse, in 1785, for the purpose of fixing a definite direction in space by the emergence of parallel rays from a material object placed in the focus of a fixed lens. This elegant instrument is nothing more than a small telescope furnished with a cross-wire in its focus, and fastened horizontally, or as nearly so as may be, on a flat iron *float*, which is made to swim on mercury, and which, of course, will, when left to itself, assume always one and the same invariable inclination to the horizon. If the cross-wires of the collimator be illuminated by a lamp, being in the focus of its object-glass, the rays from them will issue parallel, and will therefore be in a fit state to be brought to a focus by the object-glass of any other telescope, in which they will form an image *as if they came from a celestial object in their direction*, i. e. at an altitude equal to their declination. Thus the intersection of the cross of the collimator may be observed *as if it were a star*, and that, however near the two telescopes are to each other. By transferring then, the collimator *still floating* on a vessel of mercury from the one side to the other of a circle, we are furnished with two *quasi-celestial* objects, at precisely equal altitudes, on opposite sides of the centre; and if these be observed in succession with the telescope of the circle, bringing its cross to bisect the image of the cross of the collimator (for which end the wires of the latter cross are purposely set 45° inclined to the horizon), the difference of the readings on its limb will be twice the zenith distance of either; whence, as in the last article, the horizontal or zenith point is immediately determined. Another, and, in many respects, preferable form of the floating collimator, in which the telescope is *vertical*, and whereby the *zenith* point is directly ascertained, is described in the Phil. Trans. 1828, p, 257, by the same author.

Fig. 23.



(179.) By far the neatest and most delicate application of the *principle of collimation* of Rittenhouse, however, is suggested by Benzenberg, which affords at once, and by a single observation, an exact knowledge of the *nadir* point of an astronomical circle. In this combination, the

Fig. 24.



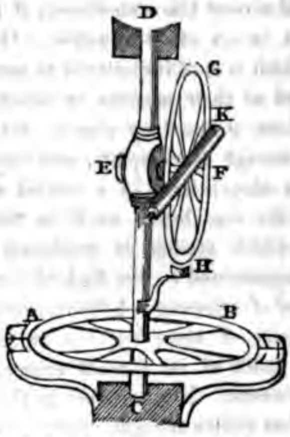
f the circle is its own collimator. The object observed is the rsectional cross of the wires in its own focus reflected in mertelescope, be brought to exact coincidence with it, and time may be allowed to examine and correct this coincidence, if not at first accurately hit, which is the case in no other situation. Generally speaking, all angular magnitudes which it is of importance to ascertain exactly, should, if possible, be observed at their maxima or minima of increase or diminution; because at these points they remain not perceptibly changed during a time long enough to complete, and even, in many cases, to repeat and verify, our observations in a careful and leisurely manner. The angle which, in the case before us, is in this predicament, is the altitude of the star, which attains its maximum or minimum on the meridian, and which is measured on the limb of the mural circle.

(181.) The purposes of astronomy, however, require that an observer should possess the means of observing any object not directly on the meridian, but at any point of its diurnal course, or wherever it may present itself in the heavens. Now, a point in the sphere is determined by reference to two great circles at right angles to each other; or of two circles, one of which passes through the pole of the other. These, in the language of geometry, are *co-ordinates* by which its situation is ascertained: for instance,—on the earth, a place is known if we know its longitude and latitude;—in the starry heavens, if we know its right ascension and declination;—in the visible hemisphere, if we know its azimuth and altitude, &c.

(182.) To observe an object at any point of its diurnal course, we must possess the means of directing a telescope to it; which, therefore, must be capable of motion in two planes at right angles to each other; and the amount of its angular motion in each must be measured on two circles *co-ordinate* to each other, whose planes must be parallel to those in which the telescope moves. The practical accomplishment of this condition is effected by making the axis of one of the circles penetrate that of the other at right angles. The pierced axis turns on fixed supports, while the other has no connection with any external support, but is sustained entirely by that which it penetrates, which is strengthened and enlarged at the point of penetration to receive it. The annexed figure exhibits the simplest form of such a combination, though very far indeed from the best in point of mechanism. The two circles are *read off* by verniers, or microscopes; the one attached to the fixed support which carries the principal axis, the other to an arm projecting from that axis. Both circles also are susceptible of being clamped, the clamps being attached to the same ultimate bearing with which the apparatus for reading off is connected.

(183.) It is manifest that such a combination, however its principal

Fig. 25.



nted (provided that its direction be invariable,) will enable us
n the situation of any object with respect to the observer's
angles reckoned upon two great circles in the visible hemi-
of which has for its poles the prolongations of the principal
vanishing points of a system of lines parallel to it, and the

(185.) In this position the apparatus assumes the name of an *equatorial*, or, as it was formerly called, a *parallactic* instrument. It is a most convenient instrument for all such observations as require an object to be kept long in view, because, being once set upon the object, it can be followed as long as we please by a *single motion*, i. e. by merely turning the whole apparatus round on its polar axis. For since, when the telescope is set on a star, the angle between its direction and that of the polar axis is equal to the polar distance of the star, it follows, that when turned about its axis, without altering the position of the telescope on the circle G H, the point to which it is directed will always lie in the small circle of the heavens coincident with the star's diurnal path. In many observations this is an inestimable advantage, and one which belongs to no other instrument. The equatorial is also used for determining the place of an unknown by comparison with that of a known object, in a manner to be described in the fifth chapter. The adjustments of the equatorial are somewhat complicated and difficult. They are best performed in this manner: — 1st, Follow the pole star round its whole diurnal course, by which it will become evident whether the polar axis is directed above or below, to the right or to the left, of the true pole,—and correct it accordingly (without any attempt, during this process, to correct the errors, if any, in the position of the declination axis). 2dly, after the polar axis is thus brought into adjustment, place the plane of the declination circle in or near the meridian; and, having there secured it, observe the transits of several known stars of widely different declinations. If the intervals between these transits correspond to the known differences of right ascensions of the stars, we may be sure that the telescope describes a true meridian, and that, therefore, the declination axis is truly perpendicular to the polar one; — if not, the deviation of the intervals from this law will indicate the direction and amount of the deviation of the axis in question, and enable us to correct it.[1]

(186.) A very great improvement has, within a few years of the present time, been introduced into the construction of the equatorial instrument. It consists in applying a clockwork movement to turn the whole instrument round upon its polar axis, and so to follow the diurnal motion of any celestial object, without the necessity of the observer's manual intervention. The driving power is the descent of a weight which communi-

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[1] See Littrow on the Adjustment of the Equatorial (Mem. Ast. Soc. vol. ii. p. 45), where formulæ are given for ascertaining the amount and direction of all the misadjustments simultaneously. But the practical observer, who wishes to avoid bewildering himself by doing two things at once, had better proceed as recommended in the text.

n to a train of wheelwork, and thus, ultimately, to the polar
at the same time, its *too swift* descent is controlled and regu-
e exact and uniform rate required to give that axis one turn in
by connecting it with a regulating clock, or (which is found
in practice) by exhausting all the superfluous energy of the
wer, by causing it to overcome a regulated friction. Artists
ucceeded in obtaining a perfectly smooth, uniform, and *regula-*
which, when so applied, serves to retain any object on which the
ay be set, commodiously, in the centre of the field of view for
rs in succession, leaving the attention of the observer undis-
having a mechanical movement to direct, and with both his
berty.

The other position in which such a compound apparatus as we
ibed in art. 182 may be advantageously mounted, is that in
principal axis occupies a vertical position, and the one circle, A
ently corresponds to the celestial horizon, and the other, G H,
l circle of the heavens. The angles measured on the former
re *azimuths*, or differences of azimuth, and those of the latter
nces, or altitudes, according as the graduation commences from
point of its limb, or from one 90° distant from it. It is there-
n by the name of an *azimuth and altitude instrument*. The

at the moment of the first observation above the *eastern*. At this point let the motion be arrested, and the horizontal circle be again read off. The difference of the readings will be the azimuthal arc described in the interval. Now, it is evident that when the altitudes of any star are equal on either side of the meridian, its *azimuths*, whether reckoned both from the north or both from the south point of the horizon, must also be equal, —consequently the north or south point of the horizon must bisect the azimuthal arc thus determined, and will therefore become known.

(189.) This method of determining the north and south points of a horizontal circle is called the "method of equal altitudes," and is of great and constant use in practical astronomy. If we note, at the moments of the two observations, the time, by a clock or chronometer, the instant half-way between them will be the moment of the star's meridian passage, which may thus be determined without a transit; and, *vice versâ*, the error of a clock or chronometer may by this process be discovered. For this last purpose, it is not necessary that our instrument should be provided with a horizontal circle at all. Any means by which altitudes can be measured will enable us to determine the moments when the same star arrives at *equal* altitudes in the eastern and western halves of its diurnal course; and, these once known, the instant of meridian passage and the error of the clock become also known.

(190.) Thus also a meridian line may be drawn and a *meridian mark* erected. For the readings on the north and south points on the limb of the horizontal circle being known, the vertical circle may be brought exactly into the plane of the meridian, by setting it to that precise reading. This done, let the telescope be depressed to the north horizon, and let the point intersected there by its cross-wires be noted, and a mark erected there, and let the same be done for the south horizon. The line joining these points is a meridian line, passing through the centre of the horizontal circle. The marks may be made secure and permanent if required.

(191.) One of the chief purposes to which the altitude and azimuth circle is applicable is the investigation of the amount and laws of refraction. For, by following with it a circumpolar star which passes the zenith, and another which grazes the horizon, through their whole diurnal course, the exact *apparent* form of their diurnal orbits, or the ovals into which their circles are distorted by refraction, can be traced; and their deviation from circles, being at every moment given by the nature of the observation *in the direction in which the refraction itself takes place* (i. e. in altitude), is made a matter of direct observation.

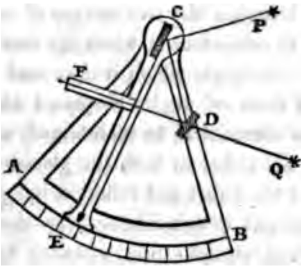
(192.) The *zenith sector* and the *theodolite* are peculiar modifications of the altitude and azimuth instrument. The former is adapted for the

observation of stars in or near the zenith, by giving a great
he vertical axis, and suppressing all the circumference of the
cle, except a few degrees of its lower part, by which a great
adius, and a consequent proportional enlargement of the divi-
arc, is obtained. The latter is especially devoted to the mea-
rizontal angles between terrestrial objects, in which the telescope
ires to be elevated more than a few degrees, and in which,
the vertical circle is either dispensed with, or executed on a
le, and with less delicacy; while, on the other hand, great care
d on securing the exact perpendicularity of the plane of the
motion, by resting its horizontal axis on two supports like the
transit-instrument, which themselves are firmly bedded on the
he horizontal circle, and turn with it.

The next instrument we shall describe is one by whose aid the
tance of any two objects may be measured, or the altitude of
e determined, either by measuring its distance from the visible
ch as the sea-offing, allowing for its dip), or from its own reflec-
surface of mercury. It is the sextant, or quadrant, commonly
lley's, from its reputed inventor, though the priority of inven-
s undoubtedly to Newton, whose claims to the gratitude of the
re thus doubled, by his having furnished at once the only

at the extremity E of which a vernier is placed to read off the divisions of the limb. On the radius A C is set a telescope F, through which any object, Q, may be seen by *direct* rays which pass through the unsilvered portion of the glass D, while another object, P, is seen through the same telescope, by rays, which, after reflection at C, have been thrown upon the silvered part of D, and are thence directed by a second reflection into the telescope. The two images so formed will both be seen in the field of view at once, and by moving the radius C E will (if the reflectors be truly perpendicular to the plane of the circle) meet and pass over, without obliterating each other. The motion, however, is arrested when they meet, and at this point the angle included between the direction C P of one object, and F Q of the other, is twice the angle E C A included between the fixed and moveable radii C A, C E. Now, the graduations of the limb being purposely made only half as distant as would correspond to degrees, the arc, A E, when read off, as if the graduations were *whole* degrees, will, in fact, read double its real amount, and therefore the numbers so read off will express, not the angle E C A, but its double, the angle subtended by the objects.

Fig. 26.



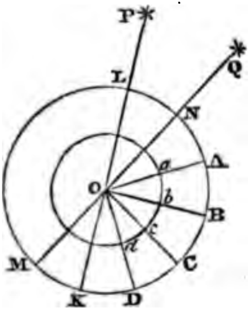
(195.) To determine the exact distances between the stars by direct observation is comparatively of little service; but in nautical astronomy the measurement of their distances from the moon, and of their altitudes, is of essential importance; and as the sextant requires no fixed support, but can be held in the hand, and used on ship-board, the utility of the instrument becomes at once obvious. For altitudes at sea, as no level, plumb-line, or artificial horizon can be used, the sea-offing affords the only resource; and the image of the star observed, seen by reflection, is brought to coincide with the boundary of the sea seen by direct rays. Thus the altitude above the sea-line is found; and this corrected for the *dip of the horizon* (art. 23) gives the true altitude of the star. On land, an artifi-

n may be used (art. 173), and the consideration of dip is ren-
ecessary.

The adjustments of the sextant are simple. They consist in
two reflectors, the one on the revolving radius C E, the other
ed one C B, so as to have their planes perpendicular to the
he circle, and parallel to each other, when the reading of the
is zero. This adjustment in the latter respect is of little
s its effect is to produce a *constant* error, whose amount is
ertained by bringing the two images of one and the same star
stant object to coincidence; when the instrument *ought* to read
f it does not, the angle which it *does* read is the zero correction
be subtracted from *all* angles measured with the sextant. The
ustments are essential to be maintained, and are performed by
vs, by whose aid either or both the glasses may be *tilted* a little
another until the direct and reflected images of a vertical *line* (a
) can be brought to coincidence over *their whole extent*, so as
single unbroken straight line, whatever be the position of the
rm, in the middle of the field of view of the telescope, whose
efully adjusted by the optician to parallelism with the plane of

In practice it is usual to leave only the reflector D on the fixed
ustable, that on the moveable being set to great nicety by the

Fig. 27.



able on O, the common axis of two circles, A M L and *a b c*, of which the former, A M L, is absolutely fixed in the plane of the objects, and carries the graduations, and the latter is freely moveable on the axis. The telescope is attached permanently to the latter circle, and moves with it. An arm O *a* A carries the index, or vernier, which reads off the graduated limb of the fixed circle. This arm is provided with two clamps, by which it can be temporarily connected with either circle, and detached at pleasure. Suppose, now, the telescope directed to P. Clamp the index arm O A to the *inner* circle, and unclamp it from the outer, and read off. Then carry the telescope round to the other object Q. In so doing, the inner circle, and the index-arm which is clamped to it, will also be carried round, over an arc A B, on the graduated limb of the outer, equal to the angle P O Q. Now clamp the index to the outer circle, and unclamp the inner, and read off: the difference of readings will of course measure the angle P O Q; but the result will be liable to two sources of error — that of *graduation* and that of observation, both which it is our object to get rid of. To this end transfer the telescope back to P, *without* unclamping the arm from the outer circle; *then*, having made the bisection of P, clamp the arm to *b*, and unclamp it from B, and again transfer the telescope to Q, by which the arm will now be carried with it to C, over a second arc, B C, equal to the angle P O Q. Now again read off; then will the difference between this reading and the *original* one measure *twice* the angle P O Q, affected with *both* errors of observation, but only with *the same error of graduation as before*. Let this process be repeated as often as we please (suppose ten times); then will the final arc A B C D read off on the circle be ten times the required angle, affected by the joint errors of all the ten observations, but only by the same constant error of graduation, which depends on the initial and final readings off alone. Now the errors of observation, when numerous, tend to

destroy one another; so that, if sufficiently multiplied, their
ill disappear from the result. There remains, then, only the
ror of graduation, which comes to be divided in the final result
ber of observations, and is therefore diminished in its influence
of its possible amount, or to less if need be. The abstract
advantage of this principle seem to be counterbalanced in
some unknown cause, which, probably, must be sought for in
lamping.

icrometers are instruments (as the name imports[1]) for measur-
reat precision, small angles, not exceeding a few minutes, or at
le degree. They are very various in construction and principle,
however, depending on the exceeding delicacy with which space
divided by the turns and parts of a turn of fine screws. Thus
rallel wire micrometer, two parallel threads (spider's lines are
used) stretched on sliding frames, one or both moveable by

Fig. 28.

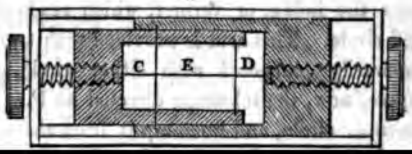
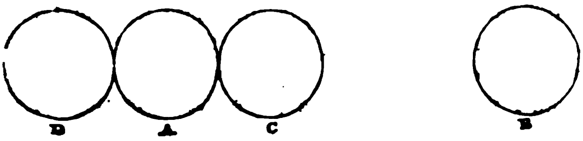


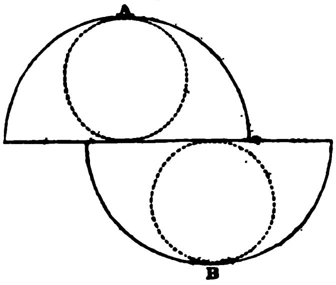
Fig. 29.



dependent (by some mechanical movement) on the will of the observer, and in any required direction from one another. As these can, therefore, be made to approach to or recede from each other at pleasure, they may be brought in the first place to approach till they touch one another on one side, as at A C, and then being made by continuing the motion to cross and touch on the opposite side, as A D, it is evident that the quantity of movement required to produce the change from one contact to the other, *if uniform*, will *measure* the double diameter of the object A.

(201.) Innumerable optical combinations may be devised to operate such duplication. The chief and most important (from its recent applications,) is the *heliometer*, in which the image is divided by bisecting *the object-glass of the telescope*, and making its two halves, set in separate brass frames, slide laterally on each other, as A B, the motion being produced and measured by a screw. Each half, by the laws of optics, forms its own image (somewhat blurred, it is true, by diffraction,[1]) in its own axis; and thus two equal and similar images are formed side by side in the focus of the eye-piece, which may be made to approach and recede by the motion of the screw, and thus afford the means of measurement as above described.

Fig. 30.



(202.) Double refraction through crystallized media affords another ccomplishing the same end. Without going into the intrica- difficult branch of optics, it will suffice to state that objects ough certain crystals (as Iceland spar, or quartz) appear images equally distinct being formed, whose angular distance other varies from nothing (or perfect coincidence,) up to a it, *according to the direction with respect to a certain fixed crystal,* called its optical axis. Suppose, then, to take the se, that the eye-lens of a telescope, instead of glass, were such a crystal (say of quartz, which may be worked as well or glass,) and *of a spherical form,* so as to offer no difference d about on its centre, other than the inclination of its optical visual ray. Then when that axis coincides with the line of of the object-glass, one image only will be seen, but when evolve on an axis perpendicular to that line, two will arise, adually out from each other, and thus originating the desired . In this contrivance, the angular amount of the rotation of affords the necessary datum for determining the separation of

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[1] This might be cured, though at an expense of light, by limiting each half to a circular space by diaphragms, as represented by the dotted lines.

Of all methods which have been proposed, however, the d most unobjectionable would appear to be the following. It on a circular plate of parallel glass, or otherwise sustained, close to and concentric with the other by a framework of metal so light as to intercept but a small portion of the light which passes on the outside (as in the annexed figure,) where the dotted lines represent the radii sustaining one, and the undotted those carrying the other disc. The whole must be so mounted as to allow one disc to revolve in its own plane behind the other, fixed, and to allow the amount of rotation to be read off. It is evident, then, that when the deviations produced by the two discs conspire, a total deviation of 10′ will be effected on all the light which has passed through them; that when they oppose each other, the rays will emerge undeviated, and that in intermediate positions a deviation varying from 0 to 10′, and calculable from the angular rotation of the one disc on the other, will arise. Now, let this combination be applied at such a point of the cone of rays, between the object-glass and its focus, that the discs shall occupy exactly half the area of its section. Then will half the light of the object lens pass undeviated — the other half deviated, as above described; and thus a duplication of image, variable and measureable (as required for micrometric measurement) will occur. If the object-glass be not very large, the most convenient point of its application will be externally before it, in which case the diameter of the discs will be to that of the object-glass as 707 : 1000; or (allowing for the spokes) about as 7 to 10.

(204.) The *Position Micrometer* is simply a straight thread or wire which is carried round by a smooth revolving motion, in the common focus of the object and eye-glasses, in a plane perpendicular to the axis of the telescope. It serves to determine the situation with respect to some fixed line in the field of view, of the line joining any two objects or points of an object seen in that field — as two stars, for instance, near enough to be seen at once. For this purpose the moveable thread is placed so as to cover both of them, or stand, as may best be judged, parallel to their line of junction. And its angle, with the fixed one, is then read off upon a small divided circle exterior to the instrument. When such a micrometer is applied (as it most commonly is) to an equatorially mounted telescope, the *zero* of its position corresponds to a direction of the wire, such as, prolonged, will represent a circle of declination in the heavens — and the "angles of position" so read off are reckoned invariably from one point, and in one direction, viz., *north*, *following*, *south*, *preceding*; so that 0° position corresponds to the situation of an object exactly *north* of that assumed as a centre of reference,— 90° to a situation exactly *eastward* or *following*; 180° exactly *south*; and 270° exactly *west*, or *preceding* in the order of diurnal movement.

CHAPTER IV.

OF GEOGRAPHY.

URE OF THE EARTH.—ITS EXACT DIMENSIONS.—ITS FORM—
EQUILIBRIUM MODIFIED BY CENTRIFUGAL FORCE.—VARIA-
GRAVITY ON ITS SURFACE.—STATICAL AND DYNAMICAL
ES OF GRAVITY.—THE PENDULUM.—GRAVITY TO A SPHE-
OTHER EFFECTS OF THE EARTH'S ROTATION.—TRADE WINDS.
RMINATION OF GEOGRAPHICAL POSITIONS.—OF LATITUDES.
NGITUDES.—CONDUCT OF A TRIGONOMETRICAL SURVEY.—
S.—PROJECTIONS OF THE SPHERE.—MEASUREMENT OF
S BY THE BAROMETER.

GEOGRAPHY is not only the most important of the practical into the various forms of mountain, table land, and valley; neither should the form of the bed of the ocean, regarded as a continuation of the surface of the land beneath the water, be left out of consideration: we know, it is true, very little of it; but this is an ignorance rather to be lamented, and, if possible, remedied, than acquiesced in, inasmuch as there are many very important branches of inquiry which would be greatly advanced by a better acquaintance with it.

(207.) With regard to the figure of the earth *as a whole*, we have already shown that, speaking loosely, it may be regarded as spherical; but the reader who has duly appreciated the remarks in art. 22 will not be at a loss to perceive that this result, concluded from observations not susceptible of much exactness, and embracing very small portions of the surface at once, can only be regarded as a first approximation, and may require to be materially modified by entering into minutiæ before neglected, or by increasing the delicacy of our observations, or by including in their extent larger areas of its surface. For instance, if it should turn out (as it will), on minuter inquiry, that the true figure is somewhat elliptical, or flattened, in the manner of an orange, having the diameter which coincides with the axis about $\frac{1}{300}$th part shorter than the diameter of its equatorial circle; — this is so trifling a deviation from the spherical form that, if a model of such proportions were turned in wood, and laid before us on a table, the nicest eye or hand would not detect the flattening, since the difference of diameters, in a globe of fifteen inches, would amount only to $\frac{1}{20}$th of an inch. In all common parlance, and for all ordinary purposes, then, it would still be called a globe; while, nevertheless, by careful measurement, the difference would not fail to be noticed; and, speaking strictly, it would be termed, not a globe, but an oblate ellipsoid, or spheroid, which is the name appropriated by geometers to the form above described.

(208.) The sections of such a figure by a plane are not circles, but ellipses; so that, on such a shaped earth, the horizon of a spectator would nowhere (except at the poles) be exactly circular, but somewhat elliptical. It is easy to demonstrate, however, that its deviation from the circular form, arising from so slight an "*ellipticity*" as above supposed, would be quite imperceptible, not only to our eye-sight, but to the test of the dip-sector; so that by that mode of observation we should never be led to notice so small a deviation from perfect sphericity. How we are led to this conclusion, as a practical result, will appear, when we have explained the means of determining with accuracy the dimensions of the whole, or any part of the earth.

(209.) As we cannot grasp the earth, nor recede from it far enough to view it at once as a whole, and compare it with a known standard of measure in any degree commensurate to its own size, but can only creep about upon it, and apply our diminutive measures to comparatively small parts of its vast surface in succession, it becomes necessary to supply, by geometrical reasoning, the defect of our physical powers, and from a delicate and careful measurement of such small parts to conclude the form and dimensions of the whole mass. This would present little difficulty, if we were sure the earth was strictly a sphere, for the proportion of the circumference of a circle to its diameter being known (viz. that of 3·1415926 to 1·0000000), we have only to ascertain the length of the entire circumference of any great circle, such as a meridian, in miles, feet, or any other standard units, to know the diameter in units of the same kind. Now, the circumference of the whole circle is known as soon as we know the exact length of any aliquot part of it, such as 1° or $\frac{1}{360}$th part; and this, being not more than about seventy miles in length, is not beyond the limits of very exact measurement, and could, in fact, be measured (if we knew its exact termination at each extremity) within a very few feet, or, indeed, inches, by methods presently to be particularized.

(210.) Supposing, then, we were to begin measuring with all due nicety from any station, in the exact direction of a meridian and go measuring on, till by some indication we were informed that we had accomplished an exact *degree* from the point we set out from, our problem would then be at once resolved. It only remains, therefore, to inquire by what indications we can be sure, 1st, that we *have* advanced *an exact degree;* and, 2dly, that we have been measuring in the *exact direction of a great circle.*

(211.) Now the earth has no landmarks on it to indicate degrees, nor traces inscribed on its surface to guide us in such a course. The compass, though it affords a tolerable guide to the mariner or the traveller, is far too uncertain in its indications, and too little known in its laws, to be of any use in *such* an operation. We must, therefore, look outwards, and refer our situation on the surface of our globe to natural marks, *external* to it, and which are of equal permanence and stability with the earth itself. Such marks are afforded by the stars. By observations of their meridian altitudes, performed at any station, and from their known polar distances, we conclude the height of the pole; and since the altitude of the pole is equal to the latitude of the place (art. 119) the same observations give the latitudes of any stations where we may establish the requisite instruments. When our latitude then, is found to have diminished a degree, we know that, *provided we have kept to the meridian*, we have described one three hundred and sixtieth part of the earth's circumference.

(212.) The direction of the meridian may be secured at every instant by the observations described in art. 162, 188; and although local difficulties may oblige us to deviate in our measurement from this exact direction, yet if we keep a strict account of the amount of this deviation, a very simple calculation will enable us to *reduce* our observed measure to its *meridional* value.

(213.) Such is the principle of that most important geographical operation, the measurement of an arc of the meridian. In its detail, however, a somewhat modified course must be followed. An observatory cannot be mounted and dismounted at every step; so that we cannot identify and measure an exact degree *neither more nor less*. But this is of no consequence, provided we know with equal precision *how much*, more or less, we have measured. In place, then, of measuring this precise aliquot part, we take the more convenient method of measuring from one good observing station to another, *about* a degree, or two or three degrees, as the case may be, or indeed any determinate angular interval apart, and determining by astronomical observation the precise difference of latitudes between the stations.

(214.) Again, it is of great consequence to avoid in this operation every source of uncertainty, because an error committed in the length of a single degree will be multiplied 360 times in the circumference, and nearly 115 times in the diameter of the earth concluded from it. Any error which may affect the astronomical determination of a star's altitude will be especially influential. Now, there is still too much uncertainty and fluctuation in the amount of refraction at moderate altitudes, not to make it especially desirable to avoid this source of error. To effect this, we take care to select for observation, at the extreme stations, some star which passes through or near the zeniths of both. The amount of refraction, within a few degrees of the zenith, is very small, and its fluctuations and uncertainty, in point of quantity, so excessively minute as to be utterly inappreciable. Now, it is the same thing whether we observe the *pole* to be raised or depressed a degree, or the *zenith distance* of a star when on a meridian to have changed by the same quantity (fig. art. 128). If at one station we observe any star to pass through the zenith, and at the other to pass one degree south or north of the zenith, we are sure that the geographical latitudes, or the altitudes of the pole at the two stations, must differ by the same amount.

(215.) Granting that the terminal points of one degree can be ascertained, its *length* may be measured by the methods which will be presently described, as we have before remarked, to within a very few feet. Now, the error which may be committed in fixing each of these terminal points cannot exceed that which may be committed in the observation of

distance of a star properly situated for the purpose in question.
with proper care, can hardly exceed half a second. Supposing
ne possibility of ten feet of error in the length of each degree
ured arc of five degrees, and of half a second in each of the
nces of one star, observed at the northern and southern sta-
lastly, suppose all these errors to conspire, so as to tend all of
ve a result greater, or all less, than the truth, it will appear,
asy proportion, that the whole amount of error which would
tailed on an estimate of the earth's diameter, as concluded
a measure, would not exceed 1147 yards, or about two thirds
nd this is ample allowance.

his, however, supposes that the form of the earth is that of a
ere, and, in consequence, the lengths of its degrees in all parts
qual. But, when we come to compare the measures of meri-
made in various parts of the globe, the results obtained,
ey agree sufficiently to show that the supposition of a spherical
ot *very* remote from the truth, yet exhibit discordances far
n what we have shown to be attributable to error of observation,
render it evident that the hypothesis, in strictness of its word-
nable. The following table exhibits the lengths of arcs of the
astronomically determined as above described), expressed in

Country.	Latitude of Middle of Arc.	Arc measured.	Measured Length in Feet.	Mean Length of the Degree at the Middle Latitude in Feet.
Sweden,[*] A B	+66° 20′ 10″·0	1° 37′ 19″·6	593277	365744
Sweden, A	+66 19 37	0 57 30·4	351832	367086
Russia, A	+58 17 37	3 35 5·2	1309742	365368
Russia, B	+56 3 55·5	8 2 28·9	2937439	365291
Prussia, B	+54 58 26·0	1 30 29·0	551073	365420
Denmark, B	+54 8 13·7	1 31 53·3	559121	365087
Hanover, A B	+52 32 16·6	2 0 57·4	736425	365300
England, A	+52 35 45	3 57 13·1	1442953	364971
England, B	+52 2 19·4	2 50 23·5	1036409	364951
France, A	+46 52 2	8 20 0.3	3040605	364872
France, A B	+44 51 2·5	12 22 12·7	4509832	364572
Rome, A	+42 59 —	2 9 47	787919	364262
America, A	+39 12 —	1 28 45·0	538100	363786
India, A B	+16 8 21·5	15 57 40·7	5794598	363044
India, A B	+12 32 20·8	1 34 56·4	574318	362956
Peru, A B	— 1 31 0·4	3 7 3·5	1131050	363626
Cape of Good Hope, A	—33 18 30	1 13 17·5	445506	364713
Cape of Good Hope, B	—35 43 20·0	3 34 34·7	1301993	364060

It is evident from a mere inspection of the second and fifth columns of this table, that *the measured length of a degree increases with the latitude*, being greatest near the poles, and least near the equator. Let us now consider what interpretation is to be put upon this conclusion, as regards the form of the earth.

(217.) Suppose we held in our hands a model of the earth smoothly turned in wood, it would be, as already observed, so nearly spherical, that neither by the eye nor the touch, unassisted by instruments, could we detect any deviation from that form. Suppose, too, we were debarred from measuring directly across from surface to surface in different directions with any instrument, by which we might at once ascertain whether

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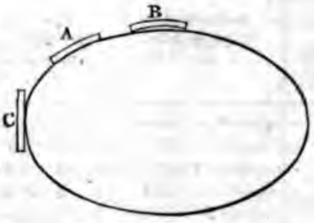
[*] The astronomers by whom these measurements were executed were as follows: —

Sweden, A B — Svanberg.
Sweden, A — Maupertuis.
Russia, A — Struve.
Russia, B — Struve, Tenner.
Prussia — Bessel, Bayer.
Denmark — Schumacher.
Hanover — Gauss.
England — Roy, Kater.
France, A — Lacaille, Cassini.
France, A B — Delambre, Mechain.
Rome — Boscovich.
America — Mason and Dixon.
India, 1st — Lambton.
India, 2d — Lambton, Everest.
Peru — Lacondamine, Bouguer.
Cape of Good Hope, A — Lacaille.
Cape of Good Hope, B — Maclear.

—*Astr. Nachr.* 574

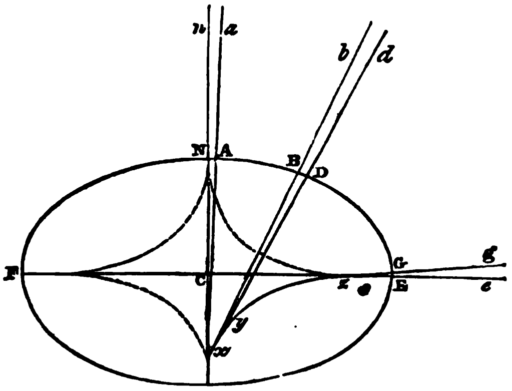
Fig. 32.



er were longer than another; how, then, we may ask, are we
h whether it is a true sphere or not? It is clear that we have
e, but to endeavour to discover, by some nicer means than
pection or feeling, whether the convexity of its surface is the
very part; and if not, where it is greatest, and where least.
hen, a thin plate of metal to be cut into a concavity at its edge,
tly to fit the surface at A: let this now be removed from A,
d successively to several other parts of the surface, taking care
plane always on a great circle of the globe, as here represented.
ve find any position, B, in which the light can enter in the
ween the globe and plate, or any other, C, where the latter tilts

(219.) Let N A B D E F represent a meridional section of the earth, C its centre, and N A, B D, G E, arcs of a meridian, each corresponding to one degree of difference of latitude, or to one degree of variation in the meridian altitude of a star, as referred to the horizon of a spectator travelling along the meridian. Let n N, a A, b B, d D, g G, e E, be the respective directions of the *plumb-line* at the stations N, A, B, D, G, E, of which we will suppose N to be at the pole and E at the equator; then will the tangents to the surface at these points respectively be perpendicular to these directions; and, consequently, if each pair, viz. n N and a A, b B and d D, g G and e E, be prolonged till they intersect each other (at the points x, y, z), the angles N x A, B y D, G z E, will each be one degree, and, therefore, all equal; so that the small curvilinear arcs N A, B D, G E, may be regarded as arcs of circles of one degree each, described about x, y, z, as centres. These are what in geometry are called *centres of curvature*, and the radii x N or x A, y B or y D, z G or z E, represent *radii of curvature*, by which the curvatures at those points are determined and measured. Now, as the arcs of different circles, which subtend equal angles at their respective centres, are in the direct proportion of their radii, and as the arc N A is greater than B D, and that again than G E, it follows that the radius N x must be greater than B y, and B y than E z. Thus it appears that the mutual intersections of the plumb-lines will not, as in the sphere, all coincide in one point C, the centre, but will be arranged along a certain curve, x y z (which will be rendered more evident by considering a number of intermediate stations).

Fig. 33.



rve geometers have given the name of the *evolute* of the curve
E, from whose centres of curvature it is constructed.
n the flattening of a round figure at two opposite points, and
rance at points rectangularly situated to the former, we recog-
istinguishing feature of the elliptic form. Accordingly, the
implest supposition that we can make respecting the nature of
n, since it is proved not to be a circle, is, that it is an ellipse,
o, having N S, the axis of the earth, for its shorter, and E F,
rial diameter, for its longer axis; and that the form of the
face is that which would arise from making such a curve revolve
horter axis N S. This agrees well with the general course of
e of the degree in going from the equator to the pole. In the
radius of curvature at E, the extremity of the longer axis is
nd at that of the shorter axis, the greatest it admits, and the
evolute agrees with that here represented.[1] Assuming, then,
n ellipse, the geometrical properties of that curve enable us to
proportion between the lengths of its axes which shall corre-
ny proposed rate of variation in its curvature, as well as to fix
absolute lengths, corresponding to any assigned length of the
given latitude. Without troubling the reader with the inves-
which may be found in any work on the conic sections,) it will
t to state results which have been arrived at by the most sys-

These conclusions are based on the consideration of those 13 arcs, to which the letter A is annexed[1], and of one other arc of 1° 7′ 31″·1, measured in Piedmont by Plana and Carlini, whose discordance with the rest, owing to local causes hereafter to be explained, arising from the exceedingly mountainous nature of the country, render the propriety of so employing it very doubtful. Be that as it may, the strikingly near accordance of the two sets of dimensions is such as to inspire the greatest confidence in both. The measurement at the Cape of Good Hope by Lacaille, also used in this determination, has always been regarded as unsatisfactory, and has recently been demonstrated by Mr. Maclear to be erroneous to a considerable extent. The omission of the former, and the substitution for the latter, of the far preferable result of Mr. Maclear's second measurement would induce, however, but a trifling change in the final result.

(221.) Thus we see that the rough diameter of 8000 miles we have hitherto used, is rather too great, the excess being about 100 miles, or $\frac{1}{80}$th part. As convenient numbers to remember, the reader may bear in mind, that in our latitude there are just as many thousands of feet in a degree of the meridian as there are days in the year (365): that, speaking loosely, a degree is about 70 British statute miles, and a second about 100 feet; that the equatorial circumference of the earth is a little less than 25,000 miles (24,899), and the ellipticity or polar flattening amounts to one 300th part of the diameter.

(222.) The two sets of results above stated are placed in juxtaposition, and the particulars given more in detail than may at first sight appear consonant, either with the general plan of this work, or the state of the reader's presumed acquaintance with the subject. But it is of importance that he should early be made to see how, in astronomy, results in admirable concordance emerge from data accumulated from totally different quarters, and how local and accidental irregularities in the data themselves become neutralized and obliterated by their impartial geometrical treatment. In the cases before us, the modes of calculation followed are widely different, and in each the mass of figures to be gone through to arrive at the result, enormous.

(223.) The supposition of an elliptic form of the earth's section through the axis is recommended by its simplicity, and confirmed by comparing the numerical results we have just set down with those of actual measurement. When this comparison is executed, discordances, it is true, are observed, which, although still too great to be referred to error of

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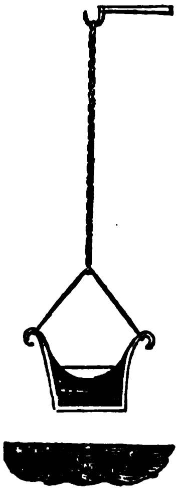
[1] In those which have both A and B, the numbers used by Mr. Airy differ slightly from Bessel's, which are those we have preferred.

nt, are yet so small, compared to the errors which would result
herical hypothesis, as completely to justify our regarding the
ellipsoid, and referring the observed deviations to either local
ral, to comparatively small causes.

ow, it is highly satisfactory to find that the general elliptical
practically proved to exist, is precisely what *ought theoretically*
om the rotation of the earth on its axis. For, let us suppose
sphere, at rest, of uniform materials throughout, and exter-
ed with an ocean of equal depth in every part. Under such
es it would obviously be in a state of *equilibrium;* and the
s surface would have no tendency to run one way or the other.
ow, a quantity of its materials were taken from the polar
d piled up all around the equator, so as to produce that dif-
the polar and equatorial diameters of 26 miles which we know
It is not less evident that a mountain ridge or equatorial *conti-*
would be thus formed, from which the water would run down
ted part at the poles. However solid matter might rest where
ed, the liquid part, at least, would not remain there, any more
vere thrown on the side of a hill. The consequence therefore,
he formation of two great polar seas, hemmed in all round by
land. Now, this is by no means the case in nature. The
pies indifferently all latitudes, with no more partiality to the

Fig. 88.



cord hangs perpendicularly. The surface of the water, instead of remaining horizontal, will become concave, as in the figure. The centrifugal force generates a tendency in *all* the water to leave the axis, and press towards the circumference; it is, therefore, urged against the pail, and forced up its sides, till the excess of height, and consequent increase of pressure downwards, just counterbalances its centrifugal force, and a *state of equilibrium* is attained. The experiment is a very easy and instructive one, and is admirably calculated to show how the *form of equilibrium* accommodates itself to varying circumstances. If, for example, we allow the rotation to cease by degrees, as it becomes slower we shall see the concavity of the water regularly diminish; the elevated outward portion will descend, and the depressed central rise, while all the time a perfectly *smooth* surface is maintained, till the rotation is exhausted, when the water resumes its horizontal state.

(226.) Suppose, then, a globe, of the size of the earth, at rest, and covered with a uniform ocean, were to be set in rotation about a certain axis, at first very slowly, but by degrees more rapidly, till it turned round once in twenty-four hours; a centrifugal force would be thus generated, whose general tendency would be to urge the water at every point of the surface to *recede* from the *axis*. A rotation might, indeed, be conceived so swift as to *flirt* the whole ocean from the surface, like water from a mop. But this would require a far greater velocity than what we now speak of. In the case supposed, the *weight* of the water would still keep it *on* the earth; and the tendency to recede from the axis *could* only be satisfied, therefore, by the water leaving the poles, and flowing towards the equator; there heaping itself up in a ridge, just as the water in our pail accumulates against the side; and being retained in opposition to its weight, or natural tendency towards the centre, by the pressure thus caused. This, however, could not take place without laying dry the polar portions of the land in the form of immensely protuberant continents; and the difference of our supposed cases, therefore, is this: — in the former, a great equatorial continent and polar seas would be formed; in the latter, protuberant land would appear at the poles, and a zone of ocean be disposed around the equator. This would be the first or immediate effect. Let us now see what would afterwards happen, in the two cases, if things were allowed to take their natural course.

(227.) The sea is constantly beating on the land, grinding it down, and scattering its worn-off particles and fragments, in the state of mud and pebbles, over its bed. Geological facts afford abundant proof that the existing continents have all of them undergone this process, even more than once, and been entirely torn in fragments, or reduced to powder, and submerged and reconstructed. Land, in this view of the subject, loses its attribute of fixity. As a mass it might hold together in opposition to forces which the water freely obeys; but in its state of successive or simultaneous degradation, when disseminated through the water, in the state of sand or mud, it is subject to all the impulses of that fluid. In the lapse of time, then, the protuberant land in both cases would be destroyed, and spread over the bottom of the ocean, filling up the lower parts, and tending continually to remodel the surface of the solid nucleus, in correspondence with the *form of equilibrium* in both cases. Thus, after a sufficient lapse of time, in the case of an earth at rest, the equatorial continent, thus forcibly constructed, would again be levelled and transferred to the polar excavations, and the spherical figure be so at length restored. In that of an earth in rotation, the polar protuberances would gradually be cut down and disappear, being transferred to the equator (as being *then* the *deepest sea*), till the earth would assume by degrees the form we observe it to have — that of a flattened or *oblate* ellipsoid.

(228.) We are far from meaning here to trace the process *by which* the earth really assumed its actual form; all we intend is, to show that this is the form to which, under the conditions of a rotation on its axis, it must *tend;* and which it would attain, even if originally and (so to speak) perversely constituted otherwise.

(229.) But, further, the dimensions of the earth and the time of its rotation being known, it is easy thence to calculate the exact amount of the centrifugal force,[1] which, at the equator, appears to be $\frac{1}{289}$th part of the force or weight by which all bodies, whether solid or liquid, tend to fall towards the earth. By this fraction of its weight, then, the sea at the equator is *lightened*, and thereby rendered susceptible of being supported on a higher level, or more remote from the the centre than at the poles, where no such counteracting force exists; and where, in consequence, the water may be considered as *specifically heavier*. Taking this principle as a guide, and combining it with the laws of gravity (as developed by Newton, and as hereafter to be more fully explained), mathematicians have been enabled to investigate, *à priori*, what would be the figure of equilibrium of such a body, constituted internally as we have reason to believe the earth to be; covered wholly or partially with a fluid; and revolving uniformly in twenty-four hours; and the result of this inquiry is found to agree very satisfactorily with what experience shows to be the case. From their investigations it appears that the form of equilibrium is, in fact, no other than an oblate ellipsoid, of a degree of ellipticity very nearly identical with what is observed, and which would be no doubt accurately so, did we know, with precision, the internal constitution and materials of the earth.

(230.) The confirmation thus incidentally furnished, of the hypothesis of the earth's rotation on its axis, cannot fail to strike the reader. A deviation of its figure from that of a sphere was not contemplated among the original reasons for adopting that hypothesis, which was assumed solely on account of the easy explanation it offers of the apparent diurnal motion of the heavens. Yet we see that, once admitted, it draws with it, as a necessary consequence, this other remarkable phenomenon, of which no other satisfactory account could be rendered. Indeed, so direct is their connection, that the ellipticity of the earth's figure was discovered and demonstrated by Newton to be a consequence of its rotation, and its amount actually calculated by him, long before any measurement had suggested such a conclusion. As we advance with our subject, we shall find the same simple principle branching out into a whole train of singular and important consequences, some obvious enough, others which at first seem entirely unconnected with it, and which, until traced by Newton up to this their origin, had ranked among the most inscrutable arcana of astronomy, as well as among its grandest phenomena.

(231.) Of its more obvious consequences, we may here mention one which falls naturally within our present subject. If the earth really revolve on its axis, this rotation must generate a centrifugal force (see art. 225,) the effect of which must of course be to counteract a certain portion of the *weight* of every body situated at the equator, as compared with its weight at the poles, or in any intermediate latitudes. Now, this is fully confirmed by experience. There is actually observed to exist a difference in the *gravity*, or downward tendency, of one and the same body, when conveyed successively to stations in different latitudes. Experiments made with the greatest care, and in every accessible part of the globe, have fully demonstrated the fact of a regular and progressive increase in the weights of bodies corresponding to the increase of latitude, and fixed its amount and the law of its progression. From these it appears, that the extreme amount of this variation of gravity, or the difference between the equatorial and polar weights of one and the same mass of matter, is 1 part in 194 of its whole weight, the rate of increase in travelling from the equator to the pole being *as the square of the sine of the latitude.*

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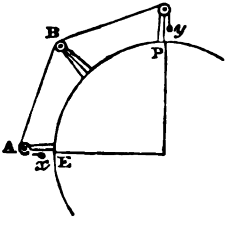
[1] Newton's Principia, iii. Prop. 19.

(232.) The reader will here naturally inquire, what is *meant* by speaking of the same body as having different weights at different stations; and, how such a fact, if true, can be ascertained. When we weigh a body by a balance or a steelyard, we do but counteract its weight by the equal weight of another body under the very same circumstances; and if both the body weighed and its counterpoise be removed to another station, their gravity, if changed at all, will be changed equally, so that they will still continue to counterbalance each other. A difference in the intensity of gravity could, therefore, never be detected by these means; nor is it in *this* sense that we assert that a body weighing 194 pounds at the equator will weigh 195 at the pole. If counterbalanced in a scale or steelyard at the former station, an additional pound placed in one or other scale at the latter would inevitably sink the beam.

(233.) The meaning of the proposition may be thus explained: — Conceive a weight x suspended at the equator by a string without weight

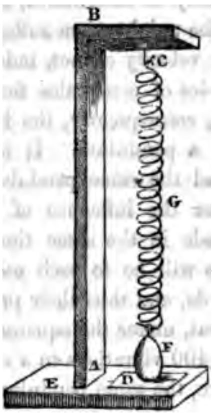
Fig. 35.



passing over a pulley, A, and conducted (supposing such a thing possible) over other pulleys, such as B, round the earth's convexity, till the other end hung down at the pole, and there sustained the weight *y*. If, then, the weights *x* and *y* were such as, at any one station, equatorial or polar, would exactly counterpoise each other on a balance, or when suspended side by side over a single pulley, they would not counterbalance each other in this supposed situation, but the polar weight *y* would preponderate; and to restore the equipoise the weight *x* must be increased by $\frac{1}{194}$th part of its quantity.

(234.) The means by which this variation of gravity may be shown to exist, and its amount measured, are twofold (like all estimations of mechanical power,) statical and dynamical. The former consists in putting the gravity of a weight in equilibrium, not with that of another weight, but with a natural power of a different kind not liable to be affected by local situation. Such a power is the elastic force of a spring. Let A B C be a strong support of brass standing on the foot A E D cast in one piece with it, into which is let a smooth plate of agate, D, which can be adjusted to perfect horizontality by a level. At C let a spiral

Fig. 36.



spring G be attached, which carries at its lower end a weight F, polished and convex below. The length and strength of the spring must be so adjusted that the weight F shall be sustained by it just to swing clear of contact with the agate plate in the highest latitude at which it is intended to use the instrument. Then, if small weights be added cautiously, it may be made to descend till it *just grazes* the agate, a contact which can be made with the utmost imaginable delicacy. Let these weights be noted; the weight F detached; the spring G carefully lifted off its hook, and secured, for travelling, from rust, strain, or disturbance, and the whole apparatus conveyed to a station in a lower latitude. It will then be found, on remounting it, that, although loaded with the same additional weights as before, the weight F will no longer have power enough to stretch the spring to the extent required for producing a similar contact. More weights will require to be added; and the additional quantity necessary will, it is evident, measure the difference of gravity between the two stations, as exerted on the whole quantity of pendent matter, *i. e.* the sum of the weight F and *half* that of the spiral spring itself. Granting that a spiral spring can be constructed of such strength and dimensions that a weight of 10,000 grains, including its own, shall produce an elongation of 10 inches without permanently straining it, one additional grain will produce a further extension of $\frac{1}{1000}$th of an inch, a quantity which cannot possibly be mistaken in such a contact as that in question. Thus we should be provided with the means of measuring the power of gravity at any station to within $\frac{1}{10000}$th of its whole quantity.

(235.) The other, or dynamical process, by which the force urging any given weight to the earth may be determined, consists in ascertaining the velocity imparted by it to the weight when suffered to fall freely in a given time, as one second. This velocity cannot, indeed, be directly measured; but indirectly, the principles of mechanics furnish an easy and certain means of deducing it, and, consequently, the intensity of gravity, by observing the oscillations of a pendulum. It is proved from mechanical principles[2], that, if one and the same pendulum be made to oscillate at different stations, or under the influence of different forces, and the numbers of oscillations made in the same time in each case be counted, the intensities of the forces will be to each other as the squares of the numbers of oscillations made, and thus their proportion becomes known. For instance, it is found that, under the equator, a pendulum of a certain form and length makes 86,400 vibrations in a mean solar day; and that, when transported to London, the same pendulum makes 86,535 vibrations

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[1] Whether the process above described could ever be so far perfected and refined as to become a substitute for the use of the pendulum must depend on the degree of permanence and uniformity of action of springs, on the constancy or variability of the effect of temperature on their elastic force, on the possibility of transporting them, absolutely unaltered, from place to place, &c. The great advantages, however, which such an apparatus and mode of observation would possess, in point of convenience, cheapness, portability, and expedition, over the present laborious, tedious, and expensive process, render the attempt well worth making.

[2] Newton's Principia, ii. Prop. 24, Cor. 3.

in the same time. Hence we conclude, that the intensity of the force urging the pendulum downwards at the equator is to that at London as $(86{,}400)^2$ to $(86{,}535)^2$, or as 1 to 1·00315; or, in other words, that a mass of matter weighing in London 100,000 pounds, exerts the same pressure on the ground, or the same effort to crush a body placed below it, that 100,315 *of the same pounds* transported to the equator would exert there.

(236.) Experiments of this kind have been made, as above stated, with the utmost care and minutest precaution to ensure exactness in all accessible latitudes; and their general and final result has been, to give $\frac{1}{194}$ for the fraction expressing the difference of gravity at the equator and poles. Now, it will not fail to be noticed by the reader, and will, probably, occur to him as an objection against the explanation here given of the fact by the earth's rotation, that this differs materially from the fraction $\frac{1}{289}$ expressing the centrifugal force at the equator. The difference by which the former fraction exceeds the latter is $\frac{1}{590}$, a small quantity in itself, but still far too large, compared with the others in question, not to be distinctly accounted for, and not to prove fatal to this explanation if it will not render a strict account of it.

(237.) The mode in which this difference arises affords a curious and instructive example of the indirect influence which mechanical causes often exercise, and of which astronomy furnishes innumerable instances. The rotation of the earth gives rise to the centrifugal force; the centrifugal force produces an ellipticity in the form of the earth itself; and this very ellipticity of form modifies its power of attraction on bodies placed at its surface, and thus gives rise to the difference in question. Here, then, we have the same cause exercising at once a direct and an indirect influence. The amount of the former is easily calculated, that of the latter with far more difficulty, by an intricate and profound application of geometry, whose steps we cannot pretend to trace in a work like the present, and can only state its nature and result.

(238.) The weight of a body (considered as undiminished by a centrifugal force) is the effect of the earth's attraction on it. This attraction, as Newton has demonstrated, consists, not in a tendency of all matter to any one particular centre, but in a disposition of every particle of matter in the universe to press towards, and if not opposed to approach to, every other. The attraction of the earth, then, on a body placed on its surface, is not a simple but a complex force, resulting from the separate attractions of all its parts. Now, it is evident, that if the earth were a perfect sphere, the attraction exerted by it on a body any where placed on its surface, whether at its equator or pole, must be exactly alike, — for the simple reason of the exact symmetry of the sphere in every direction. It is not less evident that, the earth being elliptical, and this symmetry or similitude of all its parts not existing, the same result cannot be expected. A body placed at the equator, and a similar one at the pole of a flattened ellipsoid, stand in a different geometrical relation to the mass as a whole. This difference, without entering further into particulars, may be expected to draw with it a difference in its forces of attraction on the two bodies. Calculation confirms this idea. It is a question of purely mathematical investigation, and has been treated with perfect clearness and precision by Newton, Maclaurin, Clairaut, and many other eminent geometers; and the result of their investigations is to show that, owing to the elliptic form of the earth alone, and independent of the centrifugal force, its attraction ought to increase the weight of a body in going from the equator to the pole by almost exactly $\frac{1}{590}$th part; which, together with $\frac{1}{289}$th due to the centrifugal force, make up the whole quantity, $\frac{1}{194}$th, observed.

(239.) Another great geographical phenomenon, which owes its existence to the earth's rotation, is that of the trade-winds. These mighty currents in our atmosphere, on which so important a part of navigation depends, arise from, 1st, the unequal exposure of the earth's surface to the sun's rays, by which it is unequally heated in different latitudes; and, 2dly, from that general law in the constitution of all fluids, in virtue of which they occupy a larger bulk, and become specifically lighter when hot than when cold. These causes, combined with the earth's rotation from west to east, afford an easy and satisfactory explanation of the magnificent phenomena in question.

(240.) It is a matter of observed fact, of which we shall give the explanation farther on, that the sun is constantly vertical over some one or other part of the earth between two parallels of latitude, called the tropics, respectively 23½° north, and as much south of the equator; and that the whole of that zone or belt of the earth's surface included between the tropics, and equally divided by the equator, is, in consequence of the great altitude attained by the sun in its diurnal course, maintained at a much higher temperature than those regions to the north and south which lie nearer the poles. Now, the heat thus acquired by the earth's surface is communicated to the incumbent air, which is thereby expanded, and rendered specifically lighter than the air incumbent on the rest of the globe. It is therefore, in obedience to the general laws of hydrostatics, displaced and buoyed up from the surface, and its place occupied by colder, and therefore heavier air, which glides in, on both sides, along the surface, from the regions beyond the tropics; while the displaced air, thus raised above its due level, and unsustained by any lateral pressure, flows over, as it were, and forms an upper current in the contrary direction, or towards the poles; which, being cooled in its course, and also sucked down to supply the deficiency in the extra-tropical regions, keeps up thus a continual circulation.

(241.) Since the earth revolves about an axis passing through the poles, the equatorial portion of its surface has the greatest velocity of rotation, and all other parts less in the proportion of the radii of the circles of latitude to which they correspond. But as the air, when relatively and apparently at rest on any part of the earth's surface, is only so because in reality it participates in the motion of rotation proper to that part, it follows that when a mass of air near the poles is transferred to the region near the equator by any impulse urging it directly towards that circle, in every point of its progress towards its new situation it must be found deficient in rotatory velocity, and therefore unable to keep up with the speed of the new surface over which it is brought. Hence, the currents of air which set in towards the equator from the north and south must, as they glide along the surface, at the same time lag, or hang back, and *drag upon* it in the direction *opposite* to the earth's rotation, *i. e.* from east to west. Thus these currents, which but for the rotation would be simply northerly and southerly winds, acquire, from this cause, a *relative* direction towards the west, and assume the character of permanent north-easterly and south-easterly winds.

(242.) Were any considerable mass of air to be *suddenly* transferred from beyond the topics to the equator, the difference of the rotatory velocities proper to the two situations would be so great as to produce not merely a wind, but a tempest of the most destructive violence. But this is not the case: the advance of the air from the north and south is gradual, and all the while the earth is continually acting on, and by the friction of its surface accelerating its rotatory velocity. Supposing its progress towards the equator to cease at any point, this cause would almost immediately communicate to it the deficient motion of rotation, after which it would revolve quietly with the earth, and be at relative rest. We have only to call to mind the comparative *thinness* of the coating which the atmosphere forms around the globe (art. 35), and the immense *mass* of the latter, compared with the former (which it exceeds at least 100,000,000 times), to appreciate fully the absolute *command* of any extensive territory of the earth over the atmosphere immediately incumbent on it, in point of motion.

(243.) It follows from this, then, that as the winds on both sides approach the equator, their easterly tendency must diminish.[1] The lengths of the diurnal circles increase very slowly in the immediate vicinity of the equator, and for several degrees on either side of it hardly change at all. Thus the friction of the surface has more time to act in accelerating the velocity of the air, bringing it towards a state of *relative* rest, and diminishing thereby the relative set of the currents from east to west, which, on the other hand, is feebly, and, at length, not at all reinforced by the cause which originally produced it. Arrived, then, at the equator, the trades must be expected to lose their easterly character altogether. But not only this but the northern and southern currents here meeting and opposing, will mutually destroy each other, leaving only such preponderancy as may be due to a difference of local causes acting in the two hemispheres,—which in some regions around the equator may lie one way, in some another.

(244.) The result, then, must be the production of two great tropical belts, in the northern of which a constant north-easterly, and in the southern a south-easterly, wind must prevail, while the winds in the equatorial belt, which separates the two former, should be comparatively calm and free from any steady prevalence of easterly character. All these consequences are agreeable to observed fact, and the system of aërial currents above described constitutes in reality what is understood by the regular *trade winds*.

(245.) The constant friction thus produced between the earth and atmosphere in the regions near the equator must (it may be objected) by degrees reduce and at length destroy the rotation of the whole mass. The laws of dynamics, however, render such a consequence, generally, impossible; and it is easy to see, in the present case, where and how the compensation takes place. The heated equatorial air, while it rises and flows over towards the poles, carries with it the rotatory velocity due to its equatorial situation into a higher latitude, where the earth's surface has less motion. Hence, as it travels northward or southward, it will *gain* continually more and more on the surface of the earth in its diurnal motion, and assume constantly more and more a *westerly* relative direction; and when at length it returns to the surface, in its circulation, which it must do more or less in all the interval between the tropics and the poles, it will act on it by its friction as a powerful south-west wind in the northern hemisphere, and a north-west in the southern, and restore to it the impulse taken up from it at the equator. We have here the origin of the south-west and westerly gales so prevalent in our latitudes, and of the almost universal westerly winds in the North Atlantic, which are, in fact, nothing else than a part of the general system of the re-action of the trades, and of the process by which the equilibrium of the earth's motion is maintained under their action.[1]

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[1] See Captain Hall's "Fragments of Voyages and Travels," 2d series, vol. i. p. 162, where this is very distinctly, and, so far as I am aware, for the first time, reasoned out.

(246.) In order to construct a map or model of the earth, and obtain a knowledge of the distribution of sea and land over its surface, the forms of the outlines of its continents and islands, the courses of its rivers and mountain chains, and the relative situations, with respect to each other, of those points which chiefly interest us, as centres of human habitation, or from other causes, it is necessary to possess the means of determining correctly the situation of any proposed station on its surface. For this two elements require to be known, the latitude and longitude, the former assigning its distance from the poles or the equator, the latter, the meridian on which that distance is to be reckoned. To these, in strictness, should be added, its height above the sea level; but the consideration of this had better be deferred, to avoid complicating the subject.

(247.) The latitude of a station on a sphere would be merely the length of an arc of the meridian, intercepted between the station and the nearest point of the equator, reduced into degrees. (See art. 88.) But as the earth is elliptic, this mode of conceiving latitudes becomes inapplicable, and we are compelled to resort for our definition of latitude to a generalization of that property (art. 119,) which affords the readiest means of determining it by observation, and which has the advantage of being independent of the figure of the earth, which, after all, is not *exactly* an ellipsoid, or any known geometrical solid. The latitude of a station, then, is the altitude of the elevated pole, and is, therefore, astronomically determined by those methods already explained for ascertaining that important element. In consequence, it will be remembered that, to make a *perfectly* correct map of the whole, or any part of the earth's surface, equal differences of latitude are not represented by exactly equal intervals of surface.

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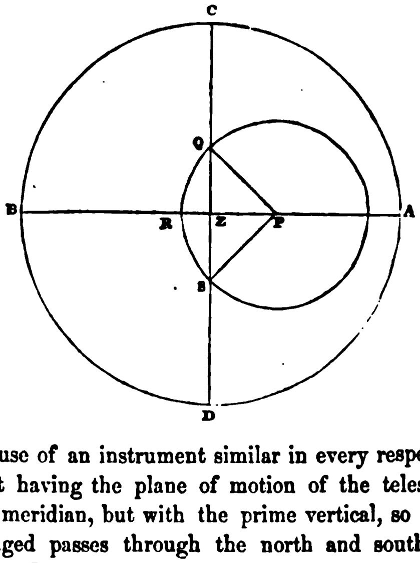
[1] As it is our object merely to illustrate the mode in which the earth's rotation affects the atmosphere on the great scale, we omit all consideration of local periodical winds, such as monsoons, &c.

It seems worth inquiry, whether hurricanes in tropical climates may not arise from portions of the upper currents prematurely diverted downwards before their relative velocity has been sufficiently reduced by friction on, and gradual mixing with, the lower strata; and so dashing upon the earth with that tremendous velocity which gives them their destructive character, and of which hardly any rational account has yet been given. But it by no means follows that this must always be the case. In general, a rapid transfer, either way, in latitude, of any mass of air which local or temporary causes might carry *above the immediate reach of the friction of the earth's surface*, would give a fearful exaggeration to its velocity. Wherever such a mass should strike the earth, a hurricane might arise; and should two such masses encounter in mid air, a tornado of any degree of intensity on record might easily result from their combination.

(248.) For the purposes of geodesical[1] measurements and trigonometrical surveys, an exceedingly correct determination of the latitudes of the most important stations is required. For this purpose, therefore, the zenith sector (an instrument capable of great precision) is most commonly used to observe stars passing the meridian near the zenith, whose declinations have become known by previous long series of observations at fixed observatories, and which are therefore called standard or fundamental stars. Recently a method[2] has been employed with great success, which consists in the use of an instrument similar in every respect to the transit instrument, but having the plane of motion of the telescope not coincident with the meridian, but with the prime vertical, so that its axis of rotation prolonged passes through the north and south points of the horizon. Let A B C D be the celestial hemisphere projected on the horizon, P the pole, Z the zenith, A B the meridian, C D the prime vertical, Q R S part of the diurnal circle of a star passing near the zenith, whose polar distance P R is but little greater than the co-latitude of the place, or the arc P Z, between the zenith and pole (art. 112.) Then the moments of this star's arrival on the prime vertical at Q and S

Fig. 87.



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[1] Γη, the earth; δεσις (from δεω, to bind,) a joining or connection (of parts.)

[2] Devised originally by Römer. Revived or re-invented by Bessel.—*Astr. Nachr.* No. 40.

will, if the instrument be correctly adjusted, be those of its crossing the middle wire in the field of view of the telescope (art. 160.) Consequently the interval between these moments will be the time of the star passing from Q to S, or the measure of the diurnal arc Q R S, which corresponds to the angle Q P S at the pole. This angle, therefore, becomes known *by the mere observation of an interval of time*, in which it is not even necessary to know the error of the clock, and in which, when the star passes near the zenith, so that the interval in question is small, even the *rate* of the clock, or its gain or loss on true sidereal time, may be neglected. Now the angle Q P S, or its half Q P R, and P Q the polar distance of the star, being known, P Z the zenith distance of the pole can be calculated by the resolution of the right-angled spherical triangle P Z Q; and thus the co-latitude (and of course the latitude) of the place of observation becomes known. The advantages gained by this mode of observation are, 1st, that no readings of a divided arc are needed, so that errors of graduation and reading are avoided: 2dly, that the arc Q R S is very much greater than its versed sine R Z, so that the difference R Z between the latitude of the place and the declination of the star is given by the observation of a magnitude very much greater than itself, or is, as it were, observed on a greatly enlarged scale. In consequence, a very minute error is entailed on R Z by the commission of even a considerable one in Q R S: 3dly, that in this mode of observation all the merely instrumental errors which affect the ordinary use of the transit instrument are either uninfluential or eliminated by simply reversing the axis.

(249.) To determine the latitude of a station, then, is easy. It is otherwise with its longitude, whose exact determination is a matter of more difficulty. The reason is this: — as there are no meridians marked upon the earth, any more than parallels of latitude, we are obliged in this case, as in the case of the latitude, to resort to marks external to the earth, *i. e.* to the heavenly bodies, for the objects of our measurement; but with this difference in the two cases — to observers situated at stations on the same *meridian* (*i. e.* differing in latitude) the heavens present different aspects at *all* moments. The portions of them which become visible in a complete diurnal rotation are not the same, and stars which are common to both describe circles differently inclined to their horizons, and differently divided by them, and attain different altitudes. On the other hand, to observers situated on the same *parallel* (*i. e.* differing only in longitude) the heavens present the same aspects. Their visible portions are the same; and the same stars describe circles equally inclined, and similarly divided by their horizons, and attain the same altitudes. In the former

is, in the latter there is *not*, any thing in the appearance of the
atched through a whole diurnal rotation, which indicates a dif-
locality in the observer.

But not two observers, at different points of the earth's surface,
t *the same instant* the same celestial hemisphere visible. Sup-
our ideas, an observer stationed at a given point of the equator,
the moment when he noticed some bright star to be in his
therefore on his meridian, he should be suddenly transported,
nt of time, round one quarter of the globe in a *westerly* direction,
t that he will no longer have the same star vertically above
ll now appear to him to be just rising, and he will have to wait
efore it again comes to his zenith, *i. e.* before the earth's rota-
west to east *carries him back again* to the line joining the star
rth's centre from which he set out.

The difference of the cases, then, may be thus stated, so as to
y to the astronomical solution of the problem of the longitude.
e of stations differing only in latitude, the same star comes to
ian at the same *time*, but at different *altitudes*. In that of
ffering only in longitude, it comes to the meridian at the same
ut at different *times*. Supposing, then, that an observer is in
of any means by which he can certainly ascertain the *time* of a

indicate $0^h\ 0^m\ 0^s$, when a certain point in the heavens, called the equinox, is on the meridian of his station. This is the *epoch* of his sidereal time; which is, therefore, entirely a *local* reckoning. It gives no information to say that an event happened at such and such an hour of sidereal time, unless we particularize the station to which the sidereal time meant appertains. Just so it is with mean or common time. This is also a local reckoning, having for its epoch *mean noon*, or the average of all the times throughout the year, when the sun is on the meridian *of that particular place to which it belongs*; and, therefore, in like manner, when we date any event by mean time, it is necessary to name the place, or particularize *what* mean time we intend. On the other hand, a date by equinoctial time is absolute, and requires no such explanatory addition.

(253.) The astronomer sets and regulates his sidereal clock by observing the meridian passages of the more conspicuous and well-known stars. Each of these holds in the heavens a certain determinate and known place with respect to that imaginary point called the equinox, and by noting the times of their passage in succession by his clock he knows when the equinox passed. At that moment his clock ought to have marked $0^h\ 0^m\ 0^s$; and if it did not, he knows and can correct its error, and by the agreement or disagreement of the errors assigned by each star he can ascertain whether his clock is correctly regulated to go twenty-four hours in one diurnal period, and if not, can ascertain and allow for its rate. Thus, although his clock may not, and indeed cannot, either be set correctly, or go truly, yet by applying its *error* and *rate* (as they are technically termed), he can correct its indications, and ascertain the exact sidereal times corresponding to them, and proper to his locality. This indispensable operation is called getting his *local time*. For simplicity of explanation, however, we shall suppose the clock a perfect instrument; or, which comes to the same thing, its error and rate applied at every moment it is consulted, and included in its indications.

(254.) Suppose, now, of two observers, at distant stations, A and B, each, independently of the other, to set and regulate his clock to the true sidereal time of his station. It is evident that if one of these clocks could be taken up without deranging its going, and set down by the side of the other, they would be found, on comparison, to differ by the exact difference of their local epochs; that is, by the time occupied by the equinox, or by any star, in passing from the meridian of A to that of B; in other words, by their difference of longitude, expressed in sidereal hours, minutes, and seconds.

(255.) A pendulum clock cannot be thus taken up and transported from place to place without derangement, but a chronometer may. Sup-

the observer at B to use a chronometer instead of a clock, he
dily transfer of the instrument to the other station, procure a
parison of sidereal times, and thus obtain his longitude from A.
if he employ a clock, yet by comparing it first with a good
r, and then transferring the latter instrument for comparison
ther clock, the same end will be accomplished, provided the
e chronometer can be depended on.

Were chronometers perfect, nothing more complete and conve-
this mode of ascertaining differences of longitude could be
An observer, provided with such an instrument, and with a por-
it, or some equivalent method of determining the local time at
station, might, by journeying from place to place, and observing
n passages of stars at each, (taking care not to alter his chro-
r let it run down,) ascertain their differences of longitude with
d precision. In this case, the same time-keeper being used at
n, if, at one of them, A, it mark true sidereal time, at any
will be just so much sidereal time in error as the difference of
of A and B is equivalent to: in other words, the longitude of
will appear as the error of the time-keeper on the local time of
travel westward, then his chronometer will appear continually
though it really goes correctly. Suppose, for instance, he set

its centre; and if in the opposite direction, one turn *less* than if he had remained upon one point of its surface: in the former case, then, he will have witnessed one alternation of day and night more, in the latter one less, than if he had trusted to the rotation of the earth alone to carry him round. As the earth revolves from west to east, it follows that a westward direction of his journey, by which he counteracts its rotation, will cause him to lose a day, and an eastward direction, by which he conspires with it, to gain one. In the former case, all his days will be longer; in the latter, shorter than those of a stationary observer. This contingency has actually happened to circumnavigators. Hence, also, it must necessarily happen that distant settlements, *on the same meridian*, will differ a day in their usual reckoning of time, according as they have been colonized by settlers arriving in an eastward or in a westward direction,—a circumstance which may produce strange confusion when they come to communicate with each other. The only mode of correcting the ambiguity, and settling the disputes which such a difference may give rise to, consists in having recourse to the equinoctial date, which can never be ambiguous.

(258.) Unfortunately for geography and navigation, the chronometer, though greatly and indeed wonderfully improved by the skill of modern artists, is yet far too imperfect an instrument to be relied on implicitly. However such an instrument may preserve its uniformity of rate for a few hours, or even days, yet in long absences from home the chances of error and accident become so multiplied, as to destroy all security of reliance on even the best. To a certain extent this may, indeed, be remedied by carrying out several, and using them as checks on each other; but, besides the expense and trouble, this is only a palliation of the evil — the great and fundamental,— as it is the only one to which *the determination of longitudes by time-keepers* is liable. It becomes necessary, therefore, to resort to other means of communicating from one station to another a knowledge of its local time, or of propagating from some principal station, as a centre, its local time as a universal standard with which the local time at any other, however situated, may be at once compared, and thus the longitudes of all places be referred to the meridian of such central point.

(259.) The simplest and most accurate method by which this object can be accomplished, when circumstances admit of its adoption, is that by telegraphic signal. Let A and B be two observatories, or other stations, provided with accurate means of determining *their respective local times*, and let us first suppose them visible from each other. Their clocks being regulated, and their errors and rates ascertained and applied, let a signal be made at A, of some sudden and definite kind, such as the flash of gun-

e explosion of a rocket, the sudden extinction of a bright light,
r which admits of no mistake, and can be seen at great dis-
he moment of the signal being made must be noted by *each*
his respective clock or watch, as if it were the transit of a star,
tronomical phenomenon, and the error and rate of the clock at
n being applied, the local time of the signal at each is deter-
nsequently, when the observers communicate their observations
nal to each other, since (owing to the almost instantaneous
n of light) it must have been seen at the same *absolute* instant
e difference of their local times, and therefore of their longitudes,
nown. For example, at A the signal is observed to happen at
id. time at A, as obtained by applying the error and rate to the
n by the clock at A, when the signal was seen there. At B the
l was seen at $5^h\ 4^m\ 0^s$, sid. time at B, similarly deduced from
oted by the clock at B, by applying *its* error and rate. Conse-
e difference of their local epochs is $4^m\ 0^s$, which is also their
of longitudes in time, or 1° 0′ 0″ in hour angle.

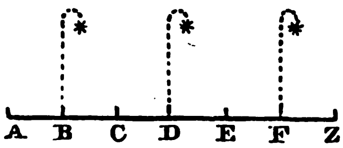
The accuracy of the final determination may be increased by
d observing several signals at stated intervals, each of which
omparison of times, and the mean of all which is, of course,
e depended on than the result of any single comparison. By
the error introduced by the comparison of clocks may be re-

mination of differences of longitude can be desired. The differences of longitude between the observatories of New York, Washington, and Philadelphia, have been very recently determined in this manner by the astronomers at those observatories.

(263.) Where no such electric communication exists, however, the interval between observing stations may be increased by causing the signals to be made not at one of them, but at an intermediate point; for, provided they are seen by both parties, it is a matter of indifference where they are exhibited. Still the interval which could be thus embraced would be very limited, and the method in consequence of little use, but for the following ingenious contrivance, by which it can be extended to any distance, and carried over any tract of country, however difficult.

(264.) This contrivance consists in establishing, between the extreme stations, whose difference of longitude is to be ascertained, and at which the local times are observed, a chain of intermediate stations, alternately destined for signals and for observers. Thus, let A and Z be the extreme stations. At B let a signal station be established, at which rockets, &c. are fired at stated intervals. At C let an observer be placed, provided with a chronometer; at D, another signal station; at E, another observer and chronometer; till the whole line is occupied by stations so arranged, that the signal at B can be seen from A and C; those at D, from C and E; and so on. Matters being thus arranged, and the errors and rates of the clocks at A and Z ascertained by astronomical observation, let a signal be made at B, and observed at A and C, and the times noted. Thus the difference between A's clock and C's chronometer becomes known. After a short interval (five minutes for instance) let a signal be made at D, and observed by C and E. Then will the difference between their respective chronometers be determined; and the difference between the former and the clock at A being already ascertained, the difference between the clock A and chronometer E is therefore known. This, however, supposes that the intermediate chronometer C has kept true sidereal time, or at least a known rate, in the interval between the signals. Now this interval is purposely made so very short, that no instrument of any pretensions to

Fig. 38.



an possibly produce an appreciable amount of error in its lapse
ns from its usual *rate*. Thus the time propagated from A to
onsidered as handed over, without gain or loss (save from error
ion), to E. Similarly, by the signal made at F, and observed
, the time so transmitted to E is forwarded on to Z; and thus
the clocks at A and Z are compared. The process may be
s often as is necessary to destroy error by a mean of results;
the line of stations is numerous, by keeping up a succes-
nals, so as to allow each observer to note alternately those on
which is easily pre-arranged, many comparisons may be kept
ong the line at once, by which time is saved, and other advan-
ned.[1] In important cases the process is usually repeated on
hts in succession.

In place of artificial signals, natural ones, when they occur
definite for observation, may be equally employed. In a clear
umber of those singular meteors, called shooting stars, which
erved, is often very great, especially on the 9th and 10th of
d some other days, as November 12 and 13; and as they are
heir appearance and disappearance, and from the great height
ey have been ascertained to take place are visible over exten-
s of the earth's surface, there is no doubt that they may be

(266.) Another species of natural signal, visible at once over a whole terrestrial hemisphere, is afforded by the eclipses of Jupiter's satellites, of which we shall speak more at large when we come to treat of those bodies. Every such eclipse is an event which possesses one great advantage in its applicability to the purpose in question, viz. that the time of its happening, at any fixed station, such as Greenwich, can be *predicted* from a long course of previous recorded observation and calculation thereon founded, and that this prediction is sufficiently precise and certain, to stand in the place of a corresponding observation. So that an observer at any other station wherever, who shall have observed one or more of these eclipses, and ascertained his local time, instead of waiting for a communication with Greenwich, to inform him at what moment the eclipse took place there, may use the *predicted Greenwich time* instead, and thence, at once, and on the spot, determine his longitude. This mode of ascertaining longitudes is, however, as will hereafter appear, not susceptible of great exactness, and should only be resorted to when others cannot be had. The nature of the observation also is such that it cannot be made at sea[1]; so that, however useful to the geographer, it is of no advantage to navigation.

(267.) But such phenomena as these are of only occasional occurrence; and in their intervals, and when cut off from all communication with any fixed station, it is indispensable to possess some means of determining longitudes, on which not only the geographer may rely for a knowledge of the exact position of important stations on land in remote regions, but on which the navigator can securely stake, at every instant of his adventurous course, the lives of himself and comrades, the interests of his country, and the fortunes of his employers. Such a method is afforded by Lunar Observations. Though we have not yet introduced the reader to the phenomena of the moon's motion, this will not prevent us from giving here the exposition of the principle of the lunar method; on the contrary, it will be highly advantageous to do so, since by this course we shall have to deal with the naked principle, apart from all the peculiar sources of difficulty with which the lunar theory is encumbered, but

[1] To accomplish this is still a desideratum. Observing chairs, suspended with studious precaution for ensuring freedom of motion, have been resorted to, under the vain hope of mitigating the effect of the ship's oscillation. The opposite course seems more promising, viz. to merely deaden the motion by a somewhat stiff suspension (as by a coarse and rough cable), and by friction strings attached to weights running through loops (not pulleys) fixed in the wood-work of the vessel. At least, such means have been found by the author of singular efficacy in increasing personal comfort in the suspension of a cot.

n fact, completely extraneous to the *principle* of its applica-
problem of the longitudes, which is quite elementary.

f there were in the heavens a clock furnished with a dial-plate
which always marked Greenwich time, the longitude of any
ld be at once determined, so soon as the *local time* was known,
ng it with this clock. Now, the offices of the dial-plate and
clock are these: — the former carries a set of marks upon it,
ion is known; the latter, by passing over and among these
rm us, by the place it holds with respect to them, what it is
what time has elapsed since a certain moment when it stood at
lar spot.

n a clock the marks on the dial-plate are uniformly distributed
he circumference of a circle, whose centre is that on which the
ve with a uniform motion. But it is clear that we should, with
nty, though with much more trouble, tell what o'clock it were,
s on the dial-plate were *un*equally distributed,— if the hands
tric, and their motion *not* uniform,— provided we knew, 1st,
ntervals round the circle at which the hour and minute marks
l; which would be the case if we had them all registered in a
the results of previous careful measurement: — 2dly, if we
xact amount and direction of excentricity of the centre of mo-

relative position, the notice of a few nights, or even hours, will satisfy the commencing student, and this is all that at present we require.

(271.) There is only one circumstance wanting to make our analogy complete. Suppose the hands of our clock, instead of moving *quite close* to the dial-plate, were considerably elevated above, or distant in front of of it. Unless, then, in viewing it, we kept our eye just in the line of their centre, we should not see them exactly thrown or *projected* upon their proper places on the dial. And if we were either unaware of this cause of optical change of place, this *parallax* — or negligent in not taking it into account — we might make great mistakes in reading the time, by referring the hand to the wrong mark, or incorrectly appreciating its distance from the right. On the other hand, if we took care to note, in every case when we had occasion to observe the time, the exact position of the eye, there would be no difficulty in ascertaining and allowing for the precise influence of this cause of apparent displacement. Now, this is just what obtains with the apparent motion of the moon among the stars. The former (as will appear) is comparatively near to the earth — the latter immensely distant; and in consequence of our not occupying the centre of the earth, but being carried about on its surface, and constantly changing place, there arises a *parallax*, which displaces the moon apparently among the stars, and must be allowed for before we can tell the true place she would occupy if seen from the centre.

(272.) Such a clock as we have described might, no doubt, be considered a very bad one; but if it were our *only* one, and if incalculable interests were at stake on a perfect knowledge of time, we should justly regard it as most precious, and think no pains ill bestowed in studying the laws of its movements, or in facilitating the means of *reading* it correctly. Such, in the parallel we are drawing, is the lunar theory, whose object is to reduce to regularity, the indications of this strangely irregular-going clock, to enable us to predict, long beforehand, and with absolute certainty, *whereabouts* among the stars, at every hour, minute, and second, in every day of every year, in Greenwich local time, the moon *would* be seen from the earth's centre, and *will* be seen from every accessible point of its surface; and such is the *lunar method* of longitudes. The moon's apparent angular distance from all those principal and conspicuous stars which lie in its course, as seen from the earth's centre, are computed and tabulated with the utmost care and precision in almanacks published under national control. No sooner does an observer, in any part of the globe, at sea or on land, measure its actual distance from any one of those *standard stars* (whose places in the heavens have been ascertained for the purpose with the most anxious solicitude,) than

fact, performed that comparison of his local time with the
of every observatory in the world, which enables him to as-
difference of longitude from one or all of them.

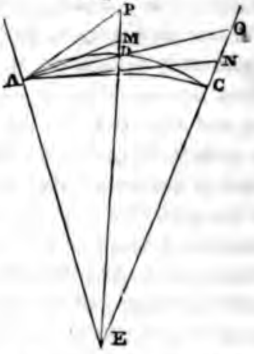
The latitudes and longitudes of any number of points on the
face may be ascertained by the methods above described; and
ing down a sufficient number of principal points, and filling in
ediate spaces by local surveys, might maps of countries be
. In practice, however, it is found simpler and easier to
particular nation into a series of great triangles, the angles
re stations conspicuously visible from each other. Of these
he *angles* only are measured by means of the *theodolite*, with
ion of *one side* only of *one triangle*, which is called *a base*,
is measured with every refinement which ingenuity can devise
command. This *base* is of moderate extent, rarely surpassing
n miles, and purposely selected in a perfectly horizontal plane,
conveniently adapted to the purposes of measurement. Its
een its two extreme points (which are dots on plates of gold
let into massive blocks of stone, and which are, or at least
e, in all cases preserved with almost religious care, as monu-
rds of the highest importance,) is then measured, with every
to ensure precision,' and its position with respect to the
s well as the geographical positions of its extremities, carefully

the country, by whose connection its whole surface may be covered, as it were, with a network of triangles. Now, it is evident that the angles of the triangle A, B, C being observed, and one of its sides, A, B, measured, the other two sides, A C, B C, may be calculated by the rules of trigonometry; and thus each of the sides A C and B C becomes in its turn a *base* capable of being employed as known sides of other triangles. For instance, the angles of the triangles A C G and B C F being known by observation, and their sides A C and B C, we can thence calculate the lengths A G, C G, and B F, C F. Again, C G and C F being known and the included angle G C F, G F may be calculated, and so on. Thus may all the stations be accurately determined and laid down, and as this process may be carried on to any extent, a map of the whole country may be thus constructed, and filled in to any degree of detail we please.

(275.) Now, on this process there are two important remarks to be made. The first is, that it is necessary to be careful in the selection of stations, so as to form triangles free from any *very* great inequality in their angles. For instance, the triangle K B F would be a very improper one to determine the situation of F from observations at B and K, because the angle F being very acute, a small error in the angle K would produce a great one in the place of F *upon the line B F.* Such *ill-conditioned* triangles, therefore, must be avoided. But if this be attended to, the accuracy of the determination of the calculated sides will not be much short of that which would be obtained by actual measurement (were it practicable); and, therefore, as we recede from the base on all sides as a centre, it will speedily become practicable to use *as bases*, the sides of *much larger* triangles, such as G F, G H, H K, &c.; by which means the next step of the operation will come to be carried on on a much larger scale, and embrace far greater intervals, than it would have been safe to do (for the above reason) in the immediate neighbourhood of the base. Thus it becomes easy to divide the whole face of a country into *great triangles* of from 80 to 100 miles in their sides (according to the nature of the ground), which, being once well determined, may be afterwards, by a second series of subordinate operations, broken up into smaller ones, and these again into others of a still minuter order, till the final filling in is brought within the limits of personal survey and draftsmanship, and till a map is constructed, with any required degree of detail.

(276.) The next remark we have to make is, that all the triangles in question are not, rigorously speaking, *plane*, but *spherical* — existing on the surface of a sphere, or rather, to speak correctly, of an ellipsoid. In very small triangles, of six or seven miles in the side, this may be neglected, as the difference is imperceptible; but in the larger ones it ken into consideration. It is evident that, as every object used g the telescope of a theodolite has some certain *elevation*, not the *soil*, but above the level of the *sea*, and as, moreover, tions differ in every instance, a *reduction to the horizon* of all ed angles would appear to be required. But, in fact, by the n of the theodolite (art. 192), which is nothing more than an d azimuth instrument, this reduction is *made* in the very act

Fig. 40.



tion, is to regard the network of triangles with which the country is covered, as the bases of an assemblage of pyramids converging to the centre of the earth. The theodolite gives us *the true measures of the angles included by the planes of these pyramids;* and the surface of an imaginary sphere on the level of the sea intersects them in an assemblage of spherical triangles, above whose angles, in the radii prolonged, the real stations of observation are raised, by the superficial inequalities of mountain and valley. The operose calculations of spherical trigonometry which this consideration would seem to render necessary for the reductions of a survey, are dispensed with in practice by a very simple and easy rule, called the *rule for the spherical excess*, which is to be found in most works on trigonometry. If we would take into account the ellipticity of the earth, it may also be done by appropriate processes of calculation, which, however, are too abstruse to dwell upon in a work like the present.

(278.) Whatever process of calculation we adopt, the result will be a reduction to the level of the sea, of all the triangles, and the consequent determination of the geographical latitude and longitude of every station observed. Thus we are at length enabled to construct maps of countries; to lay down the outlines of continents and islands; the courses of rivers; the places of cities, towns and villages; the direction of mountain ridges, and the places of their principal summits; and all those details which, as they belong to physical and statistical, rather than to astronomical geography, we need not here dilate on. A few words, however, will be necessary respecting maps, which are used as well in astronomy as in geography.

(279.) A map is nothing more than a representation, upon a plane, of some portion of the surface of a sphere, on which are traced the particulars intended to be expressed, whether they be continuous outlines or points. Now, as a spherical surface[1] can by no contrivance be extended or projected into a plane, without undue enlargement or contraction of some parts in proportion to others; and as the system adopted in so extending or projecting it will decide *what* parts shall be enlarged or relatively contracted, and in what proportions; it follows, that when large portions of the sphere are to be mapped down, a great difference in their representations may subsist, according to the system of projection adopted.

(280.) The projections chiefly used in maps, are the *orthographic*, *stereographic*, and *Mercator's*. In the *orthographic* projection, every point of the hemisphere is referred to its diametral plane or base, by a

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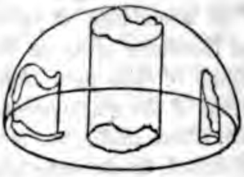
[1] We here neglect the ellipticity of the earth, which, for such a purpose as map-making, is too trifling to have any material influence.

Fig. 41.



lar let fall on it, so that the representation of the hemisphere
ed on its base, is such as would actually appear to an eye placed
ite distance from it. It is obvious, from the annexed figure,
is projection only the central portions are represented of their
, while all the exterior is more and more distorted and crowded
we approach the edges of the map. Owing to this cause, the
ic projection, though very good for small portions of the globe,
service for large ones.

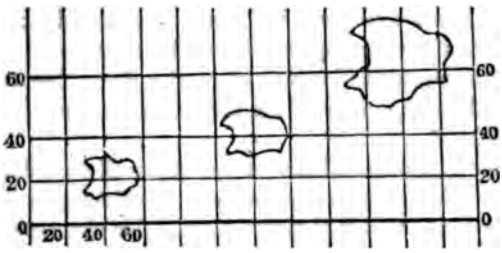
The *stereographic* projection is in great measure free from this
o understand this projection, we must conceive an eye to be
E, one extremity of a diameter, E C B, of the sphere, and to
concave surface of the sphere, every point of which, as P, is
the diametral plane A D F, perpendicular to E B by the

circles in the projection. Thus the circle X is projected into x. Only great circles passing through the vertex B are projected into straight lines traversing the centre C: thus, B P A is projected into C A.

2dly. Every very small triangle, G H K, on the sphere, is represented by a *similar* triangle, $g\ h\ k$, in the projection. This is a very valuable property, as it insures a general similarity of appearance in the map to the reality in all its parts, and enables us to project at least a hemisphere in a single map, without any violent distortion of the configurations on the surface from their real forms. As in the orthographic projection, the borders of the hemisphere are unduly crowded together; in the stereographic, their projected dimensions are, on the contrary, somewhat enlarged in receding from the centre.

(283.) Both these projections may be considered *natural* ones, inasmuch as they are really perspective representations of the surface on a plane. Mercator's is entirely an artificial one, representing the sphere as it cannot be seen from any one point, but as it might be seen by an eye carried successively over every part of it. In it, the degrees of *longitude*, and those of *latitude*, bear always to each other their due proportion: the equator is conceived to be extended out into a straight line, and the meridians are straight lines at right angles to it, as in the figure. Altogether, the general character of maps on this projection is not very dissimilar to what would be produced by referring every point in the globe to a circumscribing cylinder, by lines drawn from the centre, and then unrolling the cylinder into a plane. Like the stereographic projection, it gives a true representation, as to *form*, of every particular small part, but varies greatly in point of *scale* in its different regions; the polar portions in particular being extravagantly enlarged; and the whole map, even of a single hemisphere, not being comprisable within any finite limits.

Fig. 43.



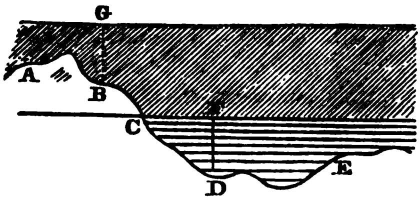
(284.) We shall not, of course, enter here into any geographical details; but one result of maritime discovery on the great scale is, so to

sive enough to call for mention as an astronomical feature.
continents and seas are laid down on a globe (and since the
f Australia and the recent addition to our antarctic knowledge
Land by Sir J. C. Ross, we are sure that no very extensive
nd remain unknown), we find that it is possible so to divide
nto two hemispheres, that one shall contain *nearly all the land*;
eing almost entirely sea. It is a fact, not a little interesting
men, and, combined with our insular station in that great high-
ions, the Atlantic, not a little explanatory of our commercial
that London[1] occupies nearly the centre of the terrestrial hemi-
stronomically speaking, the fact of this divisibility of the
an oceanic and a terrestrial hemisphere is important, as demon-
a want of absolute equality in the density of the solid mate-
two hemispheres. Considering the whole mass of land and
a state of *equilibrium*, it is evident that the half which pro-
st of necessity be *buoyant;* not, of course, that we mean to
be lighter than *water*, but, as compared with the whole globe,
egree heavier than that fluid. We leave to geologists to draw
premises their own conclusions (and we think them obvious
s to the internal constitution of the globe, and the immediate
the forces which sustain its continents at their actual elevation;

case our sounding is real and tangible; in the other, an imaginary one, measured by the length of the column of quicksilver the superincumbent air is capable of counterbalancing.

(286.) Suppose that instead of air, the earth and ocean were covered with oil, and that human life could subsist under such circumstances. Let A B C D E be a continent, of which the portion A B C projects

Fig. 44.



above the water, but is covered by the oil, which also floats at an uniform depth on the whole ocean. Then if we would know the depth of any point D below the sea-level, we let down a plummet from F. But, if we would know the height of B above the same level, we have only to send up a float from B to the surface of the oil; and *having done the same at C, a point at the sea level, the difference of the two float lines gives the height in question.*

(287.) Now, though the atmosphere differs from oil in not having a positive *surface* equally definite, and in not being capable of carrying up any float adequate to such an use, yet it possesses all the properties of a fluid really essential to the purpose in view, and this in particular,— that, over the whole surface of the globe, its *strata of equal density* supposed in a state of equilibrium, are parallel to the surface of equilibrium, or to what *would be* the surface of the sea, if *prolonged under the continents,* and therefore each or any of them has all the characters of a definite surface to measure from, provided it can be ascertained and identified. Now, the height at which, at any station B, the mercury in a barometer is supported, informs us *at once* how much of the atmosphere is incumbent on B, or, in other words, *in what stratum* of the general atmosphere (indicated by its density) B is situated: whence we are enabled finally to conclude, by mechanical reasoning,[1] at what height above the sea-level *that degree of density* is to be found over the whole surface of the globe. Such is the principle of the application of the barometer to the measurement of heights. For details, the reader is referred to other works.[2]

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[1] Newton's Princip. ii. Prop. 22.

[2] Biot, Astronomie Physique, vol. iii. For tables, see the work of Biot cited. Also those of Oltmann, annually published by the French board of longitudes in their Annuaire; and Mr. Baily's collection of Astronomical Tables and Formulæ.

We will content ourselves here with a general caution against
t dependence on barometric measurements, except as a differ-
ess, at stations not too remote from each other. They rely in
cation on the assumption of a state of equilibrium in the atmo-
ata over the whole globe — which is very far from being their
e (art. 37.) Winds, especially steady and general currents
over extensive continents, undoubtedly *tend* to produce some
conformity in the curvature of these strata to the *general* form
l-surface, and therefore to give an undue elevation to the mer-
mn at some points. On the other hand, the existence of
n the earth's surface where a permanent depression of the
prevails to the astonishing extent of nearly an inch, has been
ved by the observations of Ermann in Siberia and of Ross in
tic Seas, and is probably a result of the same cause, and may
ed as complementary to an undue habitual elevation in other

Possessed of a knowledge of the height of stations above the
ay connect all stations at the same altitude by level lines, the
which will be the outline of the sea-coast; and the rest will
the successive coast-lines which would take place were the sea
regular and equal accessions of level over the whole world, till

CHAPTER V.

OF URANOGRAPHY.

CONSTRUCTION OF CELESTIAL MAPS AND GLOBES BY OBSERVATIONS OF RIGHT ASCENSION AND DECLINATION. — CELESTIAL OBJECTS DISTINGUISHED INTO FIXED AND ERRATIC. — OF THE CONSTELLATIONS. — NATURAL REGIONS IN THE HEAVENS. — THE MILKY WAY. — THE ZODIAC. — OF THE ECLIPTIC. — CELESTIAL LATITUDES AND LONGITUDES. — PRECESSION OF THE EQUINOXES. — NUTATION. — ABERRATION. — REFRACTION. — PARALLAX. — SUMMARY VIEW OF THE URANOGRAPHICAL CORRECTIONS.

(290.) The determination of the relative situations of objects in the heavens, and the construction of maps and globes which shall truly represent their mutual configurations as well as of catalogues which shall preserve a more precise numerical record of the position of each, is a task at once simpler and less laborious than that by which the surface of the earth is mapped and measured. Every star in the great constellation which appears to revolve above us, constitutes, so to speak, a celestial station; and among these stations we may, as upon the earth, triangulate, by measuring with proper instrument their angular distances from each other, which, cleared of the effect of refraction, are then in a state for laying down on charts, as we would the towns and villages of a country: and this without moving from our place, at least for all the stars which rise above our horizon.

(291.) Great exactness might, no doubt, be attained by this means, and excellent celestial charts constructed; but there is a far simpler and easier, and at the same time, infinitely more accurate course laid open to us if we take advantage of the earth's rotation on its axis, and by observing each celestial object as it passes our meridian, refer it separately and independently to the celestial equator, and thus ascertain its place on the surface of an imaginary sphere, which may be conceived to revolve with it, and on which it may be considered as projected.

(292.) The right ascension and declination of a point in the heavens

to the longitude and latitude of a station on the earth; and
f a star on the celestial sphere is determined, when the former
re known, just as that of a town on a map, by knowing the
he great advantages which the method of meridian observation
ver that of triangulation from star to star, are, then, 1st, That
star is observed in that point of its diurnal course, when it is
nd least displaced by refraction. 2dly, That the instruments
he transit and meridian circle) are the simplest and least liable
derangement of any used by astronomers. 3dly, That all the
s can be made systematically, in regular succession, and with
ntages; there being here no question about advantageous or
geous triangles, &c. And, lastly, That, by adopting this
very quantities which we should otherwise have to calculate
d tedious operations of spherical trigonometry, and which are
the formation of a catalogue, are made the objects of imme-
urement. It is almost needless to state, then, that this is the
pted by astronomers.

To determine the right ascension of a celestial object, all that
y is to observe the moment of its meridian passage with a
rument, by a clock regulated to exact sidereal time, or reduced
applying its known error and rate. The *rate* may be obtained

obtained by celestial observation. 2dly, And more directly, by observation of their *polar distances* on the mural circle, as explained in art. 170, which is independent of any previous determination of the latitude of the station; neither, however, in this case, does observation give directly and immediately the *exact* declinations. The observations require to be corrected, first for refraction, and moreover for those minute causes of inequality which have been just alluded to in the case of right ascensions.

(296.) In this manner, then, may the places, one among the other, of all celestial objects be ascertained, and maps and globes constructed. Now here arises a very important question. How far are these places permanent? Do these stars and the greater luminaries of heaven preserve for ever one invariable connection and relation of place *inter se*, as if they formed part of a solid though invisible firmament; and, like the great natural land-marks on the earth, preserve immutably the same distances and bearings each from the other? If so, the most rational idea we could form of the universe would be that of an earth at absolute rest in the centre, and a hollow crystalline sphere circulating round it, and carrying sun, moon, and stars along in its diurnal motion. If not, we must dismiss all such notions, and inquire individually into the distinct history of each object, with a view to discovering the laws of its peculiar motions, and whether any and what other connection subsists between them.

(297.) So far is this, however, from being the case, that observations, even of the most cursory nature, are sufficient to show that some, at least, of the celestial bodies, and those the most conspicuous, are in a state of continual change of place among the rest. In the case of the moon, indeed, the change is so rapid and remarkable, that its alteration of situation with respect to such bright stars as may happen to be near it may be noticed any fine night in a few hours; and if noticed on two successive nights, cannot fail to strike the most careless observer. With the sun, too, the change of place among the stars is constant and rapid; though, from the invisibility of stars to the naked eye in the day-time, it is not so readily recognised, and requires either the use of telescopes and angular instruments to measure it, or a longer continuance of observation to be struck with it. Nevertheless, it is only necessary to call to mind its greater meridian altitude in summer than in winter, and the fact that the stars which come into view at night (and which are therefore situated in an hemisphere opposite to that occupied by the sun, and having that luminary for its centre) vary with the season of the year, to perceive that a great change must have taken place in that interval in its relative situation with respect to all the stars. Besides the sun and moon, too, there are several other bodies, called planets, which, for the most part, appear to the naked eye only as the largest and most brilliant stars, and which offer the same phenomenon of a constant change of place among the stars; now approaching, and now receding from, such of them as we may refer them to as marks; and, some in longer, some in shorter periods, making, like the sun and moon, the complete tour of the heavens.

(298.) These, however, are exceptions to the general rule. The innumerable multitude of the stars which are distributed over the vault of the heavens form a constellation, which preserves, not only to the eye of the casual observer, but to the nice examination of the astronomer, a uniformity of aspect which, when contrasted with the perpetual change in the configurations of the sun, moon, and planets, may well be termed invariable. It is true, indeed, that, by the refinement of exact measurements prosecuted from age to age, some small changes of apparent place, attributable to no illusion and to no *terrestrial* cause, have been detected in many of them. Such are called, in astronomy, the *proper motions* of the stars. But these are so excessively slow, that their accumulated amount (even in those stars for which they are greatest) has been insufficient, in the whole duration of astronomical history, to produce any obvious or material alteration in the appearance of the starry heavens.

(299.) This circumstance, then, establishes a broad distinction of the heavenly bodies into two great classes; — the fixed, among which (unless in a course of observations continued for many years) no change of mutual situation can be detected; and the erratic, or wandering — (which is implied in the word planet[1]) — including the sun, moon, and planets, as well as the singular class of bodies termed comets, in whose apparent places among the stars, and among each other, the observation of a few days, or even hours, is sufficient to exhibit an indisputable alteration.

(300.) Uranography, then, as it concerns the fixed celestial bodies (or, as they are usually called, the *fixed stars*), is reduced to a simple marking down of their relative places on a globe or on maps; to the insertion on that globe, in its due place in the great constellation of the stars, of the pole of the heavens, or the vanishing point of parallels to the earth's axis; and of the equator and place of the equinox: points and circles these, which, though artificial, and having reference entirely to our earth, and therefore subject to all changes (if any) to which the earth's axis may be liable, are yet so convenient in practice, that they have obtained an admission (with some other circles and lines), sanctioned by usage, in all globes and planispheres. The reader, however, will take care to keep

[1] Πλανητης, a wanderer.

them separate in his mind, and to familiarize himself with the idea rather of *two* or more celestial globes, superposed and fitting on each other, on one of which — a real one — are inscribed the stars; on the others those imaginary points, lines, and circles, which astronomers have devised for their own uses, and to aid their calculations; and to accustom himself to conceive in the latter or artificial spheres a capability of being shifted in any manner upon the surface of the other; so that, should experience demonstrate (as it does) that these artificial points and lines are brought, by a slow motion of the earth's axis, or by other *secular variations* (as they are called), to coincide, at very distant intervals of times, with different stars, he may not be unprepared for the change, and may have no confusion to correct in his notions.

(301.) Of course we do not here speak of those uncouth figures and outlines of men and monsters, which are usually scribbled over celestial globes and maps, and serve, in a rude and barbarous way, to enable us to talk of groups of stars, or districts in the heavens, by names which, though absurd or puerile in their origin, have obtained a currency from which it would be difficult to dislodge them. In so far as they have really (as some have) any slight resemblance to the figures called up in imagination by a view of the more splendid "constellations," they have a certain convenience; but as they are otherwise entirely arbitrary, and correspond to no *natural* subdivisions or groupings of the stars, astronomers treat them lightly, or altogether disregard them,[1] except for briefly *naming* remarkable stars, as α Leonis, β Scorpii, &c. &c., by letters of the Greek alphabet attached to them. The reader will find them on any celestial charts or globes, and may compare them with the heavens, and there learn for himself their position.

(302.) There are not wanting, however, *natural* districts in the heavens, which offer great peculiarities of character, and strike every observer: such is the *milky way*, that great luminous band, which stretches, every evening, all across the sky, from horizon to horizon, and which, when traced with diligence, and mapped down, is found to form a zone completely encircling *the whole sphere*, almost in a great circle, which is neither an *hour* circle, nor coincident with any other of our astronomical *grammata*. It is divided in one part of its course, sending off a kind of branch, which unites again with the main body, after remaining distinct for about 150 degrees, within which it suffers an interruption in its continuity. This remarkable belt has maintained, from the earliest ages, the same relative situation among the stars; and, when examined through powerful telescope, is found (wonderful to relate!) *to consist entirely of stars scattered by millions*, like glittering dust, on the black ground of the general heavens. It will be described more particularly in the subsequent portion of this work.

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[1] This disregard is neither supercilious nor causeless. The constellations seem to have been almost purposely named and delineated to cause as much confusion and inconvenience as possible. Innumerable snakes twine through long and contorted areas of the heavens, where no memory can follow them; bears, lions, and fishes, large and small, northern and southern, confuse all nomenclature, &c. A better system of constellations might have been a material help as an artificial memory.

(303.) Another remarkable region in the heavens is the *zodiac*, not from any thing peculiar in its own constitution, but from its being the area within which the apparent motions of the sun, moon, and all the greater planets are confined. To trace the path of any one of these, it is only necessary to ascertain, by continued observation, its places at successive epochs, and entering these upon our map or sphere in sufficient number to form a series, not too far disjoined, to connect them by lines from point to point, as we mark out the course of a vessel at sea by mapping down its place from day to day. Now when this is done, it is found, first, that the apparent path, or track, of the sun on the surface of the heavens, is no other than an exact great circle of the sphere which is called the *ecliptic*, and which is inclined to the equinoctial at an angle of about 23° 28′, intersecting it at two opposite points, called the equinoctial points, or equinoxes, and which are distinguished from each other by the epithets vernal and autumnal; the vernal being that at which the sun crosses the equinoctial from south to north; the autumnal, when it quits the northern and enters the southern hemisphere. Secondly, that the moon and all the planets pursue paths which, in like manner, encircle the whole heavens, but are not, like that of the sun, great circles exactly returning into themselves and bisecting the sphere, but rather spiral curves of much complexity, and described with very unequal velocities in their different parts. They have all, however, this in common, that the *general direction* of their motions is the same with that of the sun, viz. from *west to east*, that is to say, the contrary to that in which both they and the stars appear to be carried by the diurnal motion of the heavens; and, moreover, that they never deviate far from the ecliptic on either side, crossing and recrossing it at regular and equal intervals of time, and confining themselves within a zone, or belt (the *zodiac* already spoken of), extending (with certain exceptions among the smaller planets) not further than 8° or 9° on either side of the ecliptic.

(304.) It would manifestly be useless to map down on globes or charts the apparent paths of any of those bodies which never retrace the same course, and which, therefore, demonstrably, must occupy at some one moment or other of their history, every point in the area of that zone of the heavens within which they are circumscribed. The apparent complication of their movements arise (that of the moon excepted) from our viewing them from a station which is itself in motion, and would disappear, could we shift our point of view and observe them from the sun. On the other hand the apparent motion of the sun is presented to us under its least involved form, and is studied, from the station we occupy, to the greatest advantage. So that, independent of the importance of that luminary to us in other respects, it is by the investigation of the laws of its motions in the first instance that we must rise to a knowledge of those of all the other bodies of our system.

(305.) The ecliptic, which is its apparent path among the stars, is traversed by it in the period called the *sidereal year*, which consists of $365^d\ 6^h\ 9^m\ 9{\cdot}6^s$, reckoned in mean solar time or $366^d\ 6^h\ 9^m\ 9{\cdot}6^s$ reckoned in sidereal time. The reason of this difference (and it is this which constitutes the origin of the difference between solar and sidereal time) is, that as the sun's apparent annual motion *among* the stars is performed in a contrary direction to the apparent *diurnal* motion of both sun and stars, it comes to the same thing as if the diurnal motion of the sun were so much *slower* than that of the stars, or as if the sun lagged behind them in its daily course. When this has gone on for a whole year, the sun will have fallen behind the stars by a whole circumference of the heavens — or, in other words — in a year the sun will have made fewer diurnal revolutions, by one, than the stars. So that the same interval of time which is measured by $366^d\ 6^h$, &c. of sidereal time, will be called 365 days, 6 hours, &c., if reckoned in mean solar time. Thus, then, is the proportion between the mean solar and sidereal time established, which, reduced into a decimal fraction, is that of 1·00273791 to 1. The measurement of time by these different standards may be compared to that of space by the standard feet, or ells of two different nations; the proportions of which, once settled and borne in mind, can never become a source of error.

(306.) The position of the ecliptic among the stars may, for our present purpose, be regarded as invariable. It is true that this is not strictly the case; and on comparing together its position at present with that which it held at the most distant epoch at which we possess observations, we find evidences of a small change, which theory accounts for, and whose nature will be hereafter explained; but that change is so excessively slow, that for a great many successive years, or even for whole centuries, this circle may be regarded, for most ordinary purposes, as holding the same position in the sidereal heavens.

(307.) The *poles of the ecliptic*, like those of any other great circle of the sphere, are opposite points on its surface, equidistant from the

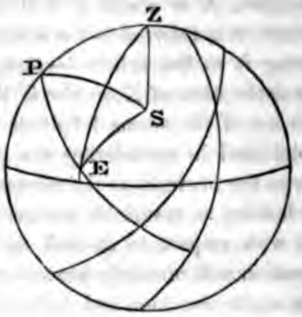
very direction. They are of course not coincident with those
inoctial, but removed from it by an angular interval equal to
ion of the ecliptic to the equinoctial (23° 28′), which is called
ity of the ecliptic. In the next figure, if P *p* represent the
south poles (by which when used without qualification we al-
the poles of *the equinoctial*), and E A Q V the equinoctial,
he ecliptic, and K *k*, its poles — the spherical angle Q V S is
ty of the ecliptic, and is equal in angular measure to P K or
ve suppose the sun's apparent motion to be in the direction
V will be the *vernal* and A the *autumnal equinox.* S and W,
ints at which the ecliptic is most distant from the equinoctial,
solstices, because, when arrived there, the sun ceases to recede
uator, and (in that sense, so far as its motion in declination is
to stand still in the heavens. S, the point where the sun has
t *northern* declination, is called the *summer*, and W, that where
st south, the *winter* solstice. These epithets obviously have
in the dependence of the seasons on the sun's declination,
be explained in the next chapter. The circle E K P Q *k p*,
es through the poles of the ecliptic and equinoctial, is called
l colure; and a meridian drawn through the equinoxes, P V
uinoctial colure. #

arc of the ecliptic intercepted between the vernal equinox and this circle, its *longitude*. In the figure, X is a star, P X R a circle of declination drawn through it, by which it is referred to the equinoctial, and K X T a circle of latitude referring it to the ecliptic — then, as V R is the right ascension, and R X the declination, of X, so also is V T its longitude, and T X its latitude. The use of the terms longitude and latitude, in this sense, seems to have originated in considering the ecliptic as forming a kind of natural equator to the heavens, as the terrestrial equator does to the earth — the former holding an invariable position with respect to the stars, as the latter does with respect to stations on the earth's surface. The force of this observation will presently become apparent.

(309.) Knowing the right ascension and declination of an object, we may find its longitude and latitude, and *vice versâ*. This is a problem of great use in physical astronomy — the following is its solution: — In our last figure, E K P Q, the solstitial colure is of course 90° distant from V, the vernal equinox, which is one of its poles — so that V R (the right ascension) being given, and also V E, the arc E R, and its measure, the spherical angle E P R, or K P X, is known. In the spherical triangle K P X, then, we have given, 1st, The side P K, which, being the distance of the poles of the ecliptic and equinoctial, is equal to the obliquity of the ecliptic; 2d, The side P X, the *polar distance*, or the complement of the declination R X; and, 3d, the included angle K P X; and therefore, by spherical trigonometry, it is easy to find the other side K X, and the remaining angles. Now K X is the complement of the required latitude X T, and the angle P K X being known, and P K V being a right angle (because S V is 90°), the angle X K V becomes known. Now this is no other than the measure of the longitude V T of the object. The inverse problem is resolved by the same triangle, and by a process exactly similar.

(310.) It is often of use to know the situation of the ecliptic in the visible heavens at any instant; that is to say, the points where it cuts the horizon, and the altitude of its highest point, or, as it is sometimes called, the *nonagesimal* point of the ecliptic, as well as the longitude of this point on the ecliptic itself from the equinox. These, and all questions referable to the same data and quæsita, are resolved by the spherical triangle Z P E, formed by the zenith Z (considered as the pole of the horizon), the pole of the equinoctial P, and the pole of the ecliptic E. The sidereal time being given, and also the right ascension of the pole of the ecliptic (which is always the same, viz. $18^h\ 0^m\ 0^s$), the hour angle Z P E of that point is known. Then, in this triangle we have given P Z, the colatitude; P E, the polar distance of the pole of the ecliptic, 23° 28′, and the angle Z P E

Fig. 46.



we may find, 1st, the side Z E, which is easily seen to be
altitude of the nonagesimal point sought; and 2dly, the angle
ch is the azimuth of the pole of the ecliptic, and which, there-
added to and subtracted from 90°, gives the azimuth of the
western intersections of the ecliptic with the horizon. Lastly,
de of the nonagesimal point may be had, by calculating in the
le the angle P E Z, which is its complement.

he *angle of situation* of a star is the angle included between equinox from east to west along the former, conspires (speaking generally) with the diurnal motion, and carries it, with reference to that motion, continually in advance upon the stars: hence it has acquired the name of the *precession of the equinoxes*, because the place of the equinox among the stars, at every subsequent moment, *precedes* (with reference to the diurnal motion) that which it held the moment before. The amount of this motion by which the equinox travels backward, or retrogrades (as it is called), on the ecliptic, is 0° 0′ 50·10″ *per annum*, an extremely minute quantity, but which, by its continual accumulation from year to year, at last makes itself very palpable, and that in a way highly inconvenient to practical astronomers, by destroying, in the lapse of a moderate number of years, the arrangement of their catalogues of stars, and making it necessary to reconstruct them. Since the formation of the earliest catalogue on record, the place of the equinox has retrograded already about 30°. The period in which it performs a complete tour of the ecliptic, is 25,868 years.[1]

(313.) The immediate uranographical effect of the precession of the equinoxes is to produce a uniform *increase of longitude* in all the heavenly bodies, whether fixed or erratic. For the vernal equinox being the initial point of longitudes, as well as of right ascension, a retreat of this point on the ecliptic *tells* upon the longitudes of all alike, whether at rest or in motion, and produces, so far as its amount extends, the *appearance* of a motion in longitude common to all, *as if* the whole heavens had a slow rotation round the poles of the ecliptic in the long period above mentioned, similar to what they have in twenty-four hours round those of the equinoctial. This increase of longitude, the reader will of course observe and bear in mind, is, properly speaking, neither a real nor an apparent *movement of the stars.* It is a purely technical result, arising from the gradual shifting of the zero point from which longitudes are reckoned. Had a fixed star been chosen as the origin of longitudes, they would have been invariable.

(314.) To form a just idea of this curious astronomical phenomenon, however, we must abandon, for a time, the consideration of the ecliptic, as tending to produce confusion in our ideas; for this reason, that the stability of the ecliptic itself among the stars is (as already hinted, art. 306) only approximate, and that in consequence its intersection with the equinoctial is liable to a certain amount of change, arising from *its* fluctuation, which mixes itself with what is due to the principal uranographical cause of the phenomenon. This cause will become at once appa-

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[1] Incipiunt magni procedere menses. — VIRGIL, *Pollio*.

ead of regarding the equinox, we fix our attention on the pole
noctial, or the vanishing point of the earth's axis.
'he place of this point among the stars is easily determined at
by the most direct of all astronomical observations,—those
eridian or mural circle. By this instrument we are enabled to
every moment the exact distance of the polar point from any
ore stars, and therefore to lay it down, by triangulating from
, with unerring precision, on a chart or globe, without the
nce to the position of the ecliptic, or to any other circle not
onnected with it. Now, when this is done with proper dili-
exactness, it results that, although for short intervals of time,
w days, the place of the pole may be regarded as not sensibly
t in reality it is in a state of constant, although extremely
n; and, what is still more remarkable, this motion is not
it compounded of one principal uniform, or nearly uniform,
other smaller and subordinate periodical fluctuations: the
ng rise to the phenomena of *precession;* the latter to another
enomenon called *nutation.* These two phenomena, it is true,
oretically speaking, to one and the same general head, and are
onnected together, forming part of a great and complicated
nsequences flowing from the earth's rotation on its axis: but

stars, in a small circle round the pole of the ecliptic. Now this cannot happen without producing corresponding changes in the apparent diurnal motion of the sphere, and the aspect which the heavens must present at very remote periods of history. The pole is nothing more than the vanishing point of the earth's axis. As this point, then, has such a motion as we have described, it necessarily follows that the earth's *axis* must have a conical motion, in virtue of which it points successively to every part of the small circle in question. We may form the best idea of such a motion by noticing a child's peg-top, when it spins not upright, or that amusing toy the te-to-tum, which, when delicately executed, and nicely balanced, becomes an elegant philosophical instrument, and exhibits, in the most beautiful manner, the whole phenomenon. The reader will take care not to confound the variation of the *position of the earth's axis in space* with a mere shifting of the imaginary line about which it revolves, in its interior. The whole earth participates in the motion, and goes along with the axis as if it were really a bar of iron driven through it. That such is the case is proved by the two great facts: 1st, that the latitudes of places on the earth, or their geographical situation with respect to the poles, have undergone no perceptible change from the earliest ages. 2dly, that the sea maintains its level, which could not be the case if the motion of the axis were not accompanied with a motion of the whole mass of the earth.[1]

(318.) The visible effect of precession on the aspect of the heavens consists in the *apparent* approach of some stars and constellations to the pole and recess of others. The bright star of the Lesser Bear, which we call the pole star, has not always been, nor will always continue to be, our cynosure: at the time of the construction of the earliest catalogues it was 12° from the pole — it is now only 1° 24′, and will approach yet nearer, to within half a degree, after which it will again recede, and slowly give place to others, which will succeed in its companionship to the pole. After a lapse of about 12,000 years, the star α Lyræ, the brightest in the northern hemisphere, will occupy the remarkable situation of a pole star approaching within about 5° of the pole.

(319.) At the date of the erection of the Great Pyramid of Gizeh, which precedes by 3970 years (say 4000) the present epoch, the longitudes of all the stars were less by 55° 45′ than at present. Calculating

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[1] Local changes of the sea level, arising from purely geological causes, are easily distinguished from that general and systematic alteration which a shifting of the axis of rotation would give rise to.

latum¹ the place of the pole of the heavens among the stars, it
nd to fall near α Draconis; its distance from that star being 3°
This being the most conspicuous star in the immediate neigh-
vas therefore the pole star at that epoch. And the latitude of
ng just 30° north, and consequently the altitude of the north
also 30°, it follows that the star in question must have had at
ulmination, at Gizeh, an altitude of 26° 15′ 35″. Now it is a
fact, ascertained by the late researches of Col. Vyse, that of
yramids still existing at Gizeh, six (including all the largest)
arrow passages by which alone they can be entered, (all which
on the northern faces of their respective pyramids) inclined to
n downwards at angles as follows.

Pyramid of Cheops	26°	41
Pyramid of Cephren	25	55
Pyramid of Mycerinus	26	2
	27	0
	27	12
	28	0
Mean -	26	47

two pyramids at Abousseir also, which alone exist in a state of an *apparent* approach and recess of all the stars in the heavens to the pole in the same period. Since, also, the place of the equinox on the ecliptic is determined by the place of the pole in the heavens, the same cause will give rise to a small alternate advance and recess of the equinoctial points, by which, in the same period, both the longitudes and the right ascensions of the stars will be also alternately increased and diminished.

(822.) Both these motions, however, although here considered separately, subsist jointly; and since, while in virtue of the nutation, the pole is describing its little ellipse of 18″·5 in diameter, it is carried by the greater and regularly progressive motion of precession over so much of its circle round the pole of the ecliptic as corresponds to nineteen years,—that is to say, over an angle of nineteen times 50″·1 round the centre (which, in a small circle of 23° 28′ in diameter, corresponds to 6′ 20″, as seen from the centre of the sphere): the path which it will pursue in virtue of the two motions, subsisting jointly, will be neither an ellipse nor an exact circle, but a gently undulated ring like that in the figure (where, however, the undulations are much exaggerated). (See *fig.* to art. 825.)

(823.) These movements of precession and nutation are common to all the celestial bodies, both fixed and erratic; and this circumstance makes it impossible to attribute them to any other cause than a real motion of the earth's axis such as we have described. Did they only affect the stars, they might, with equal plausibility, be urged to arise from a *real* rotation of the starry heavens, as a solid shell, round an axis passing through the poles of the ecliptic in 25,868 years, and a real ecliptic gyration of *that* axis in nineteen years: but since they also affect the sun, moon, and planets, which, having motions independent of the general body of the stars, cannot without extravagance be supposed *attached to* the celestial concave,[1] this idea falls to the ground; and there only remains, then, a real motion in the earth by which they *can* be accounted for. It will be shown in a subsequent chapter that they are necessary consequences of the rotation of the earth, combined with its elliptical figure, and the unequal attraction of the sun and moon on its polar and equatorial regions.

(824.) Uranographically considered, as affecting the apparent places of the stars, they are of the utmost importance in practical astronomy. When we speak of the right ascension and declination of a celestial object, it becomes necessary to state what *epoch* we intend, and whether we mean the *mean* right ascension — cleared, that is, of the periodical fluctuation in its amount, which arises from nutation, or the *apparent* right ascension, which, being reckoned from the actual place of the vernal equinox, is affected by the periodical advance and recess of the equinoctial point produced by nutation — and so of the other elements. It is the practice of astronomers to *reduce*, as it is termed, all their observations, both of right ascension and declination, to some common and convenient epoch — such as the beginning of the year for temporary purposes, or of the decade, or the century for more permanent uses, by subtracting from them the whole effect of precession in the interval; and, moreover, to divest them of the influence of nutation by investigating and subducting the amount of change, both in right ascension and declination, due to the displacement of the pole from the centre to the circumference of the little ellipse above mentioned. This last process is technically termed correcting or *equating* the observation for nutation; by which latter word is always understood, in astronomy, the getting rid of a periodical cause of fluctuation, and presenting a result, not as it *was* observed, but as it would have been observed, had that cause of fluctuation had no existence.

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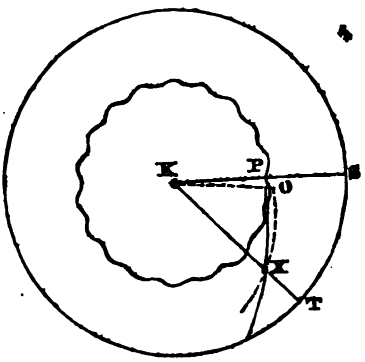
[1] This argument, cogent as it is, acquires additional and decisive force from the law of nutation, which is dependent on the position, for the time, of the lunar orbit. If we attribute it to a real motion of the celestial sphere, we must then maintain that sphere to be kept in a constant state of tremor by the motion of the moon.

(825.) For these purposes, in the present case, very convenient formulæ have been derived, and tables constructed. They are, however, of too technical a character for this work; we shall, however, point out the manner in which the investigation is conducted. It has been shown in art. 309 by what means the right ascension and declination of an object are derived from its longitude and latitude. Referring to the figure of that article, and supposing the triangle K P X orthographically projected on the plane of the ecliptic as in the annexed figure: in the triangle K P X, K P is the obliquity of the ecliptic, K X the *co-latitude* (or complement of latitude), and the angle P K X the *co-longitude* of the object X. These are the *data* of our question, of which the second is constant, and the other two are varied by the effect of precession and nutation: and their variations (considering the minuteness of the latter effect generally, and the small number of years in comparison of the whole period of 25,868, for which we ever require to estimate the effect of the former,) are of that order which may be regarded as infinitesimal in geometry, and treated as such without fear of error. The whole question, then, is reduced to this: — In a spherical triangle K P X, in which one side K X is constant, and an angle K, and adjacent side K P vary by given infinitesimal changes of the position of P: required the changes thence arising in the other side P X, and the angle K P X. This is a very simple and easy problem of spherical geometry, and being resolved,

Fig. 47.



it gives at once the reductions we are seeking; for P X being the polar distance of the object, and the angle K P X its right ascension *plus* 90°, their variations are the very quantities we seek. It only remains, then, to express in proper form the amount of the precession and nutation in *longitude* and *latitude*, when their amount in right ascension and declination will immediately be obtained.

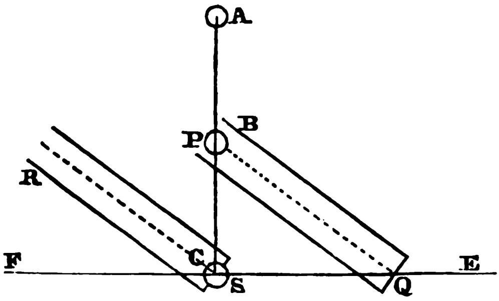
(826.) The precession in *latitude* is zero, since the latitudes of objects are not changed by it: that in longitude is a quantity proportional to the time at the rate of 50″·10 per annum. With regard to the nutation in *longitude* and *latitude*, these are no other than the abscissa and ordinate of the little ellipse in which the pole moves. The law of its motion, however, therein, cannot be understood till the reader has been made acquainted with the principal features of the moon's motion on which it depends.

(827.) Another consequence of what has been shown respecting precession and nutation is, that *sidereal time*, as astronomers use it, *i. e.* as reckoned from the transit of the equinoctial point, is *not a mean or uniformly flowing quantity*, being affected by nutation; and moreover, that so reckoned, even when cleared of the periodical fluctuation of nutation, it does not *strictly* correspond to the earth's diurnal rotation. As the sun *loses* one day in the year on the stars, by its *direct* motion in longitude; so the equinox *gains* one day in 25,868 years on them by its *retrogradation*. We ought, therefore, as carefully to distinguish between mean and apparent sidereal as between mean and apparent solar time.

(828.) Neither precession nor nutation changes the apparent places of celestial objects *inter se*. We see them, so far as these causes go, *as they* *are*, though from a station more or less unstable, as we see distant land objects correctly formed, though appearing to rise and fall when viewed from the heaving deck of a ship in the act of pitching and rolling. But there is an optical cause, independent of refraction or of perspective, which displaces them *one among the other*, and causes us to view the heavens under an aspect always to a certain slight extent false; and whose influence must be estimated and allowed for before we can obtain a precise knowledge of the place of any object. This cause is what is called the aberration of light; a singular and surprising effect arising from this, that we occupy a station not at rest but in rapid motion; and that the apparent directions of the rays of light are not the same to a spectator in motion as to one at rest. As the estimation of its effect belongs to uranography, we must explain it here, though, in so doing, we must anticipate some of the results to be detailed in subsequent chapters.

(329.) Suppose a shower of rain to fall perpendicularly in a dead calm; a person exposed to the shower, who would stand quite still and upright, would receive the drops on his hat, which would thus shelter him, but if he ran forward in any direction they would strike him in the face. The effect would be the same as if he remained still, and a wind should arise of the same velocity, and drift them against him. Suppose a ball let fall from a point A above a horizontal line E F, and that at B were placed to receive it the open mouth of an inclined hollow tube P Q; if the tube were held immoveable the ball would strike on its lower side, but if the tube were carried forward in the direction E F, with a velocity properly adjusted at every instant to that of the ball, while *preserving its inclination* to the horizon, so that when the ball in its natural descent reached C, the tube should have been carried into the position R S, it is evident that the ball would, throughout its whole descent, be found in the axis of

Fig. 48.



the tube; and a spectator referring to the tube the motion of the ball, and carried along with the former, unconscious of its motion, would fancy that the ball had been moving in the inclined direction R S of the tube's axis.

(330.) Our eyes and telescopes are such tubes. In whatever manner we consider light, whether as an advancing wave in a motionless ether, or a shower of atoms traversing space, (provided that in both cases we regard it as absolutely incapable of suffering resistance or corporeal obstruction from the particles of transparent media traversed by it,[1]) if in the interval between the rays traversing the object glass of the one or the cornea of the other (*at which moment* they acquire that convergence which directs them to a certain point *in fixed space*), and their arrival at their focus, the cross wires of the one or the retina of the other, be *slipped aside*, the point of convergence (which remains unchanged) will no longer correspond to the intersection of the wires or the central point of our visual area. The object then will *appear* displaced; and the amount of this displacement is *aberration*.

(331.) The earth is moving through space with a velocity of about 19 miles per second, in an elliptic path round the sun, and is therefore changing the direction of its motion at every instant. Light travels with a velocity of 192,000 miles per second, which, although much greater than that of the earth, is yet not *infinitely so*. Time is occupied by it in traversing any space, and in that time the earth describes a space which is to the former as 19 to 192,000, or as the tangent of 20″·5 to radius. Suppose now A P S to represent a ray of light from a star at A, and let the tube P Q be that of a telescope so inclined forward that the focus *formed* by its object-glass shall be *received* upon its cross wire, it is evident from what has been said, that the inclination of the tube must be such as to make P S : S Q :: velocity of light : velocity of the earth :: 1 : tan. 20″·5; and, therefore, the angle S P Q, or P S R, by which the axis of the telescope must deviate from the true direction of the star, must be 20″·5.

(332.) A similar reasoning will hold good when the direction of the earth's motion is not perpendicular to the visual ray. If S B be the true

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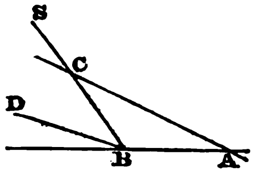
[1] This condition is indispensable. Without it we fall into all those difficulties which M. Doppler has so well pointed out in his paper on Aberration (Abhandlungen der k. boemischen Gesellschaft der Wissenschaften. Folge V. vol. iii.). If light itself, or the luminiferous ether, be corporeal, the condition insisted on amounts to a formal surrender of the dogma, either of the extension or of the impenetrability of matter; at least in the sense in which those terms have been hitherto used by metaphysicians. At the point to which science is arrived, probably few will be found disposed to maintain either the one or the other.

Fig. 48.



direction of the visual ray, and A C the position in which the telescope requires to be held in the apparent direction, we must still have the proportion B C : B A : : velocity of light: velocity of the earth : : rad.: sine of 20″·5 (for in such small angles it matters not whether we use the sines or tangents). But we have, also, by trigonometry, B C : B A : : sine of B A C: sine of A C B or C B D, which last is the apparent displacement caused by aberration. Thus it appears that the sine of the aberration, or (since the angle is extremely small) the aberration itself, is proportional to the sine of the angle made by the earth's motion in space with the visual ray, and is therefore a maximum when the line of sight is perpendicular to the direction of the earth's motion.

(333.) The uranographical effect of aberration, then, is to distort the aspect of the heavens, causing all the stars to crowd as it were directly towards that point in the heavens which is the vanishing point of all lines parallel to that in which the earth is for the time moving. As the earth moves round the sun in the plane of the ecliptic, this point must lie in that plane, 90° in advance of the earth's longitude, or 90° *behind* the sun's, and shifts of course continually, describing the circumference of the ecliptic in a year. It is easy to demonstrate that the effect on each particular star will be to make it apparently describe a small ellipse in the heavens, having for its centre the point in which the star would be seen if the earth were at rest.

(334.) Aberration then affects the apparent right ascensions and declinations of all the stars, and that by quantities easily calculable. The formulæ most convenient for that purpose, and which, systematically embracing at the same time the corrections for precession and nutation, enable the observer, with the utmost readiness, to disencumber his observations of right ascension and declination of their influence, have been constructed by Prof. Bessel, and tabulated in the appendix to the first volume of the Transactions of the Astronomical Society, where they will be found accompanied with an extensive catalogue of the places, for 1830, of the principal fixed stars, one of the most useful and best arranged works of the kind which has ever appeared.

(335.) When the body from which the visual ray emanates is itself in motion, an effect arises which is not properly speaking aberration, though it is usually treated under that head in astronomical books, and indeed confounded with it, to the production of some confusion in the mind of the student. The effect in question (which is independent of any theoretical views respecting the nature of light[1]) may be explained as follows. The ray by which we see any object is not that which it emits at the moment we look at it, but that which it *did* emit some time before, *viz.*, the time occupied by light in traversing the interval which separates it from us. The aberration of such a body then arising from the earth's velocity must be applied as a correction, not to the line joining the earth's place at the moment of observation with that occupied by the body *at the same moment*, but at that antecedent instant when the ray quitted it. Hence it is easy to derive the rule given by astronomical writers for the case of a moving object. *From the known laws of its motion and the earth's, calculate its apparent or relative angular motion in the time taken by light to traverse its distance from the earth. This is the total amount of its apparent displacement.* Its effect is to displace the body observed in a direction contrary to its apparent motion in the heavens. And it is a compound or aggregate effect consisting of two parts, one of which is the aberration, properly so called, resulting from the composition of the earth's motion with that of light, the other being what is not inaptly termed the *Equation of light*, being the allowance to be made for the *time* occupied by the light in traversing a variable space.

(336.) The complete *Reduction*, as it is called, of an astronomical observation consists in applying to the place of the observed heavenly body as read off on the instruments (supposed perfect and in perfect adjustment) five distinct and independent corrections, viz. those for refraction, parallax, aberration, precession, and nutation. Of these the correction for refraction enables us to declare what would have been the observed place, were there no atmosphere to displace it. That for parallax enables us to say from its place observed at the surface of the earth, where it would have been seen if observed from the centre. That for aberration, where it would have been observed from a motionless, instead of a moving station: while the corrections for precession and nutation refer it to fixed and determinate instead of constantly varying celestial circles. The great importance of these corrections, which pervade all astronomy, and have to be applied to every observation before it can be employed for any practical or theoretical purpose, renders this recapitulation far from superfluous.

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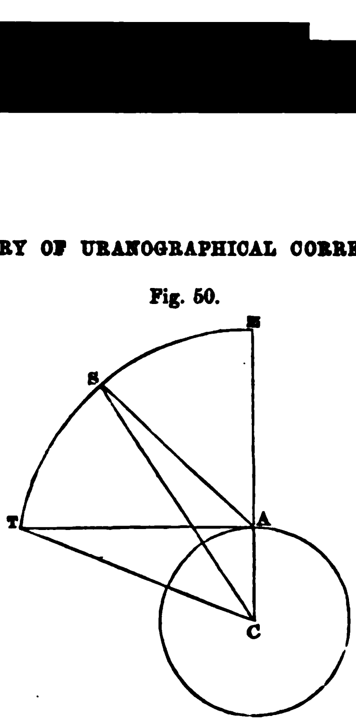
[1] The results of the undulatory and corpuscular theories of light, in the matter of aberration are, in the main, the same. We say *in the main*. There is, however, a minute difference even of numerical results. In the undulatory doctrine, the propagation of light takes place with equal velocity in all directions, whether the luminary be at rest or in motion. In the corpuscular, with an excess of velocity in the direction of the motion over that in the contrary equal to twice the velocity of the body's motion. In the cases, then, of a body moving with equal velocity directly to and directly from the earth, the aberration will be *alike* on the undulatory, but different on the corpuscular hypothesis. The utmost difference which can arise from this cause *in our system* cannot amount to above six thousandths of a second.

(337.) Refraction has been already sufficiently explained, Art. 40, and it is only, therefore, necessary here to add that in its use as an astronomical correction its amount must be applied in a contrary sense to that in which it affects the observation; a remark equally applicable to all other corrections.

(338.) The general nature of parallax or rather of parallactic motion has also been explained in Art. 80. But parallax in the uranographical sense of the word has a more technical meaning. It is understood to express that optical displacement of a body observed which is due to its being observed, not from that point which we have fixed upon as a conventional central station (from which we conceive the apparent motion would be more simple in its laws,) but from some other station remote from such conventional centre: not from the centre of the earth, for instance, but from its surface: not from the centre of the sun (which, as we shall hereafter see, is for some purposes a preferable conventional station), but from that of the earth. In the former case this optical displacement is called the *diurnal* or geocentric parallax; in the latter the *annual* or heliocentric. In either case parallax is the *correction* to be applied to the apparent place of the heavenly body, as actually seen from the station of observation, to *reduce* it to its place as it would have been seen at that instant from the conventional station.

(339.) The diurnal or geocentric parallax at any place of the earth's surface is easily calculated if we know the distance of the body, and, *vice versâ*, if we know the diurnal parallax that distance may be calculated. For supposing S the object, C the centre of the earth, A the station of observation at its surface, and C A Z the direction of a perpendicular to the surface at A, then will the object be seen from A in the direction A S, and its apparent zenith distance will be Z A S; whereas, if seen from the centre, it will appear in the direction C S, with an angular distance from the zenith of A equal to Z C S; so that Z A S — Z C S or A S C is the parallax. Now since by trignometry C S : C A :: sin C A S

Fig. 50.



= sin Z A S : sin A S C, it follows that the sine of the parallax $= \frac{\text{Radius of earth}}{\text{Distance of body}} \times \sin Z A S$.

(340.) The diurnal or geocentric parallax, therefore, at a given place, and for a given distance of the body observed, is proportional to the sine of its apparent zenith distance, and is, therefore, the greatest when the body is observed in the act of rising or setting, in which case its parallax is called its *horizontal* parallax, so that at any other zenith distance, parallax = horizontal parallax × sine of apparent zenith distance, and since A C S is always less than Z A S it appears that the application of the reduction or correction for parallax always acts in diminution of the apparent zenith distance or increase of the apparent altitude or distance from the Nadir, *i. e.* in a contrary sense to that for refraction.

(341.) In precisely the same manner as the geocentric or diurnal parallax refers itself to the zenith of the observer for its direction and quantitative rule, so the heliocentric or annual parallax refers itself for its law to the point in the heavens diametrically opposite to the place of the sun as seen from the earth. Applied as a correction, its effect takes place in a plane passing through the sun, the earth, and the observed body. Its effect is always to decrease its observed distance from that point or to increase its angular distance from the sun. And its sine is given by the relation, Distance of the observed body from the sun : distance of the earth from the sun : : sine of apparent angular distance of the body from the sun (or its apparent *elongation*) : sine of heliocentric parallax.[1]

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[1] This account of the law of heliocentric parallax is in anticipation of what follows in a subsequent chapter, and will be better understood by the student when somewhat farther advanced.

On a summary view of the whole of the uranographical correc-
divide themselves into two classes, those which *do*, and those
not, alter the apparent configurations of the heavenly bodies
The former are *real*, the latter *technical* corrections. The real
are refraction, aberration and parallax. The technical are pre-
nutation, unless, indeed, we choose to consider parallax as a
orrection introduced with a view to simplification by a better
ur point of sight.

The corrections of the first of these classes have one peculiarity
of their law, common to them all, which the student of prac-
nomy will do well to fix in his memory. *They all refer them-*
efinite apexes or points of convergence in the sphere. Thus,
n its apparent effect causes all celestial objects to draw together
e towards the zenith of the observer: geocentric parallax,
s Nadir: heliocentric, towards the place of the sun in the
berration towards that point in the celestial sphere which is
ing point of all lines parallel to the direction of the earth's
he moment, or (as will be hereafter explained) towards a point
t circle called the ecliptic, 90° behind the sun's place in that
hen applied as corrections to an observation, these directions
se to be reversed.

CHAPTER VI.

OF THE SUN'S MOTION.

APPARENT MOTION OF THE SUN NOT UNIFORM. — ITS APPARENT DIAMETER ALSO VARIABLE. — VARIATION OF ITS DISTANCE CONCLUDED. — ITS APPARENT ORBIT AN ELLIPSE ABOUT THE FOCUS. — LAW OF THE ANGULAR VELOCITY. — EQUABLE DESCRIPTION OF AREAS. — PARALLAX OF THE SUN. — ITS DISTANCE AND MAGNITUDE. — COPERNICAN EXPLANATION OF THE SUN'S APPARENT MOTION. — PARALLELISM OF THE EARTH'S AXIS. — THE SEASONS. — HEAT RECEIVED FROM THE SUN IN DIFFERENT PARTS OF THE ORBIT. — MEAN AND TRUE LONGITUDES OF THE SUN. — EQUATION OF THE CENTRE. — SIDEREAL, TROPICAL, AND ANOMALISTIC YEARS. — PHYSICAL CONSTITUTION OF THE SUN. — ITS SPOTS. — FACULÆ. — PROBABLE NATURE AND CAUSE OF THE SPOTS. — ATMOSPHERE OF THE SUN. — ITS SUPPOSED CLOUDS. — TEMPERATURE AT ITS SURFACE. — ITS EXPENDITURE OF HEAT. — TERRESTRIAL EFFECTS OF SOLAR RADIATION.

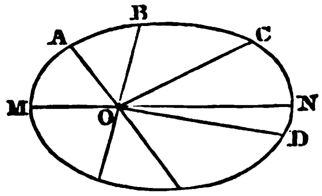
(346.) In the foregoing chapters, it has been shown that the apparent path of the sun is a great circle of the sphere, which it performs in a period of one sidereal year. From this it follows, that the line joining the earth and sun lies constantly *in one plane;* and that, therefore, whatever be the real motion from which this apparent motion arises, it must be confined to one plane, which is called the *plane of the ecliptic.*

(347.) We have already seen (art. 146) that the sun's motion in right ascension among the stars is not uniform. This is partly accounted for by the obliquity of the ecliptic, in consequence of which equal variations in longitude do not correspond to equal changes of right ascension. But if we observe the place of the sun daily throughout the year, by the transit and circle, and from these calculate the longitude for each day, it will still be found that, even in its own proper path, its apparent angular motion is far from uniform. The change of longitude in twenty-four mean solar hours *averages* 0° 59′ 8″·33; but about the 31st of December it amounts to 1° 1′ 9″·9, and about the 1st of July is only 0° 57′ 11″·5. Such are the extreme limits, and such the mean value of the sun's apparent angular velocity in its annual orbit.

(348.) This variation of its angular velocity is accompanied with a corresponding change of its distance from us. The change of distance is recognized by a variation observed to take place in its apparent diameter, when measured at different seasons of the year, with an instrument adapted for that purpose, called the *heliometer*,[1] or, by calculating from the time which its disc takes to traverse the meridian in the transit instrument. The greatest apparent diameter corresponds to the 1st of December, or to the greatest angular velocity, and measures 32′ 35″·6, the least is 31′ 31″·0; and corresponds to the 1st of July; at which epochs, as we have seen, the angular motion is also at its extreme limit either way. Now, as we cannot suppose the sun to alter its real size periodically, the observed change of its apparent size can only arise from an actual change of distance. And the sines or tangents of such small arcs being proportional to the arcs themselves, its distances from us, at the above-named epoch, must be in the inverse proportion of the apparent diameters. It appears, therefore, that the greatest, the mean, and the least distances of the sun from us are in the respective proportions of the numbers 1·01679, 1.00000, and 0·98321; and that its apparent angular velocity diminishes as the distance increases, and *vice versâ*.

(349.) It follows from this, that the real orbit of the sun, as referred to the earth supposed at rest, is not a circle with the earth in the centre. The situation of the earth within it is *excentric*, the *excentricity* amounting to 0·01679 of the mean distance, which may be regarded as our unit of measure in this inquiry. But besides this, the *form* of the orbit is not circular, but elliptic. If from any point O, taken to represent the earth, we draw a line, O A, in some fixed direction, from which we then

Fig. 51.



set off a series of angles, A O B, A O C, &c. equal to the observed longitudes of the sun throughout the year, and in these respective directions measure off from O the distances O A, O B, O C, &c. representing the distances deduced from the observed diameter, and then connect all the extremities A, B, C, &c. of these lines by a continuous curve, it is evident this will be a correct representation of the relative orbit of the sun about the earth. Now, when this is done, a deviation from the circular figure in the resulting curve becomes apparent; it is found to be evidently longer than it is broad — that is to say, elliptic, and the point O to occupy, not the *centre*, but one of the foci of the ellipse. The graphical process here described is sufficient to point out the general figure of the curve in question; but for the purposes of exact verification, it is necessary to recur to the properties of the ellipse,[1] and to express the distance of any one of its points in terms of the angular situation of that point with respect to the longer axis, or diameter of the ellipse. This, however, is readily done; and when numerically calculated, on the supposition of the excentricity, being such as above stated, a perfect coincidence is found to subsist between the distances thus computed, and those derived from the measurement of the apparent diameter.

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[1] 'Ηλιος the sun, and μετρειν to measure.

(350.) The mean distance of the earth and sun being taken for unity, the extremes are 1·01679 and 0·98321. But if we compare, in like manner, the mean or average angular velocity with the extremes, greatest and least, we shall find these to be in the proportions of 1·03386, 1·00000, and 0·96670. The variation of the sun's *angular velocity*, then, is much greater in proportion than that of its distance — fully twice as great; and if we examine its numerical expressions at different periods, comparing them with the mean value, and also with the corresponding distances, it will be found, that, by whatever fraction of its mean value the distance exceeds the mean, the angular velocity will fall short of *its* mean or average quantity by very nearly *twice* as great a fraction of the latter, and *vice versâ*. Hence we are led to conclude that the *angular velocity* is in the inverse proportion, not of the distance simply, but of its *square;* so that, to compare the daily motion in longitude of the sun, at one point, A, of its path, with that at B, we must state the proportion thus: —

$O B^2 : O A^2 ::$ daily motion at A : daily motion at B. And this is found to be exactly verified in every part of the orbit.

(351.) Hence we deduce another remarkable conclusion — viz. that if the sun be supposed really to move around the circumference of this ellipse, its actual speed cannot be uniform, but must be greatest at its least distance and less at its greatest. For, were it uniform, the apparent locity would be, of course, inversely proportional to the distance; ause the same linear change of place, being produced in the at different distances from the eye, must, by the laws of per- orrespond to apparent angular displacements inversely as those Since, then, observation indicates a more rapid law of varia- angular velocities, it is evident that mere change of distance, nied with a change of actual speed, is insufficient to account for at the increased proximity of the sun to the earth must be ed with an actual increase of its real velocity of motion along

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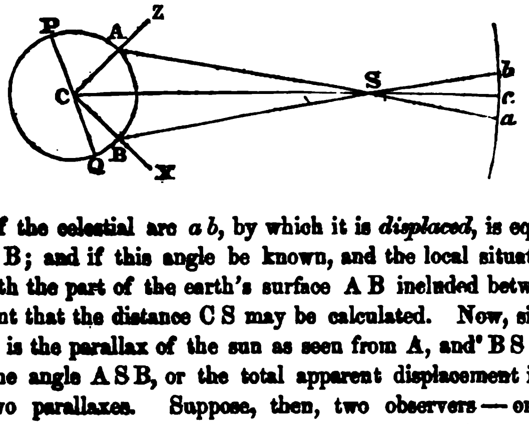
[1] See Conic Sections, by the Rev. H. P. Hamilton, or any other of the very numerous works on this subject.

This elliptic form of the sun's path, the excentric position of within it, and the unequal speed with which it is actually by the sun itself, all tend to render the calculation of its longi- theory (*i. e.* from a knowledge of the causes and nature of its fficult; and indeed impossible, so long as the *law* of its actual ntinues unknown. This *law*, however, is not immediately It does not come forward, as it were, and present itself at once, lliptic form of the orbit, by a direct comparison of angles and but requires an attentive consideration of the whole series of s registered during an entire period. It was not, therefore, uch painful and laborious calculation, that it was discovered by is 0·01679, in parts of a unit equal to the *mean distance*, or *half the longer diameter of the ellipse; i. e.* about one sixtieth part of that semidiameter; and the motion of the sun in its circumference is so regulated, that equal areas of the ellipse are passed over by the radius vector in equal times.

(355.) What we have here stated supposes no knowledge of the sun's actual distance from the earth, nor, consequently, of the actual dimensions of its orbit, nor of the body of the sun itself. To come to any conclusions on these points, we must first consider by what means we can arrive at any knowledge of the distance of an object to which we have no access. Now, it is obvious, that its *parallax* alone can afford us any information on this subject. Suppose P A B Q to represent the earth, C its centre, and S the sun, and A, B two situations of a spectator, or, which comes to the same thing, the stations of two spectators, both observing the sun S at the same instant. The spectator A will see it in the direction A S *a*, and will refer it to a point *a* in the infinitely distant sphere of the stars, while the spectator B will see it in the direction B S *b*, and refer it to *b*. The angle included between these directions, or the measure of the celestial arc *a b*, by which it is *displaced*, is equal to the angle A S B; and if this angle be known, and the local situations of A and B, with the part of the earth's surface A B included between them, it is evident that the distance C S may be calculated. Now, since A S C (art. 339) is the parallax of the sun as seen from A, and B S C as seen from B, the angle A S B, or the total apparent displacement is the sum of the two parallaxes. Suppose, then, two observers — one in the northern, the other in the southern hemisphere — at stations on the same meridian, to observe on the same day the meridian altitudes of the sun's centre. Having thence derived the apparent zenith distances, and cleared them of the effects of refraction, if the distance of the sun were equal to that of the fixed stars, the sum of the zenith distances thus found would be precisely equal to the sum of the latitudes north and south of the places of observation. For the sum in question would then be equal

Fig. 52.



to the angle Z C X, which is the meridional distance of the stations across the equator. But the effect of parallax being in both cases to increase the apparent zenith distances, their observed sum will be greater than the sum of the latitudes, by the sum of the two parallaxes, or by the angle A S B. This angle, then, is obtained by subducting the sum of the north and south latitudes from that of the zenith distances; and this once determined, the horizontal parallax is easily found, by dividing the angle so determined by the sum of the sines of the two latitudes.

(356.) If the two stations be not exactly on the same meridian (a condition very difficult to fulfil), the same process will apply, if we take care to allow for the change of the sun's actual zenith distance in the interval of time elapsing between its arrival on the meridians of the stations. This change is readily ascertained, either from tables of the sun's motion, grounded on the experience of a long course of observations, or by actual observation of its meridional altitude on several days before and after that on which the observations for parallax are taken. Of course, the nearer the stations are to each other in longitude, the less is this interval of time, and, consequently, the smaller the amount of this correction; and, therefore, the less injurious to the accuracy of the final result is any uncertainty in the daily change of zenith distance which may arise from imperfection in the solar tables, or in the observations made to determine it.

(357.) The horizontal parallax of the sun has been concluded from observations of the nature above described, performed in stations the most remote from each other in latitude, at which observatories have been instituted. It has also been deduced from other methods of a more refined nature, and susceptible of much greater exactness, to be hereafter described. Its amount so obtained, is about 8″·6. Minute as this quantity is, there can be no doubt that it is a tolerably correct approximation to the truth; and in conformity with it, we must admit the sun to be situated at a mean distance from us, of no less than 23984 times the length of the earth's radius, or about 95000000 miles.

(358.) That at so vast a distance the sun should appear to us of the size it does, and should so powerfully influence our condition by its heat and light, requires us to form a very grand conception of its actual magnitude, and of the scale on which those important processes are carried on within it, by which it is enabled to keep up its liberal and unceasing supply of these elements. As to its actual magnitude we can be at no loss, knowing its distance, and the angles under which its diameter appears to us. An object, placed at the distance of 95000000 miles, and subtending an angle of 32′ 3″, must have a real diameter of 882000 miles. Such, then, is the diameter of this stupendous globe. If we compare it with what we have already ascertained of the dimensions of our own, we shall find that in linear magnitude it exceeds the earth in the proportion 111½ to 1, and in bulk in that of 1384472 to 1.

(359.) It is hardly possible to avoid associating our conception of an object of definite globular figure, and of such enormous dimensions, with some corresponding attribute of massiveness and material solidity. That the sun is not a mere phantom, but a body having its own peculiar structure and economy, our telescopes distinctly inform us. They show us dark spots on its surface, which slowly change their places and forms, and by attending to whose situation, at different times, astronomers have ascertained that the sun revolves about an axis nearly perpendicular to the plane of the ecliptic, performing one rotation in a period of about 25 days, and in the same direction with the diurnal rotation of the earth, *i. e.* from west to east. Here, then, we have an analogy with our own globe; the slower and more majestic movement only corresponding with the greater dimensions of the machinery, and impressing us with the prevalence of similar mechanical laws, and of, at least, such a community of nature as the existence of inertia and obedience to force may argue. Now, in the exact proportion in which we invest our idea of this immense bulk with the attribute of inertia, or weight, it becomes difficult to conceive its circulation round so comparatively small a body as the earth, without, on the one hand, dragging it along, and displacing it, if bound to it by some invisible tie; or, on the other hand, if not so held to it, pursuing its course alone in space, and leaving the earth behind. If we connect two solid masses by a rod, and fling them aloft, we see them circulate about a point between them, which is their common centre of gravity; but if one of them be greatly more ponderous than the other, this common centre will be proportionally nearer to that one, and even within its surface; so that the smaller one will circulate, in fact, about the larger, which will be comparatively but little disturbed from its place.

(360.) Whether the earth move round the sun, the sun round the earth, or both round their common centre of gravity, will make no difference, so far as appearances are concerned, provided the stars be supposed sufficiently distant to undergo no sensible apparent *parallactic* displacement by the motion so attributed to the earth. Whether they are so or not must still be a matter of inquiry; and from the absence of any measurable amount of such displacement, we can conclude nothing but this, that the scale of the sidereal universe is so great, that the mutual orbit of the earth and sun may be regarded as an imperceptible point in comparison with the distance of its nearest members. Admitting, then, in conformity with the laws of dynamics, that two bodies connected with

ing about each other in free space do, in fact, revolve about their
entre of gravity, which remains immoveable by their mutual
ecomes a matter of further inquiry, *whereabouts* between them
is situated. Mechanics teach us that its place will divide their
stance in the inverse ratio of their *weights* or *masses*;[1] and
s grounded on phenomena, of which an account will be given
, inform us that this ratio, in the case of the sun and earth, is
at of 354936 to 1,—the sun being, in that proportion, more
than the earth. From this it will follow that the common point
h they both circulate is only 267 miles from the sun's centre,
$\frac{1}{300}$th part of its own diameter.

Henceforward, then, in conformity with the above statements,
he Copernican view of our system, we must learn to look upon
the comparatively motionless centre about which the earth per-
nnual elliptic orbit of the dimensions and excentricity, and with
regulated according to the law above assigned; the sun occu-
of the foci of the ellipse, and from that station quietly dissemi-
ll sides its light and heat; while the earth travelling round it,
ting itself differently to it at different times of the year and
s through the varieties of day and night, summer and winter,
enjoy.

In this annual motion of the earth, its axis preserves at all

(363.) Let, then, S represent the sun, and A, B, C, D, four positions of the earth in its orbit 90° apart, viz. A that which it has on the 21st of March, or at the time of the vernal equinox; B that of the 21st of June, or the summer solstice; C that of the 21st of September, or the autumnal equinox; and D that of the 21st of December, or the winter solstice. In each of these positions let P Q represent the axis of the earth, about which its diurnal rotation is performed without interfering with its annual motion in its orbit. Then, since the sun can only enlighten one half of the surface at once, viz. that turned towards it, the shaded portions of the globe in its several positions will represent the dark, and the bright, the enlightened halves of the earth's surface in these positions. Now, 1st, in the position A, the sun is vertically over the intersection of the equinoctial F E and the ecliptic H G. It is, therefore, in the equinox; and in this position the poles P Q, both fall on the extreme confines of the enlightened side. In this position, therefore, it is day over half the northern and half the southern hemisphere at once; and as the earth revolves on its axis, every point of its surface describes half its diurnal course in light, and half in darkness; in other words, the duration of day and night is here equal over the whole globe: hence the term *equinox*. The same holds good at the autumnal equinox on the position C.

(364.) B is the position of the earth at the time of the *northern summer* solstice. Here the north pole P, and a considerable portion of the earth's surface in its neighbourhood, as far as B, are situated *within* the enlightened half. As the earth turns on its axis in this position, therefore, the whole of that part remains constantly enlightened; therefore, at this point of its orbit, or at this season of the year, it is continual day at the north pole, and in all that region of the earth which encircles this pole as far as B — that is, to the distance of 23° 28′ from the pole, or within what is called in geography, the *arctic circle*. On the other hand, the opposite or south pole Q, with all the region comprised within the *antarctic* circle, as far as 23° 28′ from the south pole, are immersed at this season in darkness during the entire diurnal rotation, so that it is here continual night.

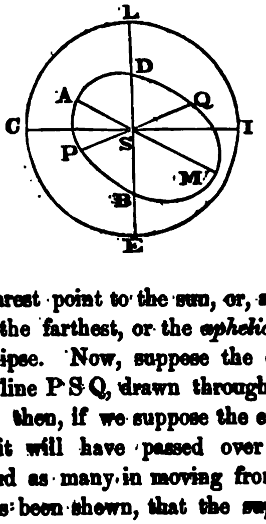
(365.) With regard to that portion of the surface comprehended between the arctic and antarctic circles, it is no less evident that the nearer any point is to the north pole, the larger will be the portion of its diurnal course comprised within the bright, and the smaller within the dark hemisphere; that is to say, the longer will be its day, and the shorter its night. Every station north of the equator will have a day of more and a night of less than twelve hours' duration, and *vice versâ*.

phenomena are exactly inverted when the earth comes to the oint D of its orbit.

Now, the temperature of any part of the earth's surface ainly on its exposure to the sun's rays. Whenever the sun is horizon of any place, that place is receiving heat; when below, th it, by the process called radiation; and the whole quantities nd parted with in the year (secondary causes apart) must ch other at every station, or the equilibrium of temperature say, the constancy which is observed to prevail in the annual f temperature as indicated by the thermometer) would not be

Whenever, then, the sun remains more than twelve hours horizon of any place, and less beneath, the general temperature ce will be above the average; when the reverse, below. As then, moves from A to B, the days growing longer, and the rter, in the northern hemisphere, the temperature of every part misphere increases, and we pass from spring to summer; while, e time, the reverse obtains in the southern hemisphere. As passes from B to C, the days and nights again approach to -the excess of temperature in the northern hemisphere above state grows less, as well as its defect in the southern; and at nal equinox C, the mean state is once more attained. From ble with what we know of the laws of the communication of heat from a luminary placed at a variable distance. Heat, like light, being equally dispersed from the sun in all directions, and being spread over the surface of a sphere continually enlarging as we recede from the centre, must, of course, diminish in intensity according to the inverse proportion of the surface of the sphere over which it is spread; that is, in the inverse proportion of the square of the distance. But we have seen (art. 350) that this is also the proportion in which the *angular velocity* of the earth about the sun varies. Hence it appears, that the *momentary supply of heat* received by the earth from the sun varies in the exact proportion of angular velocity, *i. e.* of the *momentary increase of longitude:* and from this it follows, that equal amounts of heat are received from the sun in passing over equal angles round it, in whatever part of the ellipse these angles may be situated. Let, then, S represent the sun; A Q M P the earth's orbit; A its nearest point to the sun, or, as it is called, the *perihelion* of its orbit; M the farthest, or the *aphelion;* and therefore A S M the *axis* of the ellipse. Now, suppose the orbit divided into two segments by a straight line P S Q, drawn through the sun, and anyhow situated as to direction; then, if we suppose the earth to circulate in the direction P A Q M P, it will have passed over 180° of longitude in moving from P to Q, and as many in moving from Q to P. It appears, therefore, from what has been shewn, that the supplies of heat received from the sun will be equal in the two segments, in whatever direction the line P S Q be drawn. They will, indeed, be described in unequal times; that in which the perihelion A lies in a shorter, and the other in a longer, in proportion to their unequal area: but the greater proximity of the sun in the smaller segment compensates exactly for its more rapid description, and thus an equilibrium of heat is, as it were, maintained. Were it not for this, the excentricity of the orbit would materially influence the transition of seasons. The fluctuation of distance amounts to nearly $\frac{1}{30}$th of

Fig. 54.



uantity, and, consequently, the fluctuation in the sun's direct
ver to double this, or $\frac{1}{15}$th of the whole. Now, the perihelion
is situated nearly at the place of the northern winter solstice;
e it not for the compensation we have just described, the effect
to exaggerate the difference of summer and winter in the
emisphere, and to moderate it in the northern; thus producing
lent alternation of climate in the one hemisphere, and an
perpetual spring in the other. As it is, however, no such
subsists, but an equal and impartial distribution of heat and
orded to both.

This does not prevent, however, the direct impression of the
n the height of summer,—the glow and ardour of his rays,
rfectly clear sky, at noon, in equal latitudes and under equal
ces of exposure,—from being very materially greater in the
emisphere than in the northern. One fifteenth is too considera-
ion of the whole intensity of sunshine not to aggravate in a
gree the sufferings of those who are exposed to it in thirsty
thout shelter. The accounts of these sufferings in the interior
ia, for instance, are of the most frightful kind, and would seem
ed what have ever been undergone by travellers in the northern
Africa.[1]

he surface never acquires a very elevated temperature even under the quator.

(371.) The great key to simplicity of conception in astronomy, and, ndeed, in all sciences where motion is concerned, consists in contemplaing every movement as referred to points which are either permanently ixed, or so nearly so, as that their motions shall be too small to interfere aterially with and confuse our notions. In the choice of these primary oints of reference, too, we must endeavour, as far as possible, to select uch as have simple and symmetrical geometrical relations of situation with espect to the curves described by the moving parts of the system, and rhich are thereby fitted to perform the office of natural centres — advanageous stations for the eye of reason and theory. Having learned to ttribute an orbital motion to the earth, it loses this advantage, which is ransferred to the sun, as the fixed centre about which its orbit is perormed. Precisely as, when embarrassed by the earth's diurnal motion, re have learned to transfer, in imagination, our station of observation rom its surface to its centre, by the application of the diurnal parallax; o, when we come to inquire into the movements of the planets, we shall ind ourselves continually embarrassed by the orbital motion of our point f view, unless, by the consideration of the *annual or heliocentric paralax*, we consent to refer all our observations on them to the centre of the un, or rather to the common centre of gravity of the sun, and the other odies which are connected with it in our system. Hence arises the disinction between the *geocentric* and *heliocentric* place of an object. The ormer refers its situation in space to an imaginary sphere of infinite adius, having the centre of the earth for its centre — the latter to one oncentric with the sun. Thus, when we speak of the *heliocentric longiudes* and *latitudes* of objects, we suppose the spectator situated in the sun nd referring them by circles perpendicular to the plane of the ecliptic, o the great circle marked out in the heavens by the infinite prolongation f that plane.

(372.) The point in the imaginary concave of an infinite heaven, to rhich a spectator in the sun refers the earth, must, of course, be diamerically opposite to that to which a spectator on the earth refers the sun's entre; consequently the heliocentric *latitude* of the earth is always othing, and *its heliocentric longitude always equal to the sun's geocentric ongitude* + 180°. The heliocentric equinoxes and solstices are, therefore, he same as the geocentric reversely named; and to conceive them, we ave only to imagine a plane passing through the sun's centre, parallel to he earth's equator, and prolonged infinitely on all sides. The line of inter-

his plane and the plane of the ecliptic is the line of equinoxes,
stices are 90° distant from it.

'he position of the longer axis of the earth's orbit is a point
portance. In the figure (art. 368) let E C L I be the ecliptic,
al equinox, L the autumnal (*i. e.* the points *to which the earth*
from the sun when its heliocentric longitudes are 0° *and* 180°
). Supposing the earth's motion to be performed in the direc-
I, the angle E S A, or the longitude of the perihelion, in the
was 99° 30′ 5″: we say in the year 1800, because, in point of
operation of causes hereafter to be explained, its position is
an extremely slow variation of about 12″ per annum to the
nd which in the progress of an immensely long period—of no
0984 years—carries the axis A S M of the orbit completely
whole circumference of the ecliptic. But this motion must be
for the present, as well as many other minute deviations, to be
o view when they can be better understood.

Were the earth's orbit a circle, described with a uniform
out the sun placed in its centre, nothing could be easier than
its position at any time with respect to the line of equinoxes,
itude, for we should only have to reduce to numbers the pro-
lowing; viz. One year : the time elapsed :: 360° : the arc of

the whole semi-revolution from A to M, exceed it in amount; or, in other words, the *true place* will be in advance of the *mean:* at M, one half the year will have elapsed, and one half the orbit have been described, whether it be circular or elliptic. Here, then, the *mean* and *true* places coincide; but in all the other half of the orbit, from M to A, the true place will fall short of the mean, since at M the angular motion is slowest, and the true place from this point begins to lag behind the mean — to make up with it, however, as it approaches A, where it once more overtakes it.

(375.) The quantity by which the *true* longitude of the earth differs from the *mean* longitude is called the equation of the centre, and is *additive* during all the half-year, in which the earth passes from A to M, beginning at 0° 0′ 0″, increasing to a maximum, and again diminishing to zero at M; after which it becomes subtractive, attains a maximum of subtractive magnitude between M and A, and again diminishes to 0 at A. Its maximum, both additive and subtractive, is 1° 55′ 33″·3.

(376.) By applying, then, to the earth's mean longitude, the equation of the centre corresponding to any given time at which we would ascertain its place, the true longitude becomes known; and since the sun is always seen from the earth in 180° more longitude than the earth from the sun, in this way also the sun's true place in the ecliptic becomes known. The calculation of the equation of the centre is performed by a table constructed for that purpose, to be found in all "Solar Tables." *

(377.) The maximum value of the equation of the centre depends only on the ellipticity of the orbit, and may be expressed in terms of the excentricity. *Vice versâ*, therefore, if the former quantity can be ascertained by observation, the latter may be derived from it; because, whenever the law, or numerical connection, between two quantities is known, the one can always be determined from the other. Now, by assiduous observation of the sun's transits over the meridian, we can ascertain, for every day, its exact right ascension, and thence conclude its longitude (art. 309). After this, it is easy to assign the angle by which this *observed* longitude exceeds or falls short of the mean; and the greatest amount of this excess or defect which occurs in the whole year, is the maximum equation of the centre. This, as a means of ascertaining the excentricity of the orbit, is a far more easy and accurate method than that of concluding the sun's distance by measuring its apparent diameter. The results of the two methods coincide, however, perfectly.

(378.) If the ecliptic coincided with the equinoctial, the effect of the equation of the centre, by disturbing the uniformity of the sun's apparent motion in longitude, would cause an inequality in its time of coming on

n on successive days. When the sun's centre comes to the
is *apparent noon*, and if its motion in longitude were uni-
he ecliptic coincident with the equinoctial, this would always
h *mean noon*, or the stroke of 12 on a well-regulated solar
t, independent of the want of uniformity in its motion, the
the ecliptic gives rise to another inequality in this respect; in
of which, the sun, even supposing its motion in the ecliptic
uld yet alternately, in its time of attaining the meridian, anti-
all short of the mean noon as shown by the clock. For the
ion of a celestial object forming a side of a right-angled sphe-
le, of which its longitude is the hypothenuse, it is clear that
increase of the latter must necessitate a deviation from uni-
he increase of the former.

hese two causes, then, acting conjointly, produce, in fact, a
rable fluctuation in the time as shown per clock, when the sun
s the meridian. It amounts, in fact, to upwards of half an
rent noon sometimes taking place as much as $16\frac{1}{4}$ min. before
and at others as much as $14\frac{1}{2}$ min. after. This difference
parent and mean noon is called the *equation of time*, and is
nd inserted in ephemerides for every day of the year, under
or else, which comes to the same thing, the moment, *in mean*

(381.) These *Signs* are purely technical subdivisions of the ecliptic, commencing from the actual equinox, and are not to be confounded with the *constellations* so called (and sometimes so symbolized). The constellations of the zodiac, as they now stand arranged on the ecliptic, are all a full "sign" in advance or anticipation of their symbolic cognomens thereon marked. Thus the constellation Aries actually occupies the sign Taurus ♉, the constellation Taurus, the sign Gemini ♊, and so on, the *signs* having retreated[1] among the stars (together with the equinox their origin), by the effect of precession. The bright star Spica in the constellation Virgo (α Virginis), by the observations of Hipparchus, 128 years B. C., *preceded*, or was westward of the autumnal equinox in longitude by 6°. In 1750 it followed or stood eastward of the same equinox by 20° 21′. Its place then, as referred to the ecliptic at the former epoch, would be in longitude $5^{s}\ 24^\circ\ 0'$, or in the 24th degree of the *sign* ♌, whereas in the latter epoch it stood in the 21st degree of ♍, the equinox having retreated by 26° 21′ in the interval, 1878 years, elapsed. To avoid this source of misunderstanding, the use of "signs" and their symbols in the reckoning of celestial longitudes is now almost entirely abandoned, and the ordinary reckoning (by degrees, &c. from 0 to 360) adopted in its place, and the names Aries, Virgo, &c. are becoming restricted to the constellations so called.[2]

(382.) When the sun is in either tropic, it enlightens, as we have seen, the pole on that side the equator, and shines over or beyond it to the extent of 23° 27′ 30″. The parallels of latitude, at this distance from either pole, are called the polar circles, and are distinguished from each other by the names *arctic* and *antarctic*. The regions within these circles are sometimes termed frigid zones, while the belt between the tropics is called the torrid zone, and the intermediate belts temperate zones. These last, however, are merely names given for the sake of naming; as, in fact, owing to the different distribution of land and sea in the two hemispheres, zones of *climate* are not co-terminal with zones of *latitude*.

(383.) Our seasons are determined by the apparent passages of the sun across the equinoctial, and its alternate arrival in the northern and southern hemisphere. Were the equinox invariable, this would happen at intervals precisely equal to the duration of the sidereal year; but, in fact,

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[1] *Retreated* is here used with reference to *longitude*, not to the apparent diurnal motion.

[2] When, however, the place of the sun is spoken of, the old usage prevails. Thus, if we say "the sun is in Aries," it would be interpreted to mean between 0° and 30° of longitude. So, also, "the first point of Aries" is still understood to mean the vernal, and "the first point of Libra," the autumnal equinox; and so in a few other cases.

he slow conical motion of the earth's axis described in art. 317,
x retreats on the ecliptic, and *meets* the advancing sun some-
e the whole sidereal circuit is completed. The annual retreat
inox is $50''\cdot 1$, and this arc is described by the sun in the eclip-
$19^{s}\cdot 9$. By so much *shorter*, then, is the periodical return of
s than the true sidereal revolution of the earth round the sun.
tter period, or sidereal year, is equal to $365^{d}\ 6^{h}\ 9^{m}\ 9^{s}\cdot 6$, it fol-
, that the former must be only $365^{d}\ 5^{h}\ 48^{m}\ 49^{s}\cdot 7$; and this is
eant by the *tropical* year.

We have already mentioned that the longer axis of the ellipse
by the earth has a slow motion of $11''\cdot 8$ per annum in advance.
it results, that when the earth, setting out from the perihelion,
leted one sidereal period, the perihelion will have moved
$11''\cdot 8$, which arc must be described by the earth before it can
h the perihelion. In so doing, it occupies $4^{m}\ 39^{s}\cdot 7$, and this
fore be added to the sidereal period, to give the interval between
cutive returns to the perihelion. This interval, then, is 365^{d}
$9^{s}\cdot 3$,[1] and is what is called the *anomalistic year*. All these
ve their uses in astronomy; but that in which mankind in
most interested is *the tropical year*, on which the return of
s depends, and which we thus perceive to be a compound phe-

generally comprehended under the name of perturbations and secular inequalities. Of these deviations, and their causes, we shall speak hereafter at length. It is the triumph of physical astronomy to have rendered a complete account of them all, and to have left nothing unexplained, either in the motions of the sun or in those of any other of the bodies of our system. But the nature of this explanation cannot be understood till we have developed the law of gravitation, and carried it into its more direct consequences. This will be the object of our three following chapters; in which we shall take advantage of the proximity of the moon, and its immediate connection with and dependence on the earth, to render it, as it were, a stepping-stone to the general explanation of the planetary movements. We shall conclude this by describing what is known of the physical constitution of the sun.

(386.) When viewed through powerful telescopes, provided with coloured glasses, to take off the heat, which would otherwise injure our eyes, the sun is observed to have frequently large and perfectly black spots upon it, surrounded with a kind of border, less completely dark, called a penumbra. Some of these are represented at *a*, *b*, *c*, *d*, in Plate I. fig. 2., at the end of this volume. They are, however, not permanent. When watched from day to day, or even from hour to hour, they appear to enlarge or contract, to change their forms, and at length to disappear altogether, or to break out anew in parts of the surface where none were before. In such cases of disappearance, the central dark spot always contracts into a point, and vanishes before the border. Occasionally they break up, or divide into two or more, and in those cases offer every evidence of that extreme mobility which belongs only to the fluid state, and of that excessively violent agitation which seems only compatible with the atmospheric or gaseous state of matter. The scale on which their movements take place is immense. A single second of angular measure, as seen from the earth, corresponds on the sun's disc to 461 miles; and a circle of this diameter (containing therefore nearly 167000 square miles) is the least space which can be distinctly discerned on the sun as a *visible area.* Spots have been observed, however, whose linear diameter has been upwards of 45000 miles;[1] and even, if some records are to be trusted, of very much greater extent. That such a spot should close up in six weeks' time (for they seldom last much longer), its borders must approach at the rate of more than 1000 miles a day.

(387.) Many other circumstances tend to corroborate this view of the subject. The part of the sun's disc not occupied by spots is far from bright. Its *ground* is finely mottled with an appearance of rk dots, or *pores*, which, when attentively watched, are found constant state of change. There is nothing which represents ly this appearance as the slow subsidence of some flocculent recipitates in a transparent fluid, when viewed perpendicularly : so faithfully, indeed, that it is hardly possible not to be im- th the idea of a luminous medium intermixed, but not con- ith a transparent and non-luminous atmosphere, either floating n our air, or pervading it in vast sheets and columns like flame, mers of our northern lights, directed in lines perpendicular to .

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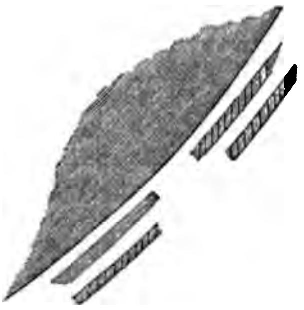
[1] Mayer, Obs. Mar. 15, 1758. "Ingens macula in sole conspiciebatur, cujus diameter = $\frac{1}{20}$ diam. solis.

Lastly, in the neighbourhood of great spots, or extensive groups rge spaces of the surface are often observed to be covered with arked curved or branching streaks, more luminous than the *faculæ*, and among these, if not already existing, spots fre- eak out. They may, perhaps, be regarded with most proba- he ridges of immense waves in the luminous regions of the sphere, indicative of violent agitation in their neighbourhood. most commonly, and best seen, towards the borders of the , and their appearance is as represented in Plate I. fig. 1.

But what *are* the spots? Many fanciful notions have been on this subject, but only one seems to have any degree of

Fig. 55.



the clouds, reflects none. (See *fig.*) The temporary removal of both the strata, but more of the upper than the lower, he supposes effected by powerful upward currents of the atmosphere, arising, perhaps, from spiracles in the body, or from local agitations.

(390.) When the spots are attentively watched, their situation on the disc of the sun is observed to change. They advance regularly towards its western limb or border, where they disappear, and are replaced by others which enter at the eastern limb, and which, pursuing their respective courses, in their turn disappear at the western. The apparent rapidity of this movement is not uniform, as it would be were the spots dark bodies passing, by an independent motion of their own, between the earth and the sun; but is swiftest in the middle of their paths across the disc, and very slow at its borders. This is precisely what would be the case supposing them to appertain to and make part of the visible surface of the sun's globe, and to be carried round by a uniform rotation of that globe on its axis, so that each spot should describe a circle parallel to the sun's equator, rendered elliptic by the effect of perspective. Their apparent paths also across the disc conform to this view of their nature, being, generally speaking, ellipses, much elongated, concentric with the sun's disc, each having one of its chords for its longer axis, and all these axes parallel to each other. At two periods of the year only do the spots appear to describe straight lines, viz. on and near to the 11th of June and the 12th of December, on which days, therefore, the plane of the circle, which a spot situated on the sun's equator describes (and consequently, the plane of that equator itself,) passes through the earth. Hence it is obvious, that the plane of the sun's equator is inclined to that of the ecliptic, and intersects it in a line which passes through the place of the earth on these days. The situation of this line, or *the line of the*

e sun's equator as it is called, is, therefore, defined by the
f the earth as seen from the sun at those epochs, which are
80° 21′ and 260° 21′ (=80° 21′+180°) being, of course,
y opposite in direction.

he inclination of the sun's axis (that of the plane of its
the ecliptic is determined by ascertaining the proportion of
nd the shorter diameter of the apparent ellipse, described by
able, well-defined spot; in order to do which, its apparent
e sun's disc must be very precisely ascertained by micrometric
epeated from day to day as long as it continues visible (usually
r 13 days, according to the magnitude of the spots, which
sh by the effect of foreshortening before they attain the actual
e disc—but the larger spots being traceable closer to the limb
aller.[1]) The *reduction* of such observations, or the conclu-
hem of the element in question, is complicated with the effect
h's motion in the interval of the observations, and with its
the ecliptic, with respect to the line of nodes. For simplicity,
pose the earth situated as it is on the 10th of March, in a line
gles to that of the nodes, *i. e.* in the heliocentric longitude
nd to remain there stationary during the whole passage of a
the disc. In this case the axis of rotation of the sun will be

through the earth, and Q a spot situated on its equator, and in that plane, and consequently in the middle of its apparent path across the disc. If the axis of rotation were perpendicular to the ecliptic, as N S, this spot would be at A, and would be seen projected on C, the centre of the sun. It is actually at Q, projected upon D, at an apparent distance C D to the *north* of the centre, which is the apparent smaller semi-axis of the ellipse described by the spot, which being known by micrometric measurement, the value of $\frac{\mathrm{C\,D}}{\mathrm{C\,N}}$ or the cosine of Q C N, the inclination of the sun's equator becomes known, C N being the apparent semi-diameter of the sun at the time. At this epoch, moreover, the northern half of the circle described by the spot is visible (the southern passing behind the body of the sun;) and the south pole *p* of the sun is within the visible hemisphere. This is the case in the whole interval from December 11th to July 12th, during which, the visual ray falls upon the southern side of the sun's equator. The contrary happens in the other half year, from July 12th to December 11th, and this is what is understood when we say that the *ascending* node (denoted ☊) of the sun's equator lies in 80° 21′ longitude—a spot on the equator passing that node being then in the act of *ascending* from the southern to the northern side of the plane of the ecliptic—such being the conventional language of astronomers in speaking of these matters.

(392.) If the observations are made at other seasons (which, however, are the less favourable for this purpose the more remote they are from the epochs here assigned); when, moreover, as in strictness is necessary, the motion of the earth in the interval of the measures is allowed for (as for a change of the point of sight); the calculations requisite to deduce the situation of the axis in space, and the duration of the revolution around it, become much more intricate, and it would be beyond the scope of this work to enter into them.[1] According to the best determinations we possess, the inclination of the sun's equator to the ecliptic is about 7° 20′ (its nodes being as above stated), and the period of rotation 25 days 7 hours 48 minutes.[2]

(393.) The region of the spots is confined, generally speaking, within about 25° on either side of the sun's equator; beyond 30° they are very

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[1] See the theory in Lalande's Astronomy, art. 3258, and the formulæ of computation in a paper by Petersen Schumacher's Nachrichten, No. 419.

[2] Bianchi (Schumacher's Nach. 483), agreeing with Laugier. Lelambre makes it $25^{d}\ 0^{h}\ 17^{m}$; Petersen, $25^{d}\ 4^{h}\ 30^{m}$. The inclination of the axis is uncertain to half a degree, and the node to several degrees. The continual changes in the spots themselves cause this uncertainty.

ous, in the manner of belts parallel to the
to physical peculiarities in certain parts of the sur
than in others to the production of the spots, o
the other, to a general influence of its rotation on i
cause of their distribution and arrangement, and
of a system of movements in the fluids which
surface bearing no remote analogy to our trade
cause arising. (See art. 239. et seq.)

(394.) The duration of individual spots is com
are formed and disappear within the limit of a si
disc—but such are for the most part small and insi
they make one or two revolutions, being recognized
by their situation with respect to the equator, their
their size, or other peculiarities, as well as by the
tween their disappearance at one limb and reappeara
a few rare cases, however, they have been watched
tions. The great spot of 1779 appeared during six
the same *group* was observed in 1840 by Sch
times.[1] It has been surmised, with considerable app
some spots, at least, are generated again and again,
time, over the same identical points of the sun's bo
example, are known to affect given localities on the
pursue definite tracks). The uncertainty which still
to the exact duration of its rotation renders it very
vincing evidence of this; nor, indeed, can it be exp
ing together into one

afforded to our globe stands in intimate connexion with those processes which are taking place on the solar surface, and to which the spots in some way or other owe their origin.

(395.) Above the luminous surface of the sun, and the region in which the spots reside, there are strong indications of the existence of a gaseous atmosphere having a somewhat imperfect transparency. When the whole disc of the sun is seen at once through a telescope magnifying moderately enough to allow it, and with a darkening glass such as to suffer it to be contemplated with perfect comfort, it is very evident that the borders of the disc are much less luminous than the centre. That this is no illusion is shown by projecting the sun's image undarkened and moderately magnified, so as to occupy a circle two or three inches in diameter, on a sheet of white paper, taking care to have it well in focus, when the same appearance will be observed. This can only arise from the circumferential rays having undergone the absorptive action of a much greater thickness of some imperfectly transparent envelope (due to greater obliquity of their passage through it) than the central. — But a still more convincing and indeed decisive evidence is offered by the phænomena attending a total eclipse of the sun. Such eclipses (as will be shown hereafter) are produced by the interposition of the dark body of the moon between the earth and sun, the moon being large enough to cover and surpass, by a very small breadth, the whole disc of the sun. Now when this takes place, were there no vaporous atmosphere capable of reflecting any light about the sun, the sky ought to appear totally dark, since (as will hereafter abundantly appear) there is not the smallest reason for believing the *moon* to have any atmosphere capable of doing so. So far, however, is this from being the case, that a bright ring or corona of light is seen, fading gradually away, as represented in Pl. I. fig. 8., which (in cases where the moon is not *centrally* superposed on the sun) is observed to be concentric with the latter, not the former body. This corona was beautifully seen in the eclipse of July 7, 1842, and with this most remarkable addition — witnessed by every spectator in Pavia, Milan, Vienna, and elsewhere: there distinct and very conspicuous *rose-coloured* protuberances (as represented in the figure cited) were seen to project beyond the dark limb of the moon, likened by some to flames, by others to mountains, but which their enormous magnitude (for to have been seen at all by the naked eye their height must have exceeded 40,000 miles), and their faint degree of illumination, clearly prove to have been cloudy masses *of the most excessive tenuity*, and which doubtless owed their support, and probably their existence, to such an atmosphere as we are now speaking of.

(396.) That the temperature at the visible surface of the sun cannot be otherwise than very elevated, much more so than any artificial heat produced in our furnaces, or by chemical or galvanic processes, we have indications of several distinct kinds: 1st, From the law of decrease of radiant heat and light, which, being inversely as the squares of the distances, it follows, that the heat received on a given area exposed at the distance of the earth, and on an equal area at the visible surface of the sun, must be in the proportion of the area of the sky occupied by the sun's apparent disc to the whole hemisphere, or as 1 to about 300000. A far less intensity of solar radiation, collected in the focus of a burning glass, suffices to dissipate gold and platina in vapour. 2dly, From the facility with which the calorific rays of the sun traverse glass, a property which is found to belong to the heat of artificial fires in the direct proportion of their intensity.[1] 3dly, From the fact, that the most vivid flames disappear, and the most intensely ignited solids appear only as black spots on the disc of the sun when held between it and the eye.[2] From the last remark it follows, that the body of the sun, however dark it may appear when seen through its spots, *may*, nevertheless, be in a state of most intense ignition. It does not, however, follow of necessity that it *must* be so. The contrary is at least physically possible. A *perfectly reflective* canopy would effectually defend it from the radiation of the luminous regions above its atmosphere, and no heat would be conducted downwards through a gaseous medium increasing rapidly in density. That the penumbral clouds *are* highly reflective, the fact of their visibility in such a situation can leave no doubt.

(397.) As the magnitude of the sun has been measured, and (as we shall hereafter see) its weight, or quantity of ponderable matter, ascertained, so also attempts have been made, and not wholly without success, from the heat actually communicated by its rays to given surfaces of material bodies exposed to their vertical action on the earth's surface, to estimate the total expenditure of heat by that luminary in a given time. The result of such experiments has been thus announced. Supposing a cylinder of ice 45 miles in diameter, to be continually darted into the sun *with the velocity of light*, and that the water produced by its fusion were continually carried off, the heat now given off constantly by radiation would then be wholly expended in its liquefaction, on the one hand, so as to leave no radiant surplus; while, on the other, the actual temperature at its surface would undergo no diminution.

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[1] By direct measurement with the *actinometer*, I find that out of 1000 calorific solar rays, 816 penetrate a sheet of plate glass 0·12 inch thick; and that of 1000 rays which have passed through one such plate, 859 are capable of passing through another. H. 1827.

[2] The ball of ignited quicklime, in Lieutenant Drummond's oxy-hydrogen lamp, gives the nearest imitation of the solar splendour which has yet been produced. The appearance of this against the sun, was, however, as described in an imperfect trial at which I was present. The experiment ought to be repeated under favourable circumstances.—*Note to the ed. of* 1833.

(398.) This immense escape of heat by radiation, we may remark, will fully explain the constant state of tumultuous agitation in which the fluids composing the visible surface are maintained, and the continual generation and filling in of *the pores*, without having recourse to internal causes. The mode of action here alluded to is perfectly represented to the eye in the disturbed subsidence of a precipitate, as described in art. 387, when the fluid from which it subsides is warm, and losing heat from its surface.

(399.) The sun's rays are the ultimate source of almost every motion which takes place on the surface of the earth. By its heat are produced all winds, and those disturbances in the electric equilibrium of the atmosphere which give rise to the phenomena of lightning, and probably also to those of terrestrial magnetism and the aurora. By their vivifying action vegetables are enabled to draw support from inorganic matter, and become, in their turn the support of animals and of man, and the sources of those great deposits of dynamical efficiency which are laid up for human use in our coal strata.[1] By them the waters of the sea are made to circulate in vapour through the air, and irrigate the land, producing springs and rivers. By them are produced all disturbances of the chemical equilibrium of the elements of nature, which, by a series of compositions and decompositions, give rise to new products, and originate a transfer of materials. Even the slow degradation of the solid constituents of the surface, in which its chief geological changes consist, is almost entirely due on the one hand to the abrasion of wind and rain, and the alternation of heat and frost; on the other to the continual beating of the sea waves, agitated by winds, the results of solar radiation. Tidal action (itself partly due to the sun's agency) exercises here a comparatively slight influence. The effect of oceanic currents (mainly originating in that influence,) though slight in abrasion, is powerful in diffusing and transporting the matter abraded; and when we consider the immense transfer of matter so produced, the increase of pressure over large spaces in the bed of the ocean, and diminution over corresponding portions of the land, we are not at a loss to perceive how the elastic power of subterranean fires, thus repressed on the one hand and relieved on the other, may break forth in points where the resistance is barely adequate to their retention, and thus bring the phenomena of even volcanic activity under the general law of solar influence.[2]

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[1] So in the edition of 1833.

[2] So in the edition of 1833.

actual combustion of
or gaseous, for the origin of the solar radiati

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[1] Electricity traversing excessively rarefied air c
doubtless, also heat. May not a continual current
circulating in the sun's immediate neighbourhood, o
and exciting, in the upper regions of its atmosphere,
however diminutive a scale, we have yet an unequivc
borealis. The possible analogy of the solar light t
distinctly insisted on by the late Sir W. Herschel, in h
be a highly curious subject of experimental inquiry, b
sheets of flame, at a distance one behind the other
brought to any required intensity,) would communica
compound ray the *penetrating* character which disting
We may also observe, that the tranquillity of the sun'
equatortal regions (if its spots be really atmospheric,) c
rotation on its axis only, but *must* arise from some cause
face of the sun, as we see the belts of Jupiter and Satu
from a cause, external to these planets, combining its
alone can produce no motions when once the form of eq

The prismatic analysis of the solar beam exhibits in th
lines," totally unlike those which belong to the light of
This may hereafter lead us to a clearer insight into its
draw any conclusions from such an indication, we mu
reaching us it has undergone the whole absorptive actio
as of the sun's. Of the latter we know nothing, and
but of the blue colour of the former we are sure; and if
absorptive) colour, the air must be expected to act on th
of other coloured media,. which often (and *especially l*
sorbed portions separated by dark intervals. It deserve
some or all the fixed lines observed by Wollaston and F
origin in our own atmosphere. Experiment
loons. on the

CHAPTER VII.

OF THE MOON. — ITS SIDEREAL PERIOD. — ITS APPARENT DIAMETER. — ITS PARALLAX, DISTANCE, AND REAL DIAMETER. — FIRST APPROXIMATION TO ITS ORBIT. — AN ELLIPSE ABOUT THE EARTH IN THE FOCUS. — ITS EXCENTRICITY AND INCLINATION. — MOTION OF ITS NODES AND APSIDES. — OF OCCULTATIONS AND SOLAR ECLIPSES GENERALLY. — LIMITS WITHIN WHICH THEY ARE POSSIBLE. — THEY PROVE THE MOON TO BE AN OPAKE SOLID. — ITS LIGHT DERIVED FROM THE SUN. — ITS PHASES. — SYNODIC REVOLUTION OR LUNAR MONTH. — OF ECLIPSES MORE PARTICULARLY. — THEIR PHENOMENA. — THEIR PERIODICAL RECURRENCE. — PHYSICAL CONSTITUTION OF THE MOON. — ITS MOUNTAINS AND OTHER SUPERFICIAL FEATURES. — INDICATIONS OF FORMER VOLCANIC ACTIVITY. — ITS ATMOSPHERE. — CLIMATE. — RADIATION OF HEAT FROM ITS SURFACE. — ROTATION ON ITS OWN AXIS. — LIBRATION. — APPEARANCE OF THE EARTH FROM IT.

(401.) THE moon, like the sun, appears to advance among the stars with a movement contrary to the general diurnal motion of the heavens, but much more rapid, so as to be very readily perceived (as we have before observed) by a few hours' cursory attention on any moonlight night. By this continual advance, which, though sometimes quicker, sometimes slower, is never intermitted or reversed, it makes the tour of the heavens in a mean or average period of $27^{d}\ 7^{h}\ 43^{m}\ 11^{s}{\cdot}5$, returning, in that time, to a position among the stars nearly coincident with that it had before, and which would be exactly so, but for reasons presently to be stated.

(402.) The moon, then, like the sun, apparently describes an orbit round the earth, and this orbit cannot be *very* different from a circle, because the apparent angular diameter of the full moon is not liable to any great extent of variation.

(403.) The distance of the moon from the earth is concluded from its horizontal parallax, which may be found either directly, by observations at remote geographical stations, exactly similar to those described in art. 355, in the case of the sun, or by means of the phænomena called occul-

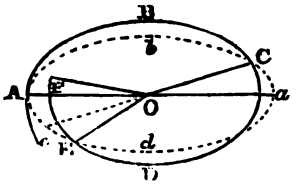
(404.) The distance of the moon's centre
station on the earth's surface, compared with its a
as measured from that station, will give its real o
the former distance is easily calculated when the
centre is known, and the apparent zenith distanc
mined by observation; for if we turn to the figure
S the moon, A the station, and C the earth's cent
the earth's radius C A, two sides of the triangle
angle C A S, which is the supplement of Z A S, t
tance, whence it is easy to find A S, the moon's d
such observations and calculations it results, that
moon is 2160 miles, or about 0·2729 of that of the
that, the bulk of the latter being considered as 1,
be 0·0204, or about $\frac{1}{49}$. The difference of the ap
moon, as seen from the earth's centre and from a
is technically called the *augmentation of the appa*
maximum occurs when the moon is in the zenith
mean angular diameter, as seen from the centre, is
= 0·545 × her horizontal parallax.

(405.) By a series of observations, such as des
continued during one or more revolutions of the
may be ascertained at every point of its orbit; and
apparent places in the heavens be observed, and re
parallax to the earth's centre, their angular interval
so that the path of the moon may then be laid do
to represent

(406.) The plane in which this orbit lies is not the ecliptic, however, but is inclined to it at an angle of 5° 8′ 48″, which is called the inclination of the lunar orbit, and intersects it in two opposite points, which are called its nodes — the *ascending node* being that in which the moon passes from the southern side of the ecliptic to the northern, and the *descending* the reverse. The points of the orbit at which the moon is nearest to, and farthest from, the earth, are called respectively its *perigee* and *apogee*, and the line joining them and the earth [illegible] the line of *apsides*.

(407.) There are, however, several remarkable circumstances which interrupt the closeness of the analogy, which cannot fail to strike the reader, between the motion of the moon around the earth, and of the earth around the sun. In the latter case, the ellipse described remains, during a great many revolutions, unaltered in its position and dimensions; or, at least, the changes which it undergoes are not perceptible but in a course of very nice observations, which have disclosed, it is true, the existence of "perturbations," but of so minute an order, that, in ordinary parlance, and for common purposes, we may leave them unconsidered. But this cannot be done in the case of the moon. Even in a single revolution, its deviation from a perfect ellipse is very sensible. It does not return to the same exact position among the stars from which it set out, thereby indicating a continual change in the *plane* of its orbit. And, in effect, if we trace by observation, from month to month, the point where it traverses the ecliptic, we shall find that the *nodes* of its orbit are in a continual state of *retreat* upon the ecliptic. Suppose O to be the earth, and A *b a d* that portion of the plane of the ecliptic which is intersected by the moon, in its alternate passages through it, from south to north, and *vice versâ;* and let A B C D E F be a portion of the moon's orbit, embracing a complete sidereal revolution. Suppose it to set out from the ascending node, A; then, if the orbit lay all in one plane, passing through O, it would have *a*, the opposite point in the ecliptic, for its descending node; after passing which, it would again ascend at A. But, in fact, its real path carries it not to *a*, but along a certain curve, A B C, to C, a

Fig. 57.



point in the ecliptic less than 180° distant from A; so that the angle A O C, or the arc of longitude described between the ascending and the descending node, is somewhat less than 180°. It then pursues its course below the ecliptic, along the curve C D E, and rises again above it, not at the point *c*, diametrically opposite to C, but at a point E, less advanced in longitude. On the whole, then, the arc described in longitude between two consecutive passages from south to north, through the plane of the ecliptic, falls short of 360° by the angle A O E; or, in other words, the ascending node appears to have retreated in one lunation on the plane of the ecliptic by that amount. To complete a sidereal revolution, then, it must still go on to describe an arc, E F, on its orbit, which will no longer, however, bring it exactly back to A, but to a point somewhat above it, or *having north latitude*.

(408.) The actual amount of this retreat of the moon's node is about 8′ 10″·64 *per diem*, on an average, and in a period of 6793·39 mean solar days, or about 18·6 years, the ascending node is carried round in a direction contrary to the moon's motion in its orbit (or from east to west) over a whole circumference of the ecliptic. Of course, in the middle of this period the position of the orbit must have been precisely reversed from what it was at the beginning. Its apparent path, then, will lie among totally different stars and constellations at different parts of this period; and this kind of spiral revolution being continually kept up, it will, at one time or other, cover with its disc every point of the heavens within that limit of latitude or distance from the ecliptic which its inclination permits; that is to say, a belt or zone of the heavens, of 10° 18′ in breadth, having the ecliptic for its middle line. Nevertheless, it still remains true that the *actual place* of the moon, in consequence of this motion, deviates in a single revolution *very little* from what it would be were the nodes at rest. Supposing the moon to set out from its node A, its latitude, when it comes to F, having completed a revolution in longitude, will not exceed 8′; which, though small in a single revolution, accumulates in its effect in a succession of many: it is to account for, and represent geometrically, this deviation, that the *motion of the nodes* is devised.

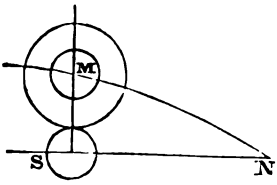
(409.) The moon's orbit, then, is not, strictly speaking, an ellipse returning into itself, by reason of the variation of the plane in which it lies, and the motion of its nodes. But even laying aside this consideration, the axis of the ellipse is itself constantly changing its direction in space, as has already been stated of the solar ellipse, but much more rapidly; making a complete revolution, in the same direction with the moon's own motion, in 3232·5753 mean solar days, or about nine years, being about 3° of angular motion in a whole revolution of the moon. This is a phenomenon known by the name of the revolution of the moon's *apsides*. Its cause will be hereafter explained. Its immediate effect is to produce a variation in the moon's distance from the earth, which is not included in the laws of exact elliptic motion. In a single revolution of the moon, this variation of distance is trifling; but in the course of many it becomes considerable, as is easily seen, if we consider that in four years and a half the position of the axis will be completely reversed, and the apogee of the moon will occur where the perigee occurred before.

(410.) The best way to form a distinct conception of the moon's motion is to regard it as describing an ellipse about the earth in the focus, and, at the same time, to regard this ellipse itself to be in a twofold state of revolution, 1st, in its own plane, by a continual advance of its axis in that plane; and 2dly, by a continual *tilting* motion of the plane itself, exactly similar to, but much more rapid than, that of the earth's equator produced by the conical motion of its axis described in art. 317.

(411.) As the moon is at a very moderate distance from us (astronomically speaking), and is in fact our nearest neighbour, while the sun and stars are in comparison immensely beyond it, it must of necessity happen, that at one time or other it must *pass over* and *occult* or *eclipse* every star and planet within the zone above described (and, as seen from the *surface* of the earth, even somewhat beyond it, by reason of parallax, which may throw it apparently nearly a degree either way from its place as seen from the centre, according to the observer's station). Nor is the sun itself exempt from being thus hidden, whenever any part of the moon's disc, in this her tortuous course, comes to *overlap* any part of the space occupied in the heavens by that luminary. On these occasions is exhibited the most striking and impressive of all the occasional phenomena of astronomy, an *eclipse of the sun*, in which a greater or less portion, or even in some rare conjunctures the whole, of its disc is obscured, and, as it were, obliterated, by the superposition of that of the moon, which appears upon it as a circularly-terminated black spot, producing a temporary diminution of daylight, or even nocturnal darkness, so that the stars appear as if at midnight. In other cases, when, at the moment that the moon is centrally superposed on the sun, it so happens that her distance from the earth is such as to render her angular diameter less than the sun's, the very singular phenomenon of an *annular solar eclipse* takes place, when the edge of the sun appears for a few minutes as a narrow ring of light, projecting on all sides beyond the dark circle occupied by the moon in its centre.

(412.) A solar eclipse can only happen when the sun and moon are *in* *conjunction*, that is to say, have the *same*, or nearly the same, position in the heavens, or the same longitude. It appears by art. 409 that this condition can only be fulfilled at the time of a *new moon*, though it by no means follows, that at *every* conjunction there *must* be an eclipse of the sun. If the lunar orbit coincided with the ecliptic, this would be the case, but as it is inclined to it at an angle of upwards of 5°, it is evident that the conjunction, or equality, of longitudes, may take place when the moon is in the part of her orbit too remote from the ecliptic to permit the discs to meet and overlap. It is easy, however, to assign the limits within which an eclipse is possible. To this end we must consider, that, by the effect of parallax, the moon's *apparent* edge may be thrown in *any* direction, according to a spectator's geographical station, by *any* amount not exceeding the horizontal parallax. Now, this comes to the same (so far as the possibility of an eclipse is concerned) as if the apparent diameter of the moon, seen from the earth's centre, were dilated by twice its horizontal parallax; for if, when so dilated, it can touch or overlap the sun, there *must* be an eclipse at *some* part or other of the earth's surface. If, then, at the moment of the nearest conjunction, the geocentric distance of the centres of the two luminaries do not exceed the sum of their semidiameters and of the moon's horizontal parallax, there will be an eclipse. This sum is, at its maximum, about 1° 34′ 27″. In the spherical triangle S N M, then, in which S is the sun's centre, M the moon's, S N the ecliptic, M N the moon's orbit, and N the node, we may

Fig. 58.



suppose the angle N S M a right angle, S M = 1° 34′ 27″, and the angle M N S = 5° 8′ 48″, the inclination of the orbit. Hence we calculate S N, which comes out 16° 58′. If, then, at the moment of the new moon, the moon's node is farther from the sun in longitude than this limit, there can be no eclipse; if within, there may, and probably will, at some part or other of the earth. To ascertain precisely whether there will or not, and, if there be, how great will be the part eclipsed, the solar and lunar tables must be consulted, the place of the node and the semidiameters exactly ascertained, and the local parallax, and apparent augmentation of the moon's diameter due to the difference of her distance from the observer and from the centre of the earth (which may amount to a sixtieth part of her horizontal diameter), determined; after which it is easy, from the above considerations, to calculate the amount overlapped of the two discs, and their moment of contact.

(413.) The calculation of the occultation of a star depends on similar considerations. An occultation is *possible*, when the moon's course, as seen from the earth's centre, carries her within a distance from the star equal to the sum of her semidiameter and horizontal parallax, and it *will happen at any particular spot*, when her apparent path, as seen from that spot, carries her centre within a distance equal to the sum of her *augmented* semidiameter and *actual* parallax. The details of these calculations, which are somewhat troublesome, must be sought elsewhere.[1]

(414.) The phenomenon of a solar eclipse and of an occultation are highly interesting and instructive in a physical point of view. They teach us that the moon is an opaque body, terminated by a real and sharply defined surface intercepting light like a solid. They prove to us, also, that at those times when we cannot *see* the moon, she really exists, and pursues her course, and that when we see her only as a crescent, however narrow, the whole globular body *is there*, filling up the deficient outline, though unseen. For occultations take place indifferently at the dark and bright, the visible and invisible outline, whichever happens to be towards the direction in which the moon is moving; with this only difference, that a star *occulted* by the bright limb, if the phenomenon be watched with a telescope, gives notice, by its gradual approach to the visible edge, when to expect its disappearance, while, if occulted at the dark limb, if the moon, at least, be more than a few days old, it is, as it were, extinguished in mid-air, without notice or visible cause for its disappearance, which, as it happens *instantaneously*, and without the slightest previous diminution of its light, is always surprising; and, if the star be a large and bright one, even startling from its suddenness. The re-appearance of the star, too, when the moon has passed over it, takes place in those cases when the bright side of the moon is foremost, not at the concave outline of the crescent, but at the invisible outline of the complete circle, and is scarcely less surprising, from its suddenness, than its disappearance in the other case.

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[1] Woodhouse's Astronomy, vol. i. See also Trans. Ast. Soc. vol. i. p. 325.

[2] There is an optical illusion of a very strange and unaccountable nature which has often been remarked in occultations. The star appears to advance actually *upon* and *within* the edge of the disc before it disappears, and that sometimes to a considerable depth. I have never myself witnessed this singular effect, but it rests on most unequivocal testimony. I have called it an optical illusion; but it is *barely possible* that a star may shine on such occasions through deep fissures in the substance of the moon.

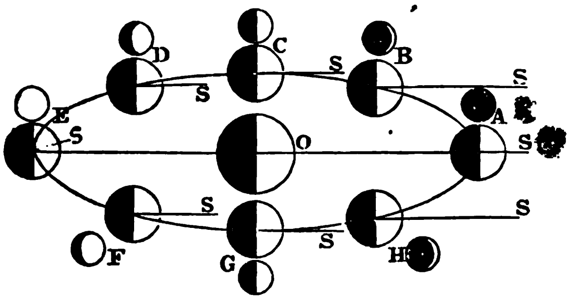
(415.) The existence of the complete circle of the disc, even when the moon is not full, does not, however, rest only on the evidence of occultations and eclipses. It may be *seen*, when the moon is crescent or waning, a few days before and after the *new moon*, with the naked eye, as a pale round body, to which the crescent seems attached, and somewhat projecting beyond its outline (which is an optical illusion arising from the greater intensity of its light.) The cause of this appearance will presently be explained. Meanwhile the fact is sufficient to show that the moon is not *inherently* luminous like the sun, but that her light is of an adventitious nature. And its crescent form, increasing regularly from a narrow semicircular line to a complete circular disc, corresponds to the appearance a globe would present, one hemisphere of which was black, the other white, when differently turned towards the eye, so as to present a greater or less portion of each. The obvious conclusion from this is, that the moon is such a globe, one half of which is brightened by the rays of some luminary sufficiently distant to enlighten the complete hemisphere, and sufficiently intense to give it the degree of splendour we see. Now, the sun alone is competent to such an effect. Its distance and light suffice; and, moreover, it is invariably observed that, when a crescent, the bright edge is *towards the sun*, and that in proportion as the moon in her monthly course becomes more and more distant from the sun, the breadth of the crescent increases, and *vice versâ*.

(416.) The sun's distance being 23984 radii of the earth, and the moon's only 60, the former is nearly 400 times the latter. Lines, therefore, drawn from the sun to every part of the moon's orbit may be regarded as very nearly parallel.[1] Suppose, now, O to be the earth, A B C D, &c. various positions of the moon in its orbit, and S the sun, at the vast distance above stated; as is shown, then, in the figure, the hemisphere of the lunar globe turned towards it (on the right) will be bright, the opposite dark, wherever it may stand in its orbit. Now, in the position A, when in conjunction with the sun, the dark part is entirely turned towards O, and the bright from it. In this case, then,

The occultations of close double stars ought to be narrowly watched, to see whether *both* individuals are thus *projected*, as well as for other purposes connected with their theory. I will only hint at one, viz. that a double star, *too close* to be seen divided with any telescope, may yet be detected to be double by the mode of its disapppearance. Should a considerable star, for instance, instead of undergoing instantaneous and complete extinction, go out by two distinct steps, following close upon each other; first losing a portion, then the whole remainder of its light, we may be sure it is a double star, though we cannot see the individuals separately.—*Note to the edit. of* 1833.

[1] The angle subtended by the moon's orbit, as seen from the sun, (in the mean state of things,) is only 17′ 12″.

Fig. 59.



the moon is not seen, it is *new* moon. When the moon has come to C, half the bright and half the dark hemisphere are presented to O, and the same in the opposite situation G: these are the first and third quarters of the moon. Lastly, when at E, the whole bright face is towards the earth, the whole dark side from it, and it is then seen wholly bright or *full moon*. In the intermediate positions B D F H, the portions of the bright face presented to O will be at first less than half the visible surface, then greater, and finally less again, till it vanishes altogether, as it comes round again to A.

(417.) These monthly changes of appearance, or *phases*, as they are called, arise, then, from the moon, an opaque body, being illuminated on one side by the sun, and reflecting from it, in all directions, a portion of the light so received. Nor let it be thought surprising that a solid substance thus illuminated should appear to *shine* and again illuminate the earth. It is no more than a white cloud does standing off upon the clear blue sky. By day, the moon can hardly be distinguished in brightness from such a cloud; and, in the dusk of the evening, clouds catching the last rays of the sun appear with a dazzling splendour, not inferior to the seeming brightness of the moon at night.[1] That the earth sends also such a light to the moon, only probably more powerful by reason of its greater apparent size[2], is agreeable to optical principles, and explains the

[1] The actual illumination of the lunar surface is not much superior to that of weathered sandstone rock in full sunshine. I have frequently compared the moon setting behind the grey perpendicular façade of the Table Mountain, illuminated by the sun just risen in the opposite quarter of the horizon, when it has been scarcely distinguishable in brightness from the rock in contact with it. The sun and moon being nearly at equal altitudes and the atmosphere perfectly free from cloud or vapour, its effect is alike on both luminaries. (H. 1848).

[2] The apparent diameter of the moon is 32′ from the earth; that of the earth seen from the moon is twice her horizontal parallax, or 1° 54′. The apparent surfaces, therefore, are as $(114)^2 : (32)^2$, or as 13 : 1 nearly.

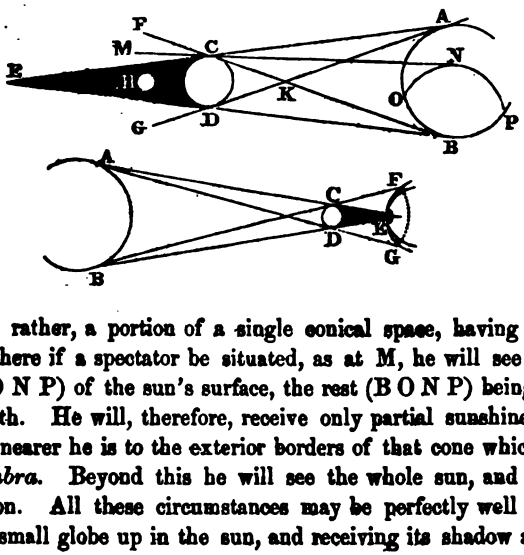
appearance of the dark portion of the young or waning moon completing its crescent (art. 413). For, when the moon is nearly new to the earth, the latter (so to speak) is nearly full to the former; it then illuminates its dark half by strong *earth-light;* and it is a portion of this, reflected back again, which makes it visible to us in the twilight sky. As the moon gains age, the earth offers it a less portion of its bright side, and the phenomenon in question dies away.

(418.) The lunar month is determined by the recurrence of its phases: it reckons from new moon to new moon; that is, from leaving its conjunction with the sun to its return to conjunction. If the sun stood still, like a fixed star, the interval between two conjunctions would be the same as the period of the moon's sidereal revolution (art. 401); but, as the sun apparently advances in the heavens in the same direction with the moon, only slower, the latter has more than a complete sidereal period to perform to come up with the sun again, and will require for it a longer time, which is the lunar month, or, as it is generally termed in astronomy, a *synodical* period. The difference is easily calculated by considering that the superfluous arc (whatever it be) is described by the sun with the velocity of $0^\circ{\cdot}98565$ *per diem, in the same time* that the moon describes that arc *plus* a complete revolution, with her velocity of $13^\circ{\cdot}17640$ *per diem;* and, the times of description being identical, the spaces are to each other in the proportion of the velocities. Let V and v be the mean angular velocities, x the superfluous arc; then $V : v :: 1 + x : x$; and $V - v : v :: 1 : x$, whence x is found, and $\frac{x}{v} =$ the time of describing x, or the difference of the sidereal and synodical periods. From these data a slight knowledge of arithmetic will suffice to derive the arc in question, and the time of its description by the moon; which being the excess of the synodic over the sidereal period, the former will be had, and will appear to be $29^d\ 12^h\ 44^m\ 2^s{\cdot}87$.

(419.) Supposing the position of the nodes of the moon's orbit to permit it, when the moon stands at A (or at the new moon), it will intercept a part or the whole of the sun's rays, and cause a solar eclipse. On the other hand, when at E (or at the full moon), the earth O will intercept the rays of the sun, and *cast a shadow* on the moon, thereby causing a lunar eclipse. And this is perfectly consonant to fact, such eclipses never happening but at the exact time of the full moon. But, what is still more remarkable, as confirmatory of the position of the earth's sphericity, this shadow, which we plainly see to enter upon and, as it were, eat away the disc of the moon, is always terminated by a circular outline, though, from the greater *size* of the circle, it is only partially seen at any one time. Now, a body which always casts a circular shadow must itself be spherical.

(420.) Eclipses of the sun are best understood by regarding the sun and moon as two independent luminaries, each moving according to known laws, and viewed from the earth: but it is also instructive to consider eclipses generally as arising from the shadow of one body thrown on another by a luminary *much larger than either.* Suppose then, A B to represent the sun, and C D a spherical body, whether earth or moon, illuminated by it. If we join and prolong A C, B D; since A B is greater than C D, these lines will meet in a point E, more or less distant from the body C D, according to its size, and within the space C E D (which represents a cone, since C D and A B are spheres), there will be a total shadow. This shadow is called the *umbra*, and a spectator situated within it can see no part of the sun's disc. Beyond the umbra are two diverging spaces (or rather, a portion of a single conical space, having K for its vertex), where if a spectator be situated, as at M, he will see a portion only (A O N P) of the sun's surface, the rest (B O N P) being obscured by the earth. He will, therefore, receive only partial sunshine; and the more, the nearer he is to the exterior borders of that cone which is called the *penumbra*. Beyond this he will see the whole sun, and be in full illumination. All these circumstances may be perfectly well shown by holding a small globe up in the sun, and receiving its shadow at different distances on a sheet of paper.

Fig. 60.



(421.) In a lunar eclipse (represented in the upper figure), the moon is seen to enter[1] the *penumbra* first, and, by degrees, get involved in the *umbra*, the former bordering the latter like a smoky haze. At this period of the eclipse, and while yet a considerable part of the moon remains unobscured, the portion involved in the umbra is invisible to the naked eye, though still perceptible in a telescope, and of a dark grey hue. But as the eclipse advances, and the enlightened part diminishes in extent, and grows gradually more and more obscured by the advance of the *penumbra*, the eye, relieved from its glare, becomes more sensible to feeble impressions of light and colour; and phenomena of a remarkable and instructive character begin to be developed. The *umbra* is seen to be very far from totally dark: and in its faint illumination it exhibits a gradation of colour, being bluish, or even (by contrast) somewhat greenish, towards the borders for a space of about 4′ or 5′ of apparent angular breadth inwards, thence passing, by delicate but rapid gradation, through rose red to a fiery or copper-coloured glow, like that of dull red-hot iron. As the eclipse proceeds this glow spreads over the whole surface of the moon, which then becomes on some occasions so strongly illuminated, as to cast a very sensible shadow, and allow the spots on its surface to be perfectly well distinguished through a telescope.

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[1] The actual contact with the penumbra is never seen; the defalcation of light comes on so very gradually that it is not till when already deeply immersed, that it is perceived to be sensibly darkened.

(422.) The cause of these singular, and sometimes very beautiful appearances, is the refraction of the sun's light in passing through our atmosphere, which at the same time becomes coloured with the hues of sunset by the absorption of more or less of the violet and blue rays, as it passes through strata nearer or more remote from the earth's surface, and therefore, more or less loaded with vapour. To show this, let A D *a* be a section of the cone of the *umbra*, and F B *h f* of the *penumbra*, through their common axis D E S, passing through the centres E S of the earth and sun, and let K M *k* be a section of these cones at a distance E M from E, equal to the radius of the moon's orbit, or 60 radii of the earth.[1] Taking this radius for unity, since E S, the distance of the sun, is 23984, and the semidiameter of the sun 111½ such units, we readily calculate D E=218, D M=158, for the distances at which the apex of the *geometrical umbra* lies behind the earth and the moon respectively. We also find for the measure of the angle E D B, 15′ 46″, and therefore D B E= 89° 44′ 14″, whence also we get M C (the linear semidiameter of the *umbra*)=0·725 (or in miles 2864), and the angle C E M, its apparent angular semidiameter as seen from E=41′ 30″. And instituting similar calculations for the geometrical *penumbra* we get M K=1·005 (3970 miles), and K E M 57′ 36″; and it may be well to remember that the doubles of these angles, or the mean angular diameters of the *umbra* and *penumbra*, are described by the moon with its mean velocity in $2^h\ 43^m$, and $3^h\ 47^m$ respectively, which are therefore the respective durations of

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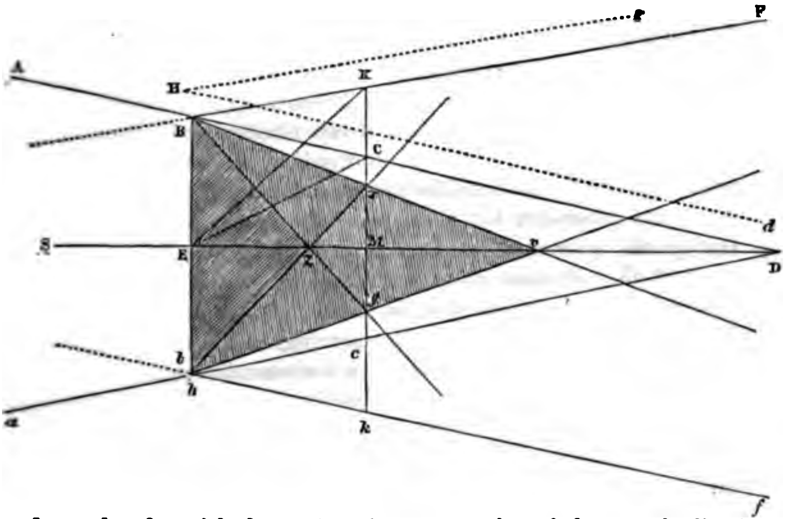
[1] The figure is unavoidably drawn out of all proportion, and the angles violently exaggerated. The reader should endeavour to draw the figure in its true proportions.

Fig. 61.



the total and partial obscuration of any one point of the moon's disc in traversing centrally the geometrical shadow.

(423.) Were the earth devoid of atmosphere, the whole of the phenomena of a lunar eclipse would consist in these partial or total obscurations. Within the space C *c* the whole of the light, and within K C and *c k* a greater or less portion of it, would be intercepted by the solid body B *b* of the earth. The refracting atmosphere, however, extends from B, *b*, to a certain unknown, but very small distance B H, *b h*, which, acting as a convex lens, of gradually (and very rapidly) decreasing density, disperses all that comparatively small portion of light which falls upon it over a space bounded externally by H *g*, parallel and very nearly coincident with B F, and internally by a line B *z*, the former representing the extreme exterior ray from the limb *a* of the sun, the latter, the extreme interior ray from the limb A. To avoid complication, however, we will trace only the courses of rays which just graze the surface at B, viz: B *z* from the upper border, A, and B *v* from the lower, *a*, of the sun. Each of these rays is bent inwards from its original course by *double* the amount of the horizontal refraction (33′) *i. e.* by 1° 6′, because, in passing from B out of the atmosphere, it undergoes a deviation equal to that at entering, and in the same direction. Instead, therefore, of pursuing the courses B D, B F, these rays respectively will occupy the positions B *z y*, B *v*, making, with the aforesaid lines, the angles D B *b*, F B *v*, each 1° 6′. Now we have found D B E = 89° 44′ 14″ and therefore

$FBE(=DBE+$ angular diam. of $\odot)=90^\circ\ 17'\ 17''$, consequently the angles EBy and EBv will be respectively $88^\circ\ 38'\ 14''$ and $89^\circ\ 11'\ 17''$ from which we conclude $Ez=42{\cdot}03$ and $Ev=88{\cdot}89$, the former falling short of the moon's orbit by 17·07, and the latter surpassing it by 28·89 radii of the earth.

(424.) The penumbra, therefore, of rays refracted at B, will be spread over the space v B y, that at H over g H d, and at the intermediate points, over similar intermediate spaces, and through this compound of superposed penumbræ the moon passes during the whole of its path through the geometrical shadow, never attaining the *absolute* umbra B z b at all. Without going into detail as to the intensity of the refracted rays, it is evident that the totality of light so thrown into the shadow is to that which the earth intercepts, as the area of a circular section of the atmosphere to that of a diametrical section of the earth itself, and, therefore, at all events but feeble. And it is still further enfeebled by actual clouds suspended in that portion of the air which forms the visible border of the earth's disc as seen from the moon, as well as by the general want of transparency caused by invisible vapour, which is especially effective in the lowermost strata, within three or four miles of the surface, and which will impart to all the rays they transmit, the ruddy hue of sunset, only *of double* the depth of tint which we admire in our glowing sunsets, by reason of the rays having to traverse twice as great a thickness of atmosphere. This redness will be most intense at the points x, y, of the moon's path through the umbra, and will thence degrade very rapidly outwardly, over the spaces x c, y C, less so inwardly, over x y. And at C, c, its hue will be mingled with the bluish or greenish light which the atmosphere scatters by irregular dispersion, or in other words by our *twilight* (art. 44). Nor will the phenomenon be uniformly conspicuous at all times. Supposing a generally and deeply clouded state of the atmosphere around the edge of the earth's disc visible from the moon (*i. e.* around that great circle of the earth, in which, at the moment the sun is in the horizon,) little or no refracted light may reach the moon.[1] Supposing that circle partly clouded and partly clear, patches of red light corresponding to the clear portions will be thrown into the umbra, and may give rise to various and changeable distributions of light on the eclipsed disc;[2] while, if entirely clear, the eclipse will be remarkable for the conspicuousness of the moon during the whole or a part of its immersion in the umbra.[3]

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[1] As in the eclipses of June 5, 1620, April 25, 1642. Lalande, Ast. 1769.

[2] As in the eclipse of Oct. 13, 1837, observed by the author.

[3] As in that of March 19, 1848, when the moon is described as giving "good light" during more than an hour after its total immersion, and some persons even doubted its being eclipsed. (Notices of R. Ast. Soc. viii. p. 132.)

(425.) Owing to the great size of the earth, the cone of its *umbra* always projects far beyond the moon; so that, if, at the time of a lunar eclipse, the moon's path be properly directed, it is sure to pass through the *umbra*. This is not, however, the case in solar eclipses. It so happens, from the adjustment of the size and distance of the moon, that the extremity of her *umbra* always falls *near* the earth, but sometimes attains and sometimes falls short of its surface. In the former case (represented in the lower figure art. 420) a black spot, surrounded by a fainter shadow, is formed, beyond which there is no eclipse on any part of the earth, but within which there may be either a total or partial one, as the spectator is within the *umbra* or *penumbra*. When the apex of the umbra falls *on* the surface, the moon at that point will appear, for an instant, to *just* cover the sun; but, when it falls short, there will be no total eclipse on any part of the earth; but a spectator, situated in or near the prolongation of the axis of the cone, will see the whole of the moon on the sun, although not large enough to cover it, *i. e.* he will witness an annular eclipse.

(426.) Owing to a remarkable enough adjustment of the periods in which the moon's *synodical* revolution, and that of her nodes, are performed, eclipses return after a certain period, very nearly in the same order and of the same magnitude. For 223 of the moon's mean synodical revolutions, or *lunations*, as they are called, will be found to occupy 6585·32 days, and nineteen complete synodical revolutions of the node to occupy 6585·78. The difference in the mean position of the node, then, at the beginning and end of 223 lunations, is nearly insensible; so that a recurrence of all eclipses within that interval must take place. Accordingly, this period of 223 lunations, or eighteen years and ten days, is a very important one in the calculation of eclipses. It is supposed to have been known to the Chaldeans, the earliest astronomers, the regular return of eclipses having been known as a physical fact for ages before their exact theory was understood. In this period there occur ordinarily 70 eclipses, 29 of the moon and 41 of the sun, visible in some part of the earth. Seven eclipses of either sun or moon at most, and two at least (both of the sun,) may occur in a year.

(427.) The commencement, duration, and magnitude of a lunar elipse are much more easily calculated than those of a solar, being independent of the position of the spectator on the earth's surface, and the same as if viewed from its centre. The common centre of the *umbra* and *penumbra* lies always in the ecliptic, at a point opposite to the sun, and the path described by the moon in passing through it is its true orbit as it stands at the moment of the full moon. In this orbit, its position, at every

nown from the lunar tables and ephemeris; and all we have,
to ascertain, is, the *moment when* the distance between the
tre and the centre of the shadow is exactly equal to the sum
idiameters of the moon and *penumbra*, or of the moon and
know when it enters upon and leaves them respectively. No
se can take place, if, at the moment of the full moon, the sun
ater angular distance from the node of the moon's orbit than
eaning by an *eclipse* the inmersion of any part of the moon in
, as its contact with the *penumbra* cannot be observed (see note
).

he dimensions of the shadow, at the place where it crosses the
h, require us to know the distances of the sun and moon at the
se are variable; but are calculated and set down, as well as
iameters, for every day, in the ephemeris, so that none of the
nting. The sun's distance is easily calculated from its elliptic
the moon's is a matter of more difficulty, by reason of the pro-
tion of the axis of the lunar orbit. (Art. 409.)

he physical constitution of the moon is better known to us
of any other heavenly body. By the aid of telescopes, we
qualities in its surface which can be no other than mountains
,—for this plain reason, we see the shadows cast by the former

Andes. The existence of such mountains is further corroborated by their appearance, as small points or islands of light beyond the extreme edge of the enlightened part, which are their tops catching the sun-beams before the intermediate plain, and which, as the light advances, at length connect themselves with it, and appear as prominences from the general edge.

(430.) The generality of the lunar mountains present a striking uniformity and singularity of aspect. They are wonderfully numerous, especially towards the Southern portion of the disc, occupying by far the larger portion of the surface, and almost universally of an exactly circular or cup-shaped form, foreshortened, however, into ellipses towards the limb; but the larger have for the most part flat bottoms within, from which rises centrally a small, steep, conical hill. They offer, in short, in its highest perfection, the true *volcanic* character, as it may be seen in the crater of Vesuvius, and in a map of the volcanic districts of the Campi Phlegræi[1] or the Puy de Dôme, but with this remarkable peculiarity, viz.: that the bottoms of many of the craters are very deeply depressed below the general surface of the moon, the internal depth being often twice or three times the external height. In some of the principal ones, decisive marks of volcanic stratification, arising from successive deposits of ejected matter, and evident indications of lava currents streaming outwards in all directions, may be clearly traced with powerful telescopes. (See Pl. V. fig. 2.[2]) In Lord Rosse's magnificent reflector, the flat bottom of the crater called Albategnius is seen to be strewed with blocks not visible in inferior telescopes, while the exterior of another (Aristillus) is all hatched over with deep gullies radiating towards its centre. What is, moreover, extremely singular in the geology of the moon is, that, although nothing having the character of seas can be traced, (for the dusky spots, which are commonly called seas, when closely examined, present appearances incompatible with the supposition of deep water,) yet there are large regions perfectly level, and apparently of a decided alluvial character.

(431.) The moon has no clouds, nor any other decisive indications of an atmosphere. Were there any, it could not fail to be perceived in the occultations of stars and the phænomena of solar eclipses, as well as in a great variety of other phænomena. The moon's diameter, for example, as measured micrometrically, and as estimated by the interval between the disappearance and reappearance of a star in an occultation, ought to differ by twice the horizontal refraction at the moon's surface. No appre-

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[1] See Breislak's map of the environs of Naples, and Desmarest's of Auvergne.

[2] From a drawing taken with a reflector of twenty feet focal length (H.)

erence being perceived, we are entitled to conclude the non-
f any atmosphere dense enough to cause a refraction of 1″ *i. e.*
1980th part of the density of the earth's atmosphere. In a
se, the existence of any sensible refracting atmosphere in the
ld enable us to trace the limb of the latter beyond the cusps,
to the sun's disc, by *a narrow, but brilliant* line of light,
to some distance along its edge. No such phænomenon is
ry faint stars ought to be extinguished before occultation, were
iable amount of vapour suspended near the surface of the moon.
s not the case; when occulted at the bright edge, indeed, the
e moon extinguishes small stars, and even at the dark limb,
in the sky caused by the near presence of the moon, renders
tion of *very* minute stars unobservable. But during the con-
f a total lunar eclipse, stars of the tenth and eleventh magni-
en to come up to the limb, and undergo *sudden* extinction as
ose of greater brightness.[1] Hence, the climate of the moon
ery extraordinary; the alternation being that of unmitigated
ng sunshine fiercer than an equatorial noon, continued for a
night, and the keenest severity of frost, far exceeding that of
winters, for an equal time. Such a disposition of things must
constant transfer of whatever moisture may exist on its surface,

regions of our atmosphere, never reaching the surface of the earth at all. Some probability is given to this by the *tendency to disappearance of clouds under the full moon*, a meteorological *fact*, (for as such we think it fully entitled to rank[1]) for which it is necessary to seek a cause, and for which no other rational explanation seems to offer. As for any other influence of the moon on the weather, we have no decisive evidence in its favour.

(483.) A circle of one second in diameter, as seen from the earth, on the surface of the moon, contains about a square mile. Telescopes, therefore, must yet be greatly improved, before we could expect to see signs of inhabitants, as manifested by edifices or by changes on the surface of the soil. It should, however, be observed, that, owing to the small density of the materials of the moon, and the comparatively feeble gravitation of bodies on her surface, muscular force would there go six times as far in overcoming the weight of materials as on the earth. Owing to the want of air, however, it seems impossible that any form of life, analogous to those on earth, can subsist there. No appearance indicating vegetation, or the slightest variation of surface, which can, in our opinion, fairly be ascribed to change of season, can any where be discerned.

(484.) The lunar summer and winter arise, in fact, from the rotation of the moon on its own axis, the period of which rotation is *exactly* equal to its sidereal revolution about the earth, and is performed in a plane 1° 30′ 11″ inclined to the ecliptic, whose *ascending* node is always precisely coincident with the *descending* node of the lunar orbit. So that the axis of rotation describes a conical surface about the pole of the ecliptic in one revolution of the node. The remarkable coincidence of the two rotations, that about the axis and that about the earth, which at first sight would seem perfectly distinct, has been asserted (but we think somewhat too hastily[2]) to be a consequence of the general laws to be explained hereafter. Be that as it may, it is the cause why we always see the same face of the moon, and have no knowledge of the other side.

(485.) The moon's rotation on her axis is uniform; but since her motion in her orbit (like that of the sun) is not so, we are enabled to look a few degrees round the equatorial parts of her visible border, on the eastern or western side, according to circumstances; or, in other words, the line joining the centres of the earth and moon fluctuates a little in its position, from its mean or average intersection with her surface, to the

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[1] From my own observation, made quite independently of any knowledge of such a tendency having been observed by others. Humboldt, however, in his personal narrative, speaks of it as well known to the pilots and seamen of Spanish America: see note at the end of the chapter (A.)

[2] See Edinburgh Review, No. 175, p. 192.

ward. And, moreover, since the axis about which she revolves
xactly perpendicular to her orbit, nor holds an invariable direc-
e, her poles come alternately into view for a small space at the
disc. These phenomena are known by the name of *librations*
nce of these two distinct kinds of libration, the same identi-
f the moon's surface is not always the centre of her disc, and
e get sight of a zone of a few degrees in breadth on all sides
er, beyond an exact hemisphere.

f there be inhabitants in the moon, the earth must present to
xtraordinary appearance of a moon of nearly 2° degrees in
xhibiting phases complementary to those which we see the
, but *immoveably fixed in their sky*, (or, at least, changing its
ace only by the small amount of the libration,) while the stars
to pass slowly beside and behind it. It will appear clouded
le spots, and belted with equatorial and tropical zones corres-
our trade-winds; and it may be doubted whether, in their per-
ge, the outlines of our continents and seas can ever be clearly
During a solar eclipse, the earth's atmosphere will become
narrow, but bright luminous ring of a ruddy colour, where it
earth, gradually passing into faint blue, encircling the whole
he dark disc of the earth, the remainder being dark and rugged

CHAPTER VIII.

OF TERRESTRIAL GRAVITY. — OF THE LAW OF UNIVERSAL GRAVITATION. — PATHS OF PROJECTILES; APPARENT — REAL — THE MOON RETAINED IN HER ORBIT BY GRAVITY. — ITS LAW OF DIMINUTION. — LAWS OF ELLIPTIC MOTION. — ORBIT OF THE EARTH ROUND THE SUN IN ACCORDANCE WITH THESE LAWS. — MASSES OF THE EARTH AND SUN COMPARED. — DENSITY OF THE SUN. — FORCE OF GRAVITY AT ITS SURFACE. — DISTURBING EFFECT OF THE SUN ON THE MOON'S MOTION.

(438.) THE reader has now been made acquainted with the chief phenomena of the motions of the earth in its orbit round the sun, and of the moon about the earth. — We come next to speak of the physical cause which maintains and perpetuates these motions, and causes the massive bodies so revolving to deviate continually from the directions they would naturally seek to follow, in pursuance of the first law of motion,[1] and bend their courses into curves concave to their centres.

(439.) Whatever attempts may have been made by metaphysical writers to reason away the connection of cause and effect, and fritter it down into the unsatisfactory relation of habitual sequence,[2] it is certain that the conception of some more real and intimate connection is quite as strongly impressed upon the human mind as that of the existence of an external world, — the vindication of whose reality has (strange to say) been regarded as an achievement of no common merit in the annals of this branch of philosophy. It is our own immediate consciousness *of effort*, when we exert force to put matter in motion, or to oppose and neutralize force, which gives us this internal conviction of *power* and *causa-*

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[1] Princip. Lex. i.

[2] See Brown "On Cause and Effect," — a work of great acuteness and subtlety of reasoning on some points, but in which the whole train of argument is vitiated by one enormous oversight; the omission, namely, of a *distinct and immediate personal consciousness of causation* in his enumeration of that *sequence of events*, by which the volition of the mind is made to terminate in the motion of material objects. I mean the consciousness of *effort*, accompanied with *intention thereby* to accomplish an end, as a thing entirely distinct from mere *desire* or *volition* on the one hand, and from mere spasmodic contraction of muscles on the other. Brown, 3d edit. Edin. 1818, p. 47. (Note to edition of 1833.)

r as it refers to the material world, and compels us to believe
ever we see material objects put in motion from a state of rest,
d from their rectilinear paths and changed in their velocities if
motion, it is in consequence of such an EFFORT *somehow*
ough not accompanied with *our* consciousness. That such an
ld be exerted with success through an interposed space, is no
ult to conceive, than that our hand should communicate motion
with which it is *demonstrably not in contact.*

All bodies with which we are acquainted, when raised into the
ietly abandoned, descend to the earth's surface in lines perpen-
it. They are therefore urged thereto by a force or effort, which
asonable to regard as the direct or indirect result of a *conscious-*
will existing *somewhere*, though beyond our power to trace,
e we term *gravity*, and whose tendency or direction, as uni-
erience teaches, is towards the earth's centre; or rather, to
tly, with reference to its spheroidal figure, perpendicular to the
still water. But if we cast a body obliquely into the air, this
though not extinguished or diminished, is materially modified
mate effect. The upward impetus we give the stone is, it is
a time destroyed, and a downward one communicated to it,
mately brings it to the surface, where it is opposed in its fur-
ce, and brought to rest. But all the while it has been conti-

deflects *her* at every instant from the tangent of her orbit, and keeps her in the elliptic path which experience teaches us she actually pursues?

(442.) If a stone be whirled round at the end of a string it will stretch the string by a *centrifugal* force, which, if the speed of rotation be sufficiently increased, will at length break the string, and let the stone escape. However strong the string, it may, by a sufficient rotary velocity of the stone, be brought to the utmost tension it will bear without breaking; and if we know what weight it is capable of carrying, the velocity necessary for this purpose is easily calculated. Suppose, now, a string to connect the earth's centre with a weight at its surface, whose strength should be just sufficient to sustain that weight suspended from it. Let us, however, for a moment imagine gravity to have no existence, and that the weight is made to revolve with the *limiting velocity* which that string can barely counteract: then will its tension be just equal to the weight of the revolving body; and any power which should continually urge the body towards the centre with a force equal to its weight would perform the office, and might supply the place of the string, if divided. Divide it then, and in its place let gravity act, and the body will circulate as before; its tendency to the centre, or *its weight*, being just balanced by its centrifugal force. Knowing the radius of the earth, we can calculate by the principles of mechanics the periodical time in which a body so balanced must circulate to keep it up; and this appears to be $1^h\ 23^m\ 22^s$.

(443.) If we make the same calculation for a body at the distance of the moon, *supposing its weight* or *gravity the same as at the earth's surface*, we shall find the period required to be $10^h\ 45^m\ 30^s$. The actual period of the moon's revolution, however, is $27^d\ 7^h\ 43^m$; and hence it is clear that the moon's velocity is not nearly sufficient to sustain it against *such* a power, supposing it to revolve in a circle, or neglecting (for the present) the slight ellipticity of its orbit. In order that a body at the distance of the moon (or the moon itself) should be capable of *keeping its distance* from the earth by the outward effort of its centrifugal force, while yet its time of revolution should be what the moon's actually is, it will appear[1] that *gravity*, instead of being as intense as at the surface, would require to be very nearly 3600 times less energetic; or, in other words, that its intensity is so enfeebled by the remoteness of the body on which it acts, as to be capable of producing in it, in the same time, only $\frac{1}{3600}$th part of the motion which it would impart to the same mass of matter at the earth's surface.

(444.) The distance of the moon from the earth's centre is a very little less than sixty times the distance from the centre to the surface, and

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[1] Newton, Princip. b. i., Prop. 4., Cor. 2.

: $60^2 : 1^2$; so that the proportion in which we must admit the
ity to be enfeebled at the moon's distance, if it be really the
retains the moon in her orbit, must be (at least in this par-
ance) that of the squares of the distances at which it is com-
w, in such a diminution of energy with increase of distance,
thing *prima facie* inadmissible. Emanations from a centre,
t and heat, do really diminish in intensity by increase of dis-
in this identical proportion; and though we cannot certainly
from this analogy, yet we do see that the power of magnetic
attractions and repulsions is actually enfeebled by distance,
more rapidly than in the simple proportion of the increased
The argument, therefore, stands thus: — On the one hand,
a real power, of whose agency we have daily experience. We
it extends to the greatest accessible heights, and far beyond;
no reason for drawing a line at any particular height, and
ting that it must cease entirely; though we have analogies to
suppose its energy may diminish as we ascend to great heights
rface, such as that of the moon. On the other hand we are
oon *is* urged towards the earth by *some* power which retains
orbit, and that the intensity of this power is such as would
to a gravity, diminished in the proportion — otherwise not
— of the squares of the distances. If gravity be *not* that

direct solution. Newton himself has shown (*Princip.* b. i. prop. 75) that, in that case, the attraction is precisely the same as if the whole matter of each sphere were collected into its centre, and the spheres were single particles there placed; so that, in this case, the general law applies in its strict wording. The effect of the trifling deviation of the earth from a spherical form is of too minute an order to need attention at present. It is, however, perceptible, and may be hereafter noticed.

(446.) The next step in the Newtonian argument is one which divests the law of gravitation of its provisional character, as derived from a loose and superficial consideration of the lunar orbit as a circle described with an average or mean velocity, and elevates it to the rank of a general and primordial relation by proving its applicability to the state of existing nature in all its circumstances. This step consists in demonstrating, as he has done[1] (*Princip.* i. 17. i. 75.), that, under the influence of such an attractive force mutually urging two spherical gravitating bodies towards each other, they will each, when moving in each other's neighbourhood, be deflected into an orbit concave towards the other, and describe, one about the other regarded as fixed, or both round their common centre of gravity, curves whose forms are limited to those figures known in geometry by the general name of conic sections. It will depend, he shows, in any assigned case, upon the particular circumstances or velocity, distance, and direction, *which* of these curves shall be described, — whether an ellipse, a circle, a parabola, or an hyperbola; but one or other it *must* be; and any one of any degree of excentricity it *may* be, according to the circumstances of the case; and, in all cases, the point to which the motion is referred, whether it be the centre of one of the spheres, or their common centre of gravity, will of necessity be the *focus* of the conic section described. He shows, furthermore (*Princip.* i. 1.), that, in every case, the *angular velocity* with which the line joining their centres moves, must be inversely proportional to the square of their mutual distance, and that equal areas of the curves described will be swept over by their line of junction in equal times.

(447.) All this is in conformity with what we have stated of the solar

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[1] We refer for these fundamental propositions, as a point of duty, to the immortal work in which they were first propounded. It is impossible for us, in this volume, to go into these investigations: even did our limits permit, it would be utterly inconsistent with our plan: a general idea, however, of their conduct will be given in the next chapter. We trust that the careful and attentive study of the Principia *in its original form* will never be laid aside, whatever be the improvements of the modern analysis as respects facility of calculation and expression. From no other quarter can a thorough and complete comprehension of the mechanism of our system, (so far as the immediate scope of that work,) be anything like so well, and we may add, so easily obtained.

novements. Their orbits are ellipses, but of different degrees
city; and this circumstance already indicates the general appli-
the principles in question.

But here we have already, by a natural and ready implication
ways the progress of generalization), taken a further and most
step, almost unperceived. We have extended the action of
he case of the earth and sun, to a distance immensely greater
of the moon, and to a body apparently quite of a different
either. Are we justified in this? or, at all events, are there
tions introduced by the change of data, if not into the general
at least into the particular interpretation, of the law of gravi-
ow, the moment we come to numbers, an obvious incongruity
When we calculate, as above, from the known distance of the
57), and from the period in which the earth circulates about it
what must be the centrifugal force of the latter by which the
tion is balanced, (and which, therefore, becomes an exact mea-
sun's attractive energy as exerted on the earth,) we find it to
ly greater than would suffice to counteract the *earth's* attrac-
equal body at that distance—greater in the high proportion of
1. It is clear, then, that if the earth be retained in its orbit
un by *solar attraction*, conformable in its rate of diminution

is reinforced, and rendered capable of resisting this almost inconceivable pressure without collapsing into smaller dimensions.

(450.) This will be more distinctly appreciated, if we estimate, as we are now prepared to do, the intensity of gravity at the sun's surface.

The attraction of a sphere being the same (art. 445) as if its whole mass were collected in its centre, will, of course, be proportional to the mass directly, and the square of the distance inversely; and, in this case, the distance is the radius of the sphere. Hence we conclude[1], that the intensities of solar and terrestrial gravity at the surfaces of the two globes are in the proportions of 27·9 to 1. A pound of terrestrial matter at the sun's surface, then, would exert a pressure equal to what 27·9 such pounds would do at the earth's. The efficacy of muscular power to overcome weight is therefore proportionally nearly 28 times less on the sun than on the earth. An ordinary man, for example, would not only be unable to sustain his own weight on the sun, but would be literally crushed to atoms under the load.[2]

(451.) Henceforward, then, we must consent to dismiss all idea of the earth's immobility, and transfer that attribute to the sun, whose ponderous mass is calculated to exhaust the feeble attractions of such comparative atoms as the earth and moon, without being perceptibly dragged from its place. Their centre of gravity lies, as we have already hinted, almost close to the centre of the solar globe, at an interval quite imperceptible from our distance; and whether we regard the earth's orbit as being performed about the one or the other centre makes no appreciable difference in any one phenomenon of astronomy.

(452.) It is in consequence of the *mutual* gravitation of all the several parts of matter, which the Newtonian law supposes, that the earth and moon, while in the act of revolving, monthly, in their mutual orbits about their common centre of gravity, yet continue to circulate, without parting company, in a greater annual orbit round the sun. We may conceive this motion by connecting two unequal balls by a stick, which, at their centre of gravity, is tied by a long string, and whirled round. Their joint *system* will circulate as one body about the common centre to which the string is attached, while yet they may go on circulating round each other in subordinate gyrations, as if the stick were quite free from any such tie, and merely hurled through the air. If the earth alone, and not the moon, gravitated to the sun, it would be dragged away, and leave the moon

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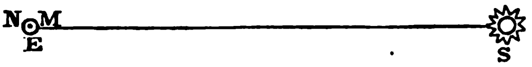
[1] Solar gravity : terrestrial : : $\frac{354936}{(440000)^2} : \frac{1}{(4000)^2}$: : 27·9 : 1; the respective radii of the sun and earth being 440000, and 4000 miles.

[2] A mass weighing 12 stone or 168lbs. on the earth, would produce a pressure of 4687 lbs. on the sun.

nd *vice versâ;* but, acting on both, they continue together
ttraction, just as the loose parts of the earth's surface continue
on it. It is, then, in strictness, not the earth or the moon
cribes an ellipse around the sun, but their common centre of
The effect is to produce a small, but very perceptible, monthly
n the sun's apparent motion as seen from the earth, which is
en into account in calculating the sun's place. The moon's
in its compound orbit round the earth and sun is an epicycloi-
ntersecting the orbit of the earth twice in every lunar month,
tely within and without it. But as there are not more than
n months in the year, and as the total departure of the moon
her way does not exceed one 400th part of the radius, this
ly to a slight undulation upon the earth's ellipse, so slight,
t if drawn in true proportion on a large sheet of paper, no eye
measurement with compasses would detect it. The real orbit
n is everywhere concave towards the sun.

Here moreover, *i. e.* in the attraction of the sun, we have the
those differences from an exact elliptic movement of the moon
thly orbit, which we have already noticed (arts. 407, 409), viz.
ograde revolution of her nodes; to the direct circulation of the
ellipse; and to all the other deviations from the laws of elliptic

Fig. 62.



above that of the earth when in the nearer point of her orbit, M, and a corresponding defect on the opposite part, N; and, in the intermediate positions, not only will a difference of *forces* subsist, but a difference of *directions* also; since however small the lunar orbit M N, it is not a point, and, therefore, the lines drawn from the sun S to its several parts cannot be regarded as strictly parallel. If, as we have already seen, the force of the sun were equally exerted, and in parallel directions on both, no disturbance of their relative situations would take place; but from the non-verification of these conditions arises a *disturbing force*, oblique to the line joining the moon and earth, which in some situations acts to *accelerate*, in others to *retard*, her elliptic orbitual motion; in some to draw the earth from the moon, in others the moon from the earth. Again, the lunar orbit, though very nearly, is yet not quite coincident with the plane of the ecliptic; and hence the action of the sun, which is very nearly parallel to the last-mentioned plane, tends to draw her somewhat *out of the plane* of her orbit, and does actually do so—producing the revolution of her nodes, and other phenomena less striking. We are not yet prepared to go into the subject of these *perturbations*, as they are called; but they are introduced to the reader's notice as early as possible, for the purpose of reassuring his mind, should doubts have arisen as to the logical correctness of our argument, in consequence of our temporary neglect of them while working our way upward to the law of gravity from a general consideration of the moon's orbit.

CHAPTER IX.

OF THE SOLAR SYSTEM.

MOTIONS OF THE PLANETS. — THEIR STATIONS AND RE-
DATIONS. — THE SUN THEIR NATURAL CENTRE OF MOTION.
RIOR PLANETS. — THEIR PHASES, PERIODS, ETC. — DIMEN-
ND FORM OF THEIR ORBITS. — TRANSITS ACROSS THE SUN.
RIOR PLANETS. — THEIR DISTANCES, PERIODS, ETC. — KEP-
AWS AND THEIR INTERPRETATION. — ELLIPTIC ELEMENTS
PLANET'S ORBIT. — ITS HELIOCENTRIC AND GEOCENTRIC
— EMPIRICAL LAW OF PLANETARY DISTANCES; — VIOLATED
CASE OF NEPTUNE. — THE ULTRA-ZODIACAL PLANETS. —
L PECULIARITIES OBSERVABLE IN EACH OF THE PLANETS.

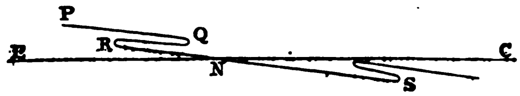
THE sun and moon are not the only celestial objects which

(457.) The apparent motions of the planets are much more irregular than those of the sun or moon. Generally speaking, and comparing their places at distant times, they all advance, though with very different *average* or *mean* velocities, in the same direction as those luminaries, *i. e.* in opposition to the apparent diurnal motion, or from west to east: all of them make the entire tour of the heavens, though under very different circumstances; and all of them, with the exception of the eight telescopic planets,—Ceres, Pallas, Juno, Vesta, Astræa, Hebe, Iris, and Flora (which may therefore be termed *ultra-zodiacal*,) — are confined in their visible paths within very narrow limits on either side the ecliptic, and perform their movements within that zone of the heavens we have called, above, the Zodiac (art. 303.)

(458.) The obvious conclusion from this is, that whatever be, otherwise, the nature and law of their motions, they are performed *nearly in the plane of the ecliptic* — that plane, namely, in which our own motion about the sun is performed. Hence it follows, that we see their evolutions, not in *plan*, but in *section;* their real angular movements and linear distances being all *foreshortened* and confounded undistinguishably, while only their deviations from the ecliptic appear of their natural magnitude, undiminished by the effect of perspective.

(459.) The apparent motions of the sun and moon, though not uniform, do not deviate very greatly from uniformity; a moderate acceleration and retardation, accountable for by the ellipticity of their orbits, being all that is remarked. But the case is widely different with the planets: sometimes they advance rapidly; then relax in their apparent speed — come to a momentary stop; and then actually reverse their motion, and run back upon their former course, with a rapidity at first increasing, then diminishing, till the reversed or retrograde motion ceases altogether. Another *station*, or moment of apparent rest or indecision, now takes place; after which the movement is again reversed, and resumes its original direct character. On the whole, however, the amount of direct motion more than compensates the retrograde; and by the excess of the former over the latter, the gradual advance of the planet from west to east is maintained. Thus, supposing the Zodiac to be unfolded into a plane surface, (or represented as in Mercator's projection,

Fig. 63.



king the ecliptic E C for its ground line,) the track of a planet
ped down by observation from day to day, will offer the appear-
R S, &c.; the motion from P to Q being direct, at Q stationary,
R retrograde, at R again stationary, from R to S direct, and

In the midst of the irregularity and fluctuation of this motion,
kable feature of uniformity is observed. Whenever the planet
ecliptic, as at N in the figure, it is said (like the moon) to be
; and as the earth necessarily lies in the plane of the ecliptic,
cannot be *apparently* or *uranographically* situated in the
rcle so called, without being *really* and *locally* situated *in that*
he visible passage of a planet through its node, then, is a phe-
ndicative of a circumstance in its real motion quite independent
tion from which we view it. Now, it is easy to ascertain, by
, when a planet passes from the north to the south side of the
e have only to convert its right ascensions and declinations into
and latitudes, and the change from north to south latitude on
ssive days will advertise us on what *day* the transition took
ile a simple proportion, grounded on the observed state of its
latitude in the interval, will suffice to fix the precise hour and
its arrival on the ecliptic. Now, this being done for several

after what has been shown in art. 449, of the immense mass of that luminary, and of the office it performs to us as a quiescent centre of our orbitual motion, nothing can be more natural than to suppose it may perform the same to other globes which, like the earth, may be revolving round it; and these globes may be visible to us by its light reflected from them, as the moon is. Now there are many facts which give a strong support to the idea that the planets are in this predicament.

(463.) In the first place, the planets really are great globes, of a size commensurate with the earth, and several of them much greater. When examined through powerful telescopes, they are seen to be round bodies, of sensible and even of considerable apparent diameter, and offering distinct and characteristic peculiarities, which show them to be solid masses, each possessing its individual structure and mechanism; and that, in one instance at least, an exceedingly artificial and complex one. (See the representations of Mars, Jupiter, and Saturn, in Plate III.) That their distances from us are great, much greater than that of the moon, and some of them even greater than that of the sun, we infer, 1st, from their being occulted by the moon, and 2dly, from the smallness of their diurnal parallax, which, even for the nearest of them, when most favourably situated, does not exceed a few seconds, and for the remote ones is almost imperceptible. From the comparison of the diurnal parallax of a celestial body, with its apparent semidiameter, we can at once estimate its real size. For the parallax is, in fact, nothing else than the apparent semidiameter of the earth as seen from the body in question (art. 339 et seq.); and, the intervening distance being the same, the real diameters must be to each other in the proportion of the apparent ones. Without going into particulars, it will suffice to state it as a general result of that comparison, that the planets are all of them incomparably smaller than the sun, but some of them as large as the earth, and others much greater.

(464.) The next fact respecting them is, that their distances from us, as estimated from the measurement of their angular diameters, are in a continual state of change, periodically increasing and decreasing within certain limits, but by no means corresponding with the supposition of regular circular or elliptic orbits described by them about the earth as a centre or focus, but maintaining a constant and obvious relation to their apparent angular distances or *elongations* from the sun. For example; the apparent diameter of Mars is greatest when in opposition (as it is called) to the sun, *i. e.* when in the opposite part of the ecliptic, or when it comes on the meridian at midnight, — being then about 18″, — but diminishes rapidly from that to about 4″, which is its apparent diameter when in *conjunction*, or when seen in nearly the same direction as that

This, and facts of a similar character, observed with respect
arent diameters of the other planets, clearly point out the sun
more than an accidental relation to their movements.

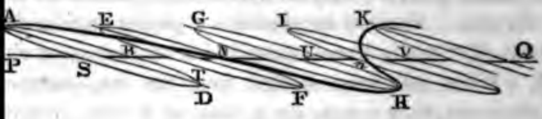
Lastly, certain of the planets, (Mercury, Venus, and Mars,)
ed through telescopes, exhibit the appearance of phases like
he moon. This proves that they are opaque bodies, shining
flected light, which can be no other than that of the sun's;
ecause there is no other source of light external to them suffi-
werful, but because the appearance and succession of the phases
are (like their visible diameters) intimately connected with
ations from the sun, as will presently be shown.

Accordingly it is found, that, when we refer the planetary move-
the sun as a centre, all that apparent irregularity which they
viewed from the earth disappears at once, and resolves itself
imple and general law, of which the earth's motion, as ex-
a former chapter, is only a particular case. In order to show
appens, let us take the case of a single planet, which we will
revolve round the sun, in a plane nearly, but not quite, coin-
the ecliptic, but passing through the sun, and of course inter-
ecliptic in a fixed line, which is the line of the planet's nodes.
must of course divide its orbit into two segments; and it is

rise before that luminary in the morning, and appear over the eastern horizon as morning stars: they do not, however, attain the same *elongation* from the sun. Mercury never attains a greater angular distance from it than about 29°, while Venus extends her excursions on either side to about 47°. When they have receded from the sun, *eastward*, to their respective distances, they remain for a time, as it were, immoveable *with respect to it*, and are carried along with it in the ecliptic with a motion equal to its own; but presently they begin to approach it, or, which comes to the same, their motion in longitude diminishes, and the sun gains upon them. As this approach goes on, their continuance above the horizon after sunset becomes daily shorter, till at length they set before the darkness has become sufficient to allow of their being seen. For a time, then, they are not seen at all, unless on very rare occasions, when they are to be observed *passing across the sun's disc as small, round, well-defined black spots*, totally different in appearance from the solar spots (art. 386.) These phenomena are emphatically called *transits* of the respective planets across the sun, and take place when the earth happens to be passing the line of their nodes while they are in that part of their orbits, just as in the account we have given (art. 412) of a solar eclipse. After having thus continued invisible for a time, however, they begin to appear on the other side of the sun, at first showing themselves only for a few minutes before sunrise, and gradually longer and longer as they recede from him. At this time their motion in longitude is rapidly retrograde. Before they attain their greatest elongation, however, they become stationary in the heavens; but their recess from the sun is still maintained by the advance of that luminary along the ecliptic, which continues to leave them behind, until, having reversed their motion, and become again *direct*, they acquire sufficient speed to commence overtaking him — at which moment they have their greatest *western* elongation: and thus is a kind of oscillatory movement kept up, while the general advance along the ecliptic goes on.

(468.) Suppose P Q to be the ecliptic, and A B D the orbit of one of these planets, (for instance, Mercury,) seen almost edgewise by an eye situated very nearly in its plane; S, the sun, its centre; and A, B, D, S, successive positions of the planet, of which B and S are in the nodes. If, then, the sun S stood apparently still in the ecliptic, the planets would simply appear to oscillate backwards and forwards from A to D, alternately passing before and behind the sun; and, if the eye happened to lie exactly in the plane of the orbit, *transiting* his disc in the former case, and being covered by it in the latter. But as the sun is not so stationary, but apparently carried along the ecliptic P Q, let it be supposed

Fig. 64.



er the spaces S T, T U, U V, while the planet in each case exe-
uarter of its period. Then will its orbit be apparently carried
the sun, into the successive positions represented in the figure;
ts real motion round the sun brings it into the respective points,
its apparent movement in the heavens will seem to have been
vavy or zigzag line A N H K. In this, its motion in longitude
been direct in the parts A N, N H, and retrograde in the parts
hile at the turns of the zigzag, as at H, it will have been sta-

The only two planets — Mercury and Venus — whose evolu-
ich as above described, are called inferior planets; their points
recess from the sun are called (as above) their *greatest* eastern
n *elongations;* and their points of nearest approach to it, their
d *superior* conjunctions,—the former when the planet passes
e earth and the sun, the latter when behind the sun.

the angle a A S, which measures their apparent interval as seen from A, being then greater than in any other situation of a upon its own circle.

(471.) Now, this angle being known by observation, we are hereby furnished with a ready means of ascertaining, at least approximately, the distance of the planet from the sun, or the radius of its orbit, supposed a circle. For the triangle S A a is right-angled at a, and consequently we have S a : S A : : sin. S A a : radius, by which proportion the radii S a, S A of the two orbits are directly compared. If the orbits were both exact circles, this would of course be a perfectly rigorous mode of proceeding: but (as is proved by the inequality of the resulting values of S a obtained at different times) this is not the case; and it becomes necessary to admit an excentricity of position, and a deviation from the exact circular form in *both* orbits, to account for this difference. Neglecting, however, at present this inequality, a mean or average value of S a may, at least, be obtained from the frequent repetition of this process in all varieties of situation of the two bodies. The calculations being performed, it is concluded that the mean distance of Mercury from the sun is about 36000000 miles; and that of Venus, similarly derived, about 68000000; the radius of the earth's orbit being 95000000.

(472.) The sidereal periods of the planets may be obtained (as before observed), with a considerable approach to accuracy, by observing their passages through the nodes of their orbits; and indeed, when a certain very minute motion of these nodes and the apsides of their orbits (similar to that of the moon's nodes and apsides, but incomparably slower) is allowed for, with a precision only limited by the imperfection of the appropriate observations. By such observations, so corrected, it appears that the sidereal period of Mercury is $87^d\ 23^h\ 15^m\ 43{\cdot}9^s$; and that of Venus, $224^d\ 16^h\ 49^m\ 8{\cdot}0^s$. These periods, however, are widely different from the intervals at which the successive appearances of the two planets at their eastern and western elongations from the sun are observed to happen. Mercury is seen at its greatest splendour as an evening star, at average intervals of about 116, and Venus at intervals of about 584 days. The difference between the *sidereal* and *synodical* revolutions (art. 418) accounts for this. Referring again to the figure of art. 470, if the sun stood still at A, while the planet advanced in its orbit, the lapse of a sidereal period, which should bring it round again to a, would also produce a similar elongation from the sun. But, meanwhile, the earth has advanced in its orbit in the same direction towards E, and therefore the next greatest elongation on the same side of the sun will happen — not in the position a A of the two bodies, but in some more advanced position, e E. The determination of this position depends on a calculation

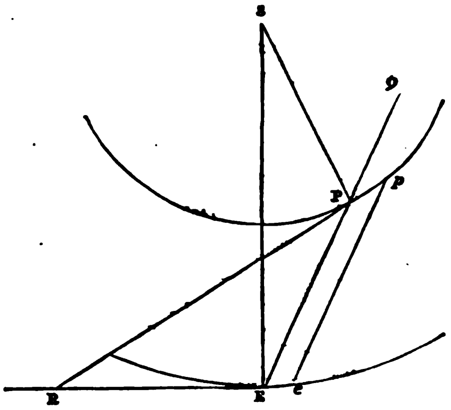
nilar to what has been explained in the article referred to; and
therefore, only state the resulting synodical revolutions of the
s, which come out respectively $115{\cdot}877^d$, and $583{\cdot}920^d$.

In this interval, the planet will have described a whole revolu-
the arc *a c e*, and the earth only the arc A C E of its orbit.
lapse, the *inferior conjunction* will happen when the earth has
ntermediate situation, B, and the planet has reached *b*, a point
e sun and earth. The greatest elongation on the opposite side
will happen when the earth has come to C, and the planet to
e line of junction C *c* is a tangent to the interior circle on the
ide from M. Lastly, the *superior* conjunction will happen
earth arrives at D, and the planet at *d* in the same line pro-
the other side of the sun. The intervals at which these phe-
ppen may easily be computed from a knowledge of the synodi-
and the radii of the orbits.

The circumferences of circles are in the proportion of their
then, we calculate the circumferences of the orbits of Mercury
s, and the earth, and compare them with the times in which
lutions are performed, we shall find that the actual velocities
h they move in their orbits differ greatly; that of Mercury
ut 109360 miles per hour, of Venus 80000, and of the earth
From this it follows that at the inferior conjunction, or at *b*

passing the earth's. Hence, the stationary points ought to lie, as it is found by observation they do, between *a* and *b*, or *c* and *b*, viz. in such a position that the obliquity of the planet's motion with respect to the line of junction shall just compensate for the excess of its velocity, and cause an equal advance of each extremity of that line, by the motion of the planet at one end, and of the earth at the other: so that, for an instant of time, the whole line shall move parallel to itself. The question thus proposed is purely geometrical, and its solution on the supposition of circular orbits is easy. Let E *e* and P *p* represent small arcs of the orbits of the earth and planet described contemporaneously, at the moment when the latter appears stationary, about S, the sun. Produce *p* P and *e* E, tangents at P and E, to meet at R, and prolong E P backwards to Q, join *e p*. Then since P E, *p e* are parallel, we have by similar triangles P *p* : E *e* :: P R : R E, and since, putting *v* and V for the respective velocities of the planet and the earth, P *p* : E *e* :: *v* : V; therefore

Fig. 66.



$$v : V :: P R : R E :: \sin. P E R : \sin. E P R$$
$$:: \cos. S E P : \cos. S P Q$$
$$:: \cos. S E P : \cos. (S E P + E S P)$$

because the angles S E R and S P R are right angles. Moreover, if *r* and R be the radii of the respective orbits, we have also

$$r : R :: \sin. S E P : \sin. (S E P + E S P)$$

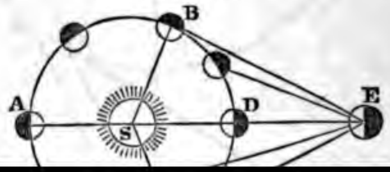
from which two relations it is easy to deduce the values of the two angles S E P and E S P; the former of which is the apparent elongation of the

the sun,[1] the latter the difference of heliocentric longitudes h and planet.

When we regard the orbits as other than circles (which they , the problem becomes somewhat complex — too much so to be d upon. It will suffice to state the results which experience d which assigns the stationary points of Mercury at from 15° elongation from the sun, according to circumstances; and of an elongation never varying much from 29°. The former con- etrograde during about 22 days; the latter, about 42.

We have said that some of the planets exhibit phases like the his is the case with both Mercury and Venus; and is readily by a consideration of their orbits, such as we have above sup- n. In fact, it requires little more than mere inspection of the exed, to show, that to a spectator situated on the earth E, an

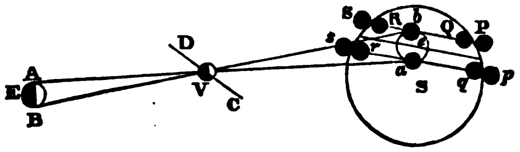
Fig. 67.



varying proportion of its visible illuminated area to its whole disc; and, 2dly, the varying angular diameter, or whole apparent magnitude of the disc itself. As it approaches its inferior conjunction from its greater elongation, the half-moon becomes a crescent, which *thins off;* but this is more than compensated, for some time, by the increasing apparent magnitude, in consequence of its diminishing distance. Thus the total light received from it goes on increasing, till at length it attains a maximum, which takes place when the planet's elongation is about 40°. #

(479.) The transits of Venus are of very rare occurrence, taking place alternately at the very unequal but regularly recurring intervals of 8, 122, 8, 105, 8, 122, &c., years in succession, and always in June or December. As astronomical phænomena, they are extremely important; since they afford the best and most exact means we possess of ascertaining the sun's distance, or its parallax. Without going into the niceties of calculation of this problem, which, owing to the great multitude of circumstances to be attended to, are extremely intricate, we shall here explain its principle, which, in the abstract, is very simple and obvious. Let E be the earth, V Venus, and S the sun, and C D the portion of Venus's relative orbit which she describes while in the act of transiting the sun's disc. Suppose A B two spectators at opposite extremities of that diameter of the earth which is perpendicular to the ecliptic, and, to avoid complicating the case, let us lay out of consideration the earth's rotation, and suppose A, B, to retain that situation during the whole time of the transit. Then, at any moment when the spectator at A sees the centre of Venus projected at *a* on the sun's disc, he at B will see it projected at *b*. If then one or other spectator could suddenly transport himself from A to B, he would see Venus suddenly displaced on the disc from *a* to *b*; and if he had any means of noting accurately the place of the points on the disc, either by micrometrical measures from its edge, or by other means, he might ascertain the angular measure of *a b* as seen from the earth. Now, since A V *a*, B V *b*, are straight lines, and therefore make equal angles on each side V, *a b* will be to A B as the distance of Venus from the sun is to its distance from the earth, or as 68 to 27, or nearly as 2½ to 1; *a b* there-

Fig. 67.



es on the sun's disc a space 2½ times as great as the earth's
and its angular measure is therefore equal to about 2½ times
apparent diameter at the distance of the sun, or (which is the
) to five times the sun's horizontal parallax (art. 298). Any
efore, which may be committed in measuring *a b*, will entail
fth of that error on the horizontal parallax concluded from it.

The thing to be ascertained, therefore, is, in fact, neither more
an the breadth of the zone P Q R S, *p q r s*, included between
e apparent paths of the centre of Venus across the sun's disc,
ntry on one side to its quitting it on the other. The whole
the observers at A, B, therefore, resolves itself into this;—to
ith all possible care and precision, each at his own station, this
ere it enters, where it quits, and what segment of the sun's disc

Now, one of the most exact ways in which (conjoined with
rometric measures) this can be done, is by noting the *time* occu-
whole transit; for the relative angular motion of Venus being,
y precisely known from the tables of her motion, and the appa-
eing very nearly a straight line, these times give us a measure
enlarged scale) of the lengths of the chords of the segments
d the sun's diameter being known also with great precision,
d sines, and therefore their difference, or the breadth of the

adequate to every purpose, by only watching favourable opportunities, and taking advantage of nicely adjusted combinations of circumstance. So important has this observation appeared to astronomers, that at the last transit of Venus, in 1769, expeditions were fitted out, on the most extensive scale, by the British, French, Russian, and other governments, to the remotest corners of the globe, for the express purpose of performing it. The celebrated expedition of Captain Cook to Otaheite was one of them. The general result of all the observations made on this most memorable occasion gives 8″5776 for the sun's horizontal parallax. The two next occurrences of this phænomenon will happen on Dec. 8, 1874, and Dec. 6, 1882.

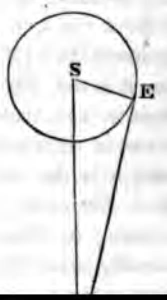
(482.) The orbit of Mercury is very elliptical, the excentricity being nearly one fourth of the mean distance. This appears from the inequality of the greatest elongations from the sun, as observed at different times, and which vary between the limits 16° 12′ and 28° 48′, and, from exact measures of such elongations, it is not difficult to show that the orbit of Venus also is slightly excentric, and that both these planets, in fact, describe ellipses, having the sun in their common focus.

(483.) Transits of Mercury over the sun's disc occasionally occur, as in the case of Venus, but more frequently; those at the ascending node in November, at the descending in May. The intervals (considering each node separately) are *usually* either 13 or 7 years, and in the order 13, 13, 13, 7, &c.; but owing to the considerable inclination of the orbit of Mercury to the ecliptic, this cannot be taken as an exact expression of the said recurrence, and it requires a period of at least 217 years to bring round the transits in regular order. One will occur in the present year (1848,) the next in 1861. They are of much less astronomical importance than that of Venus, on account of the proximity of Mercury to the sun, which affords a much less favourable combination for the determination of the sun's parallax.

(484.) Let us now consider the superior planets, or those whose orbits enclose on all sides that of the earth. That they do so is proved by several circumstances:—1st, They are not, like the inferior planets, confined to certain limits of elongation from the sun, but appear at all distances from it, even in the opposite quarter of the heavens, or, as it is called, in *opposition;* which could not happen, did not the earth at such times place itself between them and the sun: 2dly, They never appear horned, like Venus or Mercury, nor even semilunar. Those, on the contrary, which, from the minuteness of their parallax, we conclude to be the most distant from us, viz. Jupiter, Saturn, Uranus, and Neptune, never appear otherwise than round; a sufficient proof, of itself, that we see them always in a direction not very remote from that in which the

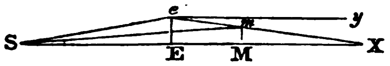
luminate them; and that, therefore, we occupy a station which
y widely removed from the centre of their orbits, or, in other
the earth's orbit is entirely enclosed within theirs, and of
ly small diameter. Only one of them, Mars, exhibits any
phase, and in its deficiency from a circular outline, never
moderately *gibbous* appearance, — the enlightened portion of
ng never less than seven-eighths of the whole. To understand
d only cast our eyes on the annexed figure, in which E is the
apparent greatest elongation from the sun S, as seen from
In this position, the angle S M E, included between the lines

Fig. 69.



of Jupiter, of Jupiter than of Saturn, of that planet than of Uranus, and of Uranus again than Neptune. The angular velocity with which a planet appears to retrograde is easily ascertained by observing its apparent place in the heavens from day to day; and from such observations, made about the time of opposition, it is easy to conclude the relative magnitudes of their orbits, as compared with the earth's, supposing their periodical times known. For, from these, their mean angular velocities are known also, being inversely as the times. Suppose, then, E *e* to be a very small portion of the earth's orbit, and M *m* a corresponding portion of that of a superior planet, described on the day of opposition, about the sun S, on which day the three bodies lie in one straight line S E M X. Then the angles E S *e* and M S *m* are given. Now, if *e m* be joined and prolonged to meet S M continued in X, the angle *e* X E, which is equal to the alternate angle X *e* Y, is evidently the retrogradation of Mars on that day, and is, therefore, also given. E *e*, therefore, and the angle E X *e*, being given in the right-angled triangle E *e* X, the side E X is easily calculated, and thus S X becomes known. Consequently, in the triangle S *m* X, we have given the side S X and the two angles *m* S X, and *m* X S, whence the other sides, S *m*, *m* X, are easily determined. Now, S *m* is no other than the radius of the orbit of the superior planet required, which in this calculation is supposed circular, as well as that of the earth; a supposition not exact, but sufficiently so to afford a satisfactory approximation to the dimensions of its orbit, and which, if the process be often repeated, in every variety of situation at which the opposition can occur, will ultimately afford an average or mean value of its diameter fully to be depended upon.

Fig. 70.



(486.) To apply this principle, however, to practice, it is necessary to know the periodic times of the several planets. These may be obtained directly, as has been already stated, by observing the intervals of their passages through the ecliptic; but, owing to the very small inclination of the orbits of some of them to its plane, they cross it so obliquely that the precise moment of their arrival on it is not ascertainable, unless by very nice observations. A better method consists in determining, from the observations of several successive days, the exact moments of their arriving *in opposition* with the sun, the criterion of which is a difference of longitudes between the sun and planet of exactly 180°. The interval

ccessive oppositions thus obtained is nearly one *synodical* pe-
would be exactly so, were the planet's orbit and that of the
circles, and uniformly described; but as that is found not to
e (and the criterion is, the *inequality* of successive synodical
so observed), the average of a great number, taken in all va-
ituation in which the oppositions occur, will be freed from the
quality, and may be taken as a *mean synodical* period. From
e considerations and by the process of calculation, indicated
the sidereal periods are readily obtained. The accuracy of this
ion will, of course, be greatly increased by embracing a long
tween the extreme observations employed. In point of fact,
al extends to nearly 2000 years in the cases of the planets
the ancients, who have recorded their observations of them in
sufficiently careful to be made use of. Their periods may, there-
garded as ascertained with the utmost exactness. Their nume-
s will be found stated, as well as the mean distances, and all
elements of the planetary orbits, in the synoptic table at the
volume, to which (to avoid repetition) the reader is once for
.

n casting our eyes down the list of the planetary distances, and
them with the periodic times, we cannot but be struck with a

cesses of trigonometry and of numerical calculation were encumbered with difficulties, of which the more recent invention of logarithmic tables has happily left us no conception, to perceive and demonstrate the real law of their connection. This connection is expressed in the following proposition: —"The squares of the periodic times of any two planets are to each other, in the same proportion as the cubes of their mean distances from the sun." Take, for example, the Earth and Mars,[1] whose periods are in the proportion of 3652564 to 6869796, and whose distance from the sun is that of 100000 to 152369; and it will be found, by any one who will take the trouble to go through the calculation, that—

$$(3652564)^2 : (6869796)^2 :: (100000)^3 : (152369)^3.$$

(488.) Of all the laws to which induction from pure observation has ever conducted man, this *third law* (as it is called) *of Kepler* may justly be regarded as the most remarkable, and the most pregnant with important consequences. When we contemplate the constituents of the planetary system from the point of view which this relation affords us, it is no longer mere analogy which strikes us — no longer a general resemblance among them, as individuals independent of each other, and circulating about the sun, each according to its own peculiar nature, and connected with it by its own peculiar tie. The resemblance is now perceived to be a true *family* likeness; they are bound up in one chain — interwoven in one web of mutual relation and harmonious agreement — subjected to one pervading influence, which extends from the centre to the farthest limits of that great system, of which all of them, the earth included, must henceforth be regarded as members. #

(489.) The laws of elliptic motion about the sun as a focus, and of the equable description of areas by lines joining the sun and planets, were originally established by Kepler, from a consideration of the observed motions of Mars; and were by him extended, analogically, to all the other planets. However precarious such an extension might then have appeared, modern astronomy has completely verified it as a matter of fact, by the general coincidence of its results with entire series of observations of the apparent places of the planets. These are found to accord satisfactorily with the assumption of a particular ellipse for each planet, whose magnitude, degree of excentricity, and situation in space, are numerically assigned in the synoptic table before referred to. It is true, that when observations are carried to a high degree of precision, and when each planet is traced through many successive revolutions, and its history car-

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[1] The expression of this law of Kepler requires a slight modification when we come to the extreme nicety of numerical calculation, for the greater planets, due to the influence of their masses. This correction is imperceptible for the Earth and Mars.

by the aid of calculations founded on these data, for many cen-
earn to regard the laws of Kepler as only *first approximations*
ch more complicated ones which actually prevail; and that to
ote observations into rigorous and mathematical accordance
other, and at the same time to retain the extremely convenient
re and relations of the ELLIPTIC SYSTEM, it becomes necessary
to a certain extent, our verbal expression of the laws, and to
numerical data or *elliptic elements* of the planetary orbits as
ely permanent, but subject to a series of extremely slow and
erceptible changes. These changes may be neglected when we
ly a few revolutions; but going on from century to century,
ally accumulating, they at length produce material departures
s from their original state. Their explanation will form the
a subsequent chapter; but for the present we must lay them
ideration, as of an order too minute to affect the general con-
th which we are now concerned. By what means astronomers
to compare the results of the elliptic theory with observation,
atisfy themselves of its accordance with nature, will be ex-
sently.

t will first, however, be proper to point out what particular
conclusion is involved in each of the three laws of Kepler,

any body, urged towards a certain central point by a force continually directed thereto, and thereby deflected into a curvilinear path, will describe about that centre equal areas in equal times; and *vice versâ*, that such equable description of areas is itself the essential criterion of a continual direction of the acting force towards the centre to which this character belongs. The first law of Kepler, then, gives us no information as to the nature or intensity of the force urging the planets to the sun; the only conclusion it involves is, that it does so urge them. It is a property of orbitual rotation under the influence of central forces *generally*, and, as such, we daily see it exemplified in a thousand familiar instances. A simple experimental illustration of it is to tie a bullet to a thin string, and, having whirled it round with a moderate velocity in a vertical plane, to draw the end of the string through a small ring, or allow it to coil itself round the finger, or round a cylindrical rod held very firmly in a horizontal position. The bullet will then approach the centre of motion in a spiral line; and the increase not only of its angular but of its linear velocity, and the rapid diminution of its periodic time when near the centre, will express, more clearly than any words, the compensation by which its uniform description of areas is maintained under a constantly diminishing distance. If the motion be reversed, and the thread allowed to uncoil, beginning with a rapid impulse, the velocity will diminish by the same degrees as it before increased. The increasing rapidity of a dancer's *pirouette*, as he draws in his limbs and straightens his whole person, so as to bring every part of his frame as near as possible to the axis of his motion, is another instance where the connection of the observed effect with the central force exerted, though equally real, is much less obvious.

(491.) The second law of Kepler, or that which asserts that the planets describe ellipses about the sun as their focus, involves, as a consequence, the *law* of solar gravitation (so be it allowed to call the force, whatever it be, which urges them towards the sun) as exerted on each individual planet, apart from all connection with the rest. A straight line, dynamically speaking, is the only path which can be pursued by a body *absolutely free*, and under the action of *no* external force. All *deflection* into a curve is evidence of the exertion of a force; and the greater the deflection in equal times, the more intense the force. Deflection from a straight line is only another word for *curvature* of path; and as a circle is characterized by the uniformity of its curvatures in all its parts—so is every other curve (as an ellipse) characterized by the particular *law* which regulates the increase and diminution of its curvature as we advance along its circumference. The deflecting force, then, which continually bends a

ly into a curve, may be ascertained, provided its direction, in
ce, and, secondly, the law of curvature of the curve itself, be
oth these enter as elements into the expression of the force.
y describe, for instance, an ellipse, under a great variety of
of the acting forces: it may glide along it, for example, as a
polished wire, bent into an elliptic form; in which case the
is always perpendicular to the wire, and the velocity is uni-
his case the *force* is directed to *no fixed* centre, and there is
description of areas at all. Or it may describe it as we may
we suspend a ball by a *very* long string, and, drawing it a
from the perpendicular, throw it *round* with a gentle impulse.
the acting force is directed to the centre of the ellipse, about
are described equally, and *to* which a force *proportional* to
(the decomposed result of terrestrial gravity) perpetually
This is at once a very easy experiment, and a very instructive
shall again refer to it. In the case before us, of an ellipse
y the action of a force directed to the *focus*, the steps of the
n of the law of force are these: 1st, The law of the areas de-
e actual *velocity* of the revolving body at every point, or the
run over by it in a given minute portion of time; 2dly, The
ature of the ellipse determines the linear amount of deflection

(492.) The third law of Kepler, which connects the distances and periods of the planets by a general rule, bears with it, as its theoretical interpretation, this important consequence, viz. that it is one and the same force, modified only by distance from the sun, which retains *all* the planets in their orbits about it. That the attraction of the sun (if such it be) is exerted upon all the bodies of our system indifferently, without regard to the peculiar materials of which they may consist, in the exact proportion of their inertiæ, or quantities of matter; that it is not, therefore, of the nature of the elective attractions of chemistry or of magnetic action, which is powerless on other substances than iron and some one or two more, but is of a more universal character, and extends equally to all the material constituents of our system, and (as we shall hereafter see abundant reason to admit) to those of other systems than our own. This law, important and general as it is, results, as the simplest of corollaries, from the relations established by Newton in the section of the *Principia* referred to (Prop. xv.), from which proposition it results, that if the earth were taken from its actual orbit, and launched anew in space at the place, in the direction, and with the velocity of any of the other planets, it would describe the very same orbit, and in the same period, which that planet actually does, a minute correction of the period only excepted, arising from the difference between the mass of the earth and that of the planet. Small as the planets are compared to the sun, some of them are not, as the earth is, mere atoms in the comparison. The strict wording of Kepler's law, as Newton has proved in his fifty-ninth proposition, is applicable only to the case of planets whose proportion to the central body is absolutely inappreciable. When this is not the case, the periodic time is shortened in the proportion of the square root of the number expressing the sun's mass or inertiæ, to that of the sum of the numbers expressing the masses of the sun and planet; and in general, whatever be the masses of two bodies revolving round each other under the influence of the Newtonian law of gravity, the square of their periodic time will be expressed by a fraction whose numerator is the cube of their mean distance, *i. e.* the greater semi-axis of their elliptic orbit, and whose denominator is the sum of their masses. When one of the masses is incomparably greater than the other, this resolves into Kepler's law; but when this is not the case, the proposition thus generalized stands in lieu of that law. In the system of the sun and planets, however, the numerical correction thus introduced into the results of Kepler's law is too small to be of any importance, the mass of the largest of the planets (Jupiter) being much less than a thousandth part of that of the sun. We shall presently, however, perceive all the importance of this generalization, when we come to speak of the satellites.

(493.) It will first, however, be proper to explain by what process of calculation the expression of a planet's elliptic orbit by its *elements* can be compared with observation, and how we can satisfy ourselves that the numerical data contained in a table of such elements for the whole system does really exhibit a true picture of it, and afford the means of determining its state at every instant of time, by the mere application of Kepler's laws. Now, for each planet, it is necessary for this purpose to know, 1st, the magnitude and form of its ellipse; 2dly, the situation of this ellipse in space, with respect to the ecliptic, and to a fixed line drawn therein; 3dly, the local situation of the planet in its ellipse at some known epoch, and its periodic time or mean angular velocity, or, as it is called, its mean motion.

(494.) The magnitude and form of an ellipse are determined by its greatest length and least breadth, or its two principal axes; but for astronomical uses it is preferable to use the semi-axis major (or half the greatest length), and the excentricity or distance of the focus from the centre, which last is usually estimated in parts of the former. Thus, an ellipse, whose length is 10 and breadth 8 parts of any scale, has for its major semi-axis 5, and for its excentricity 3 such parts; but when estimated in parts of the semi-axis, regarded as a unit, the excentricity is expressed by the fraction $\frac{3}{5}$.

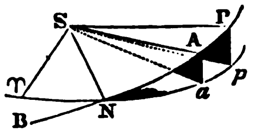
(495.) The ecliptic is the plane to which an inhabitant of the earth most naturally refers the rest of the solar system, as a sort of ground-plane; and the axis of its orbit might be taken for a line of departure in that plane or origin of angular reckoning. Were the axis *fixed*, this would be the best possible origin of longitudes; but as it has a motion (though an excessively slow one), there is, in fact, no advantage in reckoning from the axis more than from the line of the equinoxes, and astronomers therefore prefer the latter, taking account of its variation by the effect of precession, and restoring it, by calculation at every instant, to a fixed position. Now, to determine the situation of the ellipse described by a planet with respect to this plane, three *elements* require to be known: — 1st, the *inclination* of the plane of the planet's orbit to the plane of the ecliptic; 2dly, the line in which these two planes intersect each other, which of necessity passes through the sun, and whose position with respect to the line of the equinoxes is therefore given by stating its longitude. This line is called the line of the nodes. When the planet is in this line, in the act of passing from the south to the north side of the ecliptic, it is *in its ascending node*, and its longitude at that moment is the element called the *longitude of the node*. These two data determine the situation of *the plane* of the orbit; and there only remains, for the complete determination of the situation of the planet's ellipse, to know how it is placed *in* that plane, which (since its focus is necessarily in the sun) is ascertained by stating the *longitude of its perihelion*, or the place which the extremity of the axis nearest the sun occupies, when orthographically projected on the ecliptic.

(496.) The dimensions and situation of the planet's orbit thus determined, it only remains, for a complete acquaintance with its history, to determine the circumstances of its motion in the orbit so precisely fixed. Now, for this purpose, all that is needed is to know the moment of time when it is either at the perihelion, or at any other precisely determined point of its orbit, and its whole period; for these being known, the law of the areas determines the place at every other instant. This moment is called (when the perihelion is the point chosen) the *perihelion passage*, or, when some point of the orbit is fixed upon, without special reference to the perihelion, the *epoch*.

(497.) Thus, then, we have *seven* particulars or elements, which must be numerically stated, before we can reduce to calculation the state of the system at any given moment. But, these known, it is easy to ascertain the apparent positions of each planet, as it would be seen from the sun, or is seen from the earth at any time. The former is called the *heliocentric*, the latter the *geocentric*, place of the planet.

(498.) To commence with the heliocentric places. Let S represent the sun; P A N the orbit of the planet, being an ellipse, having the S in its focus, and A for its perihelion; and let *p a* N ♈ represent the projection of the orbit on the plane of the ecliptic, intersecting the line of equinoxes

Fig. 71.



S ♈ in ♈, which, therefore, is the origin of longitudes. Then will S N be the line of nodes; and if we suppose B to lie on the south, and A on the north side of the ecliptic, and the direction of the planet's motion to be from B to A, N will be the ascending node, and the angle ♈ S N the *longitude of the node.* In like manner, if P be the place of the planet at any time, and if it and the perihelion A be projected on the ecliptic, upon the points *p*, *a*, the angles ♈ S *p*, ♈ S *a*, will be the respective heliocentric longitudes of the planet and of the perihelion, the former of which is to be determined, and the latter is one of the given elements. Lastly, the angle *p* S P is the heliocentric latitude of the planet, which is also required to be known.

(499.) Now, the time being given, and also the moment of the planet's passing the perihelion, the interval, or the time of describing the portion A P of the orbit, is given, and the periodical time, and the whole area of the ellipse being known, the law of proportionality of areas to the times of their description gives the magnitude of the area A S P. From this it is a problem of pure geometry to determine the corresponding *angle* A S P, which is called the planet's *true anomaly*. This problem is of the kind called transcendental, and has been resolved by a great variety of processes, some more, some less intricate. It offers, however, no peculiar difficulty, and is practically resolved with great facility by the help of tables constructed for the purpose, adapted to the case of each particular planet.[1]

(500.) The true anomaly thus obtained, the planet's angular distance from the node, or the angle N S P, is to be found. Now, the longitudes of the perihelion and node being respectively ♈ *a* and ♈ N, which are given, their difference *a* N is also given, and the angle N of the spherical right-angled triangle A N *a*, being the inclination of the plane of the orbit to the ecliptic, is known. Hence we calculate the arc N A, or the angle N S A, which, added to A S P, gives the angle N S P required. And from this, regarded as the measure of the arc N P, forming the hypothenuse of the right-angled spherical triangle P N *p*, whose angle N, as before, is known, it is easy to obtain the other two sides, N *p* and P *p*. The latter, being the measure of the angle *p* S P, expresses the planet's heliocentric latitude; the former measures the angle N S *p*, or the planet's distance in longitude from its node, which, added to the known angle ♈ S N, the longitude of the node, gives the heliocentric longitude. This process, however circuitous it may appear, when once well understood may be gone through numerically by the aid of the usual logarithmic and trigonometrical tables, in little more time than it will have taken the reader to peruse its description.

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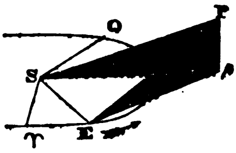
[1] It will readily be understood, that, except in the case of uniform circular motion, an equable description of *areas* about any centre is incompatible with an equable description of *angles*. The object of the problem in the text is to pass from the *area* supposed known, to the *angle*, supposed unknown: in other words, to derive the true amount of angular motion from the perihelion, or the *true anomaly* from what is technically called the mean anomaly, that is, the mean angular motion which would have been performed had the motion *in angle* been uniform instead of the motion *in area*. It happens fortunately, that this is the simplest of all problems of the transcendental kind, and can be resolved, in the most difficult case, by the rule of "false position," or trial and error, in a very few minutes. Nay, it may even be resolved instantly on inspection by a simple and easily constructed piece of mechanism, of which the reader may see a description in the Cambridge Philosophical Transactions, vol. iv. p. 425, by the author of this work.

(501.) The geocentric differs from the heliocentric place of a planet by reason of that parallactic change of apparent situation which arises from the earth's motion in its orbit. Were the planet's distances as vast as those of the stars, the earth's orbitual motion would be insensible when viewed from them, and they would always appear to us to hold the same relative situations among the fixed stars, as if viewed from the sun, *i. e.* they would then be seen in their *heliocentric* places. The difference, then, between the heliocentric and geocentric places of a planet is, in fact, the same thing with its *parallax*, arising from the earth's removal from the centre of the system and its annual motion. It follows from this, that the first step towards a knowledge of its amount, and the consequent determination of the apparent place of each planet, as referred from the earth to the sphere of the fixed stars, must be to ascertain the proportion of its linear distances from the earth and from the sun, as compared with the earth's distance from the sun, and the angular positions of all three with respect to each other.

(502.) Suppose, therefore, S to represent the sun, E the earth, and P the planet; S ♈ the line of equinoxes, ♈ E the earth's orbit, and P *p* a perpendicular let fall from the planet on the ecliptic. Then will the angle S P E (according to the general notion of parallax conveyed in art. 69) represent the parallax of the planet arising from the change of station from S to E; E P will be the apparent direction of the planet seen from E; and if S Q be drawn parallel to E *p*, the angle ♈ S Q will be the geocentric longitude of the planet, while ♈ S E represents the heliocentric longitude of the earth, ♈ S *p* that of the planet. The former of these, S E ♈, is given by the solar table; the latter, ♈ S *p*, is found by the process above described (art. 500). Moreover, S P is the radius vector of the planet's orbit, and S E that of the earth's, both of which are determined

Fig. 72.



from the known dimensions of their respective ellipses, and the places of the bodies in them at the assigned time. Lastly, the angle P S p is the planet's heliocentric latitude.

(503.) Our object, then, is, from all these data, to determine the angle ♈ S Q, and P E p, which is the geocentric latitude. The process, then, will stand as follows:—1st, In the triangle S P p, right-angled at p, given S P, and the angle P S p (the planet's radius vector and heliocentric latitude), find S p and P p; 2dly, In the triangle S E p, given S p (just found), S E (the earth's radius vector), and the angle E S p (the difference of heliocentric longitudes of the earth and planet), find the angle S p E, and the side E p. The former being equal to the alternate angle p S Q, is the parallactic removal of the planet in longitude, which, added to ♈ S p, gives its geocentric longitude. The latter, E p (which is called the *curtate distance* of the planet from the earth), gives at once the geocentric latitude, by means of the right-angled triangle P E p, of which E p and P p are known sides, and the angle P E p is the geocentric latitude sought.

(504.) The calculations required for these purposes are nothing but the most ordinary processes of plane trigonometry; and, though somewhat tedious, are neither intricate nor difficult. When executed, however, they afford us the means of comparing the places of the planets actually observed with the elliptic theory, with the utmost exactness, and thus putting it to the severest trial; and it is upon the testimony of such computations, so brought into comparison with observed facts, that we declare that theory to be a true representation of nature.

(505.) The planets Mercury, Venus, Mars, Jupiter, and Saturn, have been known from the earliest ages in which astronomy has been cultivated. Uranus was discovered by Sir W. Herschel in 1781, March 13th, in the course of a review of the heavens, in which every star visible in a telescope of a certain power was brought under close examination, when the new planet was immediately detected by its disc, under a high magnifying power. It has since been ascertained to have been observed on many previous occasions, with telescopes of insufficient power to show its disc, and even entered in catalogues as a star; and some of the observations which have been so recorded have been used to improve and extend our knowledge of its orbit. The discovery of the ultra-zodiacal planets dates from the first day of 1801, when Ceres was discovered by Piazzi, at Palermo; a discovery speedily followed by those of Juno by professor Harding, of Göttingen, in 1804; and of Pallas and Vesta, by Dr. Olbers, of Bremen, in 1802 and 1807 respectively. It is extremely remarkable that this important addition to our system had been in some sort surmised as a thing not unlikely, on the ground that the interval between the orbit of Mercury and the other planetary orbits, go on doubling as we recede from the sun, or nearly so. Thus, the interval between the orbits of the Earth and Mercury is nearly twice that between those of Venus and Mercury; that between the orbits of Mars and Mercury nearly twice that between the Earth and Mercury: and so on. The interval between the orbits of Jupiter and Mercury, however, is much too great, and would form an exception to this law, which is, however, again resumed in the case of the three planets next in order of remoteness, Jupiter, Saturn, and Uranus. It was therefore thrown out, by the late professor Bode, of Berlin,[1] as a possible surmise, that a planet not then yet discovered might exist between Mars and Jupiter: and it may easily be imagined what was the astonishment of astronomers on finding not only one, but four planets, differing greatly in all the other elements of their orbits, but agreeing very nearly, both *inter se*, and with the above stated empirical law, in respect of their mean distances from the sun. No account, *à priori* or from theory, was to be given of this singular progression, which is not, like Kepler's laws, strictly exact in numerical verification; but the circumstances we have just mentioned tended to create a strong belief that it was something beyond a mere accidental coincidence, and bore reference to the essential structure of the planetary system. It was even conjectured that the ultra-zodiacal planets are fragments of some greater planet which formerly circulated in that interval, but which has been blown to atoms by an explosion; an idea countenanced by the exceeding minuteness of these bodies which present discs; and it was argued that in that case innumerable more such fragments must exist and might come to be hereafter discovered. Whatever may be thought of such a speculation as a physical hypothesis, this conclusion has been verified to a considerable extent as a matter of fact by subsequent discovery, the result of a careful and minute examination and mapping down of the smaller stars in and near the zodiac, undertaken with that express object. Zodiacal charts of this kind, the product of the zeal and industry of many astronomers, have been constructed, in which every star down to the ninth or tenth magnitude is inserted, and these stars being compared with the actual stars of the heavens, the intrusion of any stranger within their limits cannot fail to be noticed when the comparison is systematically conducted. The discovery of Astræa, and that of Hebe by Professor Hencke, date respectively from December 8th, 1845, and July 1st, 1847; those of Iris and

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[1] The empirical law itself, as we have above stated it, is ascribed by Voiron, not to Bode (who would appear, however, at all events, to have first drawn attention to this interpretation of its interruption), but to Professor Titius of Wittemberg. (Voiron, Supplement to Bailly.)

Flora, by Mr. Hind, from August 13th and October 18th, 1847; of Metis, by Mr. Graham, from April 25, 1848; and of Hygeia, by M. De Gasparis, April 12th, 1849.

(506.) The discovery of Neptune marks in a signal manner the maturity of astronomical science. The proof, or at least the urgent presumption of the existence of such a planet, as a means of accounting (by its attraction) for certain small irregularities observed in the motions of Uranus, was afforded almost simultaneously by the independent researches of two geometers, Messrs. Adams of Cambridge and Leverrier of Paris, who were enabled, *from theory alone*, to calculate whereabouts it ought to appear in the heavens, *if visible*, the places thus independently calculated agreeing surprisingly. *Within a single degree* of the place assigned by M. Leverrier's calculations, and by him communicated to Dr. Galle of the Royal Observatory at Berlin, it was actually found by that astronomer on the very first night after the receipt of that communication, on turning a telescope on the spot, and comparing the stars in its immediate neighbourhood with those previously laid down in one of the zodiacal charts already alluded to.[1] This remarkable verification of an indication so extraordinary took place on the 23d of September, 1846.[2]

(507.) The mean distance of Neptune from the sun, however, so far from falling in with the supposed law of planetary distances above mentioned, offers a decided case of discordance. The interval between its orbit and that of Mercury, instead of being nearly double the interval between those of Uranus and Mercury, does not, in fact, exceed the latter interval by much more than half its amount. This remarkable exception may serve to make us cautious in the too ready admission of empirical laws of this nature to the rank of fundamental truths, though, as in the present instance, they may prove useful auxiliaries, and serve as stepping stones, affording a temporary footing in the path to great discoveries. The force of this remark will be more apparent when we come to explain more particularly the nature of the theoretical views which led to the discovery of Neptune itself.

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[1] Constructed by Dr. Bremiker, of Berlin. On reading the history of this noble discovery, we are ready to exclaim with Schiller—

> "Mit dem Genius steht die Natur in ewigem Bunde,
> Was der Eine verspricht liestet die Andre gewiss."

[2] Professor Challis, of the Cambridge Observatory, directing the Northumberland telescope of that Institution to the place assigned by Mr. Adams's calculations and its vicinity, on the 4th and 12th of August 1846, saw the planet on both those days, and noted its place (among those of other stars) for re-observation. He, however, postponed the *comparison* of the places observed, and, not possessing Dr. Bremiker's chart (which would have at once indicated the presence of an unmapped star,) remained in ignorance of the planet's existence as a visible object till its announcement as such by Dr. Galle.

(508.) We shall devote the rest of this chapter to an account of the physical peculiarities and probable condition of the several planets, so far as the former are known by observation, or the latter rest on probable grounds of conjecture. In this, three features principally strike us as necessarily productive of extraordinary diversity in the provisions by which, if they be, like our earth, inhabited, animal life must be supported. These are, first, the difference in their respective supplies of light and heat from the sun; secondly, the difference in the intensities of the gravitating forces which must subsist at their surfaces, or the different ratios which, on their several globes, the *inertiæ* of bodies must bear to their *weights*; and, thirdly, the difference in the nature of the materials of which, from what we know of their mean density, we have every reason to believe they consist. The intensity of solar radiation is nearly seven times greater on Mercury than on the Earth, and on Uranus 330 times less; the proportion between the two extremes being that of upwards of 2000 to 1. Let any one figure to himself the condition of our globe, were the sun to be septupled, to say nothing of the greater ratio! or were it diminished to a seventh, or to a 300th of its actual power! Again, the intensity of gravity, or its efficacy in counteracting muscular power and repressing animal activity, on Jupiter, is nearly two and a half times that on the earth, on Mars not more than one-half, on the Moon one-sixth, and on the smaller planets probably not more than one-twentieth; giving a scale of which the extremes are in the proportion of sixty to one. Lastly, the density of Saturn hardly exceeds one-eighth of the mean density of the Earth, so that it must consist of materials not much heavier than cork. Now, under the various combinations of elements so important to life as these, what immense diversity must we not admit in the conditions of that great problem, the maintenance of animal and intellectual existence and happiness, which seems, so far as we can judge by what we see around us in our own planet, and by the way in which every corner of it is crowded with living beings, to form an unceasing and worthy object for the exercise of the Benevolence and Wisdom which preside over all!

(509.) Quitting, however, the region of mere speculation, we will now show what information the telescope affords us of the actual condition of the several planets within its reach. Of Mercury we can see little more than that it is round, and exhibits phases. It is too small, and too much lost in the constant neighbourhood of the Sun, to allow us to make out more of its nature. The real diameter of Mercury is about 3200 miles: its apparent diameter varies from 5″ to 12″. Nor does Venus offer any remarkable peculiarities: although its real diameter is 7800 miles, and although it occasionally attains the considerable apparent diameter of 61″, which is larger than that of any other planet, it is yet the most difficult of them all to define with telescopes. The intense lustre of its illuminated part dazzles the sight, and exaggerates every imperfection of the telescope; yet we see clearly that its surface is not mottled over with permanent spots like the Moon; we notice in it neither mountains nor shadows, but a uniform brightness, in which sometimes we may indeed fancy, or perhaps more than fancy, brighter or obscurer portions, but can seldom or never rest fully satisfied of the fact. It is from some observations of this kind that both Venus and Mercury have been concluded to revolve on their axes in about the same time as the Earth, though in the case of Venus, Bianchini and other more recent observers have contended for a period of twenty-four times that length. The most natural conclusion, from the very rare appearance and want of permanence in the spots, is, that we do not see, as in the Moon, the real surface of these planets, but only their atmospheres, much loaded with clouds, and which may serve to mitigate the otherwise intense glare of their sunshine.

(510.) The case is very different with Mars. In this planet we frequently discern, with perfect distinctness, the outlines of what may be continents and seas. (See Plate I. *fig*. 1., which represents Mars in its gibbous state, as seen on the 16th of August, 1830, in the 20-feet reflector at Slough.) Of these, the former are distinguished by that ruddy colour which characterizes the light of this planet (which always appears red and fiery), and indicates, no doubt, an ochrey tinge in the general soil, like what the red sandstone districts on the Earth may possibly offer to the inhabitants of Mars, only more decided. Contrasted with this (by a general law in optics), the seas, as we may call them, appear greenish.[1] These spots, however, are not always to be seen equally distinct, but, *when seen*, they offer the appearance of forms considerably definite and highly characteristic,[2] brought successively into view by the rotation of the planet, from the assiduous observation of which it has even been found practicable to construct a rude chart of the surface of the planet. The variety in the spots may arise from the planet not being destitute of atmosphere and clouds; and what adds greatly to the probability of this is the appearance of brilliant white spots at its poles, — one of which appears in our figure, — which have been conjectured, with some probability, to be snow; as they disappear when they have been long exposed to the sun, and are greatest when just emerging from the long night of their polar winter, the snow line then extending to about six degrees (reckoned on a meridian of the planet) from the pole. By watching the spots during a whole night, and on successive nights, it is found that Mars has a rotation on an axis inclined about 30° 18′ to the ecliptic, and in a period of $24^h\ 37^m\ 23^s$[1] in the same direction as the Earth's, or from west to east. The greatest and least apparent diameters of Mars are 4″ and 18″, and its real diameter about 4100 miles.

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[1] I have noticed the phænomena described in the text on many occasions, but never more distinct than on the occasion when the drawing was made from which the figure in Plate I. is engraved. — *Author*.

[2] The reader will find many of these forms represented in Schumacher's *Astronomische Nachrichten*, No. 191, 434, and in the chart in No. 349, by Messrs. Beer and Mädler.

(511.) We now come to a much more magnificent planet, Jupiter, the largest of them all, being in diameter no less than 87,000 miles, and in bulk exceeding that of the Earth nearly 1300 times. It is, moreover, dignified by the attendance of four *moons*, *satellites*, or *secondary planets*, as they are called, which constantly accompany and revolve about it, as the Moon does round the Earth, and in the same direction, forming with their principal, or *primary*, a beautiful miniature system, entirely analogous to that greater one of which their central body is itself a member, obeying the same laws, and exemplifying, in the most striking and instructive manner, the prevalence of the gravitating power as the ruling principle of their motions: of these, however, we shall speak more at large in the next chapter.

(512.) The disc of Jupiter is always observed to be crossed in one certain direction by dark bands or belts presenting the appearance, in Plate III. *fig.* 2., which represents this planet as seen on the 23d of September, 1832, in the 20-feet reflector at Slough. These belts are, however, by no means alike at all times; they vary in breadth and in situation on the disc (though never in their general direction). They have even been seen broken up and distributed over the whole face of the planet; but this phænomenon is extremely rare. Branches running out from them, and subdivisions, as represented in the figure, as well as evident dark spots, are by no means uncommon; and from these, attentively watched, it is concluded that this planet revolves in the surprisingly short period of $9^h\ 55^m\ 50^s$ (sid. time), on an axis perpendicular to the direction of the belts. Now, it is very remarkable, and forms a most satisfactory comment on the reasoning by which the spheroidal figure of the Earth has been deduced from its diurnal rotation, that the outline of Jupiter's disc is evidently not circular, but elliptic, being considerably flattened in the direction of its axis of rotation. This appearance is no optical illusion, but is authenticated by micrometrical measures, which assign 107 to 100 for the proportion of the equatorial and polar diameters. And to confirm in the strongest manner, the truth of those principles on which our former conclusions have been founded, and fully to authorize their extension to this remote system, it appears, on calculation, that this is really the degree of oblateness which corresponds, on those principles, to the dimensions of Jupiter, and to the time of his rotation.

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[1] Beer and Mädler, *Astr. Nachr.* 349.

(513.) The parallelism of the belts to the equator of Jupiter, their occasional variations, and the appearances of spots seen upon them, render it extremely probable that they subsist in the atmosphere of the planet, forming tracts of comparatively clear sky, determined by currents analogous to our trade-winds, but of a much more steady and decided character, as might indeed be expected from the immense velocity of its rotation. That it is the comparatively darker body of the planet which appears in the belts is evident from this,—that they do not come up in all their strength to the edge of the disc, but fade away gradually before they reach it. (See Plate III. *fig.* 2.) The apparent diameter of Jupiter varies from 30″ to 46″.

(514.) A still more wonderful, and, as it may be termed, elaborately artificial mechanism, is displayed in Saturn, the next in order of remoteness to Jupiter, to which it is not much inferior in magnitude, being about 79,000 miles in diameter, nearly 1000 times exceeding the earth in bulk, and subtending an apparent angular diameter at the earth, of about 18″ at its mean distance. This stupendous globe, besides being attended by no less than seven satellites, or moons, is surrounded by two broad, flat, extremely thin rings, concentric with the planet and with each other; both lying in one plane, and separated by a very narrow interval from each other throughout their whole circumference, as they are from the planet by a much wider. The dimensions of this extraordinary appendage are as follows[1]:—

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[1] These dimensions are calculated from Prof. Struve's micrometric measures, Mem. Art. Soc. iii. 301, with the exception of the thickness of the ring, which is concluded from its total disappearance in 1833, in a telescope which would certainly have shown, as a visible object, a line of light one-twentieth of a second in breadth. The interval of the rings here stated is possibly somewhat too small.

	″	Miles.
Exterior diameter of exterior ring	40·095	=176,418
Interior ditto	35·289	=155,272
Exterior diameter of interior ring	34·475	=151,690
Interior ditto	26·668	=117,339
Equatorial diameter of the body	17·991	= 79,160
Interval between the planet and interior ring	4·339	= 19,090
Interval of the rings	0·408	= 1,791
Thickness of the rings not exceeding		= 250

The figure (*fig.* 3, Plate III.) represents Saturn surrounded by its rings, and having its body striped with dark belts, somewhat similar, but broader and less strongly marked than those of Jupiter, and owing, doubtless, to a similar cause.[1] That the ring is a solid opake substance is shown by its throwing its shadow on the body of the planet, on the side nearest the sun, and on the other side receiving that of the body, as shown in the figure. From the parallelism of the belts with the plane of the ring it may be conjectured that the axis of rotation of the planet is perpendicular to that plane; and this conjecture is confirmed by the occasional appearance of extensive dusky spots on its surface, which when watched, like the spots on Mars or Jupiter, indicate a rotation in $10^h\ 29^m\ 17^s$ about an axis so situated.

(515.) The axis of rotation, like that of the earth, preserves its parallelism to itself during the motion of the planet in its orbit; and the same is also the case with the ring, whose plane is constantly inclined at the same, or very nearly the same, angle to that of the orbit, and, therefore, to the ecliptic, viz. 28° 11′; and intersects the latter plane in a line, which makes at present[2] an angle with the line of equinoxes of 167° 31′. So that the *nodes of the ring* lie in 167° 31′ and 347° 31′ of longitude. Whenever, then, the planet happens to be situated in one or other of these longitudes, as at C, the plane of the ring passes through the sun, which then illuminates only the edge of it. And if the earth at that moment be in F, it will see the ring edgeways, the planet being in opposition, and therefore most favourably situated (*cæteris paribus*) for observation. Under these circumstances the ring, if seen at all, can only appear as a very narrow straight line of light projecting on either side of the body as a prolongation of its diameter. In fact, it is quite invisible in any but telescopes of extraordinary power.[3] This remarkable phenomenon takes

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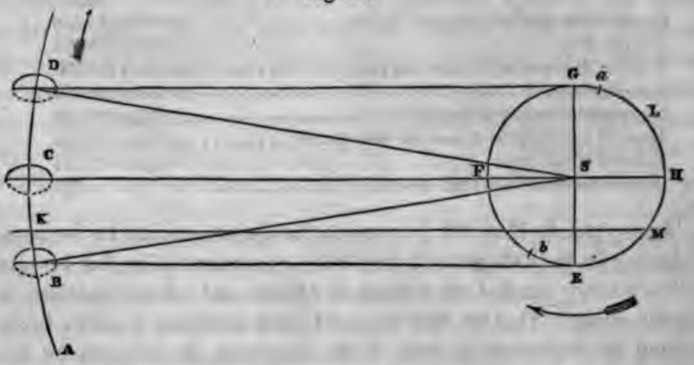
[1] The equatorial bright belt is generally well seen. The subdivision of the dark one by two narrow bright bands is seldom so distinct as represented in the plate.

[2] According to Bessel, the longitude of the node of the ring increases 46″·462 per annum. In 1800 it was 166° 53′ 8″·9.

[3] Its disappearance was *complete* when observed with a reflector eighteen inches in aperture and twenty feet in focal length, on the 29th of April, 1833, by the author.

Fig. 73.



place at intervals of fifteen years nearly (being a semi-period of Saturn in its orbit). One disappearance at least must take place whenever Saturn passes either node of its orbit; but three must frequently happen, and two are possible. To show this, suppose S to be the sun, A B C D part of Saturn's orbit situated so as to include the node of the ring (at C); E F G H the earth's orbit: S C the line of the node; E B, G D parallel to S C touching the earth's orbit in E G; and let the direction of motion of both bodies be that indicated by the arrow. Then since the ring preserves its parallelism, its plane can nowhere intersect the earth's orbit, and therefore no disappearance can take place, unless the planet be between B and D: and, on the other hand, a disappearance is possible (if the earth be rightly situated) during the whole time of the description of the arc B D. Now, since S B or S D, the distance of Saturn from the Sun, is to S E or S G, that of the Earth, as 9·54 to 1, the angle C S D or C S B = 6° 1′, and the whole angle B S D = 12° 2′, which is described by Saturn (on an average) in 359·46 days, wanting only 5·8 days of a complete year. The Earth then describes very nearly an entire revolution within the limits of time when a disappearance is possible; and since, in either half of its orbit E F G or G H F, it may equally encounter the plane of the ring, one such encounter at least is unavoidable within the time specified.

(516.) Let G *a* be the arc of the Earth's orbit described from G in 5·8 days. Then if, at the moment of Saturn's arrival at B, the Earth be at *a*, it will encounter the plane of the ring advancing parallel to itself and to B E to meet it, somewhere in the quadrant H E, as at M, after which it will be behind that plane (with reference to the direction of Saturn's motion) through all the arc M E F G up to G, where it will again overtake it at the very moment of the planet quitting the arc B D. In this state of things there will be two disappearances. If, when Saturn is at B, the Earth be anywhere in the arc *a* H E, it is equally evident that it will *meet* and pass through the advancing plane of the ring somewhere in the quadrant H E, that it will again *overtake* and pass through it somewhere in the semicircle E F G, and again meet it in some point of the quadrant G H, so that three disappearances will take place. So, also, if the Earth be at E when Saturn is at B, the motion of the Earth being at that instant directly towards B, the plane of the ring will for a short time leave it behind; but the ground so lost being rapidly regained as the earth's motion becomes oblique to the line of junction, it will soon overtake and pass through the plane in the early part of the quadrant E F, and passing on through G before Saturn arrives at D, will meet the plane again in the quadrant G H. The same will continue up to a certain point *b*, at which, if the Earth be initially situated, there will be but two disappearances — the plane of the ring there overtaking the Earth for an instant, and being immediately again left behind by it, to be again encountered by it in G H. Finally, if the initial place of the Earth (when Saturn is at B) be in the arc *b* F *a*, there will be but one passage through the plane of the ring, viz., in the semicircle G H E, the Earth being in advance of that plane throughout the whole of *b* G.

(517.) The appearances will moreover be varied according as the Earth passes from the enlightened to the unenlightened side of the ring, or *vice versâ*. If C be the ascending node of the ring, and if the under side of the paper be supposed south and the upper north of the ecliptic, then, when the Earth *meets* the plane of the ring in the quadrant H E, it passes from the bright to the dark side: where it *overtakes* it in the quadrant E F, the contrary. *Vice versâ*, when it overtakes it in F G, the transition is from the bright to the dark side, and the contrary where it meets it in G H. On the other hand, when the Earth is overtaken by the ring-plane in the interval E *b*, the change is from the bright to the dark side. When the dark side is exposed to sight, the aspect of the planet is very singular. It appears as a bright round disc, with its belts, &c., but crossed equatorially by a narrow and perfect black line. This can never of course happen when the planet is more than 6° 1′ from the node of the ring. Generally, the northern side is enlightened and visible when the heliocentric longitude of Saturn is between 173° 32′ and 341° 30′, and the southern when between 353° 32′ and 161° 30′. The greatest opening of the ring occurs when the planet is situated at 90° distance from the node of the ring, or in longitudes 77° 31′ and 257° 31′, and at these points the longer diameter of its apparent ellipse is almost exactly double the shorter.

(518.) It will naturally be asked how so stupendous an arch, if composed of solid and ponderous materials, can be sustained without collapsing and falling in upon the planet? The answer to this is to be found in a swift rotation of the ring in its own plane, which observation has detected, owing to some portion of the ring being a little less bright than others, and assigned its period at $10^h\ 32^m\ 15^s$, which, from what we know of its dimensions, and of the force of gravity in the Saturnian system, is very nearly the periodic time of a satellite revolving at the same distance as the middle of its breadth. It is the centrifugal force, then, arising from this rotation, which sustains it; and, although no observation nice enough to exhibit a difference of periods between the outer and inner rings have hitherto been made, it is more than probable that such a difference does subsist as to place each independently of the other in a similar state of equilibrium.

(519.) Although the rings are, as we have said, very nearly concentric with the body of Saturn, yet micrometrical measurements of extreme delicacy have demonstrated that the coincident is not mathematically exact, but that the centre of gravity of the rings oscillates round that of the body describing a very minute orbit, probably under laws of much complexity. Trifling as this remark may appear, it is of the utmost importance to the stability of the system of the rings. Supposing them mathematically perfect in their circular form, and exactly concentric with the planet, it is demonstrable that they would form a system in a state of *unstable equilibrium*, which the slightest external power would subvert — not by causing a rupture in the substance of the rings — but by precipitating them, *unbroken*, on the surface of the planet. For the attraction of such a ring or rings on a point or sphere excentrically within them, is not the same in all directions, but tends to draw the point or sphere towards the nearest part of the ring, or away from the centre. Hence, supposing the body to become, from any cause, ever so little excentric to the ring, the tendency of their mutual gravity is not to correct but to increase this excentricity, and to bring the nearest parts of them together. Now, external powers, capable of producing such excentricity, exist in the attractions of the satellites, as will be shown in Chap. XII.; and in order that the system may be *stable*, and possess within itself a power of resisting the first inroads of such a tendency, while yet nascent and feeble, and opposing them by an opposite or maintaining power, it has been shown that it is sufficient to admit the rings to be *loaded* in some part of their circumference, either by some minute inequality of thickness, or by some portions being denser than others. Such a load would give to the whole ring to which it was attached somewhat of the character of a heavy and sluggish satellite maintaining itself in an orbit with a certain energy sufficient to overcome minute causes of disturbance, and establish an average bearing on its centre. But even without supposing the existence of any such load,—of which, after all, we have no proof,—and granting, in its full extent, the general instability of the equilibrium, we think we perceive, in the rapid periodicity of all the causes of disturbance, a sufficient guarantee of its preservation. However homely be the illustration, we can conceive nothing more apt, in every way, to give a general conception of this maintenance of equilibrium under a constant tendency to subversion, than the mode in which a practised hand will sustain a long pole in a perpendicular position resting on the finger by a continual and almost imperceptible variation of the point of support. Be that, however, as it may, the observed oscillation of the centres of the rings about that of the planet is in itself the evidence of a perpetual contest between conservative and destructive powers — both extremely feeble, but so antagonizing one another as to prevent the latter from ever acquiring an uncontrollable ascendancy, and rushing to a catastrophe.

(520.) This is also the place to observe, that as the smallest difference of velocity between the body and the rings must infallibly precipitate the latter on the former, never more to separate, (for they would, once in contact, have attained a position of *stable equilibrium*, and be held together ever after by an immense force;) it follows, either that their motions in their common orbit round the sun must have been adjusted to each other by an external power, with the minutest precision, or that the rings must have been formed about the planet while subject to their common orbitual motion, and under the full and free influence of all the acting forces.

(521.) Several astronomers have suspected, and even consider themselves to have certainly observed, the rings of Saturn to be occasionally, at least, streaked with more or less numerous dark lines parallel to the decided black interval which separates the two rings, which latter being permanent, and seen equally and in the same part of the breadth on both sides of the ring, cannot be doubted to be a real separation.[1]

(522.) [The exterior ring of Saturn is described by many observers as rather less luminous than the interior, and the inner portion of this latter

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[1] The passage of Saturn across any considerable star would afford an admirable opportunity of testing the reality of such fissures, as it would flash in succession through them. The opportunity of watching for such occultations — when Saturn traverses the Milky-Way, for instance—should not be neglected.

observations, Mr. Dawes, at his observat
of an exquisite achromatic by Merz, of
very same fact, and even more distinctly,
darker interval between the old and new
latter into two of unequal degrees of obs
obscure than either.]

(523.) Of Uranus we see nothing but
nated disc, without rings, belts, or discern
meter is about 4″, from which it never var
ness of our orbit in comparison of its ow
35,000 miles, and its bulk 82 times that o
satellites—four at least, probably five or six
in the next chapter) offer remarkable peculi

(524.) The discovery of Neptune is so
ecliptic at present so little favourable for seei
that nothing very positive can be stated as t
two observers it has afforded strong suspicio
ring very highly inclined. And from the
M. Otto Struve, and Mr. Bond, it appears to
and very probably by two satellites—thoug
can hardly yet be considered as quite demon

(525.) If the immense distance of Ne
coming at much knowledge of its physical
ultra-zodiacal planets is no less a bar into a
of them, Pallas, has been said to have som
appearance

man placed on one of them would spring with ease 60 feet high, and sustain no greater shock in his descent than he does on the earth from leaping a yard. On such planets giants might exist; and those enormous animals, which on earth require the buoyant power of water to counteract their weight, might there be denizens of the land. But of such speculations there is no end.

(526.) We shall close this chapter with an illustration calculated to convey to the minds of our readers a general impression of the relative magnitudes and distances of the parts of our system. Choose any well levelled field or bowling-green. On it place a globe, two feet in diameter; this will represent the Sun; Mercury will be represented by a grain of mustard seed, on the circumference of a circle 164 feet in diameter for its orbit; Venus a pea, on a circle 284 feet in diameter; the Earth also a pea, on a circle of 430 feet; Mars a rather large pin's head, on a circle of 654 feet; Juno, Ceres, Vesta, and Pallas, grains of sand, in orbits of from 1000 to 1200 feet; Jupiter a moderate-sized orange, in a circle nearly half a mile across, Saturn a small orange, on a circle of four-fifths of a mile; Uranus a full-sized cherry, or small plum, upon the circumference of a circle more than a mile and a half, and Neptune a good-sized plum on a circle about two miles and a half in diameter. As to getting correct notions on this subject by drawing circles on paper, or, still worse, from those very childish toys called orreries, it is out of the question. To imitate the motions of the planets, in the above-mentioned orbits, Mercury must describe its own diameter in 41 seconds; Venus in $4^m\ 14^s$; the Earth, in 7 minutes; Mars, in $4^m\ 48^s$; Jupiter, $2^h\ 56^m$; Saturn, in $3^h\ 13^m$; Uranus, in $2^h\ 16^m$; and Neptune in $3^h\ 30^m$.

CHAPTER X.

OF THE SATELLITES.

OON, AS A SATELLITE OF THE EARTH. — GENERAL PROXIMITY
ELLITES TO THEIR PRIMARIES, AND CONSEQUENT SUBORDINA-
F THEIR MOTIONS. — MASSES OF THE PRIMARIES CONCLUDED
HE PERIODS OF THEIR SATELLITES. — MAINTENANCE OF KEP-
LAWS IN THE SECONDARY SYSTEMS. — OF JUPITER'S SATEL-
— THEIR ECLIPSES, ETC. — VELOCITY OF LIGHT DISCOVERED
EIR MEANS. — SATELLITES OF SATURN — OF URANUS — OF
E.

IN the annual circuit of the earth about the sun, it is constantly
by its satellite, the moon, which revolves round it, or rather
d their common centre of gravity; while this centre, strictly
and not either of the two bodies thus connected, moves in an

rent monthly displacement of the sun in longitude, of a parallactic kind, which is called the *menstrual equation;* whose greatest amount is, however, less than the sun's horizontal parallax, or than 8·6″.

(529.) The moon, as we have seen, is about 60 radii of the earth distant from the centre of the latter. Its proximity, therefore, to its centre of attraction, thus estimated, is much greater than that of the planets to the sun; of which Mercury, the nearest, is 84, and Uranus 2026 solar radii from *its* centre. It is owing to this proximity that the moon remains attached to the earth as a satellite. Were it much farther, the feebleness of its gravity towards the earth would be inadequate to produce that alternate acceleration and retardation in its motion about the sun, which divests it of the character of an independent planet, and keeps its movements subordinate to those of the earth. The one would outrun, or be left behind the other, in their revolutions round the sun (by reason of Kepler's third law,) according to the relative dimensions of their heliocentric orbits, after which the whole influence of the earth would be confined to producing some considerable periodical disturbance in the moon's motion, as it passed or was passed by it in each synodical revolution.

(530.) At the distance at which the moon really is from us, its gravity towards the earth is actually less than towards the sun. That this is the case, appears sufficiently from what we have already stated, that the moon's *real* path, even when between the earth and sun, is *concave towards the latter.* But it will appear still more clearly if, from the known periodic times[1] in which the earth completes its annual and the moon its monthly orbit, and from the dimensions of those orbits, we calculate the amount of deflection, in either, from their tangents, in equal very minute portions of time, as one second. These are the versed sines of the arcs described in that time in the two orbits, and these are the measures of the acting forces which produce those deflections. If we execute the numerical calculation in the case before us, we shall find 2.233 : 1 for the proportion in which the intensity of the force which retains the earth in its orbit round the sun actually exceeds that by which the moon is retained in *its* orbit about the earth.

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[1] R and *r* radii of two orbits (supposed circular,) P and *p* the periodic times; then the arcs in question (A and *a*) are to each other as $\frac{R}{P}$ to $\frac{r}{p^2}$; and since the versed sines are as the squares of the arcs directly and the radii inversely, these are to each other as $\frac{R}{p^2}$ to $\frac{r}{p^2}$; and in this ratio are the forces acting on the revolving bodies in either case.

Now the sun is 399 times more remote from the earth than the
And, as gravity increases as the squares of the distances de-
ust follow that at *equal* distances, the intensity of solar would
of terrestrial gravity in the above proportion, augmented in
ratio of the square of 400 to 1; that is, in the proportion of
1; and therefore, if we grant that the intensity of the gravi-
gy is commensurate with the mass or inertia of the attracting
re compelled to admit the mass of the earth to be no more
00 of that of the sun.

The argument is, in fact, nothing more than a recapitulation of
been adduced in Chap. VIII. (art. 448.) But it is here re-
, in order to show how the mass of a planet which is attended
more satellites can be as it were weighed against the sun, pro-
have learned from observation the dimensions of the orbits
by the planet about the sun, and by the satellites about the
d also the periods in which these orbits are respectively
It is by this method that the masses of Jupiter, Saturn,
d Neptune have been ascertained. (See Synoptic Table.)

Jupiter, as already stated, is attended by four satellites, Saturn
Uranus, certainly by four, and perhaps by six; and Neptune
more. These, with their respective *primaries* (as the central

with attention to all its details, is that of Jupiter; partly on account of the conspicuous brilliancy of its four attendants, which are large enough to offer visible and measurable discs in telescopes of great power; but more for the sake of their eclipses, which, as they happen very frequently, and are easily observed, afford signals of considerable use for the determination of terrestrial longitudes (art. 266). This method, indeed, until thrown into the back ground by the greater facility and exactness now attainable by lunar observations (art. 267) was the best, or rather the only one, which could be relied on for great distances and long intervals.

(535.) The satellites of Jupiter revolve from west to east (following the analogy of the planets and moon,) in planes very nearly, although not exactly, coincident with that of the equator of the planet, or parallel to its belts. This latter plane is inclined 3° 5′ 30″ to the orbit of the planet, and is therefore but little different from the plane of the ecliptic. Accordingly, we see their orbits projected very nearly into straight lines, in which they appear to oscillate to and fro, sometimes passing before Jupiter, and casting shadows on his disc, (which are very visible in good telescopes, like small round ink spots, the circular form of which is very evident,) and sometimes disappearing behind the body, or being eclipsed in its shadow at a distance from it. It is by these eclipses that we are furnished with accurate data for the construction of tables of the satellites' motions, as well as with signals for determining differences of longitude.

(536.) The eclipses of the satellites, in their general conception, are perfectly analogous to those of the moon, but in their detail they differ in several particulars. Owing to the much greater distance of Jupiter from the sun, and its greater magnitude, the cone of its shadow or umbra (art. 420) is greatly more elongated, and of far greater dimension, than that of the earth. The satellites are, moreover, much less in proportion to their primary, their orbits less inclined to *its* ecliptic, and (comparatively to the diameter of the planet) of smaller dimensions, than is the case with the moon. Owing to these causes, the three interior satellites of Jupiter pass through the shadow, and are totally eclipsed, every revolution; and the fourth, though, from the greater inclination of its orbit, it sometimes escapes eclipse, and *may* occasionally graze as it were the border of the shadow, and suffer partial eclipse, yet does so comparatively seldom, and, ordinarily speaking, its eclipses happen, like those of the rest, each revolution.

(537.) These eclipses, moreover, are not seen, as is the case with those of the moon, from the centre of their motion, but from a remote station, and one whose situation with respect to the line of shadow is variable.

urse, makes no difference in the *times* of the eclipses, but a
one in their visibility, and in their apparent situations with
the planet at the moments of their entering and quitting the

Suppose S to be the sun, E the earth in its orbit E F G K, J
d *a b* the orbit of one of its satellites. The cone of the shadow,
have its vertex at X, a point far beyond the orbits of all the
and the penumbra, owing to the great distance of the sun, and
uent smallness of the angle (about 6′ only) its disc subtends at
ill hardly extend, within the limits of the satellites' orbits, to
tible distance beyond the shadow, — for which reason it is not
d in the figure. A satellite revolving from west to east (in the
f the arrows) will be eclipsed when it enters the shadow at *a*,
ddenly, because, like the moon, it has a considerable diameter
the planet; so that the time elapsing from the first perceptible
ht to its total extinction will be that which it occupies in de-
out Jupiter an angle equal to its apparent diameter as seen from
of the planet, or rather somewhat more, by reason of the

Fig. 74.

particulars of the motions of the satellites, and for affording data of any material use for the calculation of terrestrial longitudes. The intervals of the eclipses, it will be observed, give the *synodic* periods of the satellites' revolutions; from which their sidereal periods must be concluded by the method in art. 418.

(539.) It is evident, from a mere inspection of our figure, that the eclipses take place to the west of the planet, when the earth is situated to the west of the line S J, *i. e.* before the opposition of Jupiter; and to the east, when in the other half of its orbit, or after the opposition. When the earth approaches the opposition, the visual line becomes more and more nearly coincident with the direction of the shadow, and the apparent place where the eclipses happen will be continually nearer and nearer to the body of the planet. When the earth comes to F, a point determined by drawing *b* F to touch the body of the planet, the *emersions* will cease to be visible, and will thenceforth, up to the time of the opposition, happen *behind* the disc of the planet. Similarly, from the opposition till the time when the earth arrives at I, a point determined by drawing *a* I tangent to the eastern limb of Jupiter, the *immersions* will be concealed from our view. When the earth arrives at G (or H) the immersion (or emersion) will happen at the very edge of the visible disc, and when between G and H (a very small space), the satellites will *pass uneclipsed behind the limb* of the planet.

(540.) Both the satellites and their shadows are frequently observed to *transit* or pass across the disc of the planet. When a satellite comes to *m*, *its* shadow will be thrown on Jupiter, and will appear to move across it as a black spot till the satellite comes to *n*. But the satellite itself will not appear to enter on the disc till it comes up to the line drawn from E to the eastern edge of the disc, and will not leave it till it attains a similar line drawn to the western edge. It appears then that the shadow will *precede* the satellite in its progress over the disc *before* the opposition of Jupiter, and *vice versâ*. In these transits of the satellites, which, with very powerful telescopes, may be observed with great precision, it frequently happens that the satellite itself is discernible *on* the disc as a bright spot if projected on a dark belt; but occasionally also as a dark spot of smaller dimensions than the shadow. This curious fact (observed by Schroeter and Harding) has led to a conclusion that certain of the satellites have occasionally *on their* own bodies, or in their atmospheres, obscure spots of great extent. We say of great extent; for the satellites of Jupiter, small as they appear to us, are really bodies of considerable size, as the following comparative table will show:[1]—

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[1] Struve, Mem. Ast. Soc. iii. 301.

	Mean apparent diameter as seen from the Earth.	Mean apparent diameter as seen from Jupiter.	Diameter in miles.	Mass.[1]
....	38″·327		87000	1·0000000
e	1·017	33′ 11″	2508	0·0000173
....	0·911	17 35	2068	0·0000232
....	1·488	18 0	3377	0·0000885
....	1·273	8 46	2890	0·0000427

ch it follows, that the first satellite appears to a spectator on
s large as our moon to us; the second and third nearly equal to
, and of somewhat more than half the apparent diameter of the
the fourth about one quarter of that diameter. So seen, they
ntly, of course, eclipse one another, and cause eclipses of the
latter visible, however, only over a very small portion of the
nd their motions and aspects with respect to each other must
petual variety and singular and pleasing interest to the inhabi-
eir primary.

Besides the eclipses and the transits of the satellites across the
may also disappear to us when not eclipsed, by passing *behind*
f the planet. Thus, when the earth is at E, the immersion of
e will be seen at a, and its emersion at b, both to the west of
, after which the satellite, still continuing its course in the di-

single eclipse, the immersion being concealed by the body, if the planet be past its opposition, the emersion if not yet arrived at it. So also of the occultation. The commencement of the occultation, or the passage of the satellite behind the disc, takes place while obscured by the shadow, before opposition, and its re-emergence after. All these particulars will be easily apparent on mere inspection of the figure (art. 536). It is only during the short time that the earth is in the arc G H (*i. e.* between the sun and Jupiter, that the cone of the shadow converging (while that of the visual rays diverges) behind the planet, permits their occultations to be completely observed both at ingress and egress, unobscured, the eclipses being then invisible.

(542.) An extremely singular relation subsists between the mean angular velocities or *mean motions* (as they are termed) of the three first satellites of Jupiter. If the mean angular velocity of the first satellite be added to twice that of the third, the sum will equal three times that of the second. From this relation it follows, that if from the mean longitude of the first added to twice that of the third, be subducted three times that of the second, the remainder will always be the same, or constant, and observation informs us that this constant is 180°, or two right angles; so that, the situations of any two of them being given, that of the third may be found. It has been attempted to account for this remarkable fact, on the theory of gravity by their mutual action; and Laplace has demonstrated, that if this relation were at any one epoch approximately true, the mutual attractions of the satellites would, in process of time, render it exactly so. One curious consequence is, that these three satellites cannot be all eclipsed at once; for, in consequence of the last-mentioned relation, when the second and third lie in the *same* direction from the centre, the first must lie on the *opposite;* and therefore, when at such a conjuncture the first is eclipsed, the other two must lie between the sun and planet, throwing its shadow on the disc, and *vice versâ.*

(543.) Although, however, for the above mentioned reason, the satellites cannot be all *eclipsed* at once, yet it may happen, and occasionally does so, that all are either eclipsed, occulted, or projected on the body, in which case they are, generally speaking, equally invisible, since it requires an excellent telescope to discern a satellite on the body, except in peculiar circumstances. Instances of the actual observations of Jupiter thus denuded of its usual attendance and offering the appearance of a solitary disc, though rare, have been more than once recorded. The first occasion in which this was noticed was by Molyneux, on November 2d, (old style)

similar observation is recorded by Sir W. Herschel as made
May 22d, 1802. The phænomenon has also been observed
allis, on April 15th, 1826; (in which case the deprivation con-
whole hours;) and lastly by Mr. H. Griesbach, on September
3.

The discovery of Jupiter's satellites, one of the first fruits of
ion of the telescope, and of Galileo's early and happy idea of
its new-found powers to the examination of the heavens, forms
most memorable epochs in the history of astronomy. The
omical solution of the great problem of "*the longitude*"—prac-
most important for the interests of mankind which has ever
ght under the dominion of strict scientific principles, dates
ly from their discovery. The final and conclusive establish-
e Copernican system of astronomy may also be considered as
o the discovery and study of this exquisite miniature system,
the laws of the planetary motions, as ascertained by Kepler,
ially that which connects their periods and distances, were
aced, and found to be satisfactorily maintained. And (as if to
historical interest on this point) it is to the observation of
ses that we owe the grand discovery of the aberration of light,
nsequent determination of the enormous velocity of that won-
ment. This we must explain now at large.

lating on the probable physical cause, he was naturally led to think of a gradual instead of an instantaneous propagation of light. This explained every particular of the observed phenomenon, but the velocity required (192000 miles per second) was so great as to startle many, and, at all events, to require confirmation. This has been afforded since, and of the most unequivocal kind, by Bradley's discovery of the aberration of light (art. 329.) The velocity of light deduced from this last phænomenon differs by less than one-eightieth of its amount from that calculated from the eclipses, and even this difference will no doubt be destroyed by nicer and more rigorously reduced observations.

(546.) The orbits of Jupiter's satellites are but little excentric, those of the two interior, indeed, have no perceptible excentricity. Their mutual action produces in them perturbations analogous to those of the planets about the sun, and which have been diligently investigated by Laplace and others. By assiduous observation it has been ascertained that they are subject to marked fluctuations in respect of brightness, and that these fluctuations happen periodically, according to their position with respect to the sun. From this it has been concluded, apparently with reason, that they turn on their axes, like our moon, in periods equal to their respective sidereal revolutions about their primary.

(547.) The satellites of Saturn have been much less studied than those of Jupiter, being far more difficult to observe. The most distant has its orbit materially inclined (no less than 12° 14')[1] to the plane of the ring, with which the orbits of all the rest nearly coincide. Nor is this the only circumstance which separates it by a marked difference of character from the system of the six interior ones, and renders it in some sort an anomalous member of the Saturnian system. Its distance from the planet's centre exceeds in the proportion of nearly three to one that of the most distant of all the rest, being no less than 64 times the radius of the globe of Saturn, a distance from the primary to which our own moon (at 60 radii) offers the only parallel. Its variation of light also in different parts of its orbit is very much greater than the case of any other secondary planet. Dominic Cassini indeed (its first discoverer, A. D. 1671) found it to disappear for nearly half its revolution, when to the east of Saturn, and though the more powerful telescopes now in use enable us to follow it round the whole of its circuit, its diminution of light is so great in the eastern half of its orbit as to render it somewhat difficult to perceive. From this circumstance (viz. from the defalcation of light occurring constantly on the same side of Saturn *as seen from the earth*, the visual ray from which is never very oblique to the direction in which the sun's light falls on it) it

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[1] Lalande, Astron. Art. 3075.

d with much certainty that this satellite revolves on its axis in
ime of rotation about the primary; as we know to be the case
noon, and as there is considerable ground for believing to be so
condaries.

The next satellite in order proceeding inwards (the first in order
ry[1]) is by far the largest and most conspicuous of all, and is
ot much inferior to Mars in size. It is the only one of the
hose theory and perturbations have been at all inquired into[2]
n to verify Kepler's law of the periodic times, which holds
tis mutandis, and under the requisite reservations, in this, as
stem of Jupiter. The three next satellites still proceeding
re very minute, and require pretty powerful telescopes to see
ile the two interior satellites which just skirt the edge of the
only be seen with telescopes of extraordinary power and per-
d under the most favourable atmospheric circumstances. At
of their discovery they were seen to thread, like beads, the
initely thin fibre of light to which the ring then seen edge-
reduced, and for a short time to advance off it at either end,
return, and hastening to their habitual concealment behind

Owing to the obliquity of the ring and of the orbits of the
to Saturn's ecliptic, there are no eclipses, occultations, or

transits of these bodies or their shadows across the disc of their primary (the interior ones excepted), until near the time when the ring is seen edgewise, and when they do take place, their observation is attended with too much difficulty to be of any practical use, like the eclipses of Jupiter's satellites, for the determination of longitudes, for which reason they have been hitherto little attended to by astronomers.

(550.) A remarkable relation subsists between the periodic times of the two interior satellites of Saturn, and those of the two next in order of distance; viz. that the period of the third (Tethys) is double that of the first (Mimas), and that of the fourth (Dione) double that of the second (Enceladus). The coincidence is exact in either case to about one-800th part of the larger period.

(551.) The satellites of Uranus require very powerful and perfect telescopes for their observation. Two are, however, much more conspicuous than the rest, and their periods and distances from the planet have been ascertained with tolerable certainty. They are the second and fourth of those set down in the synoptic table. Of the remaining four, whose existence, though announced with considerable confidence by their original discoverer, could hardly be regarded as fully demonstrated, two only have been hitherto reobserved; viz. the first of our table, interior to the two larger ones, by the independent observations of Mr. Lassell,[1] and M. Otto Struve,[2] and the fourth, intermediate between the larger ones, by the former of these astronomers. The remaining two, if future observation should satisfactorily establish their real existence, will probably be found to revolve in orbits exterior to all these.

(552.) The orbits of these satellites offer remarkable, and, indeed, quite unexpected and unexampled peculiarities. Contrary to the unbroken analogy of the whole planetary system — whether of primaries or secondaries — the planes of their orbits *are nearly perpendicular to the ecliptic*, being inclined no less than 78° 58′ to that plane, and in these orbits their motions *are retrograde;* that is to say, their positions, when projected on the ecliptic, instead of advancing *from west to east* round the centre of their primary, as is the case with every other planet and satellite, move in the opposite direction. Their orbits are nearly or quite circular, and they do not appear to have any sensible, or, at least, any rapid motion of nodes, or to have undergone any material change of inclination, in the course, at least, of half a revolution of their primary round the sun. When the earth is in the plane of their orbits, or nearly so, their apparent paths are straight lines or very elongated ellipses, in which ecome invisible, their feeble light being effaced by the superior e planet, long before they come up to its disc, so that the ob- f any eclipses or occultations they may undergo is quite out stion, with our present telescopes.

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[1] September 14th to November 9th, 1847.

[2] October 8th to December 10th, 1847.

If the observation of the satellites of Uranus be difficult, those e, owing to the immense distance of that planet, may be readily o offer still greater difficulties. Of the existence of one, dis- Mr. Lassell,[1] there can remain no doubt, it having also been y other astronomers, both in Europe and America. Accord- Otto Struve[2] its orbit is inclined to the ecliptic at the considera- f 35°; but whether, as in the case of the satellites of Uranus, on of its motion be retrograde, it is not possible to say, until it been longer observed.

8th, 1847.

Nachr. No. 629, from his own observations, September 11th to Decem- 47.

CHAPTER XI.

OF COMETS.

GREAT NUMBER OF RECORDED COMETS. — THE NUMBER OF THOSE UNRECORDED PROBABLY MUCH GREATER. — GENERAL DESCRIPTION OF A COMET. — COMETS WITHOUT TAILS, OR WITH MORE THAN ONE. — THEIR EXTREME TENUITY. — THEIR PROBABLE STRUCTURE. — MOTIONS CONFORMABLE TO THE LAW OF GRAVITY. — ACTUAL DIMENSIONS OF COMETS. — PERIODICAL RETURN OF SEVERAL. — HALLEY'S COMET. — OTHER ANCIENT COMETS PROBABLY PERIODIC. — ENCKE'S COMET. — BIELA'S. — FAYE'S. — LEXELL'S. — DE VICO'S. — BRORSEN'S. — PETERS'S. — GREAT COMET OF 1843. — ITS PROBABLE IDENTITY WITH SEVERAL OLDER COMETS. — GREAT INTEREST AT PRESENT ATTACHED TO COMETARY ASTRONOMY, AND ITS REASONS. — REMARKS ON COMETARY ORBITS IN GENERAL.

(554.) The extraordinary aspect of comets, their rapid and seemingly irregular motions, the unexpected manner in which they often burst upon us, and the imposing magnitudes which they occasionally assume, have in all ages rendered them objects of astonishment, not unmixed with superstitious dread to the uninstructed, and an enigma to those most conversant with the wonders of creation and the operations of natural causes. Even now, that we have ceased to regard their movements as irregular, or as governed by other laws than those which retain the planets in their orbits, their intimate nature, and the offices they perform in the economy of our system, are as much unknown as ever. No distinct and satisfactory account has yet been rendered of those immensely voluminous appendages which they bear about with them, and which are known by the name of their tails, (though improperly, since they often precede them in their motions,) any more than of several other singularities which they present.

(555.) The number of comets which have been astronomically observed, or of which notices have been recorded in history, is very great, amount-

ral hundreds;[1] and when we consider that in the earlier ages
ny, and indeed in more recent times, before the invention of
pe, only large and conspicuous ones were noticed; and that
ttention has been paid to the subject, scarcely a year has passed
e observation of one or two of these bodies, and that sometimes
en three have appeared at once; it will be easily supposed that
l number must be at least many thousands. Multitudes,
st escape all observation, by reason of their paths traversing
art of the heavens which is above the horizon in the daytime.
circumstanced can only become visible by the rare coincidence
clipse of the sun,—a coincidence which happened, as related
sixty-two years before Christ, when a large comet was actually
ery near the sun. Several, however, stand on record as having
t enough to be seen with the naked eye in the daytime, even
d in bright sunshine. Such were the comets of 1402, 1532,
and that of 43 B. C. which appeared during the games cele-
Augustus in honour of Venus shortly after the death of Cæsar,
the flattery of poets declared to be the soul of that hero taking
nong the divinities.

That feelings of awe and astonishment should be excited by the
l unexpected appearance of a great comet, is no way surprising;

the tail. This magnificent appendage attains occasionally an immense apparent length. Aristotle relates of the tail of the comet of 371 B. C., that it occupied a third of the hemisphere, or 60°; that of A. D. 1618 is stated to have been attended by a train no less than 104° in length. The comet of 1680, the most celebrated of modern times, and on many accounts the most remarkable of all, with a head not exceeding in brightness a star of the second magnitude, covered with its tail an extent of more than 70° of the heavens, or, as some accounts state, 90°; that of the comet of 1769 extended 97°, and that of the last *great* comet (1843) was estimated at about 65° when longest. The figure (*fig.* 2, Plate II.) is a representation of the comet of 1819 — by no means one of the most considerable, but which was, however, very conspicuous to the naked eye.

(557.) The tail is, however, by no means an invariable appendage of comets. Many of the brightest have been observed to have short and feeble tails, and a few great comets have been entirely without them. Those of 1585 and 1763 offered no vestige of a tail; and Cassini describes the comets of 1665 and 1682 as being as round[1] and as well defined as Jupiter. On the other hand, instances are not wanting of comets furnished with many tails or streams of diverging light. That of 1744 had no less than six, spread out like an immense fan, extending to a distance of nearly 30° in length. The small comet of 1823 had two, making an angle of about 160°, the brighter turned as usual from the sun, the fainter towards it, or nearly so. The tails of comets, too, are often somewhat curved, bending, in general, towards the region which the comet has left, as if moving somewhat more slowly, or as if resisted in their course.

(558.) The smaller comets, such as are visible only in telescopes, or with difficulty by the naked eye, and which are by far the most numerous, offer very frequently no appearance of a tail, and appear only as round or somewhat oval vaporous masses, more dense towards the centre, where, however, they appear to have no distinct nucleus, or anything which seems entitled to be considered as a solid body. Stars of the smallest magnitudes remain distinctly visible, though covered by what appears to be the densest portion of their substance; although the same stars would be completely obliterated by a moderate fog, extending only a few yards from the surface of the earth. And since it is an observed fact, that even those

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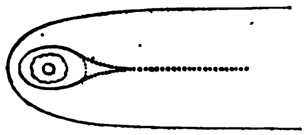
[1] This description however applies to the "disc" of the head of these comets as seen in a telescope. Cassini's expressions are, "aussi rond, aussi net, et aussi clair que Jupiter," (where it is to be observed that the latter epithet must by no means be translated *bright*). To understand this passage fully, the reader must refer to the description given farther on, of the "disc" of Halley's comet, after its perihelion passage in 1835-6.

nets which have presented the appearance of a nucleus have yet
no phases, though we cannot doubt that they shine by the re-
ar light, it follows that even these can only be regarded as great
thin vapour, susceptible of being penetrated through their
stance by the sunbeams, and reflecting them alike from their
rts and from their surfaces. Nor will any one regard this ex-
as forced, or feel disposed to resort to a phosphorescent quality
met itself, to account for the phenomena in question, when we
(what will hereafter be shown) the enormous magnitude of the
s illuminated, and the extremely small *mass* which there is
attribute to these bodies. It will then be evident that the most
tial clouds which float in the highest regions of our atmosphere,
at sunset to be drenched in light, and to glow throughout their
pth as if in actual ignition, without any shadow or dark side,
ooked upon as dense and massive bodies compared with the filmy
t spiritual texture of a comet. Accordingly, whenever powerful
have been turned on these bodies, they have not failed to dispel
n which attributes *solidity* to that more condensed part of the
ch appears to the naked eye as a nucleus; though it is true that
very minute stellar point *has* been seen, indicating the existence
body.

It is in all probability to the feeble coercion of the elastic powe a smoke, fog, or cloud, suspended in a transparent atmosphere, is evident from a fact which has been often noticed, viz.—that the portion of the tail where it comes up, and surrounds the head, is yet separate from it by an interval less luminous, as if sustained and kept off from contact by a transparent stratum, as we often see one layer of clouds over another with a considerable clear space between. These and most of the other facts observed in the history of comets, appear to indicate that the structure of a comet, as seen in section in the direction of its length, must be that of a hollow envelope, of a parabolic form, enclosing near its vertex the nucleus and head, something as represented in the annexed figure. This would account for the apparent division of the tail into two principal lateral branches, the envelope being oblique to the line of sight at its borders, and therefore a greater depth of illuminated matter being there exposed to the eye. In all probability, however, they admit great varieties of structure, and among them may very possibly be bodies of widely different physical constitution, and there is no doubt that one and the same comet at different epochs undergoes great changes, both in the disposition of its materials and in their physical state.

Fig. 75.



(561.) We come now to speak of the motions of comets. These are apparently most irregular and capricious. Sometimes they remain in sight only for a few days, at others for many months; some move with extreme slowness, others with extraordinary velocity; while not unfrequently, the two extremes of apparent speed are exhibited by the same comet in different parts of its course. The comet of 1472 described an arc of the heavens of 40° of a great circle[1] in a single day. Some pursue a direct, some a retrograde, and others a tortuous and very irregular course; nor do they confine themselves, like the planets, within any certain region of the heavens, but traverse indifferently every part. Their variations in apparent size, during the time they continue visible, are no less remarkable than those of their velocity; sometimes they make their first appearance as faint and slow moving objects, with little or no tail;

[1] 120° in extent in the former editions. But this was the arc described *in longitude*. and the comet at the time referred to had great north latitude.

rees accelerate, enlarge, and throw out from them this appen-
a increases in length and brightness till (as always happens in
they approach the sun, and are lost in his beams. After a
gain emerge, on the other side, receding from the sun with a
first rapid, but gradually decaying. It is for the most part
passing the sun, that they shine forth in all their splendour,
eir tails acquire their greatest length and developement; thus
plainly the action of the sun's rays as the exciting cause of that
ry emanation. As they continue to recede from the sun, their
inishes and the tail dies away, or is absorbed into the head,
f grows continually feebler, and is at length altogether lost
by far the greater number of cases never to be seen more.

Without the clue furnished by the theory of gravitation, the
these seemingly irregular and capricious movements might
ined for ever unresolved. But Newton, having demonstrated
lity of any conic section whatever being described about the
body revolving under the dominion of that law, immediately
he applicability of the general proposition to the case of come-
; and the great comet of 1680, one of the most remarkable on
h for the immense length of its tail and for the excessive close-
approach to the sun (within one-sixth of the diameter of that

In general, three complete observations of its right ascension and declination, with the times at which they were made, suffice for the solution of this problem, (which is, however, by no means an easy one,) and for the determination of the elements of the orbit. These consist, *mutatis mutandis*, of the same data as are required for the computation of the motion of a planet; (that is to say, the longitude of the perihelion, that of the ascending node, the inclination to the ecliptic, the semiaxis, excentricity, and time of perihelion passage, as also whether the motion is direct or retrograde;) and, once determined, it becomes very easy to compare them with the whole observed course of the comet, by a process exactly similar to that of art. 502, and thus at once to ascertain their correctness, and to put to the severest trial the truth of those general laws on which all such calculations are founded.

(564.) For the most part, it is found that the motions of comets may be sufficiently well represented by parabolic orbits,—that is to say, ellipses whose axes are of infinite length, or, at least, so very long that no appreciable error in the calculation of their motions, during all the time they continue visible, would be incurred by supposing them actually infinite. The parabola is that conic section which is the limit between the ellipse on the one hand, which returns into itself, and the hyperbola on the other, which runs out to infinity. A comet, therefore, which should describe an elliptic path, however long its axis, must *have* visited the sun before, and must again return (unless disturbed) in some determinate period,—but should its orbit be of the hyperbolic character, when once it had passed its perihelion, it could never more return within the sphere of our observation, but must run off to visit other systems, or be lost in the immensity of space. A very few comets have been ascertained to move in hyperbolas[1], but many more in ellipses. These latter, in so far as their orbits can remain unaltered by the attractions of the planets, must be regarded as permanent members of our system.

(565.) We must now say a few words on the actual dimensions of comets. The calculation of the diameters of their heads, and the lengths and breadths of their tails, offers not the slightest difficulty when once the elements of their orbits are known, for by these we know their real distances from the earth at any time, and the true direction of the tail, which we see only foreshortened. Now calculations instituted on these principles lead to the surprising fact, that the comets are by far the most voluminous bodies in our system. The following are the dimensions of some of those which have been made the subjects of such inquiry.

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[1] For example, that of 1723, calculated by Burckhardt; that of 1771, by both Burckhardt and Encke; and the second comet of 1818, by Rosenberg and Schwabe.

The tail of the great comet of 1680, immediately after its peri-
sage, was found by Newton to have no less than 20000000 of
length, and to have occupied only two days in its emission
omet's body! a decisive proof this of its being darted forth by
e force, the origin of which, to judge from the direction of the
be sought in the sun itself. Its greatest length amounted to
leagues, a length much exceeding the whole interval between
nd earth. The tail of the comet of 1769 extended 16000000
d that of the great comet of 1811, 36000000. The portion
d of this last, comprised within the transparent atmospheric en-
ich separated it from the tail, was 180000 leagues in diameter.
y conceivable, that matter once projected to such enormous dis-
uld ever be collected again by the feeble attraction of such a
comet—a consideration which accounts for the surmised pro-
minution of the tails of such as have been frequently observed.
The most remarkable of those comets which have been ascer-
nove in elliptic orbits is that of Halley, so called from the cele-
nund Halley, who, on calculating its elements from its perihelion
1682, when it appeared in great splendour, with a tail 30° in
s led to conclude its identity with the great comets of 1531 and
se elements he had also ascertained. The intervals of these

proving state of our knowledge of the methods of estimating the disturbing effect of the several planets. The last of these predictions, that of M. Lehmann, was published on the 25th of July. On the 5th of August the comet first became visible in the clear atmosphere of Rome as an exceedingly faint telescopic nebula, within a degree of its place as predicted by M. Rosenberger for that day. On or about the 20th of August it became generally visible, and, pursuing very nearly its calculated path among the stars, passed its perihelion on the 16th of November; after which, its course carrying it south, it ceased to be visible in Europe, though it continued to be conspicuously so in the southern hemisphere throughout February, March, and April, 1836, disappearing finally on the 5th of May.

(568.) Although the appearance of this celebrated comet at its last apparition was not such as might be reasonably considered likely to excite lively sensations of terror, even in superstitious ages, yet, having been an object of the most diligent attention in all parts of the world to astronomers, furnished with telescopes very far surpassing in power those which had been applied to it at its former appearance in 1759, and indeed to any of the greater comets on record, the opportunity thus afforded of studying its physical structure, and the extraordinary phænomena which it presented when so examined, have rendered this a memorable epoch in cometic history. Its first appearance, while yet very remote from the sun, was that of a small round or somewhat oval nebula, quite destitute of tail, and having a minute point of more concentrated light excentrically situated within it. It was not before the 2d of October that the tail began to be developed, and thenceforward increased pretty rapidly, being already 4° or 5° long on the 5th. It attained its greatest apparent length (about 20°) on the 15th of October. From that time, though not yet arrived at its perihelion, it decreased with such rapidity, that already on the 29th it was only 3°, and on November the 5th 2½° in length. There is every reason to believe that before the perihelion, the tail had altogether disappeared, as, though it continued to be observed at Pulkowa up to the very day of its perihelion passage, no mention whatever is made of any tail being then seen.

(569.) By far the most striking phænomena, however, observed in this part of its career, were those which, commencing simultaneously with the growth of the tail, connected themselves evidently with the production of that appendage and its projection from the head. On the 2d of October (the very day of the first observed commencement of the tail) the nucleus, which had been faint and small, was observed suddenly to have become much brighter, and to be in the act of throwing out a jet or stream of

its anterior part, or that turned *towards* the sun. This ejection
ing awhile was resumed, and with much greater apparent vio-
the 8th, and continued, with occasional intermittences, so long
itself continued visible. Both the form of this luminous ejec-
the direction in which it issued from the nucleus, meanwhile
t singular and capricious alterations, the different phases suc-
ach other with such rapidity that on no two successive nights
ppearances alike. At one time the emitted jet was single, and
vithin narrow limits of divergence from the nucleus. At others
ed a fan-shaped or swallow-tailed form, analogous to that of a
issuing from a flattened orifice: while at others again two, three,
ore jets were darted forth in different directions.[1] (See figures
plate I., fig. 4, which represent, highly magnified, the appear-
he nucleus with its jets of light, on the 8th, 9th, 10th, and 12th
r, and in which the direction of the anterior portion of the head,
nting the sun, is supposed alike in all, viz. towards the upper
engraving. In these representations the head itself is omitted,
f the figures not permitting its introduction: *e* represents the
d head as seen October 9th on a less scale.) The direction of
pal jet was observed meanwhile to oscillate to and fro on either
ine directed to the sun in the manner of a compass-needle when

a comet is powerfully excited and dilated into a vaporous state by the action of the sun's rays, escaping in streams and jets at those points of its surface which oppose the least resistance, and in all probability throwing that surface or the nucleus itself into irregular motions by its reaction in the act of so escaping, and thus altering its direction.

2dly. That this process chiefly takes place in that portion of the nucleus which is turned towards the sun; the vapour escaping chiefly in that direction.

3dly. That when so emitted, it is prevented from proceeding in the direction originally impressed upon it, by some force directed *from* the sun, drifting it back and carrying it out to vast distances behind the nucleus, forming the tail or so much of the tail as can be considered as consisting of material substance.

4thly. That this force, whatever its nature, acts unequally on the materials of the comet, the greater portion remaining unvaporized, and a considerable part of the vapour actually produced, remaining in its neighbourhood, forming the head and coma.

5thly. That the force thus acting on the materials of the tail cannot possibly be identical with the ordinary gravitation of matter, being centrifugal or repulsive, as respects the sun, and of an energy very far exceeding the gravitating force towards that luminary. This will be evident if we consider the enormous velocity with which the matter of the tail is carried backwards, in opposition both to the motion which it had as part of the nucleus, and to that which it acquired in the act of its emission, both which motions have to be destroyed in the first instance, before any movement in the contrary direction can be impressed.

6thly. That unless the matter of the tail thus repelled from the sun be retained by a peculiar and highly energetic attraction to the nucleus, differing from and exceptional to the ordinary power of gravitation, it must leave the nucleus altogether; being in effect carried far beyond the coercive power of so feeble a gravitating force as would correspond to the minute mass of the nucleus; and it is therefore very conceivable that a comet may lose, at every approach to the sun, a portion of that peculiar matter, whatever it be, on which the production of its tail depends, the remainder being of course less excitable by the solar action, and more impassive to his rays, and therefore, *pro tanto*, more nearly approximating to the nature of the planetary bodies.

(571.) After the perihelion passage, the comet was lost sight of for upwards of two months, and at its reappearance (on the 24th of January 1836) presented itself under quite a different aspect, having in the interval evidently undergone some great physical change which had operated

nsformation in its appearance. It no longer presented any
ail, but appeared to the naked eye as a hazy star of about the
th magnitude, and in powerful telescopes as a small, round,
disc, rather more than 2′ in diameter, surrounded with a
evelure or coma of much greater extent. Within the disc,
hat excentrically situated, a minute but bright nucleus appeared,
extended towards the posterior edge of the disc (or that remote
n) a short vivid luminous ray. (See fig. 4 of pl. I.) As the
ed from the sun, the coma speedily disappeared, as if absorbed
sc, which, on the other hand, increased continually in dimen-
hat with such rapidity, that in the week elapsed from January
ruary 1st, (calculating from micrometrical measures, and from
distance of the comet from the earth on those days) the actual
eal solid content of the illuminated space had dilated in the
wards of 40 to 1. And so it continued to swell out with un-
rapidity, until from this cause alone it ceased to be visible, the
becoming fainter as the magnitude increased; till at length
became undistinguishable from simple want of light to trace
this increase of dimension proceeded, the form of the disc
gradual and successive additions to its length in the direction
the sun, to that of a paraboloid, as represented in *g*, fig. 4,
e anterior curved portion preserving its planetary sharpness,

(573.) Besides the comet of Halley, several other of the great comets recorded in history have been surmised with more or less probability to return periodically, and therefore to move in elongated ellipses around the sun. Such is the great comet of 1680, whose period is estimated at 575 years, and which is considered, with the highest appearance of probability, to be identical with a magnificent comet observed at Constantinople and in Palestine, and referred by contemporary historians, both European and Chinese, to the year A. D. 1105; with that of A. D. 575, which was seen at noon-day close to the sun; with the comet of 43 B. C., already spoken of as having appeared after the death of Cæsar, and which was also observed in the day-time; and finally with two other comets, mention of which occurs in the Sibylline Oracles, and in a passage of Homer, and which are referred, as well as the obscurity of chronology and the indications themselves will allow, to the years 618 and 1194 B. C. It is to the assumed near approach of this comet to the earth about the time of the Deluge, that Whiston ascribed that overwhelming tide wave to whose agency his wild fancy ascribed that great catastrophe—a speculation, it is needless to remark, purely visionary.

(574.) Another great comet, whose return in the year actually current (1848) has been considered by more than one eminent authority in this department of astronomy[1] highly probable, is that of 1556, to the terror of whose aspect some historians have attributed the abdication of the Emperor Charles V. This comet is supposed to be identical with that of 1264, mentioned by many historians as a great comet, and observed also in China, — the conclusion in this case resting upon the coincidence of elements calculated on the observations, such as they are, which have been recorded. On the subject of this coincidence Mr. Hind has recently entered into many elaborate calculations, the result of which is strongly in favour of the supposed identity. This probability is farther increased by the fact of a comet with a tail of 40° and a head bright enough to be visible after sunrise having appeared in A. D. 975; and of two others having been recorded by the Chinese annalists in A. D. 395 and 104. It is true that if these be the same, the mean period would be somewhat short of 292 years. But the effect of planetary perturbation might reconcile even greater differences, and though up to the time of our writing no such comet has yet been observed, at least another year must elapse before its return can be pronounced hopeless.

(575.) In 1661, 1532, 1402, 1145, 891, and 243 great comets appeared—that of 1402 being bright enough to be seen at noon-day. A period of 129 years would conciliate all these appearances, and should

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[1] Pingré, Cometographie, i. 411. Lalande, Astr. 3185.

ht back the comet in 1789 or 1790 (other circumstances
That no such comet was observed about that time is no proof
not return, since, owing to the situation of its orbit, had the
passage taken place in July it might have escaped observation.
ndeed, from an elaborate discussion of the observations of 1532
came to the conclusion that these comets were not the same;
ments assigned by Olbers to the earlier of them, differ so
n those of Mechain for the same comet on the one hand, and
ell with those of the last named astronomer for the other,'
perhaps justified in regarding the question as not yet set at

We come now, however, to a class of comets of short period,
whose return there is no doubt, inasmuch as two at least of
been identified as having performed successive revolutions
sun; have had their return predicted already several times;
n each occasion scrupulously kept to their appointments. The
se is the comet of Encke, so called from Professor Encke of
o first ascertained its periodical return. It revolves in an
reat excentricity (though not comparable to that of Halley's,)
f which is inclined at an angle of about 13° 22′ to the plane
tic, and in the short period of 1211 days, or about $3\frac{1}{3}$ years.
kable discovery was made on the occasion of its fourth recorded

It will, therefore, probably fall ultimately into the sun, should it not first be dissipated altogether,—a thing no way improbable, when the lightness of its materials is considered.

(578.) By measuring the apparent magnitude of this comet at different distances from the sun, and thence, from a knowledge of its actual distance from the earth at the time, concluding its real volume, it has been ascertained to contract in bulk as it approaches to, and to expand as it recedes from, that luminary. M. Valz, who was the first to notice this fact, accounts for it by supposing it to undergo a real compression or condensation of volume arising from the pressure of an æthereal medium which he conceives to grow more dense in the sun's neighbourhood. But such an hypothesis is evidently inadmissible, since it would require us to assume the exterior of the comet to be in the nature of a skin or bag impervious to the compressing medium. The phenomenon is analogous to the increase of dimension above described as observed in the comet of Halley when in the act of receding from the sun, and is doubtless referable to a similar cause, viz. the alternate conversion of evaporable matter into the states of visible cloud and invisible gas by the alternating action of cold and heat. This comet has no tail, but offers to the view only a small ill-defined nucleus, excentrically situated within a more or less elongated oval mass of vapours, being nearest to that vertex which is towards the sun.

(579.) Another comet of short period is that of *Biela*, so called from M. Biela, of Josephstadt, who first arrived at this interesting conclusion on the occasion of its appearance in 1826. It is considered to be identical with comets which appeared in 1772, 1805, &c., and describes its very excentric ellipse about the sun in 2410 days or about 6¾ years; and in a plane inclined 12° 34′ to the ecliptic. It appeared again according to the prediction in 1832, and in 1846. Its orbit, by a remarkable coincidence, very nearly intersects that of the earth; and had the latter at the time of its passage in 1832, been a month in advance of its actual place, it would have passed through the comet,—a singular rencontre, perhaps not unattended with danger.[1]

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[1] Should calculation establish the fact of a resistance experienced also by this comet, the subject of periodical comets will assume an extraordinary degree of interest. It cannot be doubted that many more will be discovered, and by their resistance questions will come to be decided, such as the following:—What is the law of density of the resisting medium which surrounds the sun? Is it at rest or in motion? If the latter, in what direction does it move? Circularly round the sun, or traversing space? If circularly, in what plane? It is obvious that a circular or vorticose motion of the ether would *accelerate some comets and retard others*, according as their revolution was, relative to such motion, direct or retrograde. Supposing the neighbourhood of the sun to be filled with a material fluid, it is not conceivable that the circulation of the planets in

This comet is small and hardly visible to the naked eye, even
ntest. Nevertheless, as if to make up for its seeming insignifi-
he interest attaching to it in a physical point of view, it exhi-
last appearance, in 1846, a phænomenon which struck every
with amazement, as a thing without previous example in the
our system.[1] It was actually seen to separate itself into two dis-
s, which, after thus parting company, continued to journey along
hrough an arc of upwards of 70° of their apparent orbit, keeping
le within the same field of view of the telescope pointed towards
he first indication of something unusual being about to take
ht be, perhaps, referred to the 19th of December, 1845, when
appeared pear-shaped, the nebulosity being unduly elongated in
following direction.[2] But on the 13th of January, at Wash-
America, and on the 15th and subsequently in every part of
was distinctly seen to have become double; a very small and
tic body, having a nucleus of its own, being observed appended
distance of about 2′ (in arc) from its centre, and in a direction
n angle of about 328° with the meridian, running northwards
rincipal or original comet (see art. 204). From this time the
of the two comets went on progressively, though slowly. On
f January, the apparent distance of the nucleus had increased

nuclei, both had short tails, parallel in direction, and nearly perpendicular to the line of junction, but whereas at its first observation on January 18th, the new comet was extremely small and faint in comparison with the old, the difference both in point of light and apparent magnitude diminished. On the 10th of February, they were nearly equal, although the day before the moonlight had effaced the new one, leaving the other bright enough to be well observed. On the 14th and 16th, however, the new comet had gained a decided superiority of light over the old, presenting at the same time a sharp and starlike nucleus, compared by Lieut. Maury to a diamond spark. But this state of things was not to continue. Already, on the 18th, the old comet had regained its superiority, being nearly twice as bright as its companion, and offering an unusually bright and starlike nucleus. From this period the new companion began to fade away, but continued visible up to the 15th of March. On the 24th the comet was again single, and on the 22d of April both had disappeared.

(582.) While this singular interchange of light was going forwards, indications of some sort of communication between the comets were exhibited. The new or companion comet, besides its tail, extending in a direction parallel to that of the other, threw out a faint arc of light, which extended as a kind of bridge from the one to the other; and after the restoration of the original comet to its former preeminence, it, on its part, threw forth additional rays, so as to present (on the 22d and 23d February) the appearance of a comet with three faint tails, forming angles of about 120° with each other, one of which extended towards its companion.[1]

(583.) Professor Plantamour, director of the observatory of Geneva, having investigated the orbits of both these comets as separate and independent bodies, from the extensive and careful series of observations made upon them, has arrived at the conclusion that the increase of distance between the two nuclei, *at least during the interval from February 10th to March 22d*, was simply apparent, being due to the variation of distance from the earth, and to the angle under which their line of junction presented itself to the visual ray; the real distance during all that interval (neglecting small fractions) having been on an average about thirty-nine times the semi-diameter of the earth, or less than two-thirds the distance of the moon from its centre. From this it would appear, that already, at this distance, the two bodies had ceased to exercise any

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[1] These last-mentioned particulars rest on the testimony of Lieutenant Maury of Washington, who had the advantage of using a nine-inch object-glass of Munich manufacture. It does not appear that any *large* telescope was turned upon it in Europe on the dates in question.

amount of perturbative gravitation on each other; as, indeed,
probable minuteness of cometary masses we might reasonably
alculating upon the elements assigned by him[1], we find 16^{d}4
erval of their next perihelion passages. And it will be, there-
sary at their next reappearance, to look out for each comet as a
d independent body, computing its place from these elements
other had no existence. Nevertheless, as it is still perfectly
at some link of connection may subsist between them, (if, in
some unknown process the companion has not been actually
,) it will not be advisable to rely on this calculation to the
a most vigilant search throughout the whole neighbourhood of
onspicuous one, lest the opportunity should be lost of pursuing
usion the history of this strange occurrence.

A third comet of short period has still more recently been added
by M. Faye, of the observatory of Paris, who detected it on
f November 1843. A very few observations sufficed to show
rabola would satisfy the conditions of its motion, and that to
hem completely, it was necessary to assign to it an elliptic orbit
oderate excentricity. The calculations of M. Nicolai, subse-
vised, and slightly corrected by M. Leverrier, have shown that
perfect representation of its motions during the whole period

angles to the ecliptic, they must in many instances, if not actually intersect, at least pass very near to the orbits of some of the planets. We have already seen, for instance, that the orbit of Biela's comet so nearly intersects that of the earth, that an actual collision is not impossible, and indeed (supposing neither orbit variable) must in all likelihood happen in the lapse of some millions of years. Neither are instances wanting of comets having actually approached the earth within comparatively short distances, as that of 1770, which on the 1st of July of that year was witin little more than seven times the moon's distance. The same comet in 1767 passed Jupiter at a distance only one 58th of the radius of that planet's orbit, and it has been rendered extremely probable that it is to the disturbance its former orbit underwent during that appulse that we owe its appearance within our own range of vision. This exceedingly remarkable comet was found by Lexell to describe an elliptic orbit with an excentricity of 0·7858, with a periodic time of about five years and a half, and in a plane only 1° 34′ inclined to the ecliptic, having passed its perihelion on the 13th of August 1770. Its return of course was eagerly expected, but in vain, for the comet has never been seen since. Its observation on its first return in 1776 was rendered impossible by the relative situations of the perihelion and of the earth at the time, and before another revolution could be accomplished (as has since been ascertained,) viz: about the 23d of August 1779, by a singular coincidence it again approached Jupiter within one 491st part of its distance from the sun, being nearer to that planet by one-fifth than its fourth satellite. No wonder, therefore, that the planet's attraction (which at that distance would exceed that of the sun in the proportion of at least 200 to 1) should completely alter the orbit and deflect it into a curve, not one of whose elements would have the least resemblance to those of the ellipse of Lexell. It is worthy of notice that by this rencontre with the system of Jupiter's satellites, none of *their* motions suffered any perceptible derangement,—a sufficient proof of the smallness of its mass. Jupiter indeed, seems, by some strange fatality, to be constantly in the way of comets, and to serve as a perpetual stumbling-block to them.

(586.) On the 22nd of August, 1844, Signor De Vico, director of the observatory of the Collegio Romano, discovered a comet, the motions of which, a very few observations sufficed to show, deviated remarkably from a parabolic orbit. It passed its perihelion on the 2nd of September, and continued to be observed until the 7th of December. Elliptic elements of this comet, agreeing remarkably well with each other, were accordingly calculated by several astronomers; from which it appears that the period of revolution is about 1990 days, or 5½ (5·4357) years, which (supposing

disturbed in the interim) would bring it back to the perihelion
t the 13th of January, 1850. As the assemblage and con-
these elements thus computed independently, will serve better,
an any other example, to afford the student an idea of the
rithmetical certainty capable of being attained in this branch
y, difficult and complex as the calculations themselves are, and
or as *individual* observations of a body so ill-defined as the
ets are for the most part; we shall present them in a tabular
the next page: the elements being as usual; the time of peri-
age, longitude of the perihelion, that of the ascending node,
ion to the ecliptic, semiaxis and excentricity of the orbit, and
time.

net, when brightest, was visible to the naked eye, and had a
It is especially interesting to astronomers from the circum-
s having been rendered exceedingly probable by the researches
rrier, that it is identical with one which appeared in 1678 with
elements considerably changed by perturbation. This comet
remarkable, from having been concluded by Messrs. Laugier
is, to be identical with the comet of 1585 observed by Tycho
possibly also with those of 1743, 1766, and 1819.

Elliptic elements have in like manner been assigned to the
overed by M. Brorsen, on the 26th of February, 1846, which

ELEMENTS OF THE PERIODICAL COMET OF DE VICO.

COMPUTED BY	NICOLAI.	HIND.	GOLDSCHMIDT.	FAYE.	SCHUBERT.	BRUNNOW.
	1844, Sep.	1844, Sep.	1844, Sep.	1844, Sep.	1844, Sep.	1844, Sep.
Perihelion passage	2^{d}·47594	2^{d}·50412	2^{d}·45430	2^{d}·48745	2^{d}·56348	2^{d}·47402
Longitude of perihelion.........	342° 31′ 5″·5	342° 32′ 40″·1	342° 29′ 44″·9	342° 31′ 15″·5	342° 34′ 31″·5	342° 30′ 49″·6
Longitude of ☊	63 48 48·9	63 52 24·1	63 48 55·2	63 49 30·6	63 54 40·8	63 49 0·1
Inclination	2 54 45·8	2 54 27·1	2 55 1·9	2 54 45·0	2 52 51·3	2 54 50·3
Semiaxis...............................	3·09853	3·08582	3·11111	3·09946	3·02612	3·10295
Excentricity..........................	0·61716	0·61566	0·61861	0·61726	0·60866	0·61788
Period (days)........................	1992	1980	2400	1993	1923	1996

This comet is faint, and presents nothing remarkable in its appearance. Its chief interest arises from the great similarity of its *parabolic* elements to those of the comet of 1532, the place of the perihelion and node, and the inclination of the orbit, being almost identical.

(588.) Elliptic elements have also been calculated by M. D'Arrest, for a comet discovered by M. Peters, on the 26th of June, 1846, which go to assign it a place among the comets of short period, viz. $5804^{d}\cdot 3$, or very nearly 16 years. The excentricity of the orbit is 0·75672, its semi-axis 6·32066, and the inclination of its plane to that of the ecliptic 31° 2′ 14″. This comet passed its perihelion on the 1st of June, 1846.

(589.) By far the most remarkable comet, however, which has been seen during the present century, is that which appeared in the spring of 1843, and whose tail became visible in the twilight of the 17th of March in England as a great beam of nebulous light, extending from a point above the western horizon, through the stars of Eridanus and Lepus, under the belt of Orion. This situation was low and unfavourable; and it was not till the 19th that the head was seen, and then only as a faint and ill-defined nebula, very rapidly fading on subsequent nights. In more southern latitudes, however, not only the tail was seen, as a magnificent train of light extending 50° or 60° in length; but the head and nucleus appeared with extraordinary splendour, exciting in every country where it was seen the greatest astonishment and admiration. Indeed, all descriptions agree in representing it as a stupendous spectacle, such as in superstitious ages would not fail to have carried terror into every bosom. In tropical latitudes in the northern hemisphere, the tail appeared on the 3d of March, and in Van Diemen's Land, so early as the 1st, the comet having passed its perihelion on the 27th of February. Already on the 3d the head was so far disengaged from the immediate vicinity of the sun, as to appear for a short time above the horizon after sunset. On this day when viewed through a 46-inch achromatic telescope it presented a planetary disc, from which rays emerged in the direction of the tail. The tail was double, consisting of two principal lateral streamers, making a very small angle with each other, and divided by a comparatively dark line, of the estimated length of 25°, prolonged, however, on the north side by a divergent streamer, making an angle of 5° or 6° with the general direction of the axis, and traceable as far as 65° from the head. A similar though fainter lateral prolongation appeared on the south side. A fine drawing of it of this date by C. P. Smyth, Esq., of the Royal Observatory, C. G. H., represents it as highly symmetrical, and gives the idea of a vivid cone of light, with a dark axis, and nearly rectilinear sides, inclosed in a fainter cone, the sides of which curve slightly outwards. The light f the nucleus at this period is compared to that of a star of the first or second magnitude; and on the 11th, of the third; from which time it egraded in light so rapidly, that on the 19th it was invisible to the naked ye, the tail all the while continuing brilliantly visible, though much more so at a distance from the nucleus, with which, indeed, its connexion was not then obvious to the unassisted sight — a singular feature in the history of this body. The tail, subsequent to the 3d, was, generally speaking, a single straight or slightly curved broad band of light, but on the 11th it is recorded by Mr. Clerihew, who observed it at Calcutta, to have shot forth a lateral tail nearly twice as long as the regular one, but fainter, and making an angle of about 18° with its direction on the southern side. The projection of this ray (which was not seen either before or after the day in question) to so enormous a length, (nearly 100°) in a single day conveys an impression of the intensity of the forces acting to produce such a velocity of material transfer through space, such as no other natural phænomenon is capable of exciting. It is clear that *if we have to deal here with matter, such as we conceive it,* viz. *possessing inertia — at all,* it must be under the dominion of forces incomparably more energetic than gravitation.

(590.) There is abundant evidence of the comet in question having been seen in full daylight, and in the sun's immediate vicinity. It was so seen on the 28th of February, the day after its perihelion passage, by every person on board the H. E. I. C. S. Owen Glendower, then off the Cape, as a short, dagger-like object, close to the sun, a little before sunset. On the same day, at $3^h\ 6^m$ P. M., and consequently in full sunshine, the distance of the nucleus from the sun was actually measured with a sextant by Mr. Clarke, of Portland, United States, the distance, centre from centre, being then only 3° 50′ 43″. He describes it in the following terms: "The nucleus, and also every part of the tail, were as well defined as the moon on a clear day. The nucleus and tail bore the same appearance, and resembled a perfectly pure white cloud, without any variation, except a slight change near the head, just sufficient to distinguish the nucleus from the tail at that point." The denseness of the nucleus was so considerable, that Mr. Clarke had no doubt it might have been visible upon the sun's disc, had it passed between that and the observer. The length of the visible tail resulting from these measures was 59′, or not far from double the apparent diameter of the sun; and as we shall presently see that on the day in question the distance from the earth of the sun and comet must have been very nearly equal, this gives us about 1700000 miles for the linear dimensions of this, the densest portion of that appendage, making no allowance for the foreshortening, which at that time was very considerable.

(591.) The elements of this comet are among the most remarkable of any recorded. They have been calculated by several eminent astronomers, among whose results we shall specify only those which agree best; the earlier attempts to compute its path having been rendered uncertain by the difficulty attending exact observations of it in the first part of its visible career. The following are those which seem entitled to most confidence:

	Encke.	Plantamour.	Knorre.	Nicolai.	Peters.
Perihel. pass., 1843, Feb., mean time at Greenwich	27^{d}·45096	27^{d}·42935	27^{d}·39638	27^{d}·43023	27^{d}·41319
Long. of perihel.......	279° 2′ 30″	278° 18′ 3″	278° 28′ 25″	278° 36′ 33″	279° 59′ 7″
Long. of ☊..............	4 15 5	0 51 4	1 48 3	1 37 55	3 55 17
Inclination	35 12 38	35 8 56	35 35 29	35 36 29	35 15 42
Perihel. dist.............	0·00522	0·00581	0·00579	0·00558	0·00428
Motion	Retrograde.	Retrograde.	Retrograde.	Retrograde.	Retrograde.

(592.) What renders these elements so remarkable is the smallness of the perihelion distance. Of all comets which have been recorded this has made the nearest approach to the sun. The sun's radius being the sine of his apparent semi-diameter (16′ 1″ ·5) to a radius equal to the earth's mean distance$=1$, is represented on that scale by 0·00466, which falls short of 0·00534, the perihelion distance found by taking a mean of all the foregoing results, by only 0·00067, or about one seventh of its whole magnitude. The comet, therefore, approached the *luminous* surface of the sun within about the seventh part of the sun's radius! It is worth while to consider what is implied in such a fact. In the first place, the intensity both of the light and radiant heat of the sun at different distances from that luminary increase proportionally to the spherical area of the portion of the visible hemisphere covered by the sun's disc. This disc, in the case of the earth, at its mean distance has an angular diameter of 32′ 3″. At our comet in perihelio the apparent angular diameter of the sun was no less than 121° 32′. The ratio of the spherical surfaces thus occupied (as appears from spherical geometry) is that of the squares of the sines of the fourth parts of these angles to each other, or that of 1 : 47042. And in this proportion are to each other the amounts of light and heat thrown by the sun on an equal area of exposed surface on our earth and at the comet in equal instants of time. Let any one imagine the effect of so fierce a glare as that of 47000 suns such as we experience the warmth of, on the materials of which the earth's surface is composed. To form some practical idea of it we may compare it with what is recorded of Parker's great lens, whose diameter was 32½ inches and focal length six feet eight inches. The effect of this, supposing all the light and heat transmitted, and the focal concentration perfect, (both conditions very imperfectly satisfied,) would be to enlarge the sun's effective angular diameter to 23° 26′, which, compared on the same principle with a sun of 32′ in diameter, would give a multiplier of only 1915 instead of 47000. The heat to which the comet was subjected therefore surpassed that in the focus of the lens in question, on the lowest calculation, in the proportion of 24½ to 1. Yet that lens melted cornelian, agate, and rock crystal!

(593.) To this extremity of heat however the comet was exposed but for a short time. Its actual velocity in perihelio was no less than 366 miles per second, and the whole of that segment of its orbit above (*i. e.* north of) the plane of the ecliptic, and in which, as will appear from a consideration of the elements, the perihelion was situated, was described in little more than two hours; such being the whole duration of the time from the ascending to the descending node, or in which the comet had north latitude. Arrived at the descending node, its distance from the sun would be already doubled, and the radiation reduced to one fourth of its maximum amount. The comet of 1680, whose perihelion distance was 0·0062, and which therefore approached the sun's surface within one third part of his radius (more than double the distance of the comet now in question) was computed by Newton to have been subjected to an intensity of heat 2000 times that of red-hot iron,—a term of comparison indeed of a very vague description, and which modern thermotics do not recognize as affording a legitimate measure of radiant heat.[1]

(594.) Although some of the observations of this comet were vague and inaccurate, yet there seem good grounds for believing that its whole course cannot be reconciled with a parabolic orbit, and that it really describes an ellipse. Previous to any calculation, it was remarked that in the year 1668 the tail of an immense comet was seen in Lisbon, at Bologna, in Brazil, and elsewhere, occupying nearly the same situation among the stars, and at the same season of the year, viz. on the 5th of March and the following days. Its brightness was such that its reflected trace was easily distinguished on the sea. The head, when it at length came in sight, was comparatively faint and scarce discernible. No precise observations were made of this comet, but the singular coincidence of situation, season of the year, and physical resemblance, excited a strong suspicion of the identity of the two bodies, implying a period of 175 years *within a day or two more or less.* This suspicion has been converted almost into a certainty by a careful examination of what is recorded of the older comet. Locating on a celestial chart the situation of the head, concluded from the direction and appearance of the tail, when only that was seen, and its visible place, when mentioned, according to the descriptions given, it has been found practicable to derive a rough orbit from the course thus laid down: and this agrees in all its features so well with that of the modern comet as nearly to remove all doubt on the subject. Comets, moreover, are recorded to have been seen in A. D. 268, 442–3, 791, 968, 1143, 1317, 1494, which may have been returns of this, since the period above-mentioned would bring round its appearance to the years 268, 443, 618, 793, 968, 1143, 1318, and 1493, and a certain latitude must always be allowed for unknown perturbations.

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[1] A transit of this comet over the sun's disc must probably have taken place shortly after its passage through its decending node. It is greatly to be regretted that so interesting a phænomenon should have passed unobserved. Whether it be possible that some offset of its tail, darted off so late as the 7th of March, when the comet was already far south of the ecliptic, should have crossed that plane and been seen near the Pleiades, may be doubted. Certain it is, that on the evening of that day, a decidedly cometic ray was seen in the immediate neighbourhood of those stars by Mr. Nasmyth. (Ast. Soc. Notices, vol. v. p. 270.)

(595.) But this is not the only comet on record whose identity with the comet of '43 has been maintained. In 1689 a comet bearing a considerable resemblance to it was observed from the 8th to the 23d of December, and from the few and rudely observed places recorded, its elements had been calculated by Pingré, one of the most diligent inquirers into this part of astronomy.[1] From these it appears that the perihelion distance of that comet was very remarkably small, and a sufficient though indeed rough coincidence in the places of the perihelion and node tended to corroborate the suspicion. But the inclination (69°) assigned to it by Pingré appeared conclusive against it. On recomputing the elements, however, from his data, Professor Pierce has assigned to that comet an inclination widely differing from Pingré's, viz. 30° 4′[2], and quite within reasonable limits of resemblance. But how does this agree with the longer period of 175 years before assigned? To reconcile this we must suppose that these 175 years comprise at least eight returns of the comet, and that in effect a mean period of 21^{y}·875 must be allowed for its return. Now it is worth remarking that this period calculated backwards from 1843·156 will bring us upon a series of years remarkable for the appearance of great comets, many of which, as well as the imperfect descriptions we have of their appearance and situation in the heavens, offer at least no obvious contradiction to the supposition of their identity with this. Besides those already mentioned as indicated by the period of 175 years, we may specify as probable or possible intermediate returns, those of the comets of 1733 ?[1], 1689 above-mentioned, 1559 ?, 1537[2], 1515[3], 1471, 1426, 1405–6, 1383, 1361, 1340[4], 1296, 1274, 1230[5], 1208, 1098, 1056, 1034, 1012[6], 990 ?[7], 925 ?, 858 ? ?, 684[8], 552, 530[9], 421, 245 or 247[10], 180[11], 158. Should this view of the subject be the true one, we may expect its return about the end of 1864 or beginning of 1865, in which event it will be observable in the Southern Hemisphere both before and after its perihelion passage[12].

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[1] Author of the "Cométographie," a work indispensable to all who would study this interesting department of the science.

[2] United States Gazette, May 29, 1843. Considering that all the observations lie *near* the descending node of the orbit, the proximity of the comet at that time to the sun, and the loose nature of the recorded observations, no doubt almost any given inclination *might* be deduced from them. The true test in such cases is not to ascend from the old incorrect data to elements, but to descend from known and certain elements to the older data, and ascertain whether the recorded phænomena can be represented by them (perturbations included) within fair limits of interpretation. Such is the course pursued by Clausen.

(596.) M. Clausen, from the assemblage of all the observations of this comet known to him, has calculated elliptic elements which give the extraordinarily short period of 6·38 years. And in effect it has been suggested that a still further subdivision of the period of 21·875 into three of 7·292 years, would reconcile this with other remarkable comets. This seems going too far; but at all events the possibility of representing its motions by so short an ellipse will easily reconcile us to the admission of a period of 21 years. That it should only be visible in certain apparitions, and not in others, is sufficiently explained by the situation of its orbit.

(597.) We have been somewhat diffuse on the subject of this comet, for the sake of showing the degree and kind of interest which attaches to cometic astronomy in the present state of the science. In fact there is no branch of astronomy more replete with interest, and we may add more eagerly pursued at present, inasmuch as the hold which exact calculation gives us on it may be regarded as completely established; so that whatever may be concluded as to the motions of any comet which shall henceforward come to be observed, will be concluded on sure grounds and with numerical precision; while the improvements which have been introduced into the calculation of cometary perturbation, and the daily increasing familiarity of numerous astronomers with computations of this nature, enable us to trace their past and future history with a certainty, which at the commencement of the present century could hardly have been looked upon as attainable. Every comet newly discovered is at once subjected to the ordeal of a most rigorous inquiry. Its elements, roughly calculated within a few days of its appearance, are gradually approximated to as observations accumulate, by a multitude of ardent and expert computists. On the least indication of a deviation from a parabolic orbit, its elliptic elements become a subject of universal and lively interest and discussion. Old records are ransacked, with all the advantage of improved data and methods, so as to rescue from oblivion the orbits of ancient comets which present any similarity to that of the new visitor. The disturbances undergone in the interval by the action of the planets are investigated, and the past, thus brought into unbroken connexion with the present, is made to afford substantial ground for prediction of the future. A great impulse, meanwhile, has been given of late years to the discovery of comets, by the establishment in 1840[1], by his late Majesty the King of Denmark, of a prize medal, to be awarded for every such discovery, to the first observer, (the influence of which may be most unequivocally traced in the great number of these bodies which every successive year sees added to our list,) and by the circulation of notices, by special letter[2], of every such discovery (accompanied, when possible, by an ephemeris), to all observers who have shown that they take an interest in the inquiry, so as to ensure the full and complete observation of the new comet, so long as it remains within the reach of our telescopes.

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[1] P. Passage 1733·781. This great southern comet of May 17th seems too early in the year.

[2] P. P. 1536·906. In January 1537, a comet was seen in Pisces.

[3] P. P. 1515·031. A comet *predicted* the death of Ferdinand the Catholic. He died Jan. 23, 1515.

[4] P. P. 1340·031. Evidently a southern comet, and a very probable appearance.

[5] P. P. 1230·656, was perhaps a return of Halley's.

[6] P. P. 1011·906. In 1012, a very great comet in the southern part of the heavens. "Son éclat blessait les yeux." (Pingré Cométographie, from whom all these recorded appearances are taken.)

[7] P. P. 990·031. "Comète fort épouvantable," *some year* between 989 and 998.

[8] P. P. 683·781. In 684, appeared two or three comets. Dates begin to be obscure.

[9] Two distinct comets (one probably the comet of Cæsar and 1680) appeared in 530 and 531, the former observed in China, the latter in Europe.

[10] P. P. 246·281; both southern comets of the Chinese annals. The year of one or other may be wrong.

[11] P. P. 180·656. Nov. 6, A. D. 180. A southern comet of the Chinese annals.

[12] Clausen, Astron. Nachr. No. 485.

(598.) It is by no means merely as a subject of antiquarian interest, or on account of the brilliant spectacle which comets occasionally afford, that astronomers attach a high degree of importance to all that regards them. Apart even from the singularity and mystery which appertains to their physical constitution, they have become, through the medium of exact calculation, unexpected instruments of inquiry into points connected with the planetary system itself, of no small importance. We have seen that the movements of the comet of Encke, thus minutely and perseveringly traced by the eminent astronomer whose name is used to distinguish it, has afforded ground for believing in the presence of a resisting medium filling the whole of our system. Similar inquiries, prosecuted in the cases of other periodical comets, will extend, confirm, or modify our conclusions on this head. The perturbations, too, which comets experience in passing near any of the planets, may afford, and have afforded, information as to the magnitude of the disturbing masses, which could not well be otherwise obtained. Thus the approach of this comet to the planet Mercury in 1838 afforded an estimation of the mass of that planet, the more precious, by reason of the great uncertainty under which all previous determinations of that element laboured. Its approach to the same planet in the present year (1848) will be still nearer. On the 22d of November their mutual distance will be only fifteen times the moon's distance from the earth.

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[1] See the announcement of this institution in Astron. Nachr. No. 400.

[2] By Prof. Schumacher, Director of the Royal Observatory of Altona.

(599.) It is, however, in a physical point of view that these bodies offer the greatest stimulus to our curiosity. There is, beyond question, some profound secret and mystery of nature concerned in the phænomenon of their tails. Perhaps it is not too much to hope that future observation, borrowing every aid from rational speculation, grounded on the progress of physical science generally, (especially those branches of it which relate to the æthereal or imponderable elements), may ere long enable us to penetrate this mystery, and to declare whether it is really *matter*, in the ordinary acceptation of the term, which is projected from their heads with such extravagant velocity, and if not impelled, at least *directed* in its course by a reference to the sun, as its point of avoidance. In no respect is the question as to the materiality of the tail more forcibly pressed on us for consideration, than in that of the enormous sweep which it makes round the sun in perihelio, in the manner of a straight and rigid rod, in defiance of the law of gravitation, nay, even of the received laws of motion, extending (as we have seen in the comets of 1680 and 1843) from near the sun's surface to the earth's orbit, yet whirled round unbroken; in the latter case through an angle of 180° in little more than two hours. It seems utterly incredible that in such a case it is one and the same material object which is thus brandished. If there could be conceived such a thing as a *negative shadow*, a momentary impression made upon the luminiferous æther behind the comet, this would represent in some degree the conception such a phænomenon irresistibly calls up. But this of matter and force

(600.) The great number of comets whi
orbits, or orbits at least undistinguishable f
scription of that comparatively small part
bility to us, has given rise to an impression
to our system, wandering through space, an
temporary obedience to its laws during their
may be in this view, we may never have sati
On such an hypothesis, our elliptic comets
ship within the sphere of the sun's predomi
of one or other of the planets near which th
a manner as to diminish their velocity, an
elliptic motion.[1] A similar cause acting the
probability, give rise to a hyperbolic motion.
case, the comet would remain in the system, a
number of revolutions, in the latter it would
possibly be the cause of the exceedingly rare
comet as compared with elliptic ones.

(601.) All the planets without exception, a
circulate in one direction. Retrograde comet
mon occurrence, which certainly would go to
at least an independent origin. Laplace, fron
cometary orbits known in the earlier part o
cluded that the mean or average situation of th
orbits, with respect to the ecliptic, was so nea
as to afford no presumptio

and of these only two, viz. Halley's and the great comet of 1843, can be regarded as satisfactorily made out. Finally, of the 125 comets whose elements are given in the collection of Schumacher and Olbers, up to 1823, the number of retrograde comets under 10° of inclination is only 2 out of 9, and under 20°, 7 out of 23. A plane of motion, therefore, nearly coincident with the ecliptic, and a periodical return, are circumstances eminently favourable to direct revolution in the cometary as they are decisive among the planetary orbits. [Here also we may notice a very curious remark of Mr. Hind, (Ast. Nachr. No. 724,) respecting periodic comets, viz., that, so far as at present known, they divide themselves for the most part into two families, — the one having periods of about 75 years, corresponding to a mean distance about that of Uranus; the other corresponding more nearly with those of the asteroids, and with a mean distance between those small planets and Jupiter. The former group consists of four members, Halley's comet revolving in 76 years, one discovered by Olbers in 74, De Vico's 4th comet in 73, and Brorsen's 3d in 75, respectively. Examples of the latter group are to be seen in the table, p. 552, at the end of this volume. It may be added, also, that one or two of the asteroids are described as having a faint nebulous envelope about them, indicating somewhat of a cometic nature.]

PAR

OF THE LUNAR AND PLA

"Magnus ab integro sæclorum n

CHAPTE

SUBJECT PROPOUNDED.—PROBLEM OF
OF SMALL MOTIONS.—ESTIMATION
ITS GEOMETRICAL REPRESENTATION.
PARTICULAR CASES.—RESOLUTION
NENTS.—RADIAL, TRANSVERSAL, A
FORCES.—NORMAL AND TANGENTI
EFFECTS.—EFFECTS OF THE ORTHOGO
NODES.—CONDITIONS OF THEIR ADV
AN EXTERIOR PLANET DISTURBED BY
CASE.—IN EVERY CASE THE NODE
CEDES ON THE PLANE OF THE DIS
COMBINED EFFECT OF MANY SUCH DI
MOON'S NODES.—CHANGE OF INCLI
INCREASE AND DIMINUTION.—AVER
VOLUTION.—COMPENSATION IN A C
NODES.—LAGRANGE'S THEOREM OF
NATIONS OF THE PLANETARY

and theories, than suffice for the establishment and verification of the elliptic theory. These inequalities are known, in physical astronomy, by the name of *perturbations*. They arise, in the case of the primary planets, from the mutual gravitations of these planets towards each other, which derange their elliptic motions round the sun; and in that of the secondaries, partly from the mutual gravitation of the secondaries of the same system similarly deranging their elliptic motions round their common primary, and partly from the unequal attraction of the sun and planets on them and on their primary. These perturbations, although small, and, in most instances, insensible in short intervals of time, yet, when accumulated, as some of them may become, in the lapse of ages, alter very greatly the original elliptic relations, so as to render the same elements of the planetary orbits, which at one epoch represented perfectly well their movements, inadequate and unsatisfactory after long intervals of time.

(603.) When Newton first reasoned his way from the broad features of the celestial motions, up to the law of universal gravitation, as affecting all matter, and rendering every particle in the universe subject to the influence of every other, he was not unaware of the modifications which this generalization would induce upon the results of a more partial and limited application of the same law to the revolutions of the planets about the sun, and the satellites about their primaries, as their *only* centres of attraction. So far from it, his extraordinary sagacity enabled him to perceive very distinctly how several of the most important of the lunar inequalities take their origin, in this more general way of conceiving the agency of the attractive power, especially the retrograde motion of the nodes, and the direct revolution of the apsides of her orbit. And if he did not extend his investigations to the mutual perturbations of the planets, it was not for want of perceiving that such perturbations *must* exist, and *might* go the length of producing great derangements from the actual state of the system, but was owing to the then undeveloped state of the practical part of astronomy, which had not yet attained the precision requisite to make such an attempt inviting, or indeed feasible. What Newton left undone, however, his successors have accomplished; and, at this day, it is hardly too much to assert that there is not a single perturbation, great or small, which observation has become precise enough clearly to detect and place in evidence, which has not been traced up to its origin in the mutual gravitation of the parts of our system, and minutely accounted for, in its numerical amount and value, by strict calculation on Newton's principles.

(604.) Calculations of this nature require a very high analysis for their successful performance, such as is far beyond the scope and object of this work to attempt exhibiting. The reader who would master them must prepare himself for the undertaking by an extensive course of preparatory study, and must ascend by steps which we must not here even digress to point out. It will be our object, in this chapter, however, to give some general insight into the nature and manner of operation of the acting forces, and to point out what are the circumstances which, in some cases, give them a high degree of efficiency—a sort of *purchase* on the balance of the system; while, in others, with no less amount of intensity, their effective agency in producing extensive and lasting changes is compensated or rendered abortive; as well as to explain the nature of those admirable results respecting the stability of our system, to which the researches of geometers have conducted them; and which, under the form of mathematical theorems of great simplicity and elegance, involve the history of the past and future state of the planetary orbits during ages, of which, contemplating the subject in this point of view, we neither perceive the beginning nor the end.

(605.) Were there no other bodies in the universe but the sun and one planet, the latter would describe an exact ellipse about the former (or both round their common centre of gravity), and continue to perform its revolutions in one and the same orbit for ever; but the moment we add to our combination a third body, the attraction of this will draw both the former bodies out of their mutual orbits, and, by acting on them unequally, will disturb their relation to each other, and put an end to the rigorous and mathematical exactness of their elliptic motions, not only about a fixed point in space, but about one another. From this way of propounding the subject, we see that it is not the whole attraction of the newly-introduced body which produces perturbation, but *the difference* of its attractions on the two originally present.

(606.) Compared to the sun, all the planets are of extreme minuteness; the mass of Jupiter, the greatest of them all, being not more than about one 1100th part that of the sun. Their attractions on each other, therefore, are all very feeble, compared with the presiding central power, and the effects of their disturbing forces are proportionally minute. In the case of the secondaries, the chief agent by which their motions are deranged is the sun itself, whose mass is indeed great, but whose disturbing influence is immensely diminished by their near proximity to their primaries, compared to their distances from the sun, which renders the *difference* of attractions on both extremely small, compared to the whole amount. In this case the greatest part of the sun's attraction, viz. that which is common to both, is exerted to retain both primary and secondary in their common orbit about itself, and prevent their parting company.

Only the small overplus of force on one as compared with the other acts as a disturbing power. The mean value of this overplus, in the case of the moon disturbed by the sun, is calculated by Newton to amount to no higher a fraction than $\frac{1}{638000}$ of gravity at the earth's surface, or $\frac{1}{179}$ of the principal force which retains the moon in its orbit.

(607.) From this extreme minuteness of the intensities of the disturbing, compared to the principal forces, and the consequent smallness of their *momentary* effects, it happens that we can estimate each of these effects separately, as if the others did not take place, without fear of inducing error in our conclusions beyond the limits necessarily incident to a first approximation. It is a principle in mechanics, immediately flowing from the primary relations between forces and the motions they produce, that when a number of very minute forces act at once on a system, their joint effect is the sum or aggregate of their separate effects, at least within such limits, that the original relation of the parts of the system shall not have been materially changed by their action. Such effects supervening on the greater movements due to the action of the primary forces may be compared to the small ripplings caused by a thousand varying breezes on the broad and regular swell of a deep and rolling ocean, which run on as if the surface were a plane, and cross in all directions without interfering, each as if the other had no existence. It is only when their effects become accumulated in lapse of time, so as to alter the primary relations or data of the system, that it becomes necessary to have especial regard to the changes correspondingly introduced into the estimation of their momentary efficiency, by which the *rate* of the subsequent changes is affected, and periods or cycles of immense length take their origin. From this consideration arise some of the most curious theories of physical astronomy.

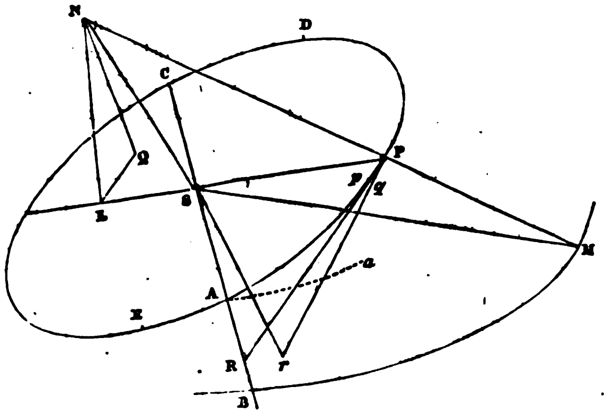
(608.) Hence it is evident, that in estimating the disturbing influence of several bodies forming a system, in which one has a remarkable preponderance over all the rest, we need not embarrass ourselves with combinations of the disturbing powers one among another, unless where immensely long periods are concerned; such as consist of many hundreds of revolutions of the bodies in question about their common centre. So that, in effect, so far as we propose to go into its consideration, the problem of the investigation of the perturbations of a system, however numerous, constituted as ours is, reduces itself to that of a system of three bodies: a predominant central body, a disturbing, and a disturbed; the two latter of which may exchange denominations, according as the motions of the one or the other are the subject of inquiry.

(609.) Both the intensity and direction of the disturbing force are continually varying, according to the relative situation of the disturbing and disturbed body with respect to the sun. If the attraction of the disturbing body M, on the central body S, and the disturbed body P, (by which designations, for brevity, we shall hereafter indicate them,) were equal, and acted in parallel lines, whatever might otherwise be its law of variation, there would be no deviation caused in the elliptic motion of P about S, or of each about the other. The case would be strictly that of art. 454; the attraction of M, so circumstanced, being at every moment exactly analogous in its effects to terrestrial gravity, which acts in parallel lines, and is equally *intense* on all bodies, great and small. But this is not the case of nature. Whatever is stated in the subsequent article to that last cited, of the disturbing effect of the sun and moon, is, *mutatis mutandis*, applicable to every case of perturbation; and it must be now our business to enter, somewhat more in detail, into the general heads of the subject there merely hinted at.

(610.) To obtain clear ideas of the manner in which the disturbing force produces its various effects, we must ascertain at any given moment, and in any relative situations of the three bodies, its direction and intensity as compared with the gravitation of P towards S, in virtue of which latter force alone P would describe an ellipse about S regarded as fixed, or rather P and S about their common centre of gravity in virtue of their mutual gravitation to each other. In the treatment of the problem of three bodies, it is convenient, and tends to clearness of apprehension, to regard one of them as fixed, and refer the motions of the others to it as to a relative centre. In the case of two planets disturbing each other's motions, the sun is naturally chosen as this fixed centre; but in that of satellites disturbing each other, or disturbed by the sun, the centre of their primary is taken as their point of reference, and the sun itself is regarded in the light of a very distant and massive satellite revolving about the primary in a *relative* orbit, equal and similar to that which the primary describes *absolutely* round the sun. Thus the generality of our language is preserved, and when, referring to any particular central body, we speak of an exterior and an interior planet, we include the cases in which the former is the sun and the latter a satellite; as, for example, in the Lunar theory. It is a principle in dynamics, that the relative motions of a system of bodies *inter se* are no way altered by impressing on all of them a common motion or motions, or a common force or forces accelerating or retarding them all equally in common directions, *i. e.* in parallel lines. Suppose, therefore, we apply to all the three bodies, S, P, and M, alike, forces equal to those with which M and P attract S, but in opposite directions. Then will the relative motions both of M and P about S be naltered; but S, being now urged by equal and opposite forces to and om both M and P, will remain at rest. Let us now consider how either f the other bodies, as P, stands affected by these newly-introduced forces, n addition to those which before acted on it. It is clear that now P will e simultaneously acted on by *four* forces; firstly, the attraction *of* S in he direction P S; secondly, an additional force, in the same direction, qual to its attraction *on* S; thirdly, the attraction *of* M in the direction M; and fourthly, a force parallel to M S, and equal to M's attraction n S. Of these, the two first, following the same law of the inverse quare of the distance S P, may be regarded as one force, precisely as if he sum of the masses of S and P were collected in S; and in virtue of heir joint action, P will describe an ellipse about S, except in so far as hat elliptic motion is disturbed by the other two forces. Thus we see hat in this view of the subject the *relative* disturbing force acting on P s no longer the mere single attraction of M, but a force resulting from he composition of that attraction with M's attraction on S transferred to in a contrary direction.

(611.) Let C P A be part of the relative orbit of the disturbed, and M B f the disturbing body, their planes intersecting in the line of nodes S A B,

Fig. 76.



nd having to each other the inclination expressed by the spherical angle A a. In M P, produced if required, take M N : M S :: M S^2 : M P^2. hen, if S M[1] be taken to represent, in quantity and direction, the accelerative attraction of M on S, M S will represent in quantity and direction the new force applied to P, parallel to that line, and N M will represent on the same scale the accelerative attraction of M on P. Consequently, the disturbing force acting on P will be the resultant of two forces applied at P, represented respectively by N M and M S, which by the laws of dynamics are equivalent to a single force represented in *quantity* and *direction* by N S, *but having* P *for its point of application.*

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[1] The reader will be careful to observe the order of the letters, where forces are epresented by lines. M S represents a force acting from M towards S, S M from S owards M.

(612.) The *line* N S, is easily calculated by trigonometry, when the relative situations and real distances of the bodies are known; and the *force* expressed by that line is directly comparable with the attractive forces of S on P by the following proportions, in which M, S, represent the masses of those bodies which are supposed to be known, and to which, at equal distances, their attractions are proportional:—

Disturbing force : M's attraction on S : : N S : S M;

M's attraction on S : S's attraction on M : : M : S;

S's attraction on M : S's attraction on P : : S P^2 : S M^2 : by compounding which proportions we collect as follows:—

Disturbing force : S's attraction on P : : M . N S . S P^2 : S . S M^3.

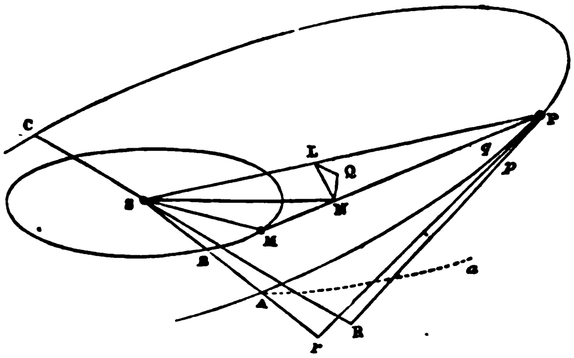
A few numerical examples are subjoined, exhibiting the results of this calculation in particular cases, chosen so as to exemplify its application under very various circumstances, throughout the planetary system. In each case the numbers set down express the proportion in which the central force retaining the disturbed body in its elliptic orbit exceeds the disturbing force, to the nearest whole number. The calculation is made for three positions of the disturbing body—viz. at its greatest, its least, and its mean distance from the disturbed.

Disturbing body.	Disturbed body.	Ratio at the greatest distance : 1.	Ratio at the mean distance : 1.	Ratio at the least distance : 1.
The Sun...........	The Moon.........	90	179	89
Jupiter	Saturn..............	354	312	128
Jupiter	The Earth	95683	147575	53268
Venus	The Earth	255208	210245	26833
Neptune	Uranus.............	57420	56592	5519
Mercury	Neptune...........	526	526	526
Jupiter	Ceres	6433	6937	1033
Saturn	Jupiter.............	20248	21579	3065

(613.) If the orbit of the disturbing body be circular, S M is invariable. In this case, N S will continue to represent the disturbing force *on the same invariable scale*, whatever may be the configuration of the three bodies with respect to each other. If the orbit of M be but little elliptic, the same will be nearly the case. In what follows throughout this chapter, except where the contrary is expressly mentioned, we shall neglect the eccentricity of the disturbing orbit.

(614.) If P be nearer to M than S is, M N is greater than M P, and N lies in M P prolonged, and therefore on the opposite side of the plane of P's orbit from that on which M is situated. The force N S therefore urges P towards M's plane, and towards a point X, situated between S and M, in the line S M. If the distance M P be equal to M S as when P is situated, suppose, at D or E, M N is also equal to M P or M S, so that N coincides with P, and therefore X with S, the disturbing forces being in these cases directed *towards* the central body. But if M P be greater than MS, M N is less than M P, and N lies between M and P, or on the same side of the plane of P's orbit that M is situated on. The force N S, therefore, applied at P, urges P towards the contrary side of that plane towards a point in the line M S produced, so that X now shifts to the farther side of S. In all cases, the disturbing force is wholly effective in the plane M P S, in which the three bodies lie.

Fig. 77.



It is very important for the student to fix distinctly and bear constantly in his mind these relations of the disturbing agency considered as a *single unresolved force*, since their recollection will preserve him from many mistakes in conceiving the mutual actions of the planets, &c. on each other. For example, in the figures here referred to, that of Art. 611 corresponds to the case of a nearer disturbed by a more distant body, as the earth by Jupiter, or the moon by the Sun; and that of the present article to the converse case: as, for instance, of Mars disturbed by the earth. Now, in this latter class of cases, whenever M P is greater than

h lie in the plane of the orbit, not in the direction we have
lial and transversal, i. e. in that of the radius vector PS and
lar to it, but in the direction of a tangent to the orbit at P,
of a normal to the curve, and at right angles to the tangent,
reason these components may be called the *tangential* and
turbing forces. When the orbit of the disturbed body is cir-
early so, this mode of resolution coincides with or differs but
the former, but, when the ellipticity is considerable, these
may deviate from the radial and transversal directions to any
s in the Newtonian mode of resolution, the effect of the one
falls wholly upon the approach and recess of the body P to
body S, and of the other wholly on the rate of description of
round S, so in this which we are now considering, the direct
one component (the normal) falls wholly on the curvature of the
point of its action, increasing that curvature when the normal
nwards, or towards the concavity of the orbit, and diminishing
the opposite direction; while, on the other hand, the tangential
is directly effective on the velocity of the disturbed body, in-
diminishing it according as its direction conspires with or
motion. It is evident enough that where the object is to trace
changes produced by the disturbing force, in *angle* and *distance*
central body, the former mode of resolution must have the

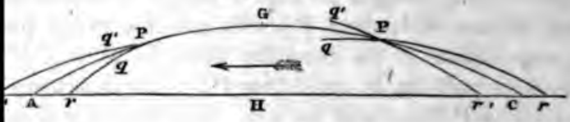
an undisturbed orbit C P. Then at P it would be moving in the direction of a tangent P R to the ellipse P A, which prolonged will intersect the plane of M's orbit somewhere in the line of nodes, as at R. Now, at P, let the disturbing force parallel to N Q act momentarily on P; then P will be deflected in the direction of that force, and instead of the arc P *p*, which it would have described in the next instant if undisturbed, will describe the arc P *q* lying in the state of things represented in Art. 611, below, and in Art. 614, above, P *p* with reference to the plane P S A. Thus, by this action of the disturbing force, the plane of P's orbit will have shifted its position in space from P S *p* (an elementary portion of the old orbit) to P S *q*, one of the new. Now the lines of nodes S A B in the former is determined by prolonging P *p* into the tangent P R, intersecting the plane M S B in R, and joining S R. And in like manner, if we prolong P *q* into the tangent P *r*, meeting the same plane in *r*, and join S *r*, this will be the new line of nodes. Thus we see that, under the circumstances expressed in the former figure, the momentary action of the orthogonal disturbing force will have caused the line of nodes to *retrograde* upon the plane of the orbit of the disturbing body, and under those represented in the latter to advance. And it is evident that the action of the other resolved portions of the disturbing force will not in the least interfere with this result, for neither of them tends either to carry P out of its former plane of motion, or to prevent its quitting it. Their influence would merely go to transfer the points of intersection of the tangents P *p* or P *q* from R or *r* to R′ or *r*′, points nearer to or farther from S than R *r*, but in the same lines.

(621.) Supposing, now, M to lie to the left instead of the right side of the line of nodes in fig. 1., P retaining its situation, and M P being less than M S, so that X shall still lie between M and S. In this situation of things (or *configuration*, as it is termed of the three bodies with respect to each other,) N will lie *below* the plane A S P, and the disturbing force will tend to raise the body P above the plane, the resolved orthogonal portion N Q in this case acting upwards. The disturbed arc P *q* will therefore lie above P *p*, and when prolonged to meet the plane M S B, will intersect it in a point *in advance* of R; so that in this configuration the node will advance upon the plane of the orbit of M, provided always that the latter orbit remains fixed, or, at least, does not itself shift its position in such a direction as to defeat this result.

(622.) Generally speaking, the node of the disturbed orbit will recede upon *any plane which we may consider as fixed*, whenever the action of the orthogonal disturbing force tends to bring the disturbed body nearer to that plane: and *vice versâ*. This will be evident on a mere inspection

exed figure, in which C A represents a semicircle of the proje
fixed plane as seen from S on the sphere of the heavens, and
of the plane of P's undisturbed orbit, the motion of P being
ction of the arrow, from C the ascending, to A the descend-
It is at once seen, by prolonging P q, P q' into arcs of great

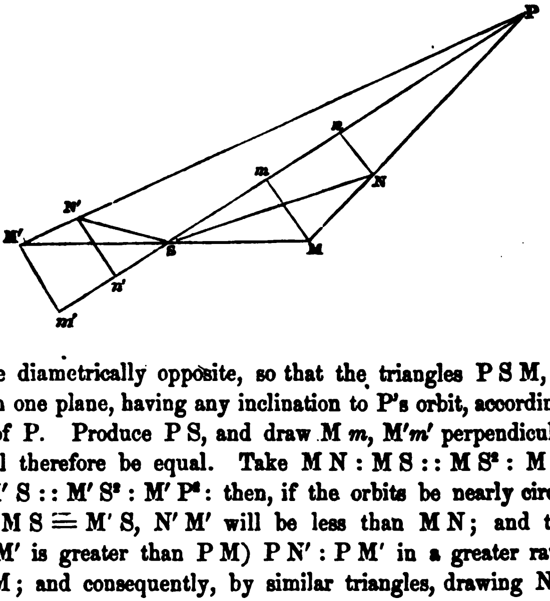
Fig. 78.



r, P r, (forwards or backwards, as the case may be) to meet
the node will have retrograded through the arc A r, or C r,
P q lies between C P A and C A, or when the perturbing force
owards the fixed plane, but will have advanced through A r' or
P q' lies above C P A, or when the disturbing impulse has
ove its old orbit or away from the fixed plane, *and this with-*
ference to whether the undisturbed orbitual motion o P *at the*
carrying it towards the plane C A *or from it,* as in the two
sented in the figure.

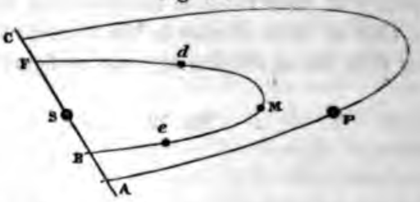
Let us now consider the mutual disturbance of two bodies M both bodies are on the same side of the line of nodes, and recede when on the opposite. They will, therefore, on an average, advance and recede during equal times (supposing the orbits nearly circular). And, therefore, if their advance were at each instant of its duration equally rapid with their recess at each corresponding instant during that phase of the movement, they would merely oscillate to and fro about a mean position, without any permanent motion in either direction. But this is not the case. The rapidity of their recess in every position favourable to recess is greater than that of their advance in the corresponding opposite position. To show this, let us consider any two configurations in which M's phases are diametrically opposite, so that the triangles P S M, P S M', shall lie in one plane, having any inclination to P's orbit, according to the situation of P. Produce P S, and draw M m, M'm' perpendicular to it, which will therefore be equal. Take $MN : MS :: MS^2 : MP^2$, and $M'N' : M'S :: M'S^2 : M'P^2$: then, if the orbits be nearly circles, and therefore $MS = M'S$, N' M' will be less than M N; and therefore (since P M' is greater than P M) P N' : P M' in a greater ratio than P N : P M; and consequently, by similar triangles, drawing N n, N' n perpendicular to P S, N' n' : M' m in a greater ratio than N n : M m, and therefore N' n' is greater than N n. Now the plane P M M' intersects P's orbit in P S, and being inclined to that orbit at the same angle through its whole extent, if from N and N' perpendiculars be conceived let fall on that orbit, these will be to each other in the proportion of N n, N' n'; and therefore the perpendicular from N' will be greater than that from N. Now since by art. 611 N' S and N S represent in quantity and direction the total disturbing forces of M' and M on P respectively, therefore these perpendiculars express (art. 615) the orthogonal disturbing

Fig. 79.



former of which tends (as above shown) to make the node
l the latter to advance; and therefore the preponderance is
pair of situations of M is in favour of a retrograde motion.
Let us next consider the case where the distance between the
ss than the semiaxis of the interior, or in which the least dis-
I from P is less than M S. Take any situation of P with

Fig. 80.



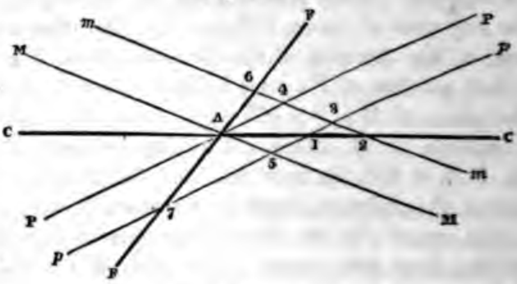
the line of nodes A C. Then two points *d* and *e*, distant by
20°, can be taken on the orbit of M equidistant from P with
se M to occupy successively every possible situation in its orbit,
its place; — then, if it were not for the existence of the arc
ch the relations of art. 624 are reversed, it would appear by
ng of that article that the motion of the node is direct when

will be reversed, yet the effect of this will be much more than counterbalanced by the more energetic and more prolonged retrograde action over *d*B; and, therefore, in this case also, on the average of every possible situation of M, the motion of the node will be retrograde.

(627.) Let us lastly consider an interior planet disturbed by an exterior. Take M D and M E (fig. of art. 611.) each equal to M S. Then first, when P is between D and the node A, being nearer than S to M, the disturbing force acts towards M's orbit on the side on which M lies, and the node recedes. It also recedes when (M retaining the same situation) P is in any part of the arc E C from E to the other node, because in that situation the direction of the disturbing force, it is true, is reversed, but that portion of P's orbit being also reversely situated with respect to the plane of M's, P is still urged towards the latter plane, but on the side opposite to M. Thus, (M holding its place) whenever P is anywhere in D A or E C, the node recedes. On the other hand, it advances whenever P is between A and E or between C and D, because, in these arcs, only one of the two determining elements (viz. the direction of the disturbing force with respect to the plane of P's orbit; and the situation of the one plane with respect to the other as to above and below) has undergone reversal. Now first, whenever M is anywhere but in the line of nodes, the sum of the arcs D A and E C exceeds a semicircle, and *that* the more, the nearer M is to a position at right angles to the line of nodes. Secondly, the arcs favourable to the recess of the node comprehend those situations in which the orthogonal disturbing force is most powerful, and *vice versâ*. This is evident, because as P approaches D or E, this component decreases, and vanishes at those points (612.) The movement of the node itself also vanishes when P comes to the node, for although in this position the disturbing orthogonal force neither vanishes nor changes its direction, yet, since at the instant of P's passing the node (A) the recess of the node is changed into an advance, it must necessarily at that point be stationary.[1] Owing to both these causes, therefore, (that the mode recedes during a longer time than it advances, and that a more energetic force acting in its recess causes it to recede more rapidly,) the retrograde motion will preponderate on the whole in each complete synodic revolution of P. And it is evident that the reasoning of this and the foregoing articles, is no way vitiated by a moderate amount of excentricity in either orbit.

It would seem, at first sight, as if a change *per saltum* took place here, but the continuity of the node's motion will be apparent from an inspection of the annexed figure, where *bad* is a portion of P's disturbed path near the node A, concave towards the plane G A. The momentary place of the moving node is determined by the inter-t is therefore a general proposition, that on the average of each ynodic revolution, the node of every disturbed planet recedes rbit of the disturbing one, or in other words, that in every pair ne node of each recedes upon the other, and of course upon any e plane which we may regard as fixed. On a plane not inter- tween them, however, the node of one orbit will advance, and other will recede. Suppose for instance, C A C to be a plane

Fig. 83.



e between P P and M M the two orbits. If $p\,p$ and $m\,m$ be ositions of the orbits, the node of P on M will have receded intersect each other in a common node. Although perfectly true, therefore, that the node of any one planet would recede on the orbit of any and each other by the individual action of that other, yet, when all act together, recess on one plane may be equivalent to advance on another, so that the motion of the node of any one orbit on a given plane, arising from their joint action, taking into account the different situations of all the planes, becomes a curiously complicated phænomenon whose law cannot be very easily expressed in words, though reducible to strict numerical statement, being, in fact, a mere geometrical result of what is above shown.

(630.) The nodes of all the planetary orbits on the *true* ecliptic, as a matter of fact, are retrograde, though they are not all so on a fixed plane, such as we may conceive to exist in the planetary system, and to be a plane of reference unaffected by their mutual disturbances. It is, however, to the ecliptic, that we are under the necessity of referring their movements from our station in the system; and if we would transfer our ideas to a fixed plane, it becomes necessary to take account of the variation of the ecliptic itself, produced by the joint action of all the planets.

(631.) Owing to the smallness of the masses of the planets, and their great distances from each other, the revolutions of their nodes are excessively slow, being in every case less than a single degree per century, and in most cases not amounting to half that quantity. It is otherwise with the moon, and that owing to two distinct reasons. First, that the disturbing force itself arising from the sun's action, (as appears from the table given in art. 612,) bears a much larger proportion to the earth's central attraction on the moon than in the case of any planet disturbed by any other. And secondly, because the synodic revolution of the moon, within which the average is struck (and always on the side of recess), is only 29½ days, a period much shorter than that of any of the planets, and vastly so than that of several among them. All this is agreeable to what has already been stated (art. 407, 408,) respecting the motion of the moon's nodes, and it is hardly necessary to mention that, when calculated, as it has been, *à priori*, from an exact estimation of all the acting forces, the result is found to coincide with perfect precision with that immediately derived from observation, so that not a doubt can subsist as to this being the real process by which so remarkable an effect is produced.

(632.) So far as the physical condition of each planet is concerned, it is evident that the position of their nodes can be of little importance. It is otherwise with the mutual inclinations of their orbits with respect to each other, and to the equator of each. A variation in the position of the ecliptic, for instance, by which its pole should shift its distance from the

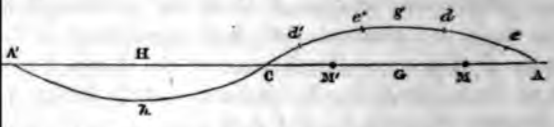
e equator, would disturb our seasons. Should the plane of the
is, for instance, ever be so changed as to bring the ecliptic to
ith the equator, we should have perpetual spring over all the
d on the other hand, should it coincide with a meridian, the
f summer and winter would become intolerable. The inquiry,
e variations of inclination of the planetary orbits *inter se*, is
ch higher practical interest than those of their nodes.

Referring to the figures of art. 610, *et seq.*, it is evident that
S P q, in which the disturbed body moves during an instant of
its quitting P, is differently inclined to the orbit of M, or to a
, from the original or undisturbed plane P S p. The difference
e position of these two planes in space is the angle between the
R and P S r, and is therefore calculable by spherical trigono-
en the angle R S r or the momentary recess of the node is
d also the inclination of the planes of the orbits to each other.
ive, then, that between the momentary change of inclination,
omentary recess of the node, there exists an intimate relation,
he research of the one is in fact bound up in that of the other.
be, perhaps, made clearer, by considering the orbit of P to be
an imaginary line, but an actual circle or elliptic hoop of some
rial, without inertia, on which, as on a wire, the body P may

arc P Q lies between the planes C A and C G A of the two orbits, and *vice versâ*, indifferently whether P be in the act of receding from the plane C A, as in the quadrant C G, or of approaching to it, as in G A, yet the same identity as to the character of the change does not subsist in respect of the inclination. The inclination of the disturbed orbit (*i. e.* of its momentary element) Pq or Pq', is measured by the spherical angle PrH or Pr'H. Now in the quadrant C G, PrH is less, and Pr'H greater than P C H; but in G A, the converse. Hence this rule:—1st, If the disturbing force urge P towards the plane of M's orbit, and the undisturbed motion of P carry it also towards that plane; and 2dly, if the disturbing force urge P from that plane, while P's undisturbed motion also carries it from it, in either case the inclination momentarily increases; but if, 3dly, the disturbing force act to, and P's motion carry it from — or if the force act from, and the motion carry it to, that plane, the inclination momentarily diminishes. Or (including all the cases under one alternative) if the action of the disturbing force and the undisturbed motion of P with reference to the plane of M's orbit be of the same character, the inclination increases; if of contrary characters, it diminishes.

(636.) To pass from the momentary changes which take place in the relations of nature to the accumulated effects produced in considerable lapses of time by the continued action of the same causes, under circumstances varied by these very effects, is the business of the integral calculus. Without going into any calculations, however, it will be easy for us to demonstrate from the principles above laid down, the leading features of this part of the planetary theory, viz. the periodic nature of the change of the inclinations of two orbits to each other, the re-establishment of their original values, and the consequent oscillation of each plane about a certain mean position. As in explaining the motion of the nodes, we will commence, as the simplest case, with that of an exterior planet disturbed by an interior one at less than half its distance from the central body. Let A C A′ be the great circle of the heavens into which M's orbit seen from S is projected, extended into a straight line, and AgChA′ the corresponding projection of the orbit of P so seen. Let M occupy some fixed situation, suppose in the semicircle A C, and let P describe a complete revolution from A through g C h to A′. Then while it is between A and g or in its first quadrant, its motion is *from* the plane of M's orbit, and at the same time the orthogonal force acts *from* that plane: the inclination, therefore, (art. 635) increases. In the second quadrant the motion of P is *to*, but the force continues to act from, the plane, and the inclination again decreases. A similar alternation takes place in its course through the quadrants Ch and hA. Thus the plane of P's orbit oscil-

Fig. 84.



d fro about its mean position twice in each revolution of P. s process if M held a fixed position at G, the forces being lly alike on either side, the extent of these oscillations would be al, and the inclination at the end of one revolution of P would isely to its original value. But if M be elsewhere, this will case, and in a single revolution of P, only a partial compensa- operated, and an overplus on the side, suppose of diminution, outstanding. But when M comes to M′, a point equidistant the other side, this effect will be precisely reversed (supposing circular). On the average of both situations, therefore, the e the same as if M were divided into two equal portions, one I and the other at M′, which will annihilate the preponderance and effect a perfect restoration. And on an average of all uations of M, the effect will in like manner be the same as if re distributed over the whole circumference of its orbit, forming configuration as to operate an almost exact compensation. But in 1 revolution of M, or 13 of P, the node itself has shifted $\frac{13}{249}$ or about $\frac{1}{19}$ of a revolution, in a direction opposite to the revolutions of M and P, so that although P has been brought back to the same configuration with respect to M, both are $\frac{1}{19}$ of a revolution in advance of the same configuration as respects the node. The compensation, therefore, will not be exact, and to make it so, this process must be gone through 19 times, at the end of which both the bodies will be restored to the same relative position, not only with respect to each other, but to the node. The fractional parts of entire revolutions, which in this explanation have been neglected, are evidently no farther influential than as rendering the compensation thus operated in a revolution of the node slightly inexact, and thus giving rise to a compound period of greater duration, at the end of which a compensation almost mathematically rigorous, will have been effected.

(639.) It is clear then, that if the orbits be circles, the lapse of a very moderate number of revolutions of the bodies will very nearly, and that of a revolution of the node almost exactly, bring about a perfect restoration of the inclinations. If, however, we suppose the orbits excentric, it is no less evident, owing to the want of symmetry in the distribution of the forces, that a perfect compensation will not be effected either in one or in any number of revolutions of P and M, independent of the motion of the node itself, as there will always be some configuration more favourable to either an increase of inclination than its opposite is unfavourable. Thus will arise a change of inclination which, were the nodes and apsides of the orbits fixed, would be always progressive in one direction until the planes were brought to coincidence. But, 1st, half a revolution of the nodes would of itself reverse the direction of this progression by making the position in question favour the opposite movement of inclination; aud, 2dly, the planetary apsides are themselves in motion with unequal velocities, and thus the configuration whose influence destroys the balance, is, itself, always shifting its place on the orbits. The variations of inclination dependent on the excentricities are therefore, like those independent of them, periodical, and being, moreover, of an order more minute (by reason of the smallness of the excentricities) than the latter, it is evident that the total variation of the planetary inclinations must fluctuate within very narrow limits. Geometers have accordingly demonstrated by an accurate analysis of all the circumstances, and an exact estimation of the acting forces, that such is the case; and this is what is meant by asserting the stability of the planetary system as to the mutual inclinations of its orbits. By the researches of Lagrange (of whose analytical conduct it is impossible here to give any idea,) the following elegant theorem has been demonstrated:—

"*If the mass of every planet be multiplied by the square root of the major axis of its orbit, and the product by the square of the tangent of its inclination to a fixed plane, the sum of all these products will be constantly the same under the influence of their mutual attraction.*" If the present situation of the plane of the ecliptic be taken for that fixed plane (the ecliptic itself being variable like the other orbits), it is found that this sum is actually very small: it must, therefore, always remain so. This remarkable theorem alone, then, would guarantee the stability of the orbits of the greater planets; but from what has above been shown of the tendency of each planet to work out a compensation on every other, it is evident that the minor ones are not excluded from this beneficial arrangement.

(640.) Meanwhile, there is no doubt that the plane of the ecliptic does actually vary by the actions of the planets. The amount of this variation is about 48″ per century, and has long been recognized by astronomers, by an increase of the latitudes of all the stars in certain situations, and their diminution in the opposite regions. Its effect is to bring the ecliptic by so much per annum nearer to coincidence with the equator; but from what we have above seen, this diminution of the obliquity of the ecliptic will not go on beyond certain very moderate limits, after which (although in an immense period of ages, being a compound cycle resulting from the joint action of all the planets,) it will again increase, and thus oscillate backward and forward about a mean position, the extent of its deviation to one side and the other being less than 1° 21′.

(641.) One effect of this variation of the plane of the ecliptic,—that which causes its nodes on a fixed plane to change,—is mixed up with the precession of the equinoxes, and undistinguishable from it, except in theory. This last-mentioned phænomenon is, however, due to another cause, analogous, it is true, in a general point of view, to those above considered, but singularly modified by the circumstances under which it is produced. We shall endeavour to render these modifications intelligible, as far as they can be made so without the intervention of analytical formulæ.

(642.) The precession of the equinoxes, as we have shown in art. 312, consists in a continual retrogradation of the node of the earth's equator on the ecliptic; and is, therefore, obviously an effect so far analogous to the general phænomenon of the retrogradation of the nodes of the orbits on each other. The immense distance of the planets, however, compared with the size of the earth, and the smallness of their masses compared to that of the sun, puts *their* action out of the question in the inquiry of its cause, and we must, therefore, look to the massive though distant sun, and to our near though minute neighbour, the moon, for its explanation. This will, accordingly, be found in their disturbing action on the redundant matter accumulated on the equator of the earth, by which its figure is rendered spheroidal, combined with the earth's rotation on its axis. It is to the sagacity of Newton that we owe the discovery of this singular mode of action.

(643.) Suppose in our figure (art. 611,) that instead of one body, P, revolving round S, there were a succession of particles not coherent, but forming a kind of fluid ring, free to change its form by any force applied. Then, while this ring revolved round S in its own plane, under the disturbing influence of the distant body M, (which now represents the moon or the sun, as P does one of the particles of the earth's equator,) two things would happen: 1st, its figure would be bent out of a plane into an undulated form, those parts of it within the arcs D A and E C being rendered more inclined to the plane of M's orbit, and those within the arcs A E, C D, less so than they would otherwise be; 2dly, the nodes of this ring, regarded as a whole, without respect to its change of figure, would retreat upon that plane.

(644.) But suppose this ring, instead of consisting of discrete molecules free to move independently, to be rigid and incapable of such flexure, like the *hoop* we have supposed in art. 633, but having inertia, then it is evident that the effort of those parts of it which tend to become more inclined will act through the medium of the ring itself (as a mechanical engine or lever) to counteract the effort of those which have *at the same instant* a contrary tendency. In so far only, then, as there exists an excess on the one or the other side will the inclination change, an average being struck at every moment of the ring's motion; just as was shown to happen in the view we have taken of the inclinations, in every complete revolution of a single disturbed body, under the influence of a fixed disturbing one.

(645.) Meanwhile, however, the nodes of the rigid ring will retrograde, the general or average tendency of the nodes of every molecule being to do so. Here, as in the other case, a struggle will take place by the counteracting efforts of the molecules contrarily disposed, propagated through the solid substance of the ring; and thus at every instant of time, an average will be struck, which being identical in its nature with that effected in the complete revolution of a single disturbed body, will, in every case, be in favour of a recess of the node, save only when the disdy, be it sun or moon, is situated in the plane of the earth's

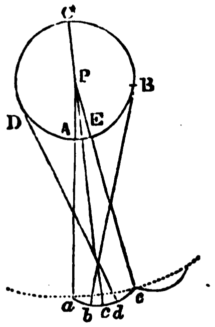
This reasoning is evidently independent of any consideratio
se which maintains the rotation of the ring; whether the par-
nall satellites retained in circular orbits under the equilibrated
ttractive and centrifugal forces, or whether they be small masses
ıs attached to a set of imaginary spokes, as of a wheel, center-
nd free only to shift their planes by a motion of those spokes
lar to the plane of the wheel. This makes no difference in
l effect; though the different velocities of rotation, which may
ed on such a system, may and will have a very great influence
e absolute and relative magnitudes of the two effects in ques-
motion of the nodes and change of inclination. This will be
rstood, if we suppose the ring *without* a rotatory motion, in which
se it is obvious that so long as M remained fixed there would
no recess of nodes at all, but only a tendency of the ring to tilt
ound a diameter perpendicular to the position of M, bringing it
e line S M.

The motion of such a ring, then, as we have been considering,
ate, so far as the recess of the nodes goes, the precession of the
only that its nodes would retrograde far more rapidly than the
recession which is excessively slow. But now conceive th

revolution this great inert mass, will have their velocity of retrogradation proportionally diminished. Thus, then, it is easy to conceive how a motion similar to the precession of the equinoxes, and, like it, characterized by extreme slowness, will arise from the causes in action.

(648.) Now a recess of the node of the earth's equator, upon a given plane, corresponds to a conical motion of its axis round a perpendicular to that plane. But in the case before us, that plane is not the ecliptic, but the moon's orbit for the time being; and it may be asked how we are to reconcile this with what is stated in art. 317 respecting the nature of the motion in question. To this we reply, that the nodes of the lunar orbit, being in a state of continual and rapid retrogradation, while its inclination is preserved nearly invariable, the point in the sphere of the heavens round which the pole of the earth's equator revolves (with that extreme slowness characteristic of the precession) is itself in a state of continual circulation round the pole of the ecliptic, with that much more rapid motion which belongs to the lunar node. A glance at the annexed figure will explain this better than words. P is the pole of the ecliptic, A the pole of the moon's orbit, moving round the small circle A B C D in 19 years; *a* the pole of the earth's equator, which at each moment of its progress has a *direction* perpendicular to the varying position of the line A *a*, and a *velocity* depending on the varying intensity of the acting causes during the period of the nodes. This velocity however being extremely small, when A comes to B, C, D, E, the line A *a* will have taken up the positions B *b*, C *c*, D *d*, E *e*, and the earth's pole *a* will thus, in one tropical revolution of the node, have arrived at *e*, having described not an exactly circular arc *a e*, but a single undulation of a wave-shape or epicycloidal curve, *a b c d e*, with a velocity alternately greater and less than its mean

Fig. 85.



d this will be repeated in every succeeding revolution of the

Now this is precisely the kind of motion which, as we have
. 325, the pole of the earth's equator really has round the pole
ptic, in consequence of the joint effects of precession and nuta-
n are thus uranographically represented. If we superadd to the
unar precession that of the solar, which alone would cause the
cribe a circle uniformly about P, this will only affect the undu-
our waved curve, by extending them in length, but will produce
n the depth of the waves, or the excursions of the earth's axis
n the pole of the ecliptic. Thus we see that the two phenomena
n and precession are intimately connected, or rather both of
tial constituent parts of one and the same phenomenon. It is
essary to state that a rigorous analysis of this great problem, by
timation of all the acting forces and summation of their ef-
fects, leads to the precise value of the co-efficients of precession
on, which observation assigns to them. The solar and lunar
the precession of the equinoxes, that is to say, those portions
uniform, are to each other in the proportion of about 2 to 5.

In the nutation of the earth's axis we have an example (the
kind which has occurred to us), of a periodical movement in
f the system, giving rise to a motion having the same precis

The system may be favourably or unfavourably constituted for such a transfer of periodic movements, or favourably in some of its parts and unfavourably in others; and accordingly as it is the one or the other, the *derivative* oscillation (as it may be termed) will be imperceptible in one case, of appreciable magnitude in another, and even more perceptible in its visible effects than the original cause in a third; of this last kind we have an instance in the moon's acceleration, to be hereafter noticed.

(651.) It so happens that our situation on the earth, and the delicacy which our observations have attained, enable us to make it as it were an instrument to *feel* these forced vibrations,—these derivative motions, communicated from various quarters, especially from our near neighbour, the moon, much in the same way as we detect, by the trembling of a board beneath us, the secret transfer of motion by which the sound of an organ-pipe is dispersed through the air, and carried down into the earth. Accordingly, the monthly revolution of the moon, and the annual motion of the sun, produce, each of them, small *nutations* in the earth's axis, whose periods are respectively half a month and half a year, each of which, in this view of the subject, is to be regarded as one portion of a period consisting of two equal and similar parts. But the most remarkable instance, by far, of this propagation of periods, and one of high importance to mankind, is that of the tides, which are forced oscillations, excited by the rotation of the earth in an ocean disturbed from its figure by the varying attractions of the sun and moon, each revolving in its own orbit, and propagating its own period into the joint phenomenon. The explanation of the tides, however, belongs more properly to that part of the general subject of perturbations which treats of the action of the radial component of the disturbing force, and is therefore postponed to a subsequent chapter.

CHAPTER XIII.

OF THE AXES, PERIHELIA, AND EXCENTRICITIES.

OF ELEMENTS IN GENERAL.—DISTINCTION BETWEEN PE-
ND SECULAR VARIATIONS.—GEOMETRICAL EXPRESSION OF
TIAL AND NORMAL FORCES.—VARIATION OF THE MAJOR
ODUCED ONLY BY THE TANGENTIAL FORCE.—LAGRANGE'S
M OF THE CONSERVATION OF THE MEAN DISTANCES AND
—THEORY OF THE PERIHELIA AND EXCENTRICITIES.—
RICAL REPRESENTATION OF THEIR MOMENTARY VARIA-
-ESTIMATION OF THE DISTURBING FORCES IN NEARLY
R ORBITS.—APPLICATION TO THE CASE OF THE MOON.—
OF THE LUNAR APSIDES AND EXCENTRICITY.—EXPERI-
ILLUSTRATION.—APPLICATION OF THE FOREGOING PRIN-
TO THE PLANETARY THEORY.—COMPENSATION IN ORBITS

the time being) of the disturbed orbit, and which tend to derange the elliptic form of the orbit, and the laws of elliptic motion in that plane. The small inclination, generally speaking, of the orbits of the planets and satellites to each other, permits us to separate these effects in theory one from the other, and thereby greatly to simplify their consideration. Accordingly, in what follows, we shall throughout neglect the mutual inclination of the orbits of the disturbed and disturbing bodies, and regard all the forces as acting and all the motions as performed in one plane.

(653.) In considering the changes induced by the mutual action of two bodies, in different aspects with respect to each other, on the magnitudes and forms of their orbits, and in their positions therein, it will be proper in the first instance to explain the conventions under which geometers and astronomers have alike agreed to use the language and laws of the elliptic system, and to continue to apply them to disturbed orbits, although those orbits so disturbed are no longer, in mathematical strictness, ellipses, or any known curves. This they do, partly on account of the convenience of conception and calculation which attaches to this system, but much more for this reason, — that it is found, and may be demonstrated from the dynamical relations of the case, that the departure of each planet from its ellipse, as determined at any epoch, is capable of being truly represented, by supposing the ellipse itself to be slowly variable, to change its magnitude and excentricity, and to shift its position and the plane in which it lies according to certain laws, while the planet all the time continues to move in this ellipse, just as it would do if the ellipse remained invariable and the disturbing forces had no existence. By this way of considering the subject, the whole effect of the disturbing forces is regarded as thrown upon the orbit, while the relations of the planet to that orbit remain unchanged. This course of procedure, indeed, is the most natural, and is in some sort forced upon us by the extreme slowness with which the variation of the elements, at least where the planets only are concerned, develop themselves. For instance, the fraction expressing the excentricity of the earth's orbit changes no more than 0.00004 in its amount in a *century;* and the place of its perihelion, as referred to the sphere of the heavens, by only 19′ 39″ in the same time. For several years, therefore, it would be next to impossible to distinguish between an ellipse so varied and one that had not varied at all; and in a single revolution, the difference between the original ellipse and the curve really represented by the varying one, is so excessively minute, that, if accurately drawn on a table, six feet in diameter, the nicest examination with microscopes, continued along the whole outlines of the two curves, would hardly detect any perceptible interval between them. Not to call a mo-

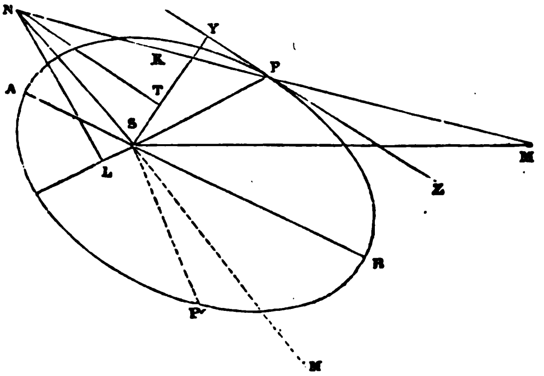
utely conforming itself to an elliptic curve, *elliptic*, would be
even granting the existence of trivial departures alternately on
on the other; though on the other hand, to neglect a varia-
continues to accumulate from age to age, till it forces itself on
would be wilful blindness.

eometers, then, have agreed, in each single revolution, or for
te interval of time, to regard the motion of each planet as
l performed according to Kepler's laws, with a reserve in
hose very small and transient fluctuations which take place
time, but at the same time to regard all the *elements* of each
a continual, though extremely slow, state of change; and, in
effects of perturbation on the system, they take account prin-
ntirely, of this change of the elements, as that upon which any
ange in the great features of the system will ultimately depend.
nd here we encounter the distinction between what are termed
ations, and such as are rapidly periodic, and are compensated
tervals. In our exposition of the variation of the inclination
bed orbit (art. 636,) for instance, we showed that, in each
ution of the disturbed body, the plane of its motion underwent
to and fro in its inclination to that of the disturbing body,
ly compensated each other; leaving, however, a portion out-
hich again is nearly compensated by the revolution of the dis-

drawn from its orbit (its slowly varying orbit,) but forthwith returns to it, to deviate presently as much the other way, while the varied orbit accommodates and adjusts itself to the average of these excursions on either side of it; and thus continues to present, for a succession of indefinite ages, a kind of medium picture of all that the planet has been doing in their lapse, in which the expression and the character is preserved; but the individual features are merged and lost. These periodic inequalities, however, are, as we have observed, by no means neglected, but it is more convenient to take account of them by a separate process, independent of the secular variations of the elements.

(657.) In order to avoid complication, while endeavouring to give the reader an insight into both kinds of variations, we shall henceforward conceive all the orbits to lie in one plane, and confine our attention to the case of two only, that of the disturbed and disturbing body, a view of the subject which (as we have seen) comprehends the case of the moon disturbed by the sun, since any one of the bodies may be regarded as fixed at pleasure, provided we conceive all its motions transferred in a contrary direction to each of the others. Let therefore A P B be the undisturbed elliptic orbit of a planet P; M a disturbing body, join M P, and supposing M K = M S. take M N : M K :: M K^2 : M P^2. Then if S N be joined, N S will represent the disturbing force of M or P, on the same scale that S M represents M's attraction on S. Suppose Z P Y a tangent at P, S Y perpendicular to it, and N T, N L perpendicular respectively to S Y and P S produced. Then will N T represent the tangential, T S the normal, N L the transversal, and L S the radial components of the disturbing force. In circular orbits or orbits only slightly elliptic, the directions

Fig. 86.



SY are nearly coincident, and the former pair of forces will
slightly from the latter. We shall here, however, take the
se, and proceed to investigate in an elliptic orbit of any degree
icity the momentary changes produced by the action of the dis-
rce in those elements on which the magnitude, situation, and
e orbit depend (*i. e.* the length and position of the major axis
xcentricity,) in the same way as in the last chapter we deter-
momentary changes of the inclination and node similarly pro-
he orthogonal force.

We shall begin with the momentary variation in the *length* of
n element of the first importance, as on it depends (art. 487)
ic time and mean angular motion of the planet, as well as the
pply of light and heat it receives in a given time from the sun,
nent or constantly progressive change in which would alter
rially the conditions of existence of living beings on its surface.
a property of elliptic motion performed under the influence of
d in conformity with Kepler's laws, that if the velocity with
anet moves at any point of its orbit be given, and also the
that point from the sun, the major axis of the orbit is thereby

It is no matter in what *direction* the planet may be moving
ment. This will influence the excentricity and the position of
but not its length. This property of elliptic motion

not on the new *direction* impressed upon it by the acting forces,—for the axis, as we have seen, is independent of that direction,—not on its change of distance from the sun, while describing the former arc,—for the elements of that arc are accommodated to it, so that one and the same axis must belong to its beginning and its end. The question, in short, whether in the next arc it shall take up a new major axis or go on with the old one will depend solely on this—whether its *velocity* has or has not undergone a change by the action of the *disturbing* force. For the central force residing in the focus can impress on it no such change of velocity as to be incompatible with the permanence of its ellipse, seeing that it is by the action of that force that the velocity is maintained in that due proportion to the distance which elliptic motion, as such, requires.

(659.) Thus we see that the momentary variation of the major axis depends on nothing but the momentary deviation from the law of elliptic velocity produced by the disturbing force, without the least regard to the direction in which that extraneous velocity is impressed, or the distance from the sun at which the planet may be situated, at the moment of its impression. Nay, we may even go farther, for, as this holds good at every instant of its motion, it will follow that after the lapse of any time, however great, the total amount of change which the axis may have undergone will be determined only by the total deviation produced by the action of the disturbing force in the velocity of the disturbed body from that which it would have had in its undisturbed ellipse, at the same distance from the centre, and that therefore the total amount of change produced in the axis in any lapse of time may be estimated, if we know at every instant the efficacy of the disturbing force to alter the velocity of the body's motion, and *that* without any regard to the alterations which the action of that force may have produced in the other elements of the motion in the same time.

(660.) Now it is not the whole disturbing force which is effective in changing P's velocity, but only its tangential component. The normal component tends merely to alter the curvature of the orbit or to deflect it into conformity with a circle of curvature of greater or lesser radius, as the case may be, and in no way to alter the velocity. Hence it appears that the *variation of the length of the axis is due entirely to the tangential force, and is quite independent on the normal.* Now it is easily shown that as the velocity increases, the axis increases (the distance remaining unaltered[1]) though not in the same exact proportion. Hence it follows

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[1] If a be the semiaxis, r the radius vector, and v the velocity of P in any point of an ellipse, a is given by the relation $v^2=\frac{2}{r}-\frac{1}{a}$, the units of velocity and force being properly assumed.

tangential disturbing force conspires with the motion of P, its
y action increases the axis of the disturbed orbit, whatever be
on of P in its orbit, and *vice versâ.*

Let A S B (fig. art. 657) be the major axis of the ellipse A P
the opposite side of A B take two points P′ and M′, similarly
ith respect to the axis with P and M on their side. Then if at
bodies equal to P and M be placed, the forces exerted by M′
S will be equal to those exerted by M on P and S, and there-
ngential disturbing force of M′ on P′ exerted in the direction
pose) will equal that exerted by M on P in the direction P M.
re (supposing it to revolve in the same direction round S as P)
arded (or accelerated, as the case may be) by *precisely* the same
hich P is accelerated (or retarded), so that the variation in the
respective orbits of P and P′ will be equal in amount, but con-
haracter. Suppose now M's orbit to be circular. Then (*if the*
imes of M *and* P *be not commensurate, so that a moderate*
revolutions may bring them back to the same precise relative
it will necessarily happen, that in the course of a very great
revolutions of both bodies, P will have been presented to M
e of the axis, at some one moment, in the same manner as at
r moment on the other. Whatever variation may have been
its axis in the one situation will have been reversed in that

from what has been said, it is clear that this will be destroyed, during M's approach to S in the other half of its orbit, so that here again, on the average of a multitude of revolutions during which P *has been* presented to M *in every situation for every distance of* M *from* S, the restoration will be effected.

(663.) If neither P's nor M's orbit be circular, and if moreover the directions of their axes be different, this reasoning, drawn from the symmetry of their relations to each other, does not apply, and it becomes necessary to take a more general view of the matter. Among the fundamental relations of dynamics, relations which presuppose no particular law of force like that of gravitation, but which express in general terms the results of the action of *force* on *matter* during *time*, to produce or change *velocity*, is one usually cited as the "Principle of the conservation of the *vis viva*," which applies directly to the case before us. This principle (or rather this theorem) declares that if a body subjected at every instant of its motion to the action of forces directed to fixed centres (no matter how numerous), and having their intensity dependent only on the distances from their respective centres of action, travel from one point of space to another, the velocity which it has on its arrival at the latter point will differ from that which it had on setting out from the former, by a quantity depending only on the different relative situations of these two points in space, without the least reference to the form of the curve in which it may have moved in passing from one point to the other, whether that curve have been described freely under the simple influence of the central forces, or the body have been compelled to glide upon it, as a bead upon a smooth wire. Among the forces thus acting may be included any constant forces, acting in parallel directions, which may be regarded as directed to fixed centres infinitely distant. It follows from this theorem, that, if the body return to the point P, from which it set out, its velocity of arrival will be the same with that of its departure; a conclusion which (for the purpose we have in view) sets us free from the necessity of entering into any consideration of the laws of the disturbing force, the change which its action may have induced in the form of the orbit of P, or the successive steps by which velocity generated at one point of its intermediate path is destroyed at another, by the reversed action of the tangential force. Now to apply this theorem to the case in question, let M be supposed to retain a fixed position during one whole revolution of P. P then is acted on, during that revolution, by three forces: 1st. by the central attraction of S directed always to S; 2nd. by that to M, always directed to M; 3rd. by a force equal to M's attraction on S; but in the direction M S, which therefore is a constant force, acting always in parallel

On completing its revolution, then, P's velocity, and therefore
axis of its orbit, will be found unaltered, at least neglecting
ively minute difference which will result from the non-arrival
lution at the *exact* point of its departure by reason of the pe-
in the orbit produced in the interim by the disturbing force,
he present we may neglect.

Now suppose M to revolve, and it will appear, by a reasoning
imilar to that of art. 662, that whatever uncompensated varia-
velocity arises in successive revolutions of P during M's recess
ll be destroyed by contrary uncompensated variations arising
approach. Or, more simply and generally thus: whatever M's
ay be, for every place which P can have, there must exist some
e of P (as P′), in which the action of M shall be precisely
Now *if the periods be incommensurable*, in an indefinite
revolutions of both bodies, for every possible combination of
M, P) there will occur, *at some time or other*, the combination
hich neutralizes the effect of the other, when carried to the
ount; so that ultimately, and when very long periods of time
ed, a complete compensation will be found to be worked out.

This supposes, however, that in such long periods the orbit of
o altered as to render the occurrence of the compensating situa-
P′) impossible. This would be the case if M's orbit were to

mean points from which they never greatly depart. And if the arrival of either of them at P, at a point *nearer* S, at the end of a complete revolution, cause an *excess* of velocity, its arrival at a more distant point will cause a deficiency, and thus, as the fluctuations of distance to and fro ultimately balance each other, so will also the excesses and defects of velocity, though in periods of enormous length, being no less than that of a complete revolution of P's apsides for the one cause of inequality, and of a complete restoration of its excentricity for the other.

(667.) The dynamical proposition on which this reasoning is based is general, and applies equally well to cases wherein the forces act in one plane, or are directed to centres anywhere situated in space. Hence, if we take into consideration the inclination of P's orbit to that of M, the same reasoning will apply. Only that in this case, upon a complete revolution of P, the variation of inclination and the motion of the nodes of P's orbit will prevent its returning to a point in the exact *plane* of its original orbit, as that of the excentricity and perihelion prevent its arrival at the same exact *distance* from S. But since it has been shown that the inclination fluctuates round a mean state from which it never departs much, and since the node revolves and makes a complete circuit, it is obvious that in a complete period of the latter the points of arrival of P at the same longitude will deviate as often and by the same quantities above as below its original point of departure from exact coincidence; and, therefore, that on the average of an infinite number of revolutions, the effect of this cause of non-compensation will also be destroyed.

(668.) It is evident, also, that the dynamical proposition in question being general, and applying equally to any *number* of fixed centres, as well as to any distribution of them in space, the conclusion would be precisely the same whatever be the number of disturbing bodies, only that the periods of compensation would become more intricately involved. We are, therefore, conducted to this most remarkable and important conclusion, viz. that the major axes of the planetary (and lunar) orbits, and, consequently, also their mean motions and periodic times, are subject to none but periodical changes; that the length of the year, for example, in the lapse of infinite ages, has no preponderating tendency either to increase or diminution,—that the planets will neither recede indefinitely from the sun, nor fall into it, but continue, so far as their mutual perturbations at least are concerned, to revolve for ever in orbits of very nearly the same dimensions as at present.

(669.) This theorem (the *Magna Charta* of our system), the discovery of which is due to Lagrange, is justly regarded as the most important, as a single result, of any which have hitherto rewarded the researches of

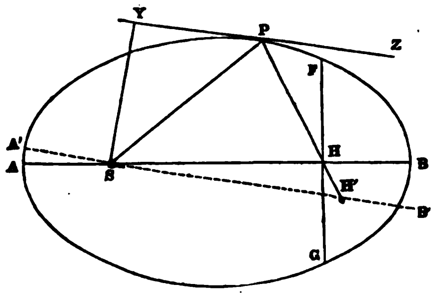
ians in this application of their science; and it is especially
remark, and follows evidently from the view here taken of it,
ld not be true but for the influence of the perturbing forces on
ents of the orbit, viz. the perihelion and excentricity, and the
and nodes; since we have seen that the revolution of the ap-
nodes, and the periodical increase and diminution of the ex-
and inclinations, are both essential towards operating that final
ete compensation which gives it a character of mathematical

We have here an instance of a perturbation of one kind
n a perturbation of another to annihilate an effect which would
ccumulate to the destruction of the system. It must, however,
mind, that it is the smallness of the excentricities of the more
planets, which gives this theorem its *practical* importance, and
es it from a mere barren speculative result. Within the limits
restoration, it is this alone which keeps the periodical fluctu-
e axis to and fro about a mean value within moderate and
limits. Although the earth might not fall into the sun, or re-
it beyond the present limits of our system, any considerable
diminution of its mean distance, to the extent, for instance, of
its actual amount, would not fail to subvert the conditions on
existence of the present race of animated beings depends
l as our system is, however, changes to anything like this ex-

such an exact adjustment to subsist at any epoch, it could not remain permanent, since by a remarkable property of perturbations of this class, which geometers have demonstrated, but the reasons of which we cannot stop to explain, any change produced on the axis of the disturbed planet's orbit is necessarily accompanied by a change *in the contrary direction* in that of the disturbing, so that the periods would recede from commensurability by the mere effect of their mutual action. Cases are not wanting in the planetary system of a certain approach to commensurability, and in one very remarkable case (that of Uranus and Neptune) of a considerably near one, not near enough, however, in the smallest degree to affect the validity of the argument, but only to give rise to inequalities of very long periods, of which more presently.[1]

(670.) The variation of the length of the axis of the disturbed orbit is due solely to the action of the tangential disturbing force. It is otherwise with that of its excentricity and of the position of its axis, or, which is the same thing, the longitude of its perihelion. Both the normal and tangential components of the disturbing force affect these elements. We shall, however, consider separately the influence of each, and, commencing, as the simplest case, with that of the tangential force;—let P be the place of the disturbed planet in its elliptic orbit A P B, whose axis at the moment is A S B and focus S. Suppose Y P Z to be a tangent to this orbit at P. Then, if we suppose A B = $2a$, the other focus of the ellipse, H, will be found by making the angle Z P H = Y P S or Y P H = 180° — Y P Z, or S P H = 180° — 2 Y P S, and taking P H = $2a$ — S P. This is evident from the nature of the ellipse, in which lines drawn from any point to the two foci make equal angles with the tangent,

Fig. 87.



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[1] 41 revolutions of Neptune are nearly equal to 81 of Uranus, giving rise to an inequality, having 6805 years for its period.

heir sum equal to the major axis. Suppose, now, th
act on P and to increase its velocity. It will therefor
that the new value assumed by a (viz. a') will be gr
e tangential force does not alter the angle of tange
new position (H′) of the upper focus, we must
ame line P H, a distance P H′ (= 2 a, — S P)
this then, and join S H′ and produce it. Then
sition of the axis, and ½ S H′ the new excentricit
st, that the new position of the perihelion A′ w
A towards the same side of the axis A B on w
tial force acts to increase the velocity, wheth
elion to aphelion, or the contrary. 2dly, That
to the action of the tangential force, the ex
between the perihelion and the perpendicul
ugh the upper focus, and diminishes when b
e perpendicular. 3dly, That for a given
iven value of the tangential force, the m
f the perihelion is a maximum when P
tion of P to the perihelion or aphelion,
lion being stationary when P is at A

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tial forces, if in estimating their *quantitative*
of the simplification introduced into their nu-
neglect of the angle P S Y, *i. e.* by the substi-
adial and transversal components. The *character*

To determine the momentary effect of the whole disturbing
, we have only to resolve it into its tangential and normal
, and estimating by these principles separately the effects of
tituent on both elements, add or substract the results according
spire or oppose each other. Or we may at once make the
H'' equal to twice the angle of deflection produced by the
e, and lay off P H''=P H+twice the variation of *a* produced
moment of time by the tangential force, and H'' will be the

The momentary velocity generated by the tangential force is
from a knowledge of that force by the ordinary principles of
and from this, the variation of the axis is easily derived.[1]
ntary velocity generated by the normal force in its own direc-
like manner caculable from a knowledge of that force, and
is by the linear velocity of P at that instant, we deduce the
locity of the tangent about P, or the momentary variation of
f tangency S P Y, corresponding.

The following *résumé* of these several results in a tabular form
ery variety of case according as P is approaching to or receding
s it is situated in the arc F A G of its orbit *about the perihe-*
the remoter arc G M F *about the aphelion*, as the tangential

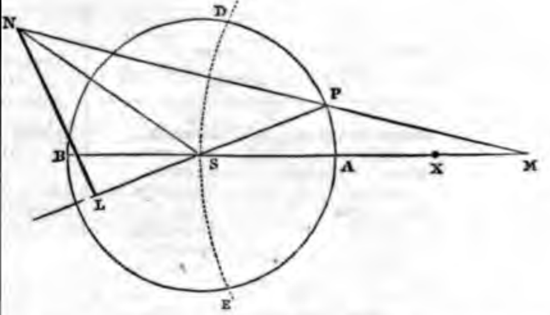
EFFECTS OF THE NORMAL DISTURBING FORCE.

Direction of P's motion.	Situation of P in orbit.	Action of Normal Force.	Effect on Elements.
Indifferent.	About Aphelion.	Inwards.	Apsides recede.
Ditto.	Ditto.	Outwards.	advance.
Ditto.	About Perihelion.	Inwards.	advance.
Ditto.	Ditto.	Outwards.	recede.
Approaching S.	Anywhere.	Inwards.	Excentr. increases.
Ditto.	Ditto.	Outwards.	decreases.
Receding from S.	Ditto.	Inwards.	decreases.
Ditto.	Ditto.	Outwards.	increases.

(674.) From the momentary changes in the elements of the disturbed orbit corresponding to successive situations of P and M, to conclude the total amount of change produced in any given time is the business of the integral calculus, and lies far beyond the scope of the present work. Without its aid, however, and by general consideration of the periodical recurrence of configurations of the same character, we have been able to demonstrate many of the most interesting conclusions to which geometers have been conducted, examples of which have already been given in the reasoning by which the permanence of the axes, the periodicity of the inclinations, and the revolutions of the nodes of the planetary orbits have been demonstrated. We shall now proceed to apply similar considerations to the motion of the apsides, and the variations of the excentricities. To this end we must first trace the changes induced on the disturbing forces themselves with the varying positions of the bodies, and here as in treating of the inclinations we shall suppose, unless the contrary is expressly indicated, both orbits to be very nearly circular, without which limitation the complication of the subject would become too embarrassing for the reader to follow, and defeat the end of explanation.

(675.) On this supposition the directions of S P and S Y, the perpendicular on the tangent at P, may be regarded as coincident, and the normal and radial disturbing forces become nearly identical in quantity, also the tangential and transversal, by the near coincidence of the points T and L (fig. art. 687). So far then as the *intensity* of the forces is concerned, it will make very little difference in which way the forces are resolved, nor will it at all materially affect our conclusions as to the effects of the normal and tangential forces, if in estimating their *quantitative* values, we take advantage of the simplification introduced into their numerical expression by the neglect of the angle P S Y, *i. e.* by the substitution for them of the radial and transversal components. The *character*

Fig. 89.

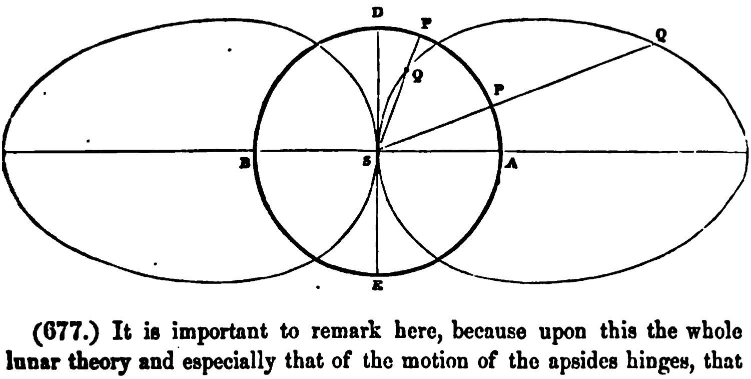


ffects depends (art. 670, 671,) on the *direction* in which the
which we shall suppose normal and tangential as before, and it
the estimation of their quantitative effects that the error in-
he neglect of this angle can fall. In the lunar orbit this angle
eds 3° 10′, and its influence on the quantitative estimation of
forces may therefore be safely neglected in a first approxima-
M N being found by the proportion $M P^2 : M S^2 :: M S :$

nishes at points intermediate between A D, D B, B E, and E A, which points, when M is considerably remote, are situated nearer to the quadrature than the syzygies.

(676.) In the lunar theory, to which we shall now proceed to apply these principles, both the geometrical representation and the algebraic expression of the disturbing forces admit of great simplification. Owing to the great distance of the sun M, at whose centre the radius of the moon's orbit never subtends an angle of more than about 8′, N P may be regarded as parallel to A B. And D S E becomes a straight line, coincident with the line of quadratures, so that V P becomes the cosine of A S P to radius S P, and $NL = NP \,.\, \sin ASP$; $LP = NP \,.\, \cos ASP$. Moreover, in this case (see the note on the last article) $NP = 3PV = 3SP \,.\, \cos ASP$; and consequently $NL = 3SP \,.\, \cos ASP \,.\, \sin ASP = \frac{3}{2} SP \,.\, \sin 2ASP$, and $LS = SP(3 \,.\, \cos ASP^2 - 1) = \frac{1}{2} SP (1 + 3 \,.\, \cos 2ASP)$ which vanishes when $\cos ASP^2 = \frac{1}{3}$, or at 64° 14′ from the syzygy. Suppose through every point of P's orbit there be drawn $SQ = 3SP \,.\, \cos ASP^2$, then will Q trace out a certain looped oval, as in the figure, cutting the orbit in four points 64° 14′ from A and B respectively, and P Q will always represent in quantity and direction the normal force acting at P.

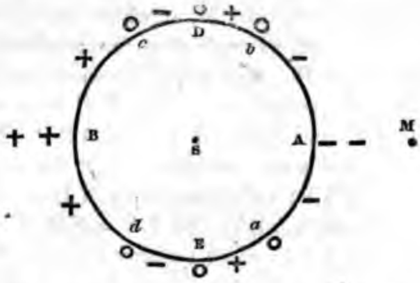
Fig. 90.



(677.) It is important to remark here, because upon this the whole lunar theory and especially that of the motion of the apsides hinges, that all the acting disturbing forces, at equal angles of elongation A S P of the moon from the sun, are *cæteris paribus* proportional to S P, the moon's distance from the earth, and are therefore greater when the moon is near its apogee than when near its perigee; the extreme proportion being that of about 28 : 25. This premised, let us first consider the effect of the normal force in displacing the lunar apsides. This we shall best be enabled to do by examining separately those cases in which the effects are

gly contrasted; viz. when the major axis of the moon's orbit is
wards the sun, and when at right angles to that direction.
, let the line of apsides be directed to the sun as in the an-
e, where A is the perigee, and take the arcs A *a*, A *b*, B *c*, B *d*

Fig. 91.



14′. Then while P is between *a* and *b* the normal force act-
rds, and the moon being near its perigee, by art. 671, the
l recede, but when between *c* and *d*, the force there acting out-

in the proportion of 28 : 25, while in the quadratures about D and E they are equal. Therefore, while the feeble movements + and — in the neighbourhood of these points destroy each other almost exactly, there will necessarily remain a considerable balance in favour of advance, in this situation of the line of apsides.

(678.) Next, suppose the apogee to lie at A, and the perigee at B. In this case it is evident that, so far as the direction of the motions of the apsides is concerned, all the conclusions of the foregoing reasoning will be reversed by the substitution of the word perigee for apogee, and *vice versâ;* and all the signs in the figure referred to will be changed. But now the most powerful forces act on the side of A, that is to say, still on the side of advance, this condition also being reversed. In either situation of the orbit, then, the apsides advance.

(679.) (Case 3.) Suppose, now, the major axis to have the situation D E, and the perigee to be on the side of D. Here, in the arc *b c* of P's motion the normal force acts inwards, and the moon is near perigee, consequently the apsides advance, but with a moderate rapidity, the maximum of the inward normal force being only half that of the outward. In the arcs A *b* and *c* B the moon is still near perigee, and the force acts outwards, but though powerfully towards A and B, yet at a constantly increasing disadvantage (art. 671.) Therefore in these arcs the apsides recede, but moderately. In *a* A and B *d* (being towards apogee) they again advance, still with a moderate velocity. Lastly, throughout the arc *d a*, being about apogee with an inward force, they recede. Here as before, if the perigee and apogee forces were equal, the advance and recess would counterbalance; but as in fact the apogee forces preponderate, there will be a balance on the entire revolution in favour of recess. The same reasoning of course holds good if the perigee be towards E. But now, between these cases and those in the foregoing articles, there is this difference, viz. that in this the dominant effect results from the inward action of the normal force in quadratures, while in the others it results from its outward, and doubly powerful action in syzygies. The recess of the apsides in their quadratures arising from the action of the normal force will therefore be less than their advance in their syzygies; and not only on this account, but also because of the much less extent of the arcs *b c* and *d a* on which the balance is mainly struck in this case, than of *a b* and *c d*, the corresponding most influential arcs in the other.

(680.) In intermediate situations of the line of apsides, the effect will be intermediate, and there will of course be a situation of them in which on an average of a whole revolution, they are stationary. This situation it is easy to see will be nearer to the line of quadratures than of syzygies, eponderance of advance will be maintained over a much more e arc than that of recess, among the possible situations which hold. On every account, therefore, the action of the normal s the lunar apsides to progress *in a complete revolution of* M nodical year, during which the motion of the sun round the e consider the earth at rest) brings the line of syzygies into all with respect to that of apsides.

Let us next consider the action of the tangential force. And (Case 1.), supposing the perigree of the moon at A, and the f her revolution to be A D B E, the tangential force *retards* through the quadrant A D, in which she *recedes* from S, there- . 670 the apsides *recede*. Through D B the force *accelerates*, moon still *recedes*, therefore they *advance*. Through B E the ds, and the moon approaches, therefore they continue to advance, throughout the quadrant E A the force accelerates, and the oaches, therefore they recede. In virtue therefore of this force, recede, during the description of the arc E A D, and advance B E, but the force being in this case as in that of the normal powerful at apogee, the latter will preponderate, and the apsides mum retarding, while the inward action of the normal at its maximum is only half the maximum of its outward. Neither is there that difference in the extent of the arcs over which the balance is struck in this, as in the other case, the action of the tangential force being inward and outward alternately over equal arcs, each a complete quadrant. Whereas, therefore, in tracing the action of the normal force, we found reason to conclude it much more effective to produce progress of the apsides in their syzygy, than in their quadrature situations, we can draw no such conclusion in that of the tangential forces: there being, as regards *that* force, a *complete symmetry*, in the four *quadrants*, while in regard of the normal force the symmetry is only a *half-symmetry* having relation to two *semicircles*.

(685.) Taking the average of many revolutions of the sun about the earth, in which it shall present itself in every possible variety of situations to the line of apsides, we see that the effect of the normal force is to produce a rapid advance in the syzygy of the apsides, and a less rapid recess in their quadrature, and on the whole, therefore, a moderately rapid general advance, while that of the tangential is to produce an equally rapid advance in syzygy, and recess in quadrature. Directly, therefore, the tangential force would appear to have no ultimate influence in causing either increase or diminution in the *mean motion* of the apsides resulting from the action of the normal force. It does so, however, indirectly, conspiring in that respect with, and greatly increasing, an indirect action of the normal force in a manner which we shall now proceed to explain.

(686.) The sun moving uniformly, or nearly so, in the same direction as P, the line of apsides when in or near the syzygy, in advancing follows the sun, and therefore remains materially longer in the neighbourhood of syzygy than if it rested. On the other hand, when the apsides are in quadrature they recede, and, moving therefore contrary to the sun's motion, remain a shorter time in that neighbourhood, than if they rested. Thus the advance, already preponderant, is made to preponderate more by its longer continuance, and the recess, already deficient, is rendered still more so by the shortening of its duration.[1] Whatever cause, then, increases directly the rapidity of both advance and recess, *though it may do both equally*, aids in this indirect process, and it is thus that the tangential force becomes effective through the medium of the progress already produced, in doing and aiding the normal force to do that which alone it would be unable to effect. Thus we have perturbation exaggerating perturbation, and thus we see what is meant by geometers, when they

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[1] Newton, Princ. i. 66. Cor. 8.

t a considerable part of the motion of the lunar apsides is due
are of the disturbing force, or, in other words, arises out of a
proximation in which the influence of the first in altering the
problem is taken into account.

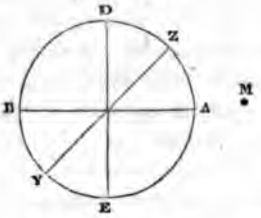
The curious and complicated effect of perturbation, described in
ticle, has given more trouble to geometers than any other part
ar theory. Newton himself had succeeded in tracing that part
ion of the apogee which is due to the direct action of the radial
finding the amount only half what observation assigns, he
have abandoned the subject in despair. Nor, when resumed
cessors, did the inquiry, for a very long period, assume a more
aspect. On the contrary, Newton's result appeared to be even
erified, and the elaborate investigations which were lavished
ubject without success, began to excite strong doubts whether
of the lunar motions could be explained at all by the New-
of gravitation. The doubt was removed, however, almost in
of its origin, by the same geometer, Clairaut, who first gave it
nd who gloriously repaired the error of his momentary hesita-
monstrating the exact coincidence between theory and observa-

(commencing with the normal force) the momentary change of excentricity will vanish at *a*, *b*, *c*, *d*, by the vanishing of that force, and at Z and Y by the effect of situation in the orbit annulling its action (art. 671). In the arcs Z *b* and Y *d* therefore the change of excentricity will be small, the acting force nowhere attaining either a great magnitude or an advantageous situation within their limits. And the force within these two arcs having the same character as to inward and outward, but being oppositely influential by reason of the approach of P to S in one of them and its recess in the other, it is evident that, so far as these arcs are concerned, a very near compensation of effects will take place, and though the apogeal arc Y *d* will be somewhat more influential, this will tell for little upon the average of a revolution.

(689.) The arcs *b* D *c* and *d* E *a* are each much less than a quadrant in extent, and the force acting inwards throughout them (which at its maximum in D and E is only half the outward force at A, B) degrades very rapidly in intensity towards either syzygy (see art. 676). Hence whether Z be between *b* *c* or *b* A, the effects of the force in these arcs will not produce very extensive changes on the excentricity, and the changes which it does produce will (for the reason already given) be opposed to each other. Although, then, the arc *a* *d* be farther from perigee than *b* *c*, and therefore the force in it is greater, yet the predominance of effect here will not be very marked, and will moreover be partially neutralized by the small predominance of an opposite character in Y *d* over Z *b*. On the other hand, the arcs *a* Z, *c* Y are both larger in extent than either of the others, and the seats of action of forces doubly powerful. Their influence, therefore, will be of most importance, and their preponderance one over the other (being opposite in their tendencies), will decide the question whether, on an average of the revolution, the excentricity shall increase or diminish. It is clear that the decision must be in favour of *c* Y, the apogeal arc, and, since in this the force is *outwards* and the moon *receding* from the earth, an *increase* of the excentricity will arise from its influence. A similar reasoning will, evidently, lead to the same conclusion were the apogee and perigee to change places, for the directions of P's motion as to approach and recess to S will be indeed reversed, but at the same time the dominant forces will have changed sides, and the arc *a* A Z will now give the character to the result. But when Z lies between A and E, as the reader may easily satisfy himself, the case will be altogether different, and the reverse conclusion will obtain. Hence the changes of excentricity emergent on the average of single revolutions from the action of the normal force will be as represented by the signs + and — in the figure above annexed.

Let us next consider the effect of the tangential force. This
n the quadrants A D, B E, and accelerates it in the alternate

Fig. 93.



the whole quadrant A D, therefore, the effect is of one charac-
rigee being less than 90° from every point in it, and in the
rant B E it is of the opposite, the apogee being so situated
Moreover, in the middle of each quadrant, the tangential
ts maximum. Now, in the other quadrants, E A and D B,
from perigeal to apogeal vicinity takes place, and the tangen-
owever powerful, has its effect annulled by situation (art. 670),

$(\frac{1}{4}+\frac{1}{12})$ part of a year will have gained on it 90°, during all which interval the apse will have been in the quadrant A E of our figure, and the excentricity continually decreasing. The decrease will then cease, but the excentricity itself will be a minimum, the sun being now at right angles to the line of apsides. Thence it will increase to a maximum when the sun has gained another 90°, and again attained the line of apsides, and so on alternately. The actual effect on the numerical value of the lunar excentricity is very considerable, the greatest and least excentricities being in the ratio of 3 to 2.[1]

(692.) The motion of the apsides of the lunar orbit may be illustrated by a very pretty mechanical experiment, which is otherwise instructive in giving an idea of the mode in which orbitual motion is carried on under the action of central forces variable according to the situation of the revolving body. Let a leaden weight be suspended by a brass or iron wire to a hook in the under side of a firm beam, so as to allow of its free motion on all sides of the vertical, and so that when in a state of rest it shall just clear the floor of the room, or a table placed ten or twelve feet beneath the hook. The point of support should be well secured from wagging to and fro by the oscillation of the weight, which should be sufficient to keep the wire as tightly stretched as it will bear, with the certainty of not breaking. Now, let a very small motion be communicated to the weight, not by merely withdrawing it from the vertical and letting it fall, but by giving it a slight impulse sideways. It will be seen to describe a regular ellipse about the point of rest as its centre. If the weight be heavy, and carry attached to it a pencil, whose point lies exactly in the direction of the string, the ellipse may be transferred to paper lightly stretched and gently pressed against it. In these circumstances, the situation of the major and minor axes of the ellipse will remain for a long time very nearly the same, though the resistance of the air and the stiffness of the wire will gradually diminish its dimensions and excentricity. But if the impulse communicated to the weight be considerable, so as to carry it out to a great angle (15° or 20° from the vertical,) this permanence of situation of the ellipse will no longer subsist. Its axis will be seen to shift its position at every revolution of the weight, advancing in the same direction with the weight's motion, by an uniform and regular progression, which at length will entirely reverse its situation, bringing the direction of the longest excursions to coincide with that in which the shortest were previously made; and so on, round the whole circle; and, in a word, imitating to the eye, very completely, the motion of the apsides of the moon's orbit.

[1] Airy, Gravitation, p. 106.

Now, if we inquire into the cause of this progression of the
will not be difficult of detection. When a weight is suspended
and drawn aside from the vertical, it is urged to the lowest
ather in a direction at every instant perpendicular to the wire)
which varies as the sine of the deviation of the wire from the
lar. Now, the sines of very small arcs are nearly in the pro-
the arcs themselves; and the more nearly, as the arcs are
If, therefore, the deviations from the vertical be so small that
glect the curvature of the spherical surface in which the weight
l regard the curve described as coincident with its projection on
l plane, it will be then moving under the same circumstances
e a revolving body attracted to a centre by a force varying
the distance; and, in this case, the curve described would be
having its centre of attraction not in the focus, but in the
d the apsides of this ellipse would remain fixed. But if the
of the weight from the vertical be considerable, the force urging
the centre will deviate in its law from the simple ratio of the
being as the *sine*, while the distances are as the *arc*. Now the
h it continues to increase as the arc increases, yet does not in-

the radius. It will require a longer continued action of the central force to do this; and before it is accomplished, more than a quadrant of its revolution must be passed over in angular motion round the centre. But this is only stating at length, and in a more circuitous manner, that fact which is more briefly and summarily expressed by saying that *the apsides of its orbit are progressive.* Nothing beyond a familiar illustration is of course intended in what is above said. The case is not an exact parallel with that of the lunar orbit, the disturbing force being simply radial, whereas in the lunar orbit a transversal force is also concerned, and even were it otherwise, only a confused and indistinct view of aspidal motion can be obtained from this kind of consideration of the curvature of the disturbed path. If we would obtain a clear one, the two foci of the instantaneous ellipse must be found from the laws of elliptic motion performed under the influence of a force directly as the distance, and the radial disturbing force being decomposed into its tangential and normal components, the momentary influence of either in altering their positions and consequently the directions and lengths of the axis of the ellipse must be ascertained. The student will find it neither a difficult nor an uninstructive exercise to work out the case from these principles, which we cannot afford the space to do.

(694.) The theory of the motion of the planetary apsides and the variation of their excentricities is in one point of view much more simple, but in another much more complicated than that of the lunar. It is simpler; because owing to the exceeding minuteness of the changes operated in the course of a single revolution, the angular position of the bodies with respect to the line of apsides is very little altered by the motion of the apsides themselves. The line of apsides neither follows up the motion of the disturbing body in its state of advance, nor *vice versâ*, in any degree capable of prolonging materially their advancing or shortening materially their receding phase. Hence no second approximation of the kind explained (in art. 686), by which the motion of the lunar apsides is so powerfully modified as to be actually doubled in amount, is at all required in the planetary theory. On the other hand, the latter theory is rendered more complicated than the former, at least in the cases of planets whose periodic times are to each other in a ratio much less than 13 to 1, by the consideration that the disturbing body shifts its position with respect to the line of apsides by a much greater angular quantity in a revolution of the disturbed body than in the case of the moon. In that case we were at liberty to suppose (for the sake of explanation), without any very egregious error, that the sun held nearly a fixed position during a single lunation. But in the case of planets whose times of revolution

ch lower ratio this cannot be permitted. In the case of Jupiter
by Saturn for example, in one sidereal revolution of Jupiter,
advanced in its orbit with respect to the line of apsides of
more than 140°, a change of direction which entirely alters
ons under which the disturbing forces act. And in the case
rior disturbed by an interior planet, the situation of the latter
t to the line of the apsides varies even more rapidly than the
the exterior or disturbed planet with respect to the central
such cases then the reasoning which we have applied to the
rbations becomes totally inapplicable; and when we take into
on also the excentricity of the orbit of the disturbing body,
e most important cases is exceedingly influential, the subject
r too complicated for verbal explanation, and can only be suc-
llowed out with the help of algebraic expression and the appli-
e integral calculus. To Mercury, Venus, and the earth indeed,
d by Jupiter, and planets superior to Jupiter, this objection to
ing in question does not apply; and in each of these cases
e are entitled to conclude that the apsides are kept in a state
sion by the action of all the larger planets of our system.

synoptic table in art. 673. If M, the disturbing body, be supposed to be successively placed in two diametrically opposite situations in its orbit, the aphelion of P will stand related to M in one of these situations precisely as its perihelion in the other. Now the orbits being so nearly circles as supposed, the distribution of the disturbing forces, whether normal or tangential, is symmetrical relative to their common diameter passing through M, or to the line of syzygies. Hence it follows that the half of P's orbit "about perihelion" (art. 673) will stand related to all the acting forces in the one situation of M, precisely as the half "about aphelion" does in the other: and also, that the half of the orbit in which P "approaches S," stands related to them in the one situation precisely as the half in which it "recedes from S" in the other. Whether as regards, therefore, the normal or tangential force, the conditions of advance or recess of apsides, and of increase or diminution of excentricities, are reversed in the two supposed cases. Hence it appears that whatever situation be assigned to M, and whatever influence it may exert on P in that situation, that influence will be annihilated in situations of M and P, diametrically opposite to those supposed, and thus, on a general average, the effect on both apsides and excentricities is reduced to nothing.

(697.) If the orbits, however, be excentric, the symmetry above insisted on in the distribution of the forces does not exist. But, in the first place, it is evident that if the excentricities be moderate, (as in the planetary orbits,) by far the larger part of the effects of the disturbing forces destroys itself in the manner described in the last article, and that it is only a residual portion, viz. that which arises from the greater proximity of the orbits at one place than at another, which can tend to produce permanent or secular effects. The precise estimation of these effects is too complicated an affair for us to enter upon; but we may at least give some idea of the process by which they are produced, and the order in which they arise. In so doing, it is necessary to distinguish between the effects of the normal and tangential forces. The effects of the former are greatest at the point of conjunction of the planets, because the normal force itself is there always at its maximum; and although, where the conjunction takes place at 90° from the line of apsides, its effect to move the apsides is nullified by situation, and when *in* that line its effect on the excentricities is similarly nullified, yet, in the situations rectangular to these, it acts to its greatest advantage. On the other hand, the tangential force vanishes at conjunction, whatever be the place of conjunction with respect to the line of apsides, and where it is at its maximum its effect is still liable to be annulled by situation. Thus it appears that the normal

ɔst influential, and mainly determines the character of the ge-
t. It is, therefore, at conjunction that the most influential
oduced, and therefore, on the long average, those conjunctions
pen about the place where the orbits are nearest will determine
l character of the effect. Now, the nearest points of approach
ipses, which have a common focus, may be variously situated
ct to the perihelion of either. It may be at the perihelion or
ɔn of the disturbed orbit, or in any intermediate position. Sup-
be at the perihelion. Then, if the disturbed orbit be *interior*
urbing, the force acts outwards, and therefore the apsides re-
xterior, the force acts inwards, and they advance. In neither
he excentricity change. If the conjunction take place at the
f the disturbed orbit, the effects will be reversed: if interme-
apsides will be less, and the excentricity more affected.

Supposing only two planets, this process would go on till the
d excentricities had so far changed as to alter the point of
proach of the orbits, so as either to accelerate or retard and
verse the motion of the apsides, and give to the variation of the
y a corresponding periodical character. But there are many

at the earth's equator, prevent any secular accumulation of small changes, by which, in the lapse of ages, its ellipticity might be materially increased or diminished. Accordingly, observation shows the *mean* excentricity of the moon's orbit to be the same now as in the earliest ages of astronomy.

(700.) The movements of the perihelia, and variations of excentricity of the planetary orbits, are interlaced and complicated together in the same manner and nearly by the same laws as the variations of their nodes and inclinations. Each acts upon every other, and every such mutual action generates its own peculiar period of circulation or compensation; and every such period, in pursuance of the principles of art. 650, is thence propagated throughout the system. Thus arise cycles upon cycles, of whose compound duration some notion may be formed, when we consider what is the length of one such period in the case of the two principal planets — Jupiter and Saturn. Neglecting the action of the rest, the effect of their mutual attraction would be to produce a secular variation in the excentricity of Saturn's orbit, from 0·08409, its *maximum*, to 0·01345, its *minimum* value: while that of Jupiter would vary between the narrow limits, 0·06036 and 0·02606 : the greatest excentricity of Jupiter corresponding to the least of Saturn, and *vice versâ*. The period in which these changes are gone through, would be 70414 years. After this example, it will be easily conceived that many millions of years will require to elapse before a complete fulfilment of the joint cycle which shall restore the whole system to its original state as far as the excentricities of its orbits are concerned.

(701.) The place of the perihelion of a planet's orbit is of little consequence to its well-being; but its excentricity is most important, as upon this (the axes of the orbits being permanent) depends the mean temperature of its surface, and the extreme variations to which its seasons may be liable. For it may be easily shown that the *mean annual amount* of light and heat received by a planet from the sun is, *cæteris paribus*, as the minor axis of the ellipse described by it. Any variation, therefore, in the excentricity, by changing the minor axis, will alter the *mean* temperature of the surface. How such a change will also influence the extremes of temperature appears from art. 368. Now it may naturally be inquired whether (in the vast cycle above spoken of, in which, at some period or other, conspiring changes may accumulate on the orbit of one planet from several quarters,) it may not happen that the excentricity of any one planet — as the earth — may become exorbitantly great, so as to subvert those relations which render it habitable to man, or to give rise to great changes, at least, in the physical comfort of his state. To this the researches of geometers have enabled us to answer in the negative. A

s been demonstrated by Lagrange between the masses, axes of
and excentricities of each planet, similar to what we have al-
d with respect to their inclinations, viz. *that if the mass of*
t be multiplied by the square root of the axis of its orbit, and
t by the square of its excentricity, the sum of all such products
the system is invariable; and as, in point of fact, this sum is
small, so it will always remain. Now, since the axes of the
liable to no secular changes, this is equivalent to saying that no
hall increase its excentricity, unless at the expense of a common
whole amount of which is, and must for ever remain, extremely

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nothing in this relation, however, taken *per se*, to secure the smaller pla-
ury, Mars, Juno, Ceres, &c.—from a catastrophe, could they accumulate
es, or any one of them, the whole amount of this *excentricity fund*. But
ver be: Jupiter and Saturn will always retain the lion's share of it. A
rk applies to the *inclination fund* of art. 639. These *funds*, be it observed,
t into debt. Every term of them is essentially positive.

CHAPTER XIV.

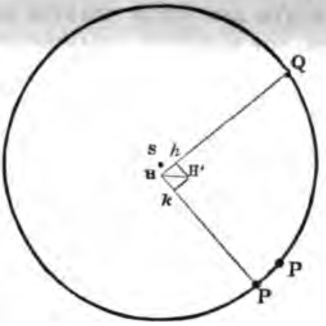
OF THE INEQUALITIES INDEPENDENT OF THE EXCENTRICITIES.—THE MOON'S VARIATION AND PARALLACTIC INEQUALITY.—ANALOGOUS PLANETARY INEQUALITIES.—THREE CASES OF PLANETARY PERTURBATION DISTINGUISHED.—OF INEQUALITIES DEPENDENT ON THE EXCENTRICITIES.—LONG INEQUALITY OF JUPITER AND SATURN.—LAW OF RECIPROCITY BETWEEN THE PERIODICAL VARIATIONS OF THE ELEMENTS OF BOTH PLANETS.—LONG INEQUALITY OF THE EARTH AND VENUS.—VARIATION OF THE EPOCH.—INEQUALITIES INCIDENT ON THE EPOCH AFFECTING THE MEAN MOTION.—INTERPRETATION OF THE CONSTANT PART OF THESE INEQUALITIES.—ANNUAL EQUATION OF THE MOON.—HER SECULAR ACCELERATION.—LUNAR INEQUALITIES DUE TO THE ACTION OF VENUS.—EFFECT OF THE SPHEROIDAL FIGURE OF THE EARTH AND OTHER PLANETS ON THE MOTIONS OF THEIR SATELLITES.—OF THE TIDES.—MASSES OF DISTURBING BODIES DEDUCIBLE FROM THE PERTURBATIONS THEY PRODUCE.—MASS OF THE MOON, AND OF JUPITER'S SATELLITES, HOW ASCERTAINED.—PERTURBATIONS OF URANUS RESULTIFG IN THE DISCOVERY OF NEPTUNE.

(702.) To calculate the actual place of a planet or the moon, in longitude and latitude at any assigned time, it is not enough to know the changes produced by perturbation in the elements of its orbit, still less to know the *secular* changes so produced, which are only the outstanding or uncompensated portions of much greater changes induced in short periods of configuration. We must be enabled to estimate the actual effect on its longitude of those periodical accelerations and retardations in the rate of its mean angular motion, and on its latitude of those deviations above and below the mean plane of its orbit, which result from the continued action of the perturbative forces, not as compensated in long periods, but as in the act of their generation and destruction in short ones. In this chapter we purpose to give an account of some of the most prominent of the *equations* or inequalities thence arising, several of which are of high historical interest, as having become known by observation previous to the discovery of their theoretical causes, and as having, by their successive explanations from the theory of gravitation, removed what were in some instances regarded as formidable objections against that theory, and afforded in all most satisfactory and triumphant verifications of it.

(703.) We shall begin with those which compensate themselves in a synodic revolution of the disturbed and disturbing body, and which are independent of any permanent excentricity of either orbit, going through their changes and effecting their compensation in orbits slightly elliptic, almost precisely as if they were circular. These inequalities result, in fact, from a circulation of the true upper focus of the disturbed ellipse about its mean place in a curve whose form and magnitude the principles laid down in the last chapter enable us to assign in any proposed case. If the disturbed orbit be circular, this mean place coincides with its centre: if elliptic, with the situation of its upper focus, as determined from the principles laid down in the last chapter.

(704.) To understand the nature of this circulation, we must consider the joint action of the two elements of the disturbing force. Suppose H to be the place of the upper focus, corresponding to any situation P of the disturbed body, and let P P′ be an infinitesimal element of its orbit, described in an instant of time. Then supposing no disturbing force to act, P P′ will be a portion of an ellipse, having H for its focus, equally whether the point P or P′ be regarded. But now let the disturbing forces act during the instant of describing P P′. Then the focus H will shift its position to H′ to find which point we must recollect, 1st. What is demonstrated (in art. 671), viz. that the effect of the normal force is to vary the position of the line P′ H so as to make the angle H P H′ equal to double the variation of the angle of tangency due to the action of that

Fig. 95.



force, without altering the distance P H: so that in virtue of the normal force alone, H would move to a point h, along the line H Q, drawn from H to a point Q, 90° in advance of P, (because S H being exceedingly small, the angle P H Q may be taken as a right angle when P S Q is so,) H *approaching* Q if the normal force act *outwards*, but receding from Q if inwards. And similarly the effect of the tangential force (art. 670) is to vary the position of H in the direction H P or P H, according as the force retards or accelerates P's motion. To find H′ then from H draw H P, H Q, to P and to a point of P's orbit 90° in advance of P. On H Q take H h, the motion of the focus due to the normal force, and on H P take H k the motion due to the tangential force; complete the parallelogram H H′, and its diagonal H H′ will be the element of the true path of H in virtue of the joint action of both forces.

(705.) The most conspicuous case in the planetary system to which the above reasoning is applicable, is that of the moon disturbed by the sun. The inequality thus arising is known by the name of the moon's variation, and was discovered so early as about the year 975 by the Arabian astronomer Aboul Wefa.[1] Its magnitude (or the extent of fluctuation to and fro in the moon's longitude which it produces) is considerable, being no less than 1° 4′, and it is otherwise interesting as being the first inequality produced by perturbation, which Newton succeeded in explaining by the theory of gravity. A good general idea of its nature may be formed by considering the direct action of the disturbing forces on the moon, supposed to move in a circular orbit. In such an orbit undisturbed, the velocity would be uniform; but the tangential force acting to accelerate her motion through the quadrants *preceding* her conjunction and opposition, and to retard it through the alternate quadrants, it is evident that the velocity will have two maxima and two minima, the former at the syzygies, the latter at the quadratures. Hence at the syzygies the velocity will exceed that which corresponds to a circular orbit, and at quadratures will fall short of it. The true orbit will therefore be less curved or more flattened than a circle in syzygies, and more curved (*i. e.* protuberant beyond a circle) in quadratures. This would be the case even were the normal force not to act. But the action of that force increases the effect in question, for at the syzygies, and as far as 64° 14′ on either side of them, it acts outwards, or in counteraction of the earth's attraction, and thereby prevents the orbit from being so much curved as it otherwise would be; while at quadratures, and for 25° 46′ on either side of them, it acts inwards, aiding the earth's attraction, and rendering that portion of the orbit more curved than it otherwise would be. Thus the joint action of both forces distorts the orbit from a circle into a flattened or elliptic form, having the longer axis in quadratures, and the shorter in syzygies; and in this orbit the moon moves faster than with her mean velocity at syzygy (*i. e.* where she is nearest the earth) and slower at quadratures where farthest. Her *angular* motion about the earth is therefore for both reasons greater in the former than in the latter situation. Hence at syzygy her true longitude seen from the earth will be in the act of gaining on her mean,—in quadratures of losing, and at some intermediate points (not very remote from the octants) will neither be gaining nor losing. But at these points, *having been gaining or losing* through the whole previous 90° the *amount of gain or loss* will have attained its maximum. Consequently at the octants the true longitude will deviate most from the mean in excess and defect, and the inequality in question is therefore *nil* at syzygies and quadratures, and attains its maxima in advance or retardation at the octants, which is agreeable to observation.

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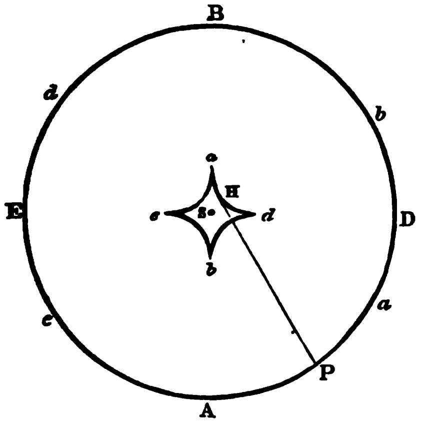
[1] Sedillot, Nouvelles Recherches pour servir à l'Histoire de l'Astronomie chez les Arabes.

(706.) Let us, however, now see what account can be rendered of this inequality by the simultaneous variations of the axis and excentricity as above explained. The tangential force, as will be recollected, is *nil* at syzygies and quadratures, and a maximum at the octants, accelerative in the quadrants E A and D B, and retarding in A D and B E. In the two former then the axis is in process of lengthening; in the two latter, shortening. On the other hand the normal force vanishes at (*a*, *b*, *d*, *e*) 64° 14′ from the syzygies. It acts outwards over *e* A *a*, *b* B *d*, and inwards over *a* D *b* and *d* E *e*. In virtue of the tangential force, then, the point H moves towards P when P is in A D, B E, and from it when in D B, E A, the motion being *nil* when at A, B, D, E, and most rapid when at the octant D, at which points, therefore, (so far as this force is concerned,) the focus H would have its mean situation. And in virtue of the normal focus, the motion of H in the direction H Q will be at its maximum of rapidity towards Q at A, or B, from Q at D or E, and *nil*, at *a*, *b*, *d*, *e*. It will assist us in following out these indications to obtain a notion of the form of the curve really described by H, if we trace separately the paths which H would pursue in virtue of either motion separately, since its true motion will necessarily result from the superposition of these partial motions, because at every instant they are at right angles to each other, and therefore cannot interfere. First, then, it is evident, from what we have said of the tangential force, that when P is at A, H is for an instant at rest, but that as P removes from A towards D, H continually approaches P along their line of junction H P, which is, therefore, at each instant a tangent to the path of H. When P is in the octant, H is at its mean

Fig. 96.

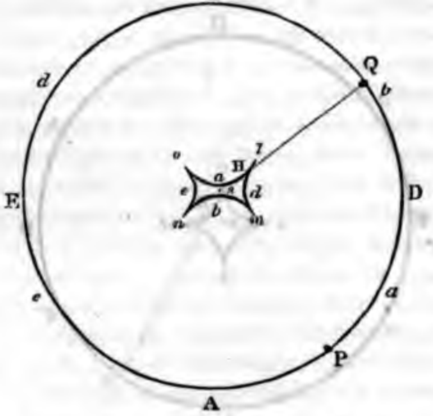


distance from P (equal to P S), and is then in the act of approaching P most rapidly. From thence to the quadrature D the movement of H towards P decreases in rapidity till the quadrature is attained, when H rests for an instant, and then begins to reverse its motion, and travel *from* P at the same rate of progress as before *towards* it. Thus it is clear that, in virtue of the tangential force alone, H would describe a four-cusped curve *a*, *d*, *b*, *e*, its direction of motion round S in this curve being opposite to that of P, so that A and *a*, D and *d*, B and *b*, E and *e*, shall be corresponding points.

(707.) Next as regards the normal force. When the moon is at A the motion of H is towards D, and is at its maximum of rapidity, but slackens as P proceeds towards D and as Q proceeds towards B. To the curve described, H Q will be always a tangent, and since at the neutral point of the normal force (or when P is 64° 14′ from A, and Q 64° 14′ from D), the motion of H is for an instant *nil* and is then reversed, the curve will have a cusp at *l* corresponding, and H will then begin to travel along the arc *l m*, while P describes the corresponding arc from neutral point to neutral point through D. Arrived at the neutral point between D and B, the motion of H along Q H will be again arrested and reversed, giving rise to another cusp at *m*, and so on. Thus, in virtue of the normal force acting alone, the path of H would be the four-cusped, elongated curve *l m n o*, described with a motion round S the reverse of P's, and having *a*, *d*, *b*, *e* for points corresponding to A, B, D, E, places of P.

(708.) Nothing is now easier than to superpose these motions. Sup-

Fig. 97.



H_2 to be the points in either curve corresponding to P, we
g to do but to set from off S, S *h* equal and parallel to S H_1
curve and from *h*, *h* H equal and parallel to S H_2 in the other.

we represent Sa_1 or Sd_1 (for they are equal) by $2a$, Sa_2 by $2b$, and Sd_2 by $2c$, then will the semiaxes of the oval $a\,d\,b\,e$, Sa and Sd be respectively $2a+2b$ and $2a+2c$, so that the excentricities of the momentary ellipses at A and D will be respectively $a+b$ and $a+c$. The total amount of the effect of the tangential force on the *axis*, in passing from syzygy to quadrature, will evidently be equal to the length of the curvilinear arc $a_1\,d_1$ (*fig.* art. 708), which is necessarily less than Sa_1+Sd_1 or $4a$. Therefore the total effect on the *semiaxis* or distance of the moon is less than $2a$, and the excess and defect of the greatest and least values of this distance thus varied above and below the mean value $S\,A=1$ (which call α) will be less than a. The moon then is moving at A *in* the *perigee* of an ellipse whose semiaxis is $1+\alpha$ and excentricity $a+b$, so that its actual distance from the earth there is $1+\alpha-a-b$, which (because α is less than a) is less than $1-b$. Again, at D it is moving *in apogee* of an ellipse whose semiaxis is $1-\alpha$ and excentricity $a+c$, so that its distance then from the earth is $1-\alpha+a+c$, which (a being greater than α) is greater than $1+c$, the latter distance exceeding the former by $2a-2\alpha+b+c$.

(710.) Let us next consider the corresponding changes induced upon the angular velocity. Now it is a law of elliptic motion that at different points of different ellipses, each differing very little from a circle, the angular velocities are to each other as the square roots of the semiaxes directly, and as the squares of the distances inversely. In this case the semiaxes at A and D are to each other as $1+\alpha$ to $1-\alpha$, or as $1:1-2\alpha$, so that their square roots are to each other as $1:1-\alpha$. Again, the distances being to each other as $1+\alpha-a-b:1-\alpha+a+c$, the inverse ratio of their squares (since α, a, b, c, are all very small quantities) is that of $1-2\alpha+2a+2c:1+2\alpha-2a-2b$, or as $1:1-4\alpha-4a-2b-2c$. The angular velocities then are to each other in a ratio compounded of these two proportions, that is in the ratio of

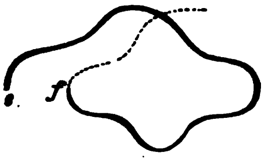
$$1\ :\ 1+3\alpha-4a-2b-2c,$$

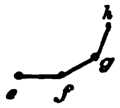
which is evidently that of a greater to a less quantity. It is obvious also, from the constitution of the second term of this ratio, that the normal force is far more influential in producing this result than the tangential.

(711.) In the foregoing reasoning the sun has been regarded as fixed. Let us now suppose it in motion (in a circular orbit), then it is evident that at equal angles of *elongation* (of P from M seen from S), equal disturbing forces, both tangential and normal, will act: only the syzygies and quadratures, as well as the neutral points of the normal force, instead of being points fixed in longitude on the orbit of the moon, will advance on that orbit with a uniform angular motion equal to the angular motion of

'he cuspidated curves $a_1 d_1 b_1 e_1$ and $a_2 d_2 b_2 e_2$, *fig.* art. 708, will,
o longer be re-entering curves; but each will have its cusps
nd as it were in the direction of the sun's motion, so as to in-
angles between them in the ratio of the synodical to the side-
ion of the moon (art. 418). And if, in like manner, the mo-
se two curves, thus separately described by H, be compounded,
g curve, though still (loosely speaking) a species of oval, will
into itself, but will make successive spiroidal convolutions about
est and nearest points being in the same ratio more than 90°
And to this movement that of the moon herself will conform,
a species of elliptic spiroid, having its least distances always in
syzygies and its greatest in that of quadratures. It is evident,
owing to the longer continued action of both forces, *i. e.* owing
ter arc over which their intensities increase and decrease by
, the branches of each curve between the cusps will be longer,
ps themselves will be more remote from S, and in the same
the dimensions of the resulting oval be enlarged, and with
amount of the inequality in the moon's motion which they

Fig. 99.





of the compound oval in which the focus really revolves, which will, therefore, be as in the annexed figure. Thus, at the end of a complete lunation, the focus will have shifted its place from e to f in a line parallel to the line of quadratures. The next revolution, and the next, the same thing would happen. Meanwhile, however, the sun has advanced in its orbit, and the line of quadratures has changed its situation by an equal angular motion. In consequence, the next terminal situation (g) of the focus will not lie in the line ef prolonged, but in a line parallel to the new situation of the line of quadratures, and this process continuing, will evidently give rise to a movement of circulation of the point e, round a mean situation in an annual period; and this, it is evident, is equivalent to an annual circulation of the central point of the compound oval itself, in a small orbit about its mean position S. Thus we see that no permanent and indefinite increase of excentricity can arise from this cause; which would be the case, however, but for the annual motion of the sun.

(718.) Inequalities precisely similar in principle to the variation and parallactic inequality of the moon, though greatly modified by the different relations of the dimensions of the orbits, prevail in all cases where planet disturbs planet. To what extent this modification is carried will be evident, if we cast our eyes on the examples given in art. 612, where it will be seen that the disturbing force in conjunction often exceeds that in opposition in a very high ratio, (being in the case of Neptune disturbing Uranus more than ten times as great.) The effect will be, that the orbit described by the centre of the compound oval about S, will be much greater relatively to the dimensions of that oval itself, than in the case of the moon. Bearing in mind the nature and import of this modification, we may proceed to inquire, apart from it, into the number and distribution of the undulations in the contour of the oval itself, arising from the alter-

direction *plus* and *minus* of the disturbing forces in the course
dic revolution. But first it should be mentioned that, in the
exterior disturbed by an interior planet, the disturbing body's
iotion exceeds that of the disturbed. Hence P, though advan-
orbit, recedes relatively to the line of syzygies, or, which comes
ie thing, the neutral points of either force overtake it in succes-
each, as it comes up to it, gives rise to a cusp in the corresponding
ve. The angles between the successive cusps will therefore be
les between the corresponding neutral points for a fixed position
he same constant ratio of the synodic to the sidereal period of P,
vever is now a ratio of less inequality. These angles then will
ted in amplitude, and, for the same reason as before, the excur-
he focus will be diminished, and the more so the shorter the
volution.

Since the cusps of either curve recur, in successive synodic
s in the same order, and at the same angular distances from
, and from the line of conjunction, the same will be true of all
ponding points in the curve resulting from their superposition.
urve, every cusp, of either constituent, will give rise to a con-
d every intercuspidal arc to a relative concavity. It is evident

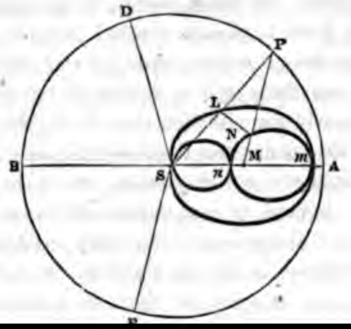
twards, or from accelerating to retarding, and *vice versa*, of the normal d tangential forces,) in a complete synodic revolution, and their distrition over the circumference.

(715.) With regard to these changes, it is necessary to distinguish ree cases, in which the perturbations of planet by planet are very disct in character. 1st. When the disturbing planet is exterior. In this e there are four neutral points of either force. Those of the tangenl force occur at the syzygies, and at the points of the disturbed orbit (which we shall call points of equidistance), equidistant from the sun and the disturbing planet (at which points, as we have shown (art. 614), the total disturbing force is always directed inwards towards the sun.) Those of the normal force occur at points intermediate between these last mentioned points, and the syzygies, which, if the disturbing planet be very distant, hold nearly the situation they do in the lunar theory, *i. e.* considerably nearer the quadratures than the syzygies. In proportion as the distance of the disturbing planet diminishes, two of these points, viz. those nearest the syzygy, approach to each other, and to the syzygy, and in the extreme case, when the dimensions of the orbits are equal, coincide with it.

(716.) The second case is that in which the disturbing planet is interior to the disturbed, but at a distance from the sun greater than half that of the latter. In this case there are four neutral points of the tangential force, and only two of the normal. Those of the tangential force occur at the syzygies, and at the points of equidistance. The force retards the disturbed body from conjunction to the first such points after conjunction, accelerates it thence to the opposition, thence again retards it to the next point of equidistance, and finally again accelerates it up to the conjunction. As the disturbing orbit contracts in dimension, the points of equidistance approach; their distance from syzygy from 60° (the extreme case) diminishing to nothing, when they coincide with each other, and with the conjunction. In the case of Saturn disturbed by Jupiter, that distance is only 23° 33′. The neutral points of the normal force lie somewhat beyond the quadratures, on the side of the opposition, and do not undergo any very material change of situation with the contraction of the disturbing orbit.

(717.) The third case is that in which the diameter of the disturbing interior orbit is less than half that of the disturbed. In this case there are only two points of evanescence for either force. Those of the tangential force are the syzygies. The disturbed planet is accelerated throughout the whole semi-revolution from conjunction to opposition, and retarded from opposition to conjunction, the maxima of acceleration and retardation occurring not far from quadrature. The neutral points of the normal ituated nearly as in the last case; that is to say, beyond the s towards the opposition. All these varieties the student will e out by simply drawing the figures, and resolving the forces in f cases, beginning with a very large and ending with a very neter of the disturbing orbit. It will greatly aid him in im- n his imagination the general relations of the subject, if he as he proceeds, for each case, the elegant and symmetrical ovals

Fig. 101.



(718.) It would be easy now, bearing these features in mind, to trace in any proposed case the form of the spiroid curve, described, as above explained, by the upper focus. It will suffice, however, for our present purpose, to remark, 1st, That between every two successive conjunctions of P and M, the same general form, the same subordinate undulations, and the same terminal displacement of the upper focus, are continually repeated. 2dly, That the motion of the focus in this curve is retrograde whenever the disturbing planet is exterior, and that in consequence the apsides of the momentary ellipse also recede, with a mean velocity such as, but for that displacement, would bring them round at the each conjunction to the same relative situation with respect to the line of syzygies. 3dly, That in consequence of this retrograde movement of the apse, the disturbed planet, apart from that consideration, would be twice in perihelio and twice in aphelio in its momentary ellipse in each synodic revolution, just as in the case of the moon disturbed by the sun — and that in consequence of this and of the undulating movement of the focus H itself, an inequality will arise, analogous, *mutatis mutandis* in each case, to the moon's variation; under which term we comprehend (not exactly in conformity to its strict technical meaning in the lunar theory) not only the principal inequality thus arising, but all its subordinate fluctuations. And on this the parallactic inequality thus violently exaggerated is superposed.

(719.) We come now to the class of inequalities which depend for their existence on an appreciable amount of *permanent* excentricity in the orbit of one or of both the disturbing and disturbed planets, in consequence of which all their conjunctions do not take place at equal distances either from the central body or from each other, and therefore that symmetry in every synodic revolution on which depends the exact restoration of both the axis and excentricity to their original values at the completion of each such revolution no longer subsists. In passing from conjunction to conjunction, then, there will no longer be effected a complete restoration of the upper focus to the same relative situation, or of the axis to the same length, which they respectively had at the outset. At the same time it is not less evident that the differences in both respects are only what remain outstanding, after the compensation of by far the greater part of the deviations to and fro from a mean state, which occur in the course of the revolution; and that they amount to but small fractions of the total excursions of the focus from its first position, or of the increase and decrease in the length of the axis effected by the direct action of the tangential force, — so small, indeed, that, unless owing to peculiar adjustments they be enabled to accumulate again and again at

conjunctions in the same direction, they would be altogether
g of any especial notice in a work of this nature. Such ad-
however, would evidently exist if the periodic times of the
e exactly commensurable; since in that case all the possible
s which could ever happen (the elements not being materially
ould take place at fixed points in longitude, the intermediate
g never visited by a conjunction. Now, of the conjunctions
buted, their relations to the lines of symmetry in the orbits
issimilar, some one must be more influential than the rest on
elements (not necessarily the *same* upon *all*). Consequently,
te cycle of conjunctions, wherein each has been visited in its
fluence of that one on the element to which it stands so espe-
ed, will preponderate over the counteracting and compensating
the rest, and thus, although in such a cycle as above specified
nd much more exact compensation will have been effected in
an in a single revolution, still that compensation will not be
ut a portion of the effect (be it to increase or to diminish the
y or the axis, or to cause the apse to advance or to recede,)
outstanding. In the next cycle of the same kind this will be
nd the result will be of the same character, and so on, till at

The excess of action, for several such triple conjunctions (7 or 8) in succession, will lie the same way, and at each of them the elements of P's orbit and its angular motion will be similarly influenced, so as to accumulate the effect upon its longitude; thus giving rise to an irregularity of considerable magnitude and very long period, which is well known to astronomers by the name of the great inequality of Jupiter and Saturn.

(721.) The arc 8° 6′ is contained $44\frac{4}{9}$ times in the whole circumference of 360°; and accordingly, if we trace round this particular conjunction, we shall find it will return to the same point of the orbit in so many times 21760 days, or in 2648 years. But the conjunction we are now considering is only one out of three. The other two will happen at points of the orbit about 123° and 246° distant, and *these points also will advance* by the same arc of 8° 6′ in 21760 days. Consequently the period of 2648 years will bring them *all* round, and in that interval each of them will pass through that point of the two orbits from which we commenced: hence *a conjunction* (one or other of the three) will happen at that point once in one third of this period, or in 883 years; and this is, therefore, the cycle in which the "great inequality" would undergo its full compensation, did the elements of the orbits continue all that time invariable. Their variation, however, is considerable in so long an interval; and, owing to this cause, the period itself is prolonged to about 918 years.

(722.) We have selected this inequality as the most remarkable instance of this kind of action on account of its magnitude, the length of its period, and its high historical interest. It had long been remarked by astronomers, that on comparing together modern with ancient observations of Jupiter and Saturn, the mean motions of these planets did not appear to be uniform. The period of Saturn, for instance, appeared to have been lengthening throughout the whole of the seventeenth century, and that of Jupiter shortening—that is to say, the one planet was constantly lagging behind, and the other getting in advance of its calculated place. On the other hand, in the eighteenth century, a process precisely the reverse seemed to be going on. It is true the whole retardations and accelerations observed were not very great; but, as their influence went on accumulating, they produced, at length, material differences between the observed and calculated places of both these planets, which as they could not then be accounted for by any theory, excited a high degree of attention, and were even, at one time, too hastily regarded as almost subversive of the Newtonian doctrine of gravity. For a long while this difference baffled every endeavour to account for it; till at length Laplace pointed

e in the near commensurability of the mean motions, as above
succeeded in calculating its period and amount.

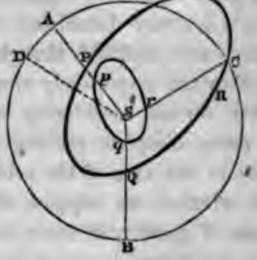
'he inequality in question amounts, at its maximum, to an al-
n and loss of about 0° 49′ in the longitude of Saturn, and a
ng loss and gain of about 0° 21′ in that of Jupiter. That an
in the one planet must necessarily be accompanied by a re-
the other, might appear at first sight self-evident, if we con-
action and reaction being equal, and in contrary directions,
nomentum Jupiter communicates to Saturn in the direction
ame momentum must Saturn communicate to Jupiter in the
P. The one, therefore, it might seem to be plausibly argued,
gged forward, whenever the other is pulled back in its orbit.
nce is correct, *so far as the general and final result goes;* but
ng by which it would, on the first glance, appear to be thus
established is fallacious, or at least incomplete. It is perfectly
hatever momentum Jupiter communicates directly to Saturn,
municates an equal momentum to Jupiter in an opposite linear
But it is not with the absolute motions of the two planets in
we are now concerned, but with the relative motion of each

far as a disturbance of their elliptic motions *about the common centre of gravity of the system is alone regarded*, whatever disturbance of velocity is generated in the one, a contrary disturbance of velocity (only in the inverse ratio of the masses and modified, though never contradicted, by the directions in which they are respectively moving), will be generated in the other. Now when we are considering only inequalities of long period comprehending many complete revolutions of both planets, and which arise from changes in the axes of the orbits, affecting their *mean motions*, it matters not whether we suppose these motions performed about the common centre of gravity, or about the sun, which never departs from that centre to any material extent (the mass of the sun being such in comparison with that of the planets, that that centre always lies within his surface.) The *mean* motion therefore, regarded as the average angular velocity during a revolution, is the same whether estimated by reference to the sun's centre, or to the centre of gravity, or, in other words, the relative mean motion referred to the sun is identical with the absolute mean motion referred to the centre of gravity.

(724.) This reasoning applies equally to every case of mutual disturbance resulting in a long inequality such as may arise from a slow and long-continued periodical increase and diminution of the axes, and geometers have accordingly demonstrated as a consequence from it, that the proportion in which such inequalities affect the *longitudes* of the two planets concerned, or the maxima of the excesses and defects of their longitudes above and below their elliptic values, thence arising, in each, are to each other in the inverse ratio of their masses multiplied by the square roots of the major axes of their orbits, and this result is confirmed by observation, and will be found verified in the instance immediately in question as nearly as the uncertainty still subsisting as to the masses of the two planets will permit.

(725.) The inequality in question, as has been observed in general, (art. 718,) would be much greater, were it not for the partial compensation which is operated in it in every triple conjunction of the planets. Suppose P Q R to be Saturn's orbit, and *p q r* Jupiter's; and suppose a conjunction to take place at P*p*, on the line S A; a second at 123° distance, on the line S B; a third at 246° distance, on S C; and the next at 368°, on S D. This last-mentioned conjunction, taking place nearly in the situation of the first, will produce nearly a repetition of the first effect in retarding or accelerating the planets; but the other two, being in the most remote situations possible from the first, will happen under entirely different circumstances as to the position of the perihelia of the orbits. Now, we have seen that a presentation of the one planet to the other in

Fig. 103.



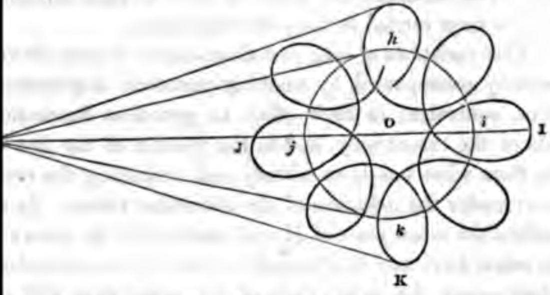
conjunction, in a variety of situations, tends to produce compensation; and, in fact, the greatest possible amount of compensation which can be produced by only three conjunctions is when they are thus equally distributed round the centre. Hence we see that it is not the whole amount of perturbation which is thus accumulated in each triple conjunction, but only that small part which is left uncompensated by the intermediate ones. The reader, who possesses already some acquaintance with the subject, will not be at a loss to perceive how this consideration is, in fact, equivalent to that part of the geometrical investigation of this inequality which leads us to seek its expression in terms of the third order, or involving the cubes and products of three dimensions of the excentricities and inclinations; and how the continual accumulation of small quantities, during long periods, corresponds to what geometers intend when they speak of small terms receiving great accessions of magnitude by the introduction of large coefficients in the process of integration.

(726.) Similar considerations apply to every case of approximate commensurability which can take place among the mean motions of any two planets. Such, for instance, is that which obtains between the mean motion of the earth and Venus,—13 times the period of Venus being very nearly equal to 8 times that of the earth. This gives rise to an extremely near coincidence of every fifth conjunction, in the same parts of each orbit (within $\frac{1}{240}$th part of a circumference,) and therefore to a correspondingly extensive accumulation of the resulting uncompensated perturbation. But, on the other hand, the part of the perturbation thus accumulated is only that which remains outstanding after passing the equalizing ordeal of five conjunctions equally distributed round the circle; or, in the language of geometers, is dependent on powers and products of the excentricities and inclinations of the fifth order. It is, therefore, extremely minute, and the whole resulting inequality, according to the elaborate calculations of Mr. Airy, to whom it owes its detection, amounts to no more than a few seconds at its maximum, while its period is no less than 240 years. This example will serve to show to what minuteness these inquiries have been carried to the planetary theory.

(727.) That variations of long period arising in the way above described are necessarily accompanied by similarly periodical displacements of the upper focus, equivalent in their effect to periodical fluctuations in the magnitude of the excentricity, and in the position of the line of apsides, is evident from what has been already said respecting the motion of the upper focus under the influence of the disturbing forces. In the case of circular orbits the *mean* place of H coincides with S the centre of the sun, but if the orbits have any independent ellipticity, this coincidence will no longer exist — and the *mean* place of the upper focus will come to be inferred from the average of all the situations which it actually holds during an entire revolution. Now the fixity of this point depends on the equality of each of the branches of the cuspidated curves, and consequent equality of excursion of the focus in each particular direction, in every successive situation of the line of conjunction. But if there be some one line of conjunction in which these excursions are greater in any one particular direction than in another, the mean place of the focus will be displaced, and if this process be repeated, that mean place will continue to deviate more and more from its original position, and thus will arise a circulation of the *mean* place *of the focus for a revolution* about another mean situation, the average of all the former mean places during a complete cycle of conjunctions. Supposing S to be the sun, O the situation the upper focus would have, had these inequalities no existence, and H K the path of the upper focus, which it pursues about O by reason of them, then it is evident that in the course of a complete cycle of the inequality in question, the excentricity will have fluctuated between the extreme limits S J and S I and the direction of the longer axis between the extreme position S H and S K, and that if we suppose $i j h k$ to be the corresponding mean places of the focus, $i j$ will be the extent of the fluctuation of the mean excentricity, and the angle $h s k$, that of the longitude of the perigee.

(728.) The periods then in which these fluctuations go through their phases are necessarily equal in duration with that of the inequality in longitude, with which they stand in connexion. But it by no means follows that their maxima all coincide. The variation of the axis to which that of the mean motion corresponds, depends on the tangential force only whose maximum is not at conjunction or opposition, but at points remote from either, while the excentricity depends both on the normal and tangential forces, the maximum of the former of which is at the conjunction

Fig. 104.



ular conjunction therefore, which is most influential on the so on the excentricity, so that it can by no means be concluded the maximum value of the axis coincides with the maximum, mum of the excentricity, or with the greatest excursion to or ine of apsides from its mean situation, all that can be safely that as either the axis or the excentricity of the one orbit of the other will vary in the opposite direction.

he primary elements of the lunar and planetary orbits, which ment under the title "Epoch of the mean longitude," the mean longitude of the planet so computed for a fixed date; as, for instance, the commencement of the year 1800, mean time at a given place. Supposing now all elements of the orbit invariable, if we were to go through this reverse process, and thus ascertain the epoch (so defined) from any number of different perfectly correct heliocentric longitudes, it is clear we should always come to the same result. One and the same "epoch" would come out from all the calculations.

(730.) Considering then the "epoch" in this light, as merely a result of this reversed process of calculation, and not as the direct result of an observation instituted for the purpose at the precise epochal moment of time, (which would be, generally speaking, impracticable,) it might be conceived subject to variation in two distinct ways, viz. dependently and independently. Dependently it must vary, as a necessary consequence of the variation of the other elements; because, if setting out from one and the same observed heliocentric longitude of the planet, we calculate back to the epoch with two different sets of intermediate elements, the one set consisting of those which it had immediately before its arrival at that longitude, the other that which it takes up immediately after (*i. e.* with an unvaried and varied system), we cannot (unless by singular accident of mutual counteraction) arrive at the same result; and the difference of the results is evidently the variation of the epoch. On the other hand, however, it cannot vary independently; for since this is the only mode in which the unvaried and varied epochs can become known, and as both result from direct processes of calculation involving only given data, the results can only differ by reason of the difference of those data. Or we may argue thus. The change in the path of the planet, and its place in that path so changed, at any future time (supposing it to undergo no further variation), are entirely owing to the change in its velocity and direction, produced by the disturbing forces at the point of disturbance; now these latter changes (as we have above seen) are *completely* represented by the momentary change in the situation of the upper focus, taken in combination with the momentary variation in the plane of the orbit; and these therefore express the total effect of the disturbing forces. There is, therefore, no direct and specific action on the epoch as an independent variable. It is simply left to accommodate itself to the altered state of things in the mode already indicated.

(731.) Nevertheless, should the effects of perturbation by inducing changes on these other elements affect the mean longitude of the planet in any other way than can be considered as properly taken account of, by the varied periodic time due to a change of axis, such effects must be re-

incident on the epoch. This is the case with a very curious
turbations which we are now to consider, and which have their
alteration of the average distance at which the disturbed body
every instant of a complete revolution, distinct from, and not
out by the variation of the major semi-axis, or *momentary*
tance" which is an imaginary magnitude, to be carefully distin-
m the averge of the actual distances now contemplated. Per-
of this class (like the moon's variation, with which they are
connected) are independent on the excentricity of the disturbed
which reason we shall simplify our treatment of this part of the
supposing that orbit to have no permanent excentricity, the
in its successive displacements merely revolving about a mean
incident with the lower. We shall also suppose M very dis-
the lunar theory.

Referring to what is said in arts. 706 and 707, and to the figures
ng those articles, and considering first the effect of the tangential
e that besides the effect of that force in changing the length
and consequently the periodic time, it causes the upper focus
ibe, in each revolution of P, a four-cusped curve, *a*, *b*, *d*, *e*,

the disturbing force. Although, therefore, in perigee at syzygies and in apogee at quadratures, the apogeal recess is much greater than the perigeal approach, inasmuch as S *d* greatly exceeds S *a*. On the average of a whole revolution, then, the recesses will preponderate, and the average distance will therefore be greater in the disturbed than in the undisturbed orbit. And it is manifest that this conclusion is quite independent of any change in the length of the axis, which the normal force has no power to produce.

(734.) But neither does the normal force operate any change of linear velocity in the disturbed body. When carried out, therefore, by the effect of that force to a greater distance from S, the angular velocity of its motion round S will be diminished: and contrariwise when brought nearer. The average of all the momentary angular motions, therefore, will decrease with the increase in that of the momentary distances; and in a higher ratio, since the angular velocity, under an equable description of areas, is inversely as the square of the distance, and the disturbing force, being (in the case supposed) directed to or from the centre, does not disturb that equable description (art. 616). Consequently, on the average of a whole revolution, the angular motion is slower, and therefore the time of completing a revolution, and returning to the same longitude, longer than in the undisturbed orbit, and *that* independent of and without any reference to the length of the momentary axis, and the "periodic time" or "mean motion" dependent thereon. We leave to the reader to follow out (as is easy to do) the same train of reasoning in the cases of planetary perturbation, when M is not very remote, and when it is interior to the disturbed orbit. In the latter case the preponderant effect changes from a retardation of angular velocity to an acceleration, and the dilatation of the average dimensions of P's orbit to a contraction.

(735.) The above is an accurate analysis, according to strict dynamical principles, of an effect which, speaking roughly, may be assimilated to an alteration of M's gravitation towards S by the mean preponderant amount of the outward and inward action of the normal forces constantly exerted — nearly as would be the case if the mass of the disturbing body were formed into a ring of uniform thickness, concentric with S, and of such diameter as to exert an action on P everywhere equal to such mean preponderant force, and in the same direction as to inwards or outwards. For it is clear that the action of such a ring on P, will be the difference of its attractions on the two points P and S, of which the latter occupies its centre, the former is excentric. Now the attraction of a ring on its centre is manifestly equal in all directions, and therefore, estimated in any one direction, is zero. On the other hand, on a point P out of its

ithin the ring, the resulting attraction will always be *outwards*,
nearest point of the ring, or directly from the centre.[1] But
thout the ring, the resulting force will act always *inwards*,
wards its centre. Hence it appears that the mean effect of
orce of the ring will be different in its direction, according as
the disturbing body is exterior or interior to that of the dis-
the former case it will act in diminution, in the latter in
on of the central gravity.

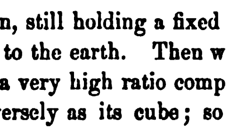
egarding, still, only the mean effect, as produced in a great
revolutions of both bodies, it is evident that such an increase
orce will be accompanied with a diminution of periodic time
e of a body revolving with a stated velocity, and *vice versâ*.
have shown, is the first and most obvious effect of the radial
disturbing force, when exactly analyzed. It alters permanently,
rtain mean amount, the distances and times of revolution of
ies composing the planetary system, from what they would be,
anet circulate about the sun uninfluenced by the attraction of
e angular motion of the interior bodies of the system being
ed less, and those of the exterior greater, than on that suppo-
e latter effect indeed might be at once concluded from this

orbit may be considered as adding to the general aggregate of the attracting matter within, which is not the less efficient for being distributed over space, and maintained in a state of circulation.

(787.) This effect, however, is one which we have no means of measuring, or even of detecting, otherwise than by calculation. For our knowledge of the periods of the planets is drawn from observations made on them in their actual state, and therefore under the influence of this, which may be regarded as a sort of *constant part* of the perturbative action. Their observed mean motions are therefore affected by the whole amount of its influence; and we have no means of distinguishing this by observation from the direct effect of the sun's attraction, with which it is blended. Our knowledge, however, of the masses of the planets assures us that it is extremely small; and this, in fact, is all which it is at all important to us to know, in the theory of their motions.

(788.) The action of the sun upon the moon, in like manner, tends, by its mean influence during many successive revolutions of both bodies, to increase permanently the moon's distance and periodic time. But this general average is not established, either in the case of the moon or planets, without a series of subordinate fluctuations, which we have purposely neglected to take account of in the above reasoning, and which obviously tend, in the average of a great multitude of revolutions, to neutralize each other. In the lunar theory, however, some of these subordinate fluctuations are very sensible to observation. The most conspicuous of these is the moon's annual equation; so called because it consists in an alternate increase and decrease in her longitude, corresponding with

Fig. 105.



the earth's situation in its annual orbit; *i. e.* to its angular distance from the perihelion, and therefore having a year instead of a month, or aliquot part of a month, for its period. To understand the mode of its production, let us suppose the sun, still holding a fixed position in longitude, to approach gradually nearer to the earth. Then will all its disturbing forces be gradually increased in a very high ratio compared with the diminution of the distance (being inversely as its cube; so that its effects of every kind are three times greater in respect of any change of distance, than

be by the simple law of proportionality). Hence, it is obvious
us H (art. 707) in the act of describing each intercuspidal arc
ve *a*, *d*, *b*, *e*, will be continually carried out farther and farther
d the curve, instead of returning into itself at the end of each
will open out into a sort of cuspidated spiral, as in the figure
Retracing now the reasoning of art. 733, as adapted to this
ngs, it will be seen that so long as this dilatation goes on, so
he difference between M's recess from S in aphelio and its
perihelio (which is equal to the difference of consecutive long
emidiameters of this curve) also continue to increase, and with
age of the distances of M from S in a whole revolution, and
y also the time of performing such a revolution. The reverse
go on as the sun again recedes. Thus it appears that, as the
hes the earth, the mean angular motion of the moon on the
a whole revolution will diminish, and the duration of each
ll therefore exceed that of the foregoing, and *vice versâ.*

he moon's orbit being supposed circular, the sun's orbitual
have no other effect than to keep the moon longer under the
every gradation of the disturbing force, than would have been
d his situation in longitude remained unaltered (art. 711.) The

were on the point of again declaring the theory of gravity inadequate to its explanation, others were for rejecting altogether the evidence on which it rested, although quite as satisfactory as that on which most historical events are credited. It was in this dilemma that Laplace once more stepped in to rescue physical astronomy from its reproach, by pointing out the real cause of the phænomenon in question, which, when so explained, is one of the most curious and instructive in the whole range of our subject,— one which leads our speculations farther into the past and future, and points to longer vistas in the dim perspective of changes which our system has undergone and is yet to undergo, than any other which observation assisted by theory has developed.

(741.) The year is not an exact number of lunations. It consists of twelve and a fraction. Supposing then the sun and moon to set out from conjunction together; at the twelfth conjunction subsequent the sun will not have returned precisely to the same point of its annual orbit, but will fall somewhat short of it, and at the thirteenth will have overpassed it. Hence in twelve lunations the gain of longitude during the first half year will be somewhat under and in thirteen somewhat over-compensated. In twenty-six it will be nearly twice as much over-compensated, in thirty-nine not quite so nearly three times as much, and so on, until, after a certain number of such multiples of a lunation have elapsed, the sun will be found half a revolution in advance, and in place of receding farther at the expiration of the next, it will have begun to approach. From this time every succeeding cycle will destroy some portion of that over-compensation, until a complete revolution of the sun in excess shall be accomplished. Thus arises a subordinate or rather supplementary inequality, having for its period as many years as is necessary to multiply the deficient arc into a whole revolution, at the end of which time a much more exact compensation will have been operated, and so on. Thus after a moderate number of years an almost perfect compensation will be effected, and if we extend our views to centuries we may consider it as quite so. Such at least would be the case if the solar ellipse were invariable. But that ellipse is kept in a continual but excessively slow state of change by the action of the planets on the earth. Its axis, it is true, remains unaltered; but its excentricity is, and has been since the earliest ages, diminishing; and this diminution will continue (there is little reason to doubt) till the excentricity is annihilated altogether, and the earth's orbit becomes a perfect circle; after which it will again open out into an ellipse, the excentricity will again increase, attain a certain moderate amount, and then again decrease. The time required for these evolutions, though calculable, has not been calculated, further than to satisfy us that it is

reckoned by hundreds or by thousands of years. It is a period,
in which the whole history of astronomy and of the human race
but as it were a point, during which all its changes are to be
as uniform. Now, it is by this variation in the excentricity of
's orbit that the secular acceleration of the moon is caused. The
tion above spoken of (even after the lapse of centuries) will now,
e only imperfectly effected, owing to this slow shifting of one
ssential data. The steps of restoration are no longer identical
equal to, those of change. The struggle up hill is not main-
n equal terms with the downward tendency. The ground is all
slowly sliding beneath the feet of the antagonists. During the
ne that the earth's excentricity is diminishing, a preponderance
to the reaction over the action; and it is not till that diminution
se, that the tables will be turned, and the process of ultimate
n will commence. Meanwhile, a minute, outstanding, and un-
ted effect in favour of acceleration is left at each recurrence, or
rrence, of the same configurations of the sun, the moon, and the
igee. These accumulate, and at length affect the moon's longi-
n extent not to be overlooked.

The phænomenon, of which we have now given an account, is reason of this will be readily apprehended, if we consider that however trifling the increase of her longitude which would arise in a single revolution, from a minute and almost infinitesimal increase of her mean angular velocity, that increase is not only repeated in each subsequent revolution, but is reinforced during each by a similar fresh accession of angular motion generated in its lapse. This process goes on so long as the angular motion continues to increase, and only begins to be reversed when lapse of time, bringing round a contrary action on the angular motion, shall have destroyed the excess of velocity previously gained, and begun to operate a retardation. In this respect, the advance gained by the moon on her undisturbed place may be assimilated, during its increase, to the space described from rest under the action of a continually accelerating force. The velocity gained in each instant is not only effective in carrying the body forward during each subsequent instant, but new velocities are every instant generated, and go on adding their cumulative effects to those before produced.

(744.) The distance of the earth from the sun, like that of the moon from the earth, may be affected in its average value estimated over long periods embracing many revolutions, in two modes, conformably to the theory above delivered. 1st, it may vary by a variation in the length of the axis major of its orbit, arising from the direct action of some tangential disturbing force on its velocity, and thereby producing a change of mean motion and periodic time in virtue of the Keplerian law of periods, which declares that the periodic times are in the sesquiplicate ratio of the mean distances. Or, 2dly, it may vary by reason of that peculiar action on the average of actual distances during a revolution, which arises from variations of excentricity and perihelion only, and which produces that sort of change in the mean motion which we have characterized as incident on the epoch. The change of mean motion thus arising, has nothing whatever to do with any variation of the major axis. It does not depend on the change of distance by the Keplerian law of *periods*, but by that of *areas*. The altered mean motion is not sub-sesquiplicate to the altered axis of the ellipse, which in fact does not alter at all, but is *sub-duplicate* to the altered *average of distances* in a revolution; a distinction which must be carefully borne in mind by every one who will clearly understand either the subject itself, or the force of Newton's explanation of it in the 6th Corollary of his celebrated 66th Proposition. In whichever mode, however, an alteration in the mean motion is effected, if we accommodate the general sense of our language to the specialities of the case, it remains true that every change in the mean motion is accompanied with a corresponding change in the mean distance.

Now we have seen (art. 726), that Venus produces in the earth
tion in longitude, of so long a period (240 years), that it can-
regarded without violence to ordinary language, otherwise than
tion of the mean motion. Of course, therefore, it follows that
t half of this long period of time, in which the earth's motion
, the distance between the sun and earth is on the increase, and

Minute as is the equation in question, and consequent altera-
ar distance, and almost inconceivably minute as is the effect
on the moon's mean angular velocity in a single lunation, yet
number of lunations (1484), during which the effect goes on
ng in one direction, causes the moon, at the moment when that
on has attained its maximum to be very sensibly in advance of
rbed place (viz. by 23″ of longitude), and after 1484 more
as much in arrear. The calculations by which this curious
been established, formidable from their length and intricacy,
the industry, as the discovery of its origin is to the sagacity,
or Hansen.

The action of Venus, just explained, is indirect, being as it
of reflection of its influence on the earth's orbit. But a very

to the law described. Each of these periodically acting forces produces its perturbative effect, according to the law of the superposition of small motions, as if the others had no existence. And if it happen, as in an immense majority of cases it does, that the cycle of any particular one of these partial forces has no relation to the periodic time of the disturbed body, so as to bring it to the same, or very nearly the same point of its orbit, or to any situation favourable to any particular form of disturbance, over and over again when the force is at its maximum; that force will, in a few revolutions, neutralize its own effect, and nothing but fluctuations of brief duration can result from its action. The contrary will evidently be the case, if the cycle of the force coincide so nearly with the cycle of the moon's anomalistic revolution, as to bring round the maximum of the force acting in one and the same direction (whether tangential or normal) either accurately, or very nearly indeed to some definite point, as, for example, the apogee of her orbit. Whatever the effect produced by such a force on the angular motion of the moon, if it be not exactly compensated in one cycle of its action, it will go on accumulating, being repeated over and over again under circumstances very nearly the same, for many successive revolutions, until at length, owing to the want of precise accuracy in the adjustment of that cycle to the anomalistic period, the maximum of the force (in the same phase of its action) is brought to coincide with a point in the orbit (as the perigee), determinative of an opposite effect, and thus, at length, a compensation will be worked out; in a time, however, so much the longer as the difference between the cycle of the force and the moon's anomalistic period is less.

(747.) Now, in fact, in the case of Venus disturbing the moon, there exists a cyclical combination of this kind. Of course the disturbing force of Venus on the moon varies with her distance from the earth, and this distance again depends on her configuration with respect to the earth and the sun, *taking into account the ellipticity of both their orbits.* Among the combinations which take their rise from this latter consideration, and which, as may easily be supposed, are of great complexity, there is a term (an exceedingly minute one), whose argument or cycle is determined by subtracting 16 times the mean motion of the earth from 18 times that of Venus. The difference is so very nearly the mean motion of the moon in her anomalistic revolution, that whereas the latter revolution is completed in $27^d\ 18^h\ 18^m\ 32{\cdot}3^s$, the cycle of the force is completed in $27^d\ 18^h\ 7^m\ 35{\cdot}6^s$, differing from the other by no more than $10^m\ 56{\cdot}7^s$, or about one 3625th part of a complete period of the moon from apogee to apogee. During half of this very long interval (that is to say, during about 136½ years), the perturbations produced by a force of this character,

he lunar perturbations, that we cannot omit some explanatory
, especially since many persons find a strange difficulty in con-
manner in which they are produced. That the sun, or moon,
its attraction heap up the waters of the ocean under it, seems
ry natural. That it should at the same time heap them up on
e side seems, on the contrary, palpably absurd. The error of
f objectors is of the same kind with that noticed in art. 723,
s in disregarding the attraction of the disturbing body on the
e earth, and looking on it as wholly effective on the superficial
ere the earth indeed absolutely fixed, held in its place by an
rce, and the water left free to move, no doubt the effect of the
power would be to produce a single accumulation vertically
isturbing body. But it is not by its whole attraction, but by
ce of its attractions on the superficial water at both sides, and
ral mass, that the waters are raised: just as in the theory of
the difference of the sun's attractions on the moon and on the
rded as moveable and as obeying that amount of attraction
e to its situation) gives rise to a *relative* tendency in the moon
rom the earth in conjunction and opposition, and to approach

form of the earth) urges P towards S, will be equivalent to a single force deviating from the direction P S towards X. Suppose P T to be the direction of this force, which, it is easy to see, will be directed towards a point in D S *produced*, at an extremely small distance below S, because of the excessive minuteness of the disturbing force compared to gravity.[1] Then if this be done at every point of the quadrant A D, it will be evident that the direction P T of the resultant force will be always that of a tangent to the small cuspidated curve *a d* at T, to which tangent the surface of the ocean at P must everywhere be perpendicular, by reason of that law of hydrostatics which requires the direction of gravity to be everywhere perpendicular to the surface of a fluid *in equilibrio*. The form of the curve D P A, to which the surface of the ocean will tend to conform itself, so as to place itself everywhere in equilibrio under two acting forces, will be that which always has P T for its radius of curvature. It will therefore be slightly less curved at D, and more so at A, being in fact no other than an ellipse, having S for its centre, *d a* for its evolute, and S A, S D for its longer and shorter semi-axes respectively; so that the whole surface (supposing it covered with water) will tend to assume, as its form of equilibrium, that of an oblongated ellipsoid, having its longer axis directed towards the disturbing body, and its shorter of course at right angles to that direction. The difference of the longer and shorter semi-axes of this ellipsoid due to the moon's attraction would be about 58 inches: that of the ellipsoid, similarly formed in virtue of the sun, about 2½ times less, or about 23 inches.

(751.) Let us suppose the moon only to act, and to have no orbitual motion; then if the earth also had no diurnal motion, the ellipsoid of equilibrium would be quietly formed, and all would be thenceforward tranquil. There is never time, however, for this spheroid to be fully formed. Before the waters can take their level, the moon has advanced in her orbit, both diurnal and monthly, (for in this theory it will answer the purpose of clearness better, if we suppose the earth's diurnal motion transferred to the sun and moon in the contrary direction,) the vertex of the spheroid has shifted on the earth's surface, and the ocean has to seek a new bearing. The effect is to produce an immensely broad and excessively flat wave (not a circulating *current*), which follows, or endeavours

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[1] According to Newton's calculation, the maximum disturbing force of the sun on the water does not exceed one 25736400th part of its gravity. That of the moon will therefore be to this fraction as the cube of the sun's distance to that of the moon's directly, and as the mass of the sun to that of the moon inversely, *i. e.* as $(400)^3 \times 0{\cdot}012517 : 354936$, which, reduced to numbers, gives, for the moon's maximum of power to disturb the waters, about one 11400000th of gravity, or somewhat less than 2½ times the sun's.

the apparent motions of the moon, and must, in fact, by the
f forced vibrations, imitate, by equal though not by *synchronous*
the periodical inequalities of that motion. When the higher
arts of this wave strike our coasts, they experience what we
nd low water.

The sun also produces precisely such a wave, whose vertex tends
he apparent motion of the sun in the heavens, and also to imi-
iodic inequalities. This solar wave co-exists with the lunar—
es superposed on it, sometimes transverse to it, so as to partly
t, according to the monthly synodical configuration of the two
This alternate mutual reinforcement and destruction of the
unar tides cause what are called the spring and neap tides—the
g their sum, the latter their difference. Although the real
either tide is, at present, hardly within the reach of exact cal-
et their proportion at any one place is probably not very remote
f the ellipticities which would belong to their respective sphe-
l an equilibrium be attained. Now these ellipticities, for the
unar spheroids, are respectively about two and five feet; so that
spring tide will be to the neap as 7 to 3, or thereabouts.

guished would be identical. But all these causes tend to create a difference, and to make that difference not alike at all ports. The observation of the establishments of harbours is a point of great maritime importance; nor is it of less consequence, theoretically speaking, to a knowledge of the true distribution of the tide-waves over the globe. In making such observations, care must be taken not to confound the time of "slack water," when the current caused by the tide ceases to flow visibly one way or the other, and that of *high* or *low water*, when the level of the surface ceases to rise or fall. These are totally distinct phænomena, and depend on entirely different causes, though it is true they may sometimes coincide in point of time. They are, it is feared, too often mistaken one for the other by practical men; a circumstance which, whenever it occurs, must produce the greatest confusion in any attempt to reduce the system of the tides to distinct and intelligible laws.

(755.) The declination of the sun and moon materially affects the tides at any particular spot. As the vertex of the tide-wave tends to place itself vertically under the luminary which produces it, when this vertical changes its point of incidence on the surface, the tide-wave must tend to shift accordingly, and thus, by monthly and annual periods, must tend to increase and diminish alternately the principal tides. The period of the moon's nodes is thus introduced into this subject; her excursions in declination in one part of that period being 29°, and in another only 17°, on either side the equator.

(756.) Geometry demonstrates that the efficacy of a luminary in raising tides is inversely proportional to the cube of its distance. The sun and moon, however, by reason of the ellipticity of their orbits, are alternately nearer to and farther from the earth than their mean distances. In consequence of this, the efficacy of the sun will fluctuate between the extremes 19 and 21, taking 20 for its mean value, and that of the moon between 43 and 59. Taking into account this cause of difference, the highest spring tide will be to the lowest neap as $59 + 21$ to $43 - 19$, or as 80 to 24, or 10 to 3. Of all the causes of differences in the height of tides however, local situation is the most influential. In some places the tide-wave, rushing up a narrow channel, is suddenly raised to an extraordinary height. At Annapolis, for instance, in the Bay of Fundy, it is said to rise 120 feet. Even at Bristol the difference of high and low water occasionally amounts to 50 feet.

(757.) It is by means of the perturbations of the planets, as ascertained by observation and compared with theory, that we arrive at a knowledge of the masses of those planets which having no satellites, offer no other hold upon them for this purpose. Every planet produces an amount of

n in the motions of every other, proportioned to its mass, and
ee of advantage or *purchase* which its situation in the system
r their movements. The latter is a subject of exact calcula-
ormer is unknown, otherwise than by observation of its effects.
ermination, however, of the masses of the planets by this
ory lends the greatest assistance to observation, by pointing out
ations most favourable for eliciting this knowledge from the
ass of superposed inequalities which affect every observed place
; by pointing out the laws of each inequality in its periodical
cay; and by showing how every particular inequality depends
nitude on the mass producing it. It is thus that the mass of
lf (employed by Laplace in his investigations, and interwoven
planetary tables) has been ascertained, by observations of the
nts produced by it in the motions of the ultra-zodiacal planets,
en insufficiently determined, or rather considerably mistaken,
too much on observations of its satellites, made long ago by
others, with inadequate instrumental means. The same con-
been arrived at, and nearly the same mass obtained, by means
urbations produced by Jupiter on Encke's comet. The error
great importance; the mass of Jupiter being by far the most

heights of spring and neap tides which, we have seen, (art. 752,) depends on the proportional influence of the two luminaries. 2dly, from the phænomenon of nutation, which, being the result of the moon's attraction alone, affords a means of calculating her mass, independent of any knowledge of the sun's. Both methods agree in assigning to our satellite a mass about one seventy-fifth that of the earth.[1]

(760.) Not only, however, has a knowledge of the perturbations produced on other bodies of our system enabled us to estimate the mass of a disturbing body already known to exist, and to produce disturbance. It has done much more, and enabled geometers to satisfy themselves of the existence, and even to indicate the situation of a planet previously unknown, with such precision, as to lead to its immediate discovery on the very first occasion of pointing a telescope to the place indicated. We have already (art. 506,) had occasion to mention in general terms this great discovery; but its importance, and its connexion with the subject before us, calls for a more specific notice of the circumstances attending it. When the regular observation of Uranus, consequent on its discovery in 1781, had afforded some certain knowledge of the elements of its orbit, it became possible to calculate backwards into time past, with a view to ascertain whether certain stars of about the same apparent magnitude, observed by Flamsteed, and since reported as *missing*, might not possibly be this planet. No less than six ancient observations of it as a supposed star were thus found to have been recorded by that astronomer, — one in 1690, one in 1712, and four in 1715. On further inquiry, it was also ascertained to have been observed by Bradley in 1753, by Mayer in 1756, and no less than twelve times by Le Monnier, in 1750, 1764, 1768, 1769, and 1771, all the time without the least suspicion of its planetary nature. The observations, however, so made, being all circumstantially registered, and made with instruments the best that their respective dates admitted, were quite available for correcting the elements of the orbit, which, as will be easily understood, is done with so much the greater precision the larger the arc of the ellipse embraced by the extreme observations employed. It was, therefore, reasonably hoped and expected, that, by making use of the data thus afforded, and duly allowing for the perturbations produced since 1690, by Saturn, Jupiter, and the inferior planets, elliptic elements would be obtained, which, taken in conjunction with those perturbations, would represent not only all the observations up to the time of executing the calculations, but also all future observations, in as satisfactory a manner as those of any of the other planets are actually represented. This expectation, however, proved delusive. M. Bouvard,

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[1] Laplace, Expos. du Syst. du Monde, pp. 285, 300.

most expert and laborious calculators of whom astronomy has
st, and to whose zeal and indefatigable industry we owe the
piter and Saturn in actual use, having undertaken the task of
similar tables for Uranus, found it impossible to reconcile
observations above mentioned with those made from 1781 to
to represent both series by means of the same ellipse and the
of perturbations. He therefore rejected altogether the ancient
grounded his computations solely on the modern, although
ot without serious misgivings as to the grounds of such a pro-
d "leaving it to future time to determine whether the difficulty
ng the two series arose from inaccuracy in the older observa-
ether it depend on some extraneous and unperceived influence
have acted on the planet."

But neither did the tables so calculated continue to represent,
recision, observations subsequently made. The "error of the
er attaining a certain amount, by which the true longitude of
in advance of the computed, and which advance was steadily
from about the year 1795 to 1822, began, about the latter
lly to diminish, till, in 1830–31, the tabular and observed
agreed. But, far from remaining in accordance, the planet,

Thus in Flamsteed's earliest observations in 1690, the dot so marked is placed above the line at 65″·9 above the line, the observed longitude being so much greater than the calculated.

(763.) If, neglecting the individual points, we draw a curve (indicated in the figure by a fine, unbroken line,) through their general course, we shall at once perceive a certain regularity in its undulations. It presents two great elevations above, and one nearly as great intermediate depression below the medial line or abscissa. And it is evident that these undulations would be very much reduced, and the errors in consequence greatly palliated, if each dot were removed in the vertical direction through a distance and in the direction indicated by the corresponding point of the curve A, B, C, D, E, F, G, H, intersecting the abscissa at points 180° distant, and making equal excursions on either side. Thus the point *a* for 1750 being removed upwards or in the direction towards *b* through a distance equal to *c b*, would be brought almost to precise coincidence with the point *c* in the abscissa. Now, this is a clear indication that a very large part of the differences in question are due, not to perturbation, but simply to error in the elements of Uranus, which have been assumed as the basis of calculation. For such excesses and defects of longitude alternating over arcs of 180° are precisely what would arise from error in the excentricity, or in the place of the perihelion, or in both. In ellipses slightly excentric, the true longitude alternately exceeds and falls short of the mean during 180° for each deviation, and the greater the excentricity, the greater these alternate fluctuations to and fro. If then the excentricity of a planet's orbit be assumed erroneously (suppose too great) the observed longitudes will exhibit a less amount of such fluctuation above and below the mean than the computed, and the difference of the two, instead of being, as it ought to be, always *nil*, will be alternately + and — over arcs of 180°. If then a difference be observed following such a law, it may arise from erroneously assumed excentricity, provided always the longitudes at which they agree (supposed to differ by 180°) be coincident with those of the perihelion and aphelion; for in elliptic motion nearly circular, these are the points where the mean and true longitudes agree, so that any fluctuation of the nature observed, if this condition be not satisfied, cannot arise from error of excentricity. Now the longitude of the perihelion of Uranus in the elements employed by Bouvard is (neglecting fractions of a degree) 168°, and of the aphelion 348°. These points, then, in our figure, fall at π and α, respectively, that is to say, nearly half way between A C, C E, E G, &c. It is evident, therefore, that it is not to error or excentricity that the fluctuation in question is mainly due.

Let us now consider the effect of an erroneous assumption of
f the perihelion. Suppose in fig. 2, Plate A, $o x$ to represent
de of a planet, and $x y$ the excess of its true above its mean
due to ellipticity. Then if R be the place of the perihelion,
T, the aphelion in longitude, y will always lie in a certain un-
urve P Q R S T, above[1] P T, between R and T, and below it
and R. Now suppose the place of the perihelion shifted for-
, or the whole curve shifted bodily forward into the situation
en at the same longitude $o x$, the excess of the true above the
itude will be $x y'$ only; in other words, this excess will have
by the quantity $y y'$ below its former amount. Take therefore
g. 3, Pl. A,) $o y = o x$ and $y y'$ always $= y y'$ in *fig.* 2, and
s constructed the curve K L M N O, the ordinate $y y'$ will
resent the effect of the supposed change of perihelion. It is
e excentricity being always supposed small), that this curve
t also of alternate superior and inferior waves of 180° each in
and the points L, N of its intersection with the axis, will
ongitudes corresponding to X, Y, intermediate between the
, q; and S, s of the original curves, that is to say (if these in-
, S s, or R r, to which both are equal, be very small,) very

1712 to 1800, would be annihilated, or at least so far reduced as hardly to exceed the ordinary errors of observation; and as respects the observation of 1690, the still outstanding difference of about 35″ would not be more than might be attributed to a not very careful observation at so early an epoch. Now the assumption of such a new line of longitudes as the correct one, is in effect equivalent to the admission of a slight amount of error in the periodic time and epoch of Uranus; for it is evident that by reckoning from the inclined instead of the horizontal line, we in effect alter all the apparent outstanding errors by an amount proportional to the time before or after the date at which the two lines intersect (viz. about 1789). As to the direction in which this correction should be made, it is obvious by inspection of the course of the dots, that if we reckon from A B, or any line parallel to it, the observed planet on the long run keeps falling more and more behind the calculated one; *i. e.* its assigned mean angular velocity by the tables is too great, and must be diminished, or its periodic time requires to be increased.

(766.) Let this increase of period be made, and in correspondence with that change let the longitudes be reckoned on *a b*, and the residual differences from that line instead of A B, and we shall have then done all that can be done in the way of reducing and palliating these differences, and that, with such success, that up to the year 1804 it might have been safely asserted that positively no ground whatever existed for suspecting any disturbing influence. But with this epoch an action appears to have commenced, and gone on increasing, producing an acceleration of the motion in longitude, in consequence of which Uranus continually gains on its elliptic place, and continued to do so till 1822, when it ceased to gain, and the excess of longitude was at its maximum, after which it began rapidly to lose ground, and has continued to do so up to the present time. It is perfectly clear, then, that in this interval some extraneous cause must have come into action which was not so before, or not in sufficient power to manifest itself by any marked effect, and that *that* cause must have ceased to act, or rather begun to reverse its action, in or about the year 1822, the reverse action being even more energetic than the direct.

(767.) Such is the phænomenon in the simplest form we are *now* able to present it. Of the various hypotheses formed to account for it, during the progress of its developement, none seemed to have any degree of rational probability but that of the existence of an exterior, and hitherto undiscovered, planet, disturbing, according to the received laws of planetary disturbance, the motion of Uranus by its attraction, or rather superposing its disturbance on those produced by Jupiter and Saturn, the only two of the old planets which exercise any sensible disturbing action on

t. Accordingly, this was the explanation which naturally, and
necessity, suggested itself to those conversant with the plane-
rbations who considered the subject with any degree of atten-
e idea, however, of setting out from the observed anomalous
, and employing them as data to ascertain the distance and situ-
he unknown body, or, in other words, to resolve the inverse
f perturbations, "*given the disturbances to find the orbit, and*
that orbit of the *disturbing planet*," appears to have occurred
o mathematicians, Mr. Adams in England and M. Leverrier in
th sufficient distinctness and hopefulness of success to induce them
its solution. Both succeeded, and their solutions, arrived at with
dependence, and by each in entire ignorance of the other's at-
re found to agree in a surprising manner when the nature and
f the problem is considered; the calculations of M. Leverrier
for the heliocentric longitude of the disturbing planet for the
1846, 326° 0′, and those of Mr. Adams (brought to the same
° 19′, differing only 3° 19′; the plane of its orbit deviating very
f it all, from that of the ecliptic.

On the day above mentioned — a day for ever memorable in the
astronomy — Dr. Galle, one of the astronomers of the Royal

determined, the mass, and all the elements of the unknown planet (supposed to revolve in the same plane and the same direction with Uranus), except its major semiaxis. This was assumed in the first instance (in conformity with "Bode's law," art. (505), and certainly at the time with a high *primâ facie* probability,) to be double that of Uranus, or 38·364 radii of the Earth's orbit. Without *some* assumption as to the value of this element, owing to the peculiar form of the analytical expression of the perturbations, the analytical investigation would have presented difficulties apparently insuperable. But besides these, it was also necessary to regard as unknown, or at least as liable to corrections of unknown magnitude of the same order as the perturbations, all the elements of Uranus itself, a circumstance whose necessity will be easily understood, when we consider that the received elements could only be regarded as provisional, and must *certainly* be erroneous, the places from which they were obtained being affected by at least some portions of the very perturbations in question. This consideration, though indispensable, added vastly both to the complication and the labour of the inquiry. The axis (and therefore the mean motion) of the one orbit, then, being known very nearly, and that of the other thus hypothetically assumed, it became practicable to express in terms, partly algebraic, partly numerical, the amount of perturbation at any instant, by the aid of general expressions delivered by Laplace in his "*Mécanique Céleste*" and elsewhere. These, then, together with the corrections due to the altered elements of Uranus itself, being applied to the tabular longitudes, furnished, when compared with those observed, a series of *equations*, in which the elements and mass of Neptune, and the corrections of those of Uranus entered as the *unknown quantities*, and by whose resolution (no slight effort of analytical skill) all their values were at length obtained. The calculations were then repeated, reducing at the same time the value of the assumed distance of the new planet, the discordances between the given and calculated results indicating it to have been assumed too large when the results were found to agree better, and the solutions to be, in fact, more satisfactory. Thus, at length, elements were arrived at for the orbit of the unknown planet, as below.

	Leverrier.	Adams.
Epoch of Elements	Jan. 1, 1847.	Oct. 6, 1846.
Mean longitude in Epoch	318° 47′ 4″	323° 2′
Semiaxis Major	36·1539	37·2474
Excentricity	0·107610	0·120615
Longitude of Perihelion	284° 45′ 8″	299° 11′
Mass (the Sun being 1)	0·00010727	0·00015003

nts of M. Leverrier were obtained from a consideration of the
s up to the year 1845, those of Mr. Adams only as far as
subsequently taking into account, however, those of the five
1845, the latter was led to conclude that the semiaxis ought
much further diminished, and that a mean distance of 33·33
hat of Uranus as 1 : 0·574) would probably satisfy all the obser-
y nearly.[1]

On the actual discovery of the planet, it was, of course, assidu-
ved, and it was soon ascertained that a mean distance, even less
Adams's last presumed value, agreed better with its motion;
ner were elements obtained from direct observation, sufficiently
te to trace back its path in the heavens for a considerable inter-
e, than it was ascertained to have been observed as a star by
n the 8th and 10th of May, 1795, the latter of the two obser-
wever, having been rejected by him as faulty, by reason of its
ent with the former (a consequence of the motion of the
the interval.) From these observations, combined with those
nulated, the elements calculated by Prof. Walker, U. S., result
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Elements · · · Jan. 1, 1847, M. noon, Greenwich.

calculated from M. Leverrier's elements for the day in question, instead of 36·1539 (the mean distance) comes out almost exactly 33; and indeed, if we consider that the excentricity assigned by those elements gives for the perihelion distance 32·2634, the longitude assigned to the perihelion brings the whole arc of the orbit (more than 83°), described in the interval from 1806 to 1847 to lie within 42° one way or the other of the perihelion, and therefore, during the whole of that interval, the hypothetical planet would be moving within limits of distance from the sun, 32·6 and 33·0. The following comparative tables of the relative situations of Uranus, the real and hypothetical planet, will exhibit more clearly than any lengthened statement, the near imitation of the motion of the former by the latter within that interval. The longitudes are heliocentric.[1]

A. D.	Uranus.	Neptune.		Leverrier.		Adams.	
	Long.	Long.	Rad. Vec.	Long.	Rad. Vec.	Long.	Rad. Vec.
1805·0	197°·8	235°·9	30·3	241°·2	33·1	246°·5	34·2
1810·0	220·9	247·0	30·3	251·1	32·8	255·9	33·7
1815·0	243·2	258·0	30·3	261·2	32·5	265·5	33·3
1820·0	264·7	268·8	30·2	271·4	32·4	275·4	33·1
1821·0	269·0	271·0	30·2	273·5	32·3	277·4	33·0
1822·0	273·3	273·2	30·2	275·6	32·3	279·5	33·0
1823·0	277·6	275·3	30·2	277·6	32·3	281·5	32·9
1824·0	281·8	277·4	30·2	279·7	32·3	283·6	32·9
1825·0	285·8	279·6	30·2	281·8	32·3	285·6	32·8
1830·0	306·1	290·5	30·1	292·1	32·3	296·0	32·8
1835·0	326·0	301·4	30·1	302·5	32·4	306·3	32·8
1840·0	345·7	312·2	30·1	312·6	32·6	316·3	32·9
1845·0	365·3	323·1	30·0	322·6	32·9	326·0	33·1
1847·0	373·3	327·6	30·0	326·5	33·1	329·3	33·2

(772.) From this comparison it will be seen that Uranus arrived at its conjunction with Neptune at or immediately before the commencement of 1822, with the calculated planet of Leverrier at the beginning of the following year 1823, and with that of Adams about the end of 1824. Both the theoretical planets, and especially that of M. Leverrier, therefore, during the whole of the above interval of time, so far as the *directions* of their attractive forces on Uranus are concerned, would act nearly on it as the true planet must have done. As regards the intensity of the relative disturbing forces, if we estimate these by the principles of art. (612) at the epochs of conjunction, and for the commencement of 1805 and nd for the respective denominators of the fractions of the sun's n Uranus regarded as unity, which express the *total* disturbing in each case, as below:

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[1] The calculations are carried only to tenths of degrees, as quite sufficient for the object in view.

			1805.	Conjunction.	1845.
n	Pierce's mass	$\frac{1}{19840}$	27540	7508	32390
	Struve's mass	$\frac{1}{14406}$	20244	5519	23810
heoretical Planet, mass		$\frac{1}{9322}$	20837	5193	19935

s here assigned to Neptune are those respectively deduced by e and M. Struve from observations of the satellite discovered ssell made with the large telescopes of Fraunhofer in the obser- Cambridge, U. S. and Pulkova respectively. These it will be iffer very considerably, as might reasonably be expected in the nicrometrical measurements of such difficulty, and it is not pos- esent to say to which the preference ought to be given. As with the mass assigned by M. Struve, an agreement on the satisfactory could not have been looked for within the interval y in question.

Subject then to this uncertainty as to the real mass of Neptune, ical planet of Leverrier must be considered as representing with uch fidelity as could possibly be expected in a research of such of four lobes, having a common point in S, viz., S W *m* X S *a* S *n* S *b* S W, and that expressing the tangential, A Z *c f* B *e d* Y A Z, consisting of four mutually intersecting loops, surrounding and touching the disturbed orbit in four points, A B *c d*. The normal force acts outward over all that part of the orbit, both in conjunction and opposition, corresponding to the portions of the lobes *m*, *n*, exterior to the disturbed orbit, and inwards in every other part. The figure sets in a clear light the great disproportion between the energy of this force near the conjunction, and in any other configuration of the planets; its exceedingly rapid degradation as P approaches the point of neutrality (whose situation is 35° 5′ on either side of the conjunction, an arc described synodically by Uranus in $16^{y}\cdot 72$); and the comparatively short duration and consequent inefficacy to produce any great amount of perturbation, of the more intense part of its inward action in the small portions of the orbit corresponding to the lobes *a*, *b*, in which the line representing the inward force exceeds the radius of the circle. It exhibits, too, with no less distinctness, the gradual developement, and rapid degradation and extinction of the tangential force from its neutral points, *c*, *d*, on either side up to the conjunction, where its action is reversed, being accelerative over the arc *d* A, and retardative over A *c*, each of which arcs has an amplitude of 71° 20′, and is described by Uranus synodically in $34^{y}\cdot 00$. The insignificance of the tangential force in the configurations remote from conjunction throughout the arc *c* B *d* is also obviously expressed by the small comparative developement of the loops *e*, *f*.

(774.) Let us now consider how the action of these forces results in the production of that peculiar character of perturbation which is exhibited in our curve, *fig*. 4, Plate A. It is at once evident that the increase of the longitude from 1800 to 1822, the cessation of that increase in 1822, and its conversion into a decrease during the subsequent interval is in complete accordance with the growth, rapid decay, extinction at conjunction, and subsequent reproduction in a reversed sense of the tangential force: so that we cannot hesitate in attributing the greater part of the perturbation expressed by the swell and subsidence of the curve between the years 1800 and 1845,—all that part, indeed, which is symmetrical on either side of 1822—to the action of the tangential force.

(775.) But it will be asked,—has then the normal force (which, on the plain showing of *fig*. 5, is nearly twice as powerful as the tangential, and which does not reverse its action, like the latter force, at the point of junction, but, on the contrary, is there most energetic,) no influence in producing the observed effects? We answer, very little, within the period to which observation had extended up to 1845. The effect of the tan-

e on the longitude is direct and immediate (art. 660), that of
indirect, consequential, and cumulative with the progress of
34.) The effect of the tangential force on *the mean motion*
through the medium of the change it produces on the axis,
nsient: the reversed action after conjunction (supposing the
lar), exactly destroying all the previous effect, and leaving the
n on the whole unaffected. In the passage through the con-
en, the tangential force produces a sudden and powerful accel-
ceeded by an equally powerful and equally sudden retardation,
, its action is completed, and no trace remains in the subse-
n of the planet that it ever existed, for its action on the peri-
excentricity is in like manner also nullified by its reversal of
But with the normal force the case is far otherwise. Its *im-*
fect on the angular motion is *nil.* It is not till it has acted
h to produce a perceptible change in the distance of the dis-
et from the sun that the angular velocity begins to be sensibly
d it is not till its whole outward action has been exerted (*i. e.*
ole interval from neutral point to neutral point) that its maxi-
in lifting the disturbed planet away from the sun has been
nd the full amount of diminution in angular velocity it is

similar in character, but far inferior in magnitude to those now under discussion.

(776.) Posterity will hardly credit that, with a full knowledge of all the circumstances attending this great discovery — of the calculations of Leverrier and Adams — of the communication of its predicted place to Dr. Galle — and of the new planet being actually found by him in that place, in the remarkable manner above commemorated; not only have doubts been expressed as to the validity of the calculations of those geometers, and the legitimacy of their conclusions, but these doubts have been carried so far as to lead the objectors to attribute the acknowledged fact of a planet previously unknown occupying that precise place in the heavens at that precise time, to sheer accident![1] What share accident may have had in the successful issue of the calculations, we presume the reader, after what has been said, will have little difficulty in satisfying himself. As regards the time when the discovery was made, much has also been attributed to fortunate coincidence. The following considerations will, we apprehend, completely dissipate this idea, if still lingering in the mind of any one at all conversant with the subject. The period of Uranus being 84·0140 years, and that of Neptune 164·6181, their synodic revolution (art. 418.), or the interval between two successive conjunctions, is 171·58 years. The late conjunction having taken place about the beginning of 1822; that next preceding must have happened in 1649, or more than 40 years before the first recorded observation of

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[1] These doubts seem to have originated partly in the great disagreement between the predicted and real elements of Neptune, partly in the near (*possibly* precise) commensurability of the mean motions of Neptune and Uranus. We conceive them however to be founded in a total misconception of the nature of the problem, which was not, from such obviously uncertain indications as the observed discordances could give, to determine as astronomical quantities the axis, excentricity and mass of the disturbing planet; but practically to discover where to look for it: when, if once found, these elements would be far better ascertained. To do this, *any axis, excentricity, perihelion, and mass, however wide of the truth,* which would represent, even roughly the amount, but *with tolerable correctness the direction* of the disturbing force during the very moderate interval when the departures from theory were really considerable, would equally serve their purposes; and with an excentricity, mass, and perihelion disposable, it is obvious that any assumption of the axis between the limits 30 and 38, nay, even with a much wider inferior limit, would serve the purpose. In his attempt to assign an inferior limit to the axis, and in the value so assigned, M. Leverrier, it must be admitted, was not successful. Mr. Adams, on the other hand, influenced by no considerations of the kind which appear to have weighed with his brother geometer, fixed ultimately (as we have seen) on an axis not very egregiously wrong. Still it were to be wished, for the satisfaction of all parties, that some one would undertake the problem *de novo*, employing formulæ not liable to the passage through infinity, which, technically speaking, hampers, or may be supposed to hamper the *continuous* application of the usual perturbational formulæ when cases of commensurability occur.

1690, to say nothing of its discovery as a planet. In 1690,
st have been effectually out of reach of any perturbative influ-
considering, and so it remained during the whole interval from
800. From that time the effect of perturbation began to be-
ble, about 1805 prominent, and in 1820 had nearly reached its
At this epoch an alarm was sounded. The maximum was
d,—the event, so important to astronomy, was still in progress
ement,—when the fact (any thing rather than a striking one)
l, and made matter of complaint. But the time for discussing
ith any prospect of success was not yet come. Every thing
n the precise determination of the epoch of the maximum,
erturbing and perturbed planet were in conjunction, and upon
increase and diminution of the perturbation itself on either
t point. Now it is always difficult to assign the time of the
of a maximum by observations liable to errors bearing a ratio
considerable to the whole quantity observed. Until the lapse
ars from 1822 it would have been impossible to have fixed that
any certainty, and as respects the law of degradation and total
itude over which the sensible perturbations extend, we are
arrived at a period when this can be said to be completely de-

PART III.

OF SIDEREAL ASTRONOMY.

CHAPTER XV.

OF THE FIXED STARS. — THEIR CLASSIFICATION BY MAGNITUDES. — PHOTOMETRIC SCALE OF MAGNITUDES. — CONVENTIONAL OR VULGAR SCALE. — PHOTOMETRIC COMPARISON OF STARS. — DISTRIBUTION OF STARS OVER THE HEAVENS. — OF THE MILKY WAY OR GALAXY. — ITS SUPPOSED FORM THAT OF A FLAT STRATUM PARTIALLY SUBDIVIDED. — ITS VISIBLE COURSE AMONG THE CONSTELLATIONS. — ITS INTERNAL STRUCTURE. — ITS APPARENTLY INDEFINITE EXTENT IN CERTAIN DIRECTIONS. — OF THE DISTANCE OF THE FIXED STARS. — THEIR ANNUAL PARALLAX. — PARALLACTIC UNIT OF SIDEREAL DISTANCE. — EFFECT OF PARALLAX ANALOGOUS TO THAT OF ABERRATION. — HOW DISTINGUISHED FROM IT. — DETECTION OF PARALLAX BY MERIDIONAL OBSERVATIONS. — HENDERSON'S APPLICATION TO α CENTAURI. — BY DIFFERENTIAL OBSERVATIONS. — DISCOVERIES OF BESSEL AND STRUVE. — LIST OF STARS IN WHICH PARALLAX HAS BEEN DETECTED. — OF THE REAL MAGNITUDES OF THE STARS. — COMPARISON OF THEIR LIGHTS WITH THAT OF THE SUN.

(777.) BESIDES the bodies we have described in the foregoing chapters, the heavens present us with an innumerable multitude of other objects, which are called generally by the name of stars. Though comprehending individuals differing from each other, not merely in brightness, but in many other essential points, they all agree in one attribute,—a high degree of permanence as to apparent relative situation. This has procured them the title of "fixed stars;" an expression which is to be understood in a comparative and not an absolute sense, it being certain that many, and probable that all, are in a state of motion, although too slow to be perceptible unless by means of very delicate observations, continued during a long series of years.

(778.) Astronomers are in the habit of distinguishing the stars into cording to their apparent brightness. These are termed magni- he brightest stars are said to be of the first magnitude; those so far short of the first degree of brightness as to make a arked distinction are classed in the second; and so on down to or seventh, which comprise the smallest stars visible to the , in the clearest and darkest night. Beyond these, however, continue the range of visibility, and magnitudes from the 8th e 16th are familiar to those who are in the practice of using nstruments; nor does there seem the least reason to assign a is progression; every increase in the dimensions and power of ts, which successive improvements in optical science have aving brought into view multitudes innumerable of objects efore; so that, for any thing experience has hitherto taught us, r of the stars may be really infinite, in the only sense in which ign a meaning to the word.

This classification into magnitudes, however, it must be ob- entirely arbitrary. Of a multitude of bright objects, differing intrinsically, both in size and in splendour, and arranged at istances from us, one must of necessity appear the brightest, one v it, and so on. An order of succession (relative, of course, to tion in any conclusions we may draw from numerical statements of the number of individuals, which have been arranged in our artificial classes antecedent to any general or definite *principle* of arrangement. In fact, astronomers have not yet agreed upon any principle by which the magnitudes may be photometrically classed *à priori*, whether for example a scale of brightnesses decreasing in geometrical progression should be adopted, each term being one-half of the preceding, or one-third, or any other ratio, or whether it would not be preferable to adopt a scale decreasing as the squares of the terms of an harmonic progression, *i. e.* according to the series $1, \frac{1}{4}, \frac{1}{9}, \frac{1}{16}, \frac{1}{25}$, &c. The former would be a purely photometric scale, and would have the apparent advantage, that the light of a star of any magnitude would bear a fixed proportion to that of the magnitude next above it, an advantage, however, merely apparent, as it is certain, from many optical facts, that the unaided eye forms very different judgments of the proportions existing between bright lights, and those between feeble ones. The latter scale involves a physical idea, that of supposing the scale of magnitudes to correspond to the appearance of a first magnitude standard star, removed successively to twice, three times, &c., its original or standard distance. Such a scale, which would make the nominal magnitude a sort of index to the *presumable* or average distance, on the supposition of an equality among the real lights of the stars, would facilitate the expression of speculative ideas on the constitution of the sidereal heavens. On the other hand, it would at first sight appear to make too small a difference between the lights in the lower magnitudes. For example, on this principle of nomenclature, the light of a star of the seventh magnitude would be thirty-six 49ths of that of one of the sixth, and of the tenth 81 hundredths of the ninth, while between the first and the second the proportion would be that of four to one. So far, however, from this being really objectionable, it falls in well with the general tenor of the optical facts already alluded to, inasmuch as the eye (in the absence of disturbing causes) does actually discriminate with greater precision between the relative intensities of feeble lights than of bright ones, so that the fraction $\frac{36}{49}$, for instance, expresses quite as great a step downwards (physiologically speaking) from the sixth magnitude, as $\frac{1}{4}$ does from the first. As the choice, therefore, so far as we can see, lies between these two scales, in drawing the lines of demarcation between what may be termed the *photometrical magnitudes* of the stars, we have no hesitation in adopting, and recommending others to adopt, the latter system in preference to the former.

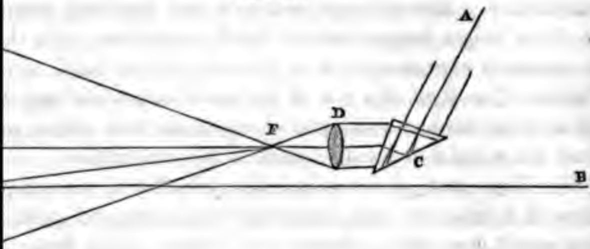
(781.) The conventional magnitudes actually in use among astronomers, so far as their usage is consistent with itself, conforms moreover

more nearly to this than to the geometrical progression. It has
n,[1] by direct photometric measurement of the light of a consi-
mber of stars, from the first to the fourth magnitude, that if M
ber expressing the magnitude of a star on the above system,
number expressing the magnitude of the same star in the loose
lar language at present conventionally or rather provisionally
far as it can be collected from the conflicting authorities of
servers, the difference between these numbers, or $M - m$, is
all the higher parts of the scale, and is less than half a mag-
414). The standard star assumed as the unit of magnitude
surements referred to, is the bright southern star α Centauri, a
hat superior to Arcturus in lustre. If we take the distance of
r unity, it follows that when removed to the distances 1·414,
4, &c., its apparent lustre would equal those of average stars
2d, 3d, &c. magnitudes, *as ordinarily reckoned*, respectively.
The difference of lustre between stars of two consecutive mag-
so considerable as to allow of many intermediate gradations
ctly well distinguished. Hardly any two stars of the first, or
nd magnitude, would be judged by an eye practised in such
s to be exactly equal in brightness. Hence, the necessity, if

α Centauri as $1^2 : (2·924)^2$, making its place in the photometric scale (so defined) 2·924. Lists of stars, northern and southern, comprehending those of the vulgar first, second and third magnitudes, with their magnitudes decimally expressed in both systems, will be found at the end of this work. The light of a star of the sixth magnitude may be roughly stated as about the hundredth part of one of the first. Sirius would make between three and four hundred stars of that magnitude.

(783.) The exact photometrical determination of the comparative intensities of light of the stars is attended with many and great difficulties, arising partly from their differences of colour; partly from the circumstance that no invariable standard of *artificial* light has yet been discovered; partly from the physiological cause above alluded to, by which the eye is incapacitated from judging correctly of the proportion of two lights, and can only decide (and that with not very great precision) as to their equality or inequality; and partly from other physiological causes. The least objectionable mode hitherto proposed would appear to be the following. A natural standard of comparison is in the first instance selected, brighter than any of the stars, so as to allow of being equalized with any of them by a *reduction* of its light optically effected, and at the same time either invariable, or at least only *so* variable that its changes can be exactly calculated and reduced to numerical estimation. Such a standard is offered by the planet Jupiter, which, being much brighter than any star, subject to no phases, and variable in light only by the variation of its distance from the sun, and which, moreover, comes in succession above the horizon at a convenient altitude, simultaneously with all the fixed stars, and, in the absence of the moon, twilight, and other disturbing causes (which fatally affect all observations of this nature), combines all the requisite conditions. Let us suppose, now, that Jupiter being at A and the star to be compared with it at B, a glass prism, C, is so placed that the light of the planet deflected by *total internal reflexion* at its base, shall emerge parallel to B E, the direction of the star's visual ray. After reflexion let it be received on a lens, D, in whose focus, F, it will form a small, bright, star-like image, capable of being viewed by an eye placed at E, so far out of the axis of the cone of diverging rays as to admit of seeing at the same time, and with the same eye, and so comparing this image with the star seen directly. By bringing the eye nearer to or further from the focus, F, the apparent brightness of the focal point will be varied in the inverse ratio of the square of the distance, E F, and therefore may be *equalized*, as well as the eye can judge of such equalities, with the star. If this be done for two stars several times alternately, and a mean of the results taken, by measuring E F, their relative bright-

Fig. 107.



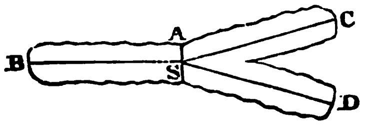
be obtained: that of Jupiter, the temporary standard of com-
eing altogether eliminated from the result.

A moderate number of well selected stars being thus photometri-
rmined by repeated and careful measurements, so as to afford an
d and graduated scale of brightness among the stars themselves,
ediate steps or grades of magnitude may be filled up, by insert-
en them, according to the judgment of the eye, other stars,
n ascending or descending *sequence*, each member of such a
being brighter than that below, and less bright than that

impartiality over the sphere: a marked preference however being observable, especially in the southern hemisphere, to a zone or belt, following the direction of a great circle passing through ϵ Orionis and α Crucis. But if we take in the whole amount visible to the naked eye, we shall perceive a great increase of number as we approach the borders of the Milky Way. And when we come to telescopic magnitudes, we find them crowded beyond imagination, along the extent of that circle, and of the branches which it sends off from it; so that in fact its whole light is composed of nothing but stars of every magnitude, from such as are visible to the naked eye down to the smallest point of light perceptible with the best telescopes.

(786.) These phænomena agree with the supposition that the stars of our firmament, instead of being scattered in all directions indifferently through space, form a stratum of which the thickness is small, in comparison with its length and breadth; and in which the earth occupies a place somewhere about the middle of its thickness, and near the point where it subdivides into two principal laminæ, inclined at a small angle to each other (art. 802). For it is certain that, to an eye so situated, the apparent density of the stars, supposing them pretty equally scattered through the space they occupy, would be least in a direction of the visual ray (as S A), perpendicular to the lamina, and greatest in that of its breadth, as S B, S C, S D; increasing rapidly in passing from one to the other direction, just as we see a slight haze in the atmosphere thickening into a decided fog-bank near the horizon, by the rapid increase of the mere length of the visual ray. Such is the view of the construction of the starry firmament taken by Sir William Herschel, whose powerful telescopes first effected a complete analysis of this wonderful zone, and demonstrated the fact of its entirely consisting of stars.[1] So crowded are they in some parts of it, that by counting the stars in a single field of his

Fig. 108.



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[1] Thomas Wright of Durham (Theory of the Universe, London, 1750) appears so early as 1734 to have entertained the same general view as to the constitution of the Milky Way and starry firmament, founded, quite in the spirit of just astronomical speculation, on a partial resolution of a portion of it with a "one-foot reflector" (a reflector one foot in focal length). See an account of this rare work by M. de Morgan in Phil. Mag. Ser. 3. xxxii. p. 241. et seq.

ie was led to conclude that 50,000 had passed under his review
wo degrees in breadth, during a single hour's observation. In
f the milky way which is situated in 10*h* 30*m* R A and between
and 150th degree of N P D, upwards of 5000 stars have been
o exist in a square degree. The immense distances at which
r regions must be situated will sufficiently account for the vast
nce of small magnitudes which are observed in it.

The course of the Milky Way as traced through the heavens
ided eye, neglecting occasional deviations and following the line
test brightness as well as its varying breadth and intensity will
nforms nearly to that of a great circle inclined at an angle of
to the equinoctial, and cutting that circle in R A 6*h* 47*m* and
so that its northern and southern poles respectively are situated
2*h* 47*m* N P D 63° and R. A. 0*h* 47*m* N P D 117°. Through-
gion where it is so remarkably subdivided (art. 186), this great
s an intermediate situation between the two great streams; with
pproximation however to the brighter and continuous stream,
 fainter and interrupted one. If we trace its course in order
cension, we find it traversing the constellation Cassiopeia, its
art passing about two degrees to the north of the star δ of that

Puppis), sending off a narrow and winding branch on the preceding side as far as γ Argûs, where it terminates abruptly. The main stream pursues its southward course to the 123d parallel of N P D, where it diffuses itself broadly and again subdivides, opening out into a wide fan-like expanse nearly 20° in breadth formed of interlacing branches, all which terminate abruptly, in a line drawn nearly through λ and γ Argûs. #

(788.) At this place the continuity of the Milky Way is interrupted by a wide gap, and where it recommences on the opposite side it is by a somewhat similar fan-shaped assemblage of branches which converge upon the bright star η Argûs. Thence it crosses the hind feet of the Centaur, forming a curious and sharply defined semicircular concavity of small radius, and enters the Cross by a very bright neck or isthmus of not more than 3 or 4 degrees in breadth, being the narrowest portion of the Milky Way. After this it immediately expands into a broad and bright mass, enclosing the stars α and β Crucis, and β Centauri, and extending almost up to α of the latter constellation. In the midst of this bright mass, surrounded by it on all sides, and occupying about half its breadth, occurs a singular dark pear-shaped vacancy, so conspicuous and remarkable as to attract the notice of the most superficial gazer, and to have acquired among the early southern navigators the uncouth but expressive appellation of the *coal-sack*. In this vacancy which is about 8° in length, and 5° broad, only one very small star visible to the naked eye occurs, though it is far from devoid of telescopic stars, so that its striking blackness is simply due to the effect of contrast with the brilliant ground with which it is on all sides surrounded. This is the place of nearest approach of the Milky Way to the South Pole. Throughout all this region its brightness is very striking, and when compared with that of its more northern course already traced, conveys strongly the impression of greater proximity, and would almost lead to a belief that our situation as spectators is separated on all sides by a considerable interval from the dense body of stars composing the Galaxy, which in this view of the subject would come to be considered as a flat ring of immense and irregular breadth and thickness, within which we are excentrically situated, nearer to the southern than to the northern part of its circuit.

(789.) At α Centauri, the Milky Way again subdivides[1], sending off a great branch of nearly half its breadth, but which thins off rapidly, at an angle of about 20° with its general direction, towards the preceding side, to η and *d* Lupi, beyond which it loses itself in a narrow and faint streamlet. The main stream passes on increasing in breadth to γ Normæ, where it makes an abrupt elbow and again subdivides into one principal and con

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[1] All the maps and globes place this subdivision at β Centauri, but erroneously.

eam of very irregular breadth and brightness on the following
complicated system of interlaced streaks and masses on the
which covers the tail of Scorpio, and terminates in a vast and
on over the whole extensive region occupied by the preceding
iuchus, extending northwards to the parallel of 103° N P D,
ich it cannot be traced; a wide interval of 14°, free from all
of nebulous light, separating it from the great branch on the
of the equinoctial of which it is usually represented as a con-

Returning to the point of separation of this great branch from
ream, let us now pursue the course of the latter. Making an
d to the following side, it passes over the stars ι Aræ, θ and ι
l γ Tubi to γ Sagittarii, where it suddenly collects into a vivid
about 6° in length and 4° in breadth, so excessively rich in
very moderate calculation makes their number exceed 100,000.
of this mass, this stream crosses the ecliptic in longitude about
proceeding along the bow of Sagittarius into Antinous has its
led by three deep concavities, separated from each other by
protuberances, of which the larger and brighter (situated
lamstead's stars 3 and 6 Aquilæ) forms the most conspicuous

(792.) We have been somewhat circumstantial in describing the course and principal features of the Via Lactea, not only because there does not occur anywhere (so far as we know) any correct account of it, but chiefly by reason of its high interest in sidereal astronomy, and that the reader may perceive how very difficult it must necessarily be to form any just conception of the real, solid form, as it exists in space, of an object so complicated, and which we see from a point of view so unfavourable. The difficulty is of the same kind which we experience when we set ourselves to conceive the real shape of an auroral arch or of the clouds, but far greater in degree, because we know the laws which regulate the formation of the latter, and limit them to certain conditions of altitude — because their motion presents them to us in various aspects, but chiefly because we contemplate them from a station considerably below their general plane, so as to allow of our mapping out some kind of ground-plan of their shape. All these aids are wanting when we attempt to map and model out the Galaxy, and beyond the obvious conclusion that its form must be, generally speaking, *flat*, and of a thickness small in comparison with its area in length and breadth, the laws of perspective afford us little further assistance in the inquiry. Probability may, it is true, here and there enlighten us as to certain features. Thus when we see, as in the coal-sack, a sharply defined oval space free from stars, insulated in the midst of a uniform band of not much more than twice its breadth, it would seem much less probable that a conical or tubular hollow traverses the whole of a starry stratum, continuously extended from the eye outwards, than that a *distant* mass of comparatively moderate thickness should be simply perforated from side to side, or that an oval vacuity should be seen foreshortened in a *distant* foreshortened area, not really exceeding two or three times its own breadth. Neither can we without obvious improbability refuse to admit that the long lateral offsets which at so many places quit the main stream and run out to great distances, are either planes seen edgeways, or the convexities of curved surfaces viewed tangentially, rather than cylindrical or columnar excrescences bristling up obliquely from the general level. And in the same spirit of probable surmise we may account for the intricate reticulations above described as existing in the region of Scorpio, rather by the accidental crossing of streaks thus originating, at very different distances, or by a cellular structure of the mass, than by real intersections. Those cirrous clouds which are often seen in windy weather, convey no unapt impression either of the kind of appearance in question, or of the structure it suggests. It is to other indications however, and chiefly to the telescopic examination of its intimate constitution, and to the law of the distribution of stars, not only

bosom, but generally over the heavens, that we must look for
ite knowledge respecting its true form and extent.

It is on observations of this latter class, and not on merely
or conjectural views, that the generalization in Art. 786, which
phænomena of the starry firmament to the system of the Ga-
eir embodying fact, is brought to depend. The process of
" the heavens was devised by Sir W. Herschel for this purpose.
d in simply counting the stars of all magnitudes which occur in
ds of view, of 15′ in diameter, visible through a reflecting tele-
8 inches aperture, and 20 feet focal length, with a magnifying
180° : the points of observation being very numerous and taken
ately in every part of the surface of the sphere visible in our
On a comparison of many hundred such "gauges" or local
ns it appears that the density of star-light (or the number of
ing on an average of several such enumerations in any one im-
ighbourhood) is least in the pole of the *Galactic circle,*[1] and
n all sides, with the *Galactic polar distance* (and that nearly
all directions) down to the Milky Way itself, where it attains
m. The progressive rate of increase in proceeding from the
first slow, but becomes more and more rapid as we approach the

(794.) These numbers fully bear out the statement in Art. 786, and even draw closer the resemblance by which that statement is there illustrated. For the rapidly increasing density of a fog-bank as the visual ray is depressed towards the plane of the horizon is a consequence not only of the mere increase in length of the foggy space traversed, but also of an actual increase of density in the fog itself in its lower strata. Now this very conclusion follows from a comparison *inter se* of the numbers above set down, as M. Struve has clearly shown from a mathematical analysis of the empirical formula, which faithfully represents their law of progression, and of which he states the result in the following table, expressing the densities of the stars at the respective distances, 1, 2, 3, &c., from the galactic plane, taking the mean density of the stars in that plane itself for unity.

Distances from the Galactic Plane.	Density of Stars.	Distances from the Galactic Plane.	Density of Stars.
0·00	1·00000	0·50	0·08646
0·05	0·48568	0·60	0·05510
0·10	0·33288	0·70	0·03079
0·20	0·23895	0·80	0·01414
0·30	0·17980	0·866	0·00532
0·40	0·13021		

The unit of distance, of which the first column of this table expresses fractional parts, is the distance at which such a telescope is capable of rendering just *visible* a star of average magnitude, or, as it is termed, its *space-penetrating* power. As we ascend therefore from the galactic plane into this kind of stellar atmosphere, we perceive that the density of its parallel strata decreases with great rapidity. At an altitude above that plane equal to only one-twentieth of the telescopic limit, it has already diminished to one-half, and at an altitude of 0·866, to hardly more than one-two-hundredth of its amount in that plane. So far as we can perceive there is no flaw in this reasoning, if only it be granted, 1st, that the level planes are continuous, and of equal density throughout; and, 2dly, *that an absolute and definite limit is set to telescopic vision, beyond which, if stars exist, they elude our sight, and are to us as if they existed not*: a postulate whose probability the reader will be in a better condition to estimate, when in possession of some other particulars respecting the constitution of the Galaxy to be described presently.

(795.) A similar course of observation followed out in the southern hemisphere, leads independently to the same conclusion as to the law of the visible distribution of stars over the southern galactic hemisphere, or that half of the celestial surface which has the south galactic pole for its centre. A system of gauges, extending over the whole surface of that taken with the same telescope, field of view and magnifying loyed in Sir William Herschel's gauges, has afforded the aveers of stars *per* field of 15′ in diameter, within the areas of cling that pole at intervals of 15°, set down in the following

es of Galactic South Polar Distance.	Average Number of Stars per Field of 15′.
0° to 15°	6·05
15 to 30	6·62
30 to 45	9·08
45 to 60	13·49
60 to 75	26·29
75 to 90	59·06

These numbers are not directly comparable with those of M. en in Art. 793, because the latter corresponds to the limiting nces, while these are the averages for the included zones. That tronomer, however, has given a table of the average gauges ap- *each degree* of north galactic polar distance,[1] from which it is culate averages for the whole extent of each zone. How near he results of this calculation for the northern hemisphere ex-

exhibiting a rapid succession of closely clustering rich patches separated by comparatively poor intervals, and indeed in some instances by spaces absolutely dark *and completely void of any star*, even of the smallest telescopic magnitude. In some places not more than 40 or 50 stars on an average occur in a "gauge" field of 15′, while in others a similar average gives a result of 400 or 500. Nor is less variety observable in the character of its different regions in respect of the magnitudes of the stars they exhibit, and the proportional numbers of the larger and smaller magnitudes associated together, than in respect of their aggregate numbers. In some, for instance, extremely minute stars, though never altogether wanting, occur in numbers so moderate as to lead us irresistibly to the conclusion that in these regions we see *fairly through* the starry stratum, since it is impossible otherwise (supposing their light not intercepted) that the numbers of the smaller magnitudes should not go on continually increasing ad infinitum. In such cases moreover the ground of the heavens, as seen between the stars, is for the most part perfectly dark, which again would not be the case, if innumerable multitudes of stars, too minute to be individually discernible, existed beyond. In other regions we are presented with the phenomenon of an almost uniform degree of brightness of the individual stars, accompanied with a very even distribution of them over the ground of the heavens, both the larger and smaller magnitudes being strikingly deficient. In such cases it is equally impossible not to perceive that we are looking *through* a sheet of stars nearly of a size, and of no great thickness compared with the distance which separates them from us. Were it otherwise we should be driven to suppose the more distant stars uniformly the larger, so as to compensate by their greater intrinsic brightness for their greater distance, a supposition contrary to all probability. In others again, and that not unfrequently, we are presented with a double phænomenon of the same kind, viz. a tissue as it were of large stars spread over another of very small ones, the intermediate magnitudes being wanting. The conclusion here seems equally evident that in such cases we look through two sidereal sheets separated by a starless interval.

(798.) Throughout by far the larger portion of the extent of the Milky Way in both hemispheres, the general blackness of the ground of the heavens on which its stars are projected, and the absence of that innumerable multitude and excessive crowding of the smallest visible magnitudes, and of glare produced by the aggregate light of multitudes too small to affect the eye singly, which the contrary supposition would appear to necessitate, must, we think, be considered unequivocal indications that its dimensions *in directions where these conditions obtain*, are not only not

t that the space-penetrating power of our telescopes suffices
ierce through and beyond it. It is but right, however, to warn
s that this conclusion has been controverted, and that by an
ot lightly to be put aside, on the ground of certain views taken
as to a defect of perfect transparency in the celestial spaces, in
hich the light of the more distant stars is enfeebled more than
on to their distance. The extinction of light thus originating,
in geometrical progression while the distance increases in
l, a limit, it is argued, is placed to the space-penetrating powers
es, far within that which distance alone apart from such obscu-
ld assign. It would lead us too far aside of the objects of a
this nature to enter upon any discussion of the grounds (partly
al) on which these views rely. It must suffice here to observe
jection alluded to, if applicable to any, is equally so to every
galaxy. We are not at liberty to argue that at one part of its
ce, our view is limited by this sort of cosmical veil which
s the smaller magnitudes, cuts off the nebulous light of distant
l closes our view in impenetrable darkness; while at another
pelled by the clearest evidence telescopes can afford to believe

correct that we look out beyond them into space, the smallest visible stars appear as such, not by reason of excessive distance, but of a real inferiority of size or brightness.

(799.) When we speak of the comparative remoteness of certain regions of the starry heavens beyond others, and of our own situation in them, the question immediately arises, what is the distance of the nearest fixed star? What is the scale on which our visible firmament is constructed? And what proportion do its dimensions bear to those of our own immediate system? To these questions astronomy has at length been enabled to afford an answer.

(800.) The diameter of the earth has served us for the base of a triangle, in the *trigonometrical survey* of our system (art. 274,) by which to calculate the distance of the sun; but the extreme minuteness of the sun's parallax (art. 357,) renders the calculation from this "ill-conditioned" triangle (art. 275,) so delicate, that nothing but the fortunate combination of favourable circumstances, afforded by the transits of Venus (art. 479,) could render its results even tolerably worthy of reliance. But the earth's diameter is too small a base for direct triangulation to the verge even of our own system (art. 526,) and we are, therefore obliged, to substitute the *annual parallax* for the diurnal, or, which comes to the same thing, to ground our calculation on the relative velocities of the earth and planets in their orbits (art. 486,) when we would push our triangulation to that extent. It might be naturally enough expected, that by this enlargement of our base to the vast diameter of the earth's orbit, the next step in our survey (art. 275,) would be made at a great advantage;—that our change of station, from side to side of it, would produce a considerable and easily measurable amount of annual parallax in the stars, and that by its means we should come to a knowledge of their distance. But, after exhausting every refinement of observation, astronomers were, up to a very late period, unable to come to any positive and coincident conclusion upon this head; and the amount of such parallax, even for the nearest fixed star examined with the requisite attention, remained mixed up with, and concealed among, the errors incidental to all astronomical determinations. The nature of these errors has been explained in the earlier part of this work, and we need not remind the reader of the difficulties which must necessarily attend the attempt to disentangle an element not exceeding a few tenths of a second, or at most a whole second, from the host of uncertainties entailed on the results of observations by them: none of them individually perhaps of greater magnitude, but embarrassing by their number and fluctuating amount. Nevertheless, by successive refinements in instrument making, and by constantly progressive approximation to the

ledge of the Uranographical corrections, that assurance had
ned, even in the earlier years of the present century, viz. that
ble in northern latitudes, to which attention had been directed,
an amount of parallax exceeding a single second of arc. It is
e to pause for a moment to consider what conclusions would
the admission of a parallax to this amount.

Radius is to the sine of 1″ as 206265 to 1. In this proportion
st must the distance of the fixed stars from the sun exceed that
from the earth. Again, the latter distance, as we have already
57,) exceeds the earth's radius in the proportion of 23984 to
therefore the earth's radius for unity, a parallax of 1″ sup-
ance of 4947059760 or nearly five thousand millions of such
lastly, to descend to ordinary standards, since the earth's radius
ken at 4000 of our miles, we find 19788239040000 or about
ions of miles for our resulting distance.

n such numbers the imagination is lost. The only mode we
nceiving such intervals at all is by the time which it would
light to traverse them. Light, as we know (art. 545,) travels
of 192000 miles per second, traversing a semidiameter of the
it in $8^m\ 13^s{\cdot}3$. It would, therefore, occupy 206265 times this

viewed through the gauging telescope, the same impression as a star of the sixth magnitude does to the naked eye. Among the infinite multitude of such stars in the remoter regions of the galaxy, it is but fair to conclude that innumerable individuals equal in intrinsic brightness to those which immediately surround us, must exist. The light of such stars, then, must have occupied upwards of 2000 years in travelling over the distance which separates them from our own system. It follows, then, that when we observe the places and note the appearance of such stars, we are only reading their history of two thousand years' anterior date, thus wonderfully recorded. We cannot escape this conclusion, but by adopting as an alternative an intrinsic inferiority of light in *all* the smaller stars of the galaxy. We shall be better able to estimate the probability of this alternative when we shall have made acquaintance with other sidereal systems, whose existence the telescope discloses to us, and whose analogy will satisfy us that the view of the subject here taken is in perfect harmony with the general tenor of astronomical facts.

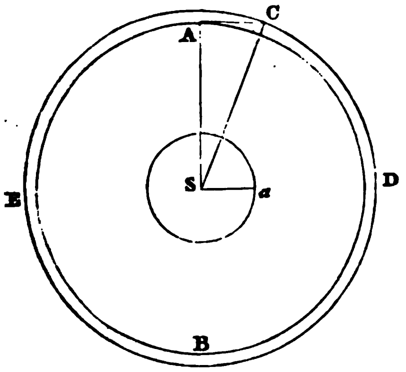
(804.) Hitherto we have spoken of a parallax of 1″ as a mere limit, below which that of any star yet examined, assuredly, or at least very probably, falls, and it is not without a certain convenience to regard this amount of parallax as a sort of unit of reference, which, connected in the reader's recollection with a parallactic unit of distance from our system of 20 billions of miles, and with a 3¼ years' journey of light, may save him the trouble of such calculations, and ourselves the necessity of covering our pages with such enormous numbers, when speaking of stars whose parallax has actually been ascertained with some approach to certainty, either by direct meridian observation, or by more refined and delicate methods. These we shall proceed to explain, after first pointing out the theoretical peculiarities which enable us to separate and disentangle its effects from those of the Uranographical corrections, and from other causes of error, which, being periodical in their nature, add greatly to the difficulty of the subject. The effects of precession and proper motion (see art. 852,) which are uniformly progressive from year to year, and that of nutation, which runs through its period in nineteen years, it is obvious enough, separate themselves at once by these characters from that of parallax; and, being known with very great precision, and being certainly independent, as regards their causes, of any individual peculiarity in the stars affected by them, whatever small uncertainty may remain respecting the numerical elements which enter into their computation (or, in mathematical language, their *co-efficients*), can give rise to no embarrassment. With regard to aberration the case is materially different. This correction affects the place of a star by a fluctuation, annual in its period, and

o far agreeing with parallax. It is also very similar in the
variation at different seasons of the year, parallax having for
e art. 343, 344,) the apparent place of the sun in the ecliptic,
ion a point in the same great circle 90° behind that place, so
the formulæ of calculation (the coefficients excepted) are the
oth, substituting only for the sun's longitude in the expression
, that longitude diminished by 90° for the other. Moreover, in
of *absolute certainty* respecting the nature of the propagation
tronomers have hitherto considered it necessary to assume, at
possibility, that the velocity of light may be to some slight
pendent on individual peculiarities in the body emitting it.[1]

If we suppose a line drawn from the star to the earth at all
the year, it is evident that this line will sweep over the surface
edingly acute, oblique cone, having for its axis the line joining
d star, and for its base the earth's annual orbit, which, for the
rpose, we may suppose circular. The star will therefore appear
each year about its mean place, regarded as fixed, and in virtue
alone, a minute ellipse, the section of this cone by the surface
stial sphere, perpendicular to the visual ray. But there is also
y in which the same fact may be represented. The apparent

virtue of both simultaneously be found in C, *i. e.* in the circumference of a circle whose radius is S C, and at a point in that circle in advance of A, the aberrational place, by the angle A S C. Now, since S A : A C :: 20·5 : 1, we find for the angle A S C 2° 47′ 35″, and for the length of

Fig. 109.



the radius S C of the circle representing the compound motion 20″·524. The difference (0″·024) between this and S C, the radius of the aberration circle, is quite imperceptible, and even supposing a quantity so minute to be capable of detection by a prolonged series of observations, it would remain a question whether it were produced by parallax or by a specific difference of aberration from the general average 20″·5 in the star itself. It is therefore to the difference of 2° 48′ between the angular situation of the displaced star in this hypothetical orbit, *i. e.* in the *arguments* (as they are called) of the joint correction (♈ S C) and that of aberration alone (♈ S A), that we have to look for the resolution of the problem of parallax. The reader may easily figure to himself the delicacy of an inquiry which turns wholly (even when stripped of all its other difficulties) on the *precise* determination of a quantity of this nature, and of such very moderate magnitude.

(806.) But these other difficulties themselves are of no trifling order. All astronomical instruments are affected by differences of temperature. Not only do the materials of which they are composed expand and contract, but the masonry and solid piers on which they are erected, nay even the very soil on which these are founded, participate in the general change from summer warmth to winter cold. Hence arise slow oscillatory movements of exceedingly minute amount, which levels and plumb-lines afford

adequate means of detecting, and which *being also annual in*
(after rejecting whatever is merely casual and momentary)
lves intimately with the matter of our inquiry. Refraction
its casual variations from night to night, which a long series
ions would eliminate, depends for its theoretical expression on
tion of the strata of our atmosphere, and the law of the dis-
heat and moisture at different elevations, which cannot be
by difference of season. No wonder then that mere meridional
s should, almost up to the present time, have proved insufficient,
ne very remarkable instance, to afford unquestionable evidence,
tory quantitative measurement of the parallel of any fixed

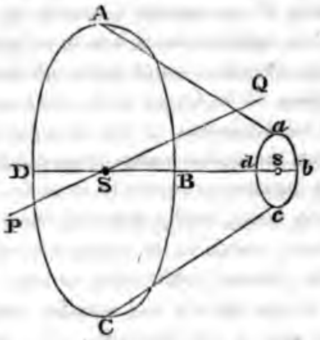
he instance referred to is that of α Centauri, one of the brightest
ny other reasons, one of the most remarkable of the southern
m a series of observations of this star, made at the Royal
y of the Cape of Good Hope in the years 1832 and 1833, by
enderson, with the mural circle of that establishment, a paral-
mount of an entire second was concluded on his reduction of
tions in question after his return to England. Subsequent
s by Mr. Maclear, partly with the same, and partly with a new

stance of its being affected by a *proper motion* (see art. 852) *i. e.* a regular and continually progressive annual displacement among the surrounding stars to the extent of more than 5″ per annum, a quantity so very much exceeding the average of similar minute annual displacements which many other stars exhibit, as to lead to a suspicion of its being actually nearer to our system. It is not a little remarkable that a similar presumption of proximity exists also in the case of α Centauri, whose unusually large proper motion of nearly 4″ per annum is stated by Professor Henderson to have been the motive which induced him to subject his observations of that star to that severe discussion which led to the detection of its parallax. M. Bessel's observations of 61 Cygni were commenced in August 1837, immediately on the establishment at the Königsberg observatory of a magnificent heliometer, the workmanship of the celebrated optician Fraunhofer, of Munich, an instrument especially fitted for the system of observation adopted; which being totally different from that of direct meridional observation, more refined in its conception, and susceptible of far greater accuracy in its practical application, we must now explain.

(809.) Parallax, proper motion, and specific aberration (denoting by the latter phrase that part of the aberration of a star's light which may be supposed to arise from its individual peculiarities, and which we have every reason to believe at all events an exceedingly minute fraction of the whole,) are the only uranographical corrections which do not necessarily affect alike the apparent places of two stars situated in, or *very nearly in*, the same visual line. Supposing then two stars at an immense distance, the one behind the other, but otherwise so situated as to appear very nearly along the same visual line, they will constitute what is called a star *optically double*, to distinguish it from a star *physically double*, of which more hereafter. Aberration (that which is common to all stars), precession, nutation, *nay, even refraction, and instrumental causes of apparent displacement, will affect them alike*, or so very nearly alike (if the minute difference of their apparent places be taken into account) as to admit of the difference being neglected, or very accurately allowed for, by an easy calculation. If then, instead of attempting to determine by observation the place of the nearer of two *very unequal* stars (which will probably be the larger) by direct observation of its right ascension and polar distance, we content ourselves with referring its place to that of its remoter and smaller companion by *differential observation*, *i. e.* by measuring only its *difference* of situation from the *latter*, we are at once relieved of the necessity of making these corrections, and from all uncertainty as to their influence on the result. And for the very same reason, errors of adjustment (art. 186), of graduation, and a host of instrumental errors, which his delicate purpose fatally affect the absolute determination of s place, are harmless when only the difference of their places, ly affected by such causes, is required to be known.

Throwing aside therefore the consideration of all these errors tions, and disregarding for the present the minute effect of

Fig. 110.



erration and the uniformly progressive effect of proper motion, exactly both the distance between two objects which can be seen together in the same field of a telescope, and the position of the line joining them with respect to the horizon, or the meridian, or any other determinate direction in the heavens. The double image micrometer, and especially the heliometer (art. 200, 201) is peculiarly adapted for this purpose. The images of the two stars formed side by side, or in the same line prolonged, however momentarily displaced by temporary refraction or instrumental tremor, *move together*, preserving their relative situation, the judgment of which is no way disturbed by such irregular movements. The heliometer also, taking in a greater range than ordinary micrometers, enables us to compare one large star with more than one adjacent small one, and to select such of the latter among many near it, as shall be most favourably situated for the detection of any motion in the large one, not participated in by its neighbours.

(812.) The star examined by Bessel has two such neighbours, both very minute, and therefore probably very distant, most favourably situated, the one (*s*) at a distance of 7′ 42″, the other (*s*′) at 11′ 46″ from the large star, and so situated, that their directions from that star make nearly a right angle with each other. The effect of parallax therefore would necessarily cause the two distances S *s* and S *s*′ to vary so as to attain their maximum and minimum values alternately at three-monthly intervals, and this is what was actually observed to take place, the one distance being always on the increase or decrease when the other was stationary (the uniform effect of proper motion being understood of course to be always duly accounted for). This alternation, though so small in amount as to indicate, as a final result, a parallax, or rather a difference of parallaxes between the large and small stars of hardly more than one third of a second, was maintained with such regularity as to leave no room for reasonable doubt as to its cause, and having been confirmed by the further continuance of these observations, and quite recently by the exact coincidence between the result thus obtained, and that deduced by M. Peters from observations of the same star at the observatory of Pulkova[1], is considered on all hands as fully established. The parallax of this star finally resulting from Bessel's observation is 0″·348, so that its distance from our system is very nearly three parallactic units. (Art. 804.)

(813.) The bright star α Lyræ has also near it, at only 43″ distance (and therefore within the reach of the parallel wire or ordinary double image micrometer) a very minute star, which has been subjected since 1835 to a severe and assiduous scrutiny by M. Struve, on the same principle of differential observation. He has thus established the existence

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[1] With the great vertical circle by Ertel.

ırable amount of parallax in the large star, less indeed than
Cygni (being only about ¼ of a second), but yet sufficient
the delicacy of his measurements) to justify this excellent
announcing the result as at least highly probable, on the
only five nights' observation, in 1835 and 1836. This pro-
e continuation of the measures to the end of 1838 and the
e, though not in this case precisely coincident, result of Mr.
vestigations have converted into a certainty. M. Struve has
f being the first to bring into practical application this method
ion, which, though proposed for the purpose, and its great ad-
inted out by Sir William Herschel so early as 1781[1], remained
ductive of any result, owing partly to the imperfection of
s for the measurement of distance, and partly to a reason
hall presently have occasion to refer to.

f the component individuals S, *s* (*fig.* art. 810,) be (as is often
ery close to each other, the parallactic variation of their *angle*
, or the extreme angle included between the lines A *a*, C *c*,
y considerable, even for a small amount of difference of paral-
en the large and small stars. For instance in the case of two
rs 15″ asunder, and otherwise favourably situated for observa-

Although the extreme minuteness of the last four of these results deprives them of much numerical reliance, it is at least certain that the parallaxes by no means follow the order of magnitudes, and this is farther shown by the fact that α Cygni, one of M. Peters's stars, shows absolutely no indications of any measurable parallax whatever.

(816.) From the distance of the stars we are naturally led to the consideration of their real magnitudes. But here a difficulty arises, which, so far as we can judge of what optical instruments are capable of effecting, must always remain insuperable. Telescopes afford us only negative information as to the apparent angular diameter of any star. The round, well-defined, planetary discs which good telescopes show when turned upon any of the brighter stars are phænomena of diffraction, dependent, though at present somewhat enigmatically, on the mutual interference of the rays of light. They are consequently, so far as this inquiry is concerned, mere optical illusions, and have therefore been termed *spurious* discs. The proof of this is that telescopes of different apertures and magnifying powers, when applied for the purpose of measuring their angular diameters, give different results, the greater aperture (even with the same magnifying power) giving the smaller disc. That the true disc of even a large and bright star can have but a very minute angular measure, appears from the fact that in the occultation of such a star by the moon, its extinction is *absolutely instantaneous*, not the smallest trace of gradual diminution of light being perceptible. The apparent or spurious disc also remains *perfectly round* and *of its full size* up to the instant of disappearance, which could not be the case were it a real object. If our sun were removed to the distance expressed by our parallactic unit (art. 804), its apparent diameter of 32′ 3″ would be reduced to only 0″·0093, or less than the hundredth of a second, a quantity which we have not the smallest reason to hope any practical improvement in telescopes will ever show as an object having distinguishable *form*.

(817.) There remains therefore only the indication which the quantity of light they send to us may afford. But here again another difficulty besets us. The light of the sun is so immensely superior in intensity to that of any star, that it is impracticable to obtain any direct comparison between them. But by using the moon as an intermediate term of comparison it may be done, not indeed with much precision, but sufficiently well to satisfy in some degree our curiosity on the subject. Now α Centauri has been directly compared with the moon by the method explained in Art. 783. By a mean of eleven such comparisons made in various states of the moon, duly reduced and making the proper allowance on photometric principles for the moon's light lost by transmission through d prism, it appears that the mean quantity of light sent to the full moon exceeds that sent by α Centauri in the proportion of . Now Wollaston, by a method apparently unobjectionable, proportion of the sun's light to that of the full moon to be 072 to 1. Combining these results, we find the light sent us to be that sent by α Centauri as 21,955,000,000, or about thousand millions to 1. Hence from the parallax assigned, at star, it is easy to conclude that its intrinsic splendour, as vith that of our sun at equal distances, is 2·3247, that of the unity.[2]

The light of Sirius is four times that of α Centauri and its pa- 0″·230 (Art. 230). This in effect ascribes to it an intrinsic equal to 63·02 times that of our sun.[3]

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n, Phil. Trans. 1829, p. 27.

of Astronomical Observations at the Cape of Good Hope, &c. Art. 278, p.

the results obtained near the quadratures of the moon (which is the sit- favourable to exactness) be used, the resulting value of the intrinsic light he sun being unity) is 4·1586. On the other hand, if only those procured moon (the worst time for observation) be employed, the result is 1·4017. s of this kind will startle no one conversant with Photometry. That a lly emits more light than our sun must, we conceive, be regarded as an

CHAPTER XVI.

VARIABLE AND PERIODICAL STARS.—LIST OF THOSE ALREADY KNOWN. —IRREGULARITIES IN THEIR PERIODS AND LUSTRE WHEN BRIGHTEST.—IRREGULAR AND TEMPORARY STARS.—ANCIENT CHINESE RECORDS OF SEVERAL. — MISSING STARS. — DOUBLE STARS. — THEIR CLASSIFICATION.— SPECIMENS OF EACH CLASS.—BINARY SYSTEMS. — REVOLUTION ROUND EACH OTHER. — DESCRIBE ELLIPTIC ORBITS UNDER THE NEWTONIAN LAW OF GRAVITY.—ELEMENTS OF ORBITS OF SEVERAL.—ACTUAL DIMENSIONS OF THEIR ORBITS.—COLOURED DOUBLE STARS. — PHÆNOMENON OF COMPLEMENTARY COLOURS. — SANGUINE STARS. —PROPER MOTION OF THE STARS. —PARTLY ACCOUNTED FOR BY A REAL MOTION OF THE SUN. — SITUATION OF THE SOLAR APEX. — AGREEMENT OF SOUTHERN AND NORTHERN STARS IN GIVING THE SAME RESULT.—PRINCIPLES ON WHICH THE INVESTIGATION OF THE SOLAR MOTION DEPENDS.—ABSOLUTE VELOCITY OF THE SUN'S MOTION. — SUPPOSED REVOLUTION OF THE WHOLE SIDEREAL SYSTEM ROUND A COMMON CENTRE.— SYSTEMATIC PARALLAX AND ABERRATION. — EFFECT OF THE MOTION OF LIGHT IN ALTERING THE APPARENT PERIOD OF A BINARY STAR.

(819.) Now, for what purpose are we to suppose such magnificent bodies scattered through the abyss of space? Surely not to illuminate *our* nights, which an additional moon of the thousandth part of the size of our own would do much better, nor to sparkle as a pageant void of meaning and reality, and bewilder us among vain conjectures. Useful, it is true, they are to man as points of exact and permanent reference; but he must have studied astronomy to little purpose, who can suppose man to be the only object of his Creator's care, or who does not see in the vast and wonderful apparatus around us provision for other races of animated beings. The planets, as we have seen, derive their light from the sun; but that cannot be the case with the stars. These doubtless, then, are themselves suns, and may, perhaps, each in its sphere, be the presiding centre round which other planets, or bodies of which we can form no conception from any analogy offered by our own system, may be circulating.

Analogies, however, more than conjectural, are not wanting to
orrespondence between the dynamical laws which prevail in the
ons of the stars and those which govern the motions of our own
Wherever we can trace the law of periodicity — the regular re-
the same phænomena in the same times — we are strongly
with the idea of rotatory or orbitual motion. Among the stars
which, though no way distinguishable from others by any appa-
e of place, nor by any difference of appearance in telescopes,
a more or less regular periodical increase and diminution of
lving in one or two cases a complete extinction and revival.
alled periodical stars. The longest known and one of the most
is the star *Omicron*, in the constellation Cetus (sometimes
Ceti), which was first noticed as variable by Fabricius in 1596.
about twelve times in eleven years, or more exactly in a period
^{h} 7^m; remains at its greatest brightness about a fortnight, being
ne occasions equal to a large star of the second magnitude;
uring about three months, till it becomes completely invisible
d eye, in which state it remains about five months: and con-
asing during the remainder of its period. Such is the general
its phases. It does not always however return to the same

tainly appears strongly to suggest the revolution round it of some opaque body, which, when interposed between us and Algol, cuts off a large portion of its light; and this is accordingly the view taken of the matter by Goodricke, to whom we owe the discovery of this remarkable fact,[1] in the year 1782; since which time the same phænomena have continued to be observed, but with this remarkable additional point of interest, viz. that the more recent observations, as compared with the earlier ones, indicate a diminution in the periodic time. The latest observations of Argelander, Heis, and Schmidt, even go to prove that this diminution is not uniformly progressive, but is actually proceeding with accelerated rapidity, which however will probably not continue, but, like other cyclical combinations in astronomy, will by degrees relax, and then be changed into an increase, according to laws of periodicity which, as well as their causes, remain to be discovered. The first minimum of this star in the year 1844 occurred on Jan. 3, at $4^h\ 14^m$ Greenwich mean time.[2]

(822.) The star δ in the constellation Cepheus is also subject to periodical variations, which, from the epoch of its first observation by Goodricke in 1784 to the present time, have been continued with perfect regularity. Its period from minimum to minimum is $5^d\ 8^h\ 47^m\ 39^s{\cdot}5$, the first or epochal minimum for 1849 falling on Jan. 2, $3^h\ 13^m\ 37^s$ M. T. at Greenwich. The extent of its variation is from the fifth to between the third and fourth magnitudes. Its increase is more rapid than its diminution, the interval between the minimum and maximum of its light being only $1^d\ 14^h$, while that from the maximum to the minimum is $3^d\ 19^h$.

(823.) The periodical star β Lyræ, discovered by Goodricke also in 1784, has a period which has been usually stated at from $6^d\ 9^h$ to $6^d\ 11^h$, and there is no doubt that in about this interval of time its light undergoes a remarkable diminution and recovery. The more accurate observations of M. Argelander however have led him to conclude[3] the true period to be $12^d\ 21^h\ 53^m\ 10^s$, and that in this period a double maximum and minimum takes place, the two maxima being nearly equal and both about

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[1] The same discovery appears to have been made nearly about the same time by Palitzch, a farmer of Prolitz, near Dresden,—a peasant by station, an astronomer by nature,—who, from his familiar acquaintance with the aspect of the heavens, had been led to notice among so many thousand stars this one as distinguished from the rest by its variation, and had ascertained its period. The same Palitzch was also the first to rediscover the predicted comet of Halley in 1759, which he saw nearly a month before any of the astronomers, who, armed with their telescopes, were anxiously watching its return. These anecdotes carry us back to the era of the Chaldean shepherds.

[2] Ast. Nach. No. 472.

[3] Astron. Nachr. No. 264. See also the valuable papers by this excellent astronomer in A. N. Nos. 417, 455, &c.

agnitude, but the minima considerably unequal, viz. 4·3 and
addition to this curious subdivision of the whole interval of
o two semi-periods, we are presented in the case of this star
er instance of slow alteration of period, which has all the ap-
f being itself periodical. From the epoch of its discovery in
e year 1840 the period was continually lengthening, but more
lowly, till at the last-mentioned epoch it ceased to increase, and
een slowly on the decrease. As an epoch for the least or ab-
imum of this star, M. Argelander's calculations enable us to
6 January $3^d\ 0^h\ 9^m\ 53^s$ G. M. T.

Another periodical star whose changes have been carefully ob-
Aquilæ or Antinoi, first pointed out by Pigott in 1784 (a year
uch discoveries) as belonging to that class. Its period is $7^d\ 4^h$
he first minimum for 1849 occurring on Jan. 2, at $19^h\ 22^m\ 55^s$
It occupies fifty-seven hours in its increase from 5m to 4·3m,
ours in its decrease.

These are all the variable stars which have been observed with
are and for a sufficient length of time to enable us to speak
ion as to their periods, epochs, and phases of brightness. But

Star.	Period.	Change of Mag.		Discovered by
	d. dec.	from	to	
η Cygni	Many years	4·5	5·6	Herschel, jr., 1842?
* Virginis R. A. (1840) = $12^h\ 3^m$; N. P. D. 82° 8′	145 days	6·7	0	Harding, 1814.
* Coronæ Bor. (B. A. C. 5236)	10½ months	6	0	Pigott, 1795.
7 Arietis (B. A. C. 581.)	5 years?	6	8	Piazzi, 1798.
η Argûs	Irregular	1	4	Burchell, 1827.
α Orionis	Ditto	1	1·2	Herschel, jr., 1836.
α Ursæ Majoris	Some years	1·2	2	Ditto, 1846.
η Ursæ Majoris	Ditto	1·2	2	Ditto, 1846.
β Ursæ Minoris	2 or 3 years?	2	2·3	Struve, 1838.
α Cassiopeiæ	225 days?	2	2·3	Herschel, jr., 1838.
α Hydræ	29 or 30 days?	2·3	3	Ditto, 1837.
* R. A. (1847) = $22^h\ 58^m\ 57^s{\cdot}9$ N. P. D. = 80° 17′ 30″	Unknown	8?	0	Hind, 1848.
* R. A. (1848) = $7^h\ 33^m\ 55^s{\cdot}2$ N. P. D. = 66° 11′ 56″	Ditto	9	0	Ditto, 1848.
* R. A. (1848) = $7^h\ 40^m\ 10^s{\cdot}3$ N. P. D. = 65° 53′ 29″	Ditto	9	0	Ditto, 1848.
Near * R. A. $22^h\ 21^m\ 0^s{\cdot}4$ (1848.) N. P. D. 100° 42′ 40″	Ditto	7·8	0	Rümker.
* R. A. (1848) $14^h\ 44^m\ 39^s{\cdot}6$ N. P. D. 101° 45′ 25″	Ditto	8	9·10	Schumacher.
δ Ursæ Majoris	Many years	2?	2·3	Matter of general remark.

N. B. In the above list the letters B. A. C. indicate the catalogue of the British Association, B. the catalogue of Bode. Numbers before the name of the constellation (as 34 Cygni) denote Flamsteed's stars. Since this table was drawn up, four additional stars, variable from the 8th or 9th magnitude to 0, have been communicated to us by Mr. Hind, whose places are as follows: (1.) R. A. $1^h\ 38^m\ 24^s$; N. P. D. 81° 9′ 39″; (2.) $4^h\ 50^m\ 42^s$, 82° 6′ 36″ (1846); (3.) $8^h\ 43^m\ 8^s$, 86° 11′ (1800); (4.) $22^h\ 12^m\ 9^s$, 82° 59′ 24″ (1800.) Mr. Hind remarks that about several variable stars some degree of haziness is perceptible at their minimum. Have they clouds revolving round them as planetary or cometary attendants? He also draws attention to the fact that the red colour predominates among variable stars generally. The double star, No. 2718 of Struve's Catalogue, R. A. $20^h\ 34^m$, P. D. 77° 54′, is stated by the author to be variable. Captain Smyth (Celestial Cycle, i. 274) mentions also 3 Leonis and 18 Leonis as variable, the former from 6^m to 0, P = 78 days, the latter from 5^m to 10^m, P = $311^d\ 23^h$, but without citing any authority. Piazzi sets down 96 and 97 Virginis and 38 Herculis as variable stars. [The blood-red star, $4^h\ 51^m\ 50^s{\cdot}9$, 102° 2′ 4″ (1850), discovered by Mr. Hind, is stated by Schmidt (Ast. Nachr. 760) to have been seen by him 6^m in Jan. 1850, and to have totally disappeared in Dec. 1850 and Jan. 1851.]

(826.) Irregularities similar to those which have been noticed in the case of ο Ceti, in respect of the maxima and minima of brightness attained in successive periods, have been also observed in several others of the stars in the foregoing list. χ Cygni, for example, is stated by Cassini to have been scarcely visible throughout the years 1699, 1700, 1701, at those times when it was expected to be most conspicuous. No. 59 Scuti is sometimes visible to the naked eye at its minimum, and sometimes not so, and its maximum is also very irregular. Pigott's variable star in Corona

M. Argelander to vary for the most part so little that the
can hardly decide on its maxima and minima, while yet after
whole years of these slight fluctuations, they suddenly become
t the star completely vanishes. The variations of α Orionis,
most striking and unequivocal in the years 1836—1840,
years since elapsed became much less conspicuous. They
(Jan. 1849) to have recommenced.

hese irregularities prepare us for other phænomena of stellar
hich have hitherto been reduced to no law of periodicity, and
oked upon, in relation to our ignorance and inexperience, as
casual; or, if periodic, of periods too long to have occurred
once within the limits of recorded observation. The phæno-
llude to are those of *Temporary Stars*, which have appeared,
to time, in different parts of the heavens, blazing forth with
ry lustre; and after remaining awhile apparently immoveable,
way, and left no trace. Such is the star which, suddenly ap-
ne time about the year 125 B.C., and which was visible in the
said to have attracted the attention of Hipparchus, and led
w up a catalogue of stars, the earliest on record. Such, too,
ar which appeared A.D. 389 near α Aquilæ, remaining for

after becoming completely invisible, re-appeared, and, after undergoing one or two singular fluctuations of light, during two years, at last died away entirely, and has not since been seen.

(829.) On the night of the 28th of April, 1848, Mr. Hind observed a star of the fifth magnitude or 5·4 (very conspicuous to the naked eye) in a part of the constellation Ophiuchus (R.A. 16^h 51^m $1^s·5$. N.P.D. 102° 39′ 14″), where, from perfect familiarity with that region, he was certain that up to the fifth of that month no star so bright as 9·10 m. previously existed. Neither has any record been discovered of a star being there observed at any previous time. From the time of its discovery it continued to diminish, without any alteration of place, and before the advance of the season rendered further observation impracticable, was nearly extinct. Its colour was ruddy, and was thought by many observers to undergo remarkable changes, an effect probably of its low situation.

(830.) The alterations of brightness in the southern star η Argûs, which have been recorded, are very singular and surprising. In the time of Halley (1677) it appeared as a star of the fourth magnitude. Lacaille, in 1751, observed it of the second. In the interval from 1811 to 1815, it was again of the fourth; and again from 1822 to 1826 of the second. On the 1st of February, 1827, it was noticed by Mr. Burchell to have increased to the first magnitude, and to equal α Crucis. Thence again it receded to the second; and so continued until the end of 1837. All at once in the beginning of 1838 it suddenly increased in lustre so as to surpass all the stars of the first magnitude except Sirius, Canopus, and α Centauri, which last star it nearly equalled. Thence it again diminished, but this time not below the first magnitude until April, 1843, when it had again increased so as to surpass Canopus, and nearly equal Sirius in splendour. "A strange field of speculation," it has been remarked, "is opened by this phænomenon. The temporary stars heretofore recorded have all become totally extinct. Variable stars, so far as they have been carefully attended to, have exhibited periodical alternations, in some degree at least regular, of splendour and comparative obscurity. But here we have a star fitfully variable to an astonishing extent, and whose fluctuations are spread over centuries, apparently in no settled period, and with no regularity of progression. What origin can we ascribe to these sudden flashes and relapses? What conclusions are we to draw as to the comfort or habitability of a system depending for its supply of light and heat on so uncertain a source?" Speculations of this kind can hardly be termed visionary, when we consider that, from what has before been said, we are compelled to admit a community of nature between the fixed stars and our

nd when we reflect that geology testifies to the fact of exten-
s having taken place at epochs of the most remote antiquity in
and temperature of our globe; changes difficult to reconcile
eration of secondary causes, such as a different distribution of
d, but which would find an easy and natural explanation in a
on of the supply of light and heat afforded primarily by the sun

he Chinese annals of Ma-touan-lin,[1] in which stand *officially*
ough rudely, remarkable astronomical phænomena, supply a
"strange stars," among which, though the greater part are
mets, some may be recognized as belonging in all probability
s of *Temporary Stars* as above characterized. Such is that
corded to have appeared in A. D. 173, between α and β *Cen-*
h (no doubt, scintillating from its low situation) exhibited
lours," and remained visible from December in that year till
next. And another which these annals assign to A. D. 1011,
would seem to be identical with a star elsewhere referred to
"which was of extraordinary brilliancy, and remained visible
hern part of the heavens during three months,"[2] a situation

especially voyagers at sea, provided with only good eyes, or moderate instruments, might employ their time to excellent advantage. It holds out a sure promise of rich discovery, and is one in which astronomers in established observatories are almost of necessity precluded from taking a part, by the nature of the observations required. Catalogues of the comparative brightness of the stars in each constellation have been constructed by Sir Wm. Herschel, with the express object of facilitating these researches, and the reader will find them, and a full account of his method of comparison, in the Phil. Trans. 1796, and subsequent years.

(833.) We come now to a class of phænomena of quite a different character, and which give us a real and positive insight into the nature of at least some among the stars, and enable us unhesitatingly to declare them subject to the same dynamical laws, and obedient to the same power of gravitation which governs our own system. Many of the stars, when examined with telescopes, are found to be double, *i. e.* to consist of two (in some cases three or more) individuals placed near together. This might be attributed to accidental proximity, did it occur only in a few instances; but the frequency of this companionship, the extreme closeness, and, in many cases, the near equality of the stars so conjoined, would alone lead to a strong suspicion of a more near and intimate relation than mere casual juxtaposition. The bright star, Castor, for example, when much magnified, is found to consist of two stars of nearly the third magnitude, within 5″ of each other. Stars of this magnitude, however, are not so common in the heavens as to render it otherwise than excessively improbable that, if scattered at random, they would fall so near. But this improbability becomes immensely increased by a consideration of the fact, that this is only one out of a great many similar instances. Mitchell, in 1767, applying the rules for the calculation of probabilities to the case of the six brightest stars in the group called the Pleiades, found the odds to be 500000 to 1 against their proximity being the mere result of a random scattering of 1500 stars (which he supposed to be the total number of stars of that magnitude in the celestial sphere[1]) over the heavens. Speculating further on this, as an indication of physical connexion rather than fortuitous assemblage, he was led to surmise the possibility, (since converted into a certainty, but at that time, antecedent to any observation) of the existence of compound stars revolving about one another, or rather about their common centre of gravity. M. Struve, pursuing the same train of thought as applied specially to the cases of double and triple

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[1] This number is considerably too small, and in consequence, Mitchell's odds in this case materially overrated. But enough will remain, if this be rectified, fully to bear out his argument. Phil. Trans. vol. 57.

ns of stars, and grounding his computations on a more perfect
n of the stars visible down to the 7th magnitude, in the part
vens visible at Dorpat, calculates that the odds are 9570 to 1
y two stars, from the 1st to the 7th magnitude, inclusive, out
le possible number of binary combinations then visible, falling,
usly scattered) within 4″ of each other. Now, the number
s of such binary combinations actually observed at the date
lculation was already 91, and many more have since been
the list. Again, he calculates that the odds against any
fortuitously scattered, falling within 32″ of a third, so as to
a triple star, is not less than 173524 to 1. Now, four such
ns occur in the heavens; viz. θ Orionis, σ Orionis, 11 Monoce-
Cancri. The conclusion of a physical connexion of some kind
therefore unavoidable.

Presumptive evidence of another kind is furnished by the fol-
sideration. Both α Centauri and 61 Cygni are "Double Stars."
ist of two individuals, nearly equal, and separated from each
n interval of about a quarter of a minute. In the case of 61
stars exceeding the 7th magnitude, there is already a primâ

which the distance between the component individuals does not exceed a single second. They are divided into classes by M. Struve (the first living authority in this department of Astronomy) according to the proximity of their component individuals. The first class comprises those only in which the distance does not exceed 1″; the second those in which it exceeds 1″, but falls short of 2″; the 3d class extends from 2″ to 4″ distance; the 4th from 4″ to 8″; the 5th from 8″ to 12″; the 6th from 12″ to 16″; the 7th from 16″ to 24″; the 8th from 24″ to 32″. Each class he again subdivides into two sub-classes of which the one under the appellation of *conspicuous* double stars (duplices lucidæ) comprehends those in which both individuals exceed the 8½ magnitude, that is to say, are *separately* bright enough to be easily seen in any moderately good telescope. All others, in which one or both the constituents are below this limit of easy visibility, are collected into another sub-class, which he terms *residuary* (*Duplices reliquæ*). This arrangement is so far convenient, that after a little practice in the use of telescopes as applied to such objects, it is easy to judge what optical power will probably suffice to resolve a star of any proposed class and either sub-class, or would at least be so if the second or residuary sub-class were further sub-divided by placing in a third sub-class "delicate" double stars, or those in which the companion star is so very minute as to require a high degree of optical power to perceive it, of which instances will presently be given.

(886.) The following may be taken as specimens of each class. They are all taken from among the lucid, or conspicuous stars, and to such of our readers as may be in possession of telescopes, and may be disposed to try them on such objects, will afford him a ready test of their degree of efficiency.

Class I., 0″ to 1″.

γ Coronæ Bor.
γ Centauri.
γ Lupi.
ε Arietis.
ζ Herculis.
η Coronæ.
η Herculis.
λ Cassiopeiæ.
λ Ophiuchi.
π Lupi.
7 Ophiuchi.
φ Draconis.
φ Ursæ Majoris.
χ Aquilæ.
ω Leonis.
Atlas Pleiadum.
4 Aquarii.
42 Comæ.
52 Arietis.
66 Piscium.

Class II., 1″ to 2″.

γ Circini.
δ Cygni.
ε Chamæleontis.
ζ Bootis.
ι Cassiopeiæ.
ι 2 Cancri.
ξ Ursæ Majoris.
π Aquilæ.
σ Coronæ Bor.
2 Camelopardi.
32 Orionis.
52 Orionis.

Class III., 2″ to 4″.

α Piscium.
β Hydræ.
γ Ceti.
γ Leonis.
γ Coronæ Aus.
γ Virginis.
δ Serpentis.
ε Bootis.
ε Draconis.
ε Hydræ.
ζ Aquarii.
ζ Orionis.
ι Leonis.
ι Trianguli.
κ Leporis.
μ Draconis.
μ Canis.
ρ Herculis.
σ Cassiopeiæ.
44 Bootis.

Class IV., 4″ to 8″.

um.
um.
Bor.

θ Phœnicis.
κ Cephei.
λ Orionis.
μ Cygni.
ξ Bootis.

ξ Cephei.
π Bootis.
ρ Capricorni.
υ Argûs.
ω Aurigæ.

μ Eridani.
70 Ophiuchi.
12 Eridani.
32 Eridani.
95 Herculis.

Class V., 8″ to 12″.

ζ Antliæ.
η Cassiopeiæ.
θ Eridani.

ι Orionis.
f Eridani.
2 Canum Ven.

Class VI., 12″ to 16″.

γ Volantis.
η Lupi.
ζ Ursæ Major.

κ Bootis.
8 Monocerotis.
61 Cygni.

Class VII., 16″ to 24″.

Ven.

θ Serpentis.
κ Coronæ Aus.
χ Tauri.

24 Comæ.
41 Draconis.
61 Ophiuchi.

Class VIII., 24″ to 32″.

κ Herculis.
κ Cephei.
ψ Draconis.

χ Cygni.
23 Orionis.

Among the most remarkable triple, quadruple, or multiple stars to be beautifully and closely double, the one pair being about 2½″, the other about 3″ asunder. Each of the stars ζ Cancri, ξ Scorpii, 11 Monocerotis, and 12 Lyncis consists of a principal star, closely double, and a smaller and more distant attendant, while θ Orionis presents the phænomenon of four brilliant principal stars, of the respective 4th, 6th, 7th, and 8th magnitudes, forming a trapezium, the longest diagonal of which is 21″.4, and accompanied by two excessively minute and very close companions (as in the annexed figure), to perceive *both* which is one of the severest tests which can be applied to a telescope.

(838.) Of the "delicate" sub-class of double stars, or those consisting of very large and conspicuous principal stars, accompanied by very minute companions, the following specimens may suffice:

α 2 Canc̆ri.
α 2 Capricorni.
α Indi.
α Lyræ.
α Polaris.
β Aquarii.
γ Hydræ.
ι Ursæ Majoris.
κ Circini.
κ Geminorum.
μ Persei.
7 Bootis.
φ Virginis.
χ Eridani.
16 Aurigæ.
94 Ceti.

(839.) To the amateur of Astronomy the double stars offer a subject of very pleasing interest, as tests of the performance of his telescopes, and by reason of the finely contrasted colours which many of them exhibit, of which more hereafter. But it is the high degree of physical interest which attaches to them, which assigns them a conspicuous place in modern Astronomy, and justifies the minute attention and unwearied diligence bestowed on the measurement of their angles of position and distances, and the continual enlargement of our catalogues of them by the discovery of new ones. It was, as we have seen, under an impression that such combinations, if diligently observed, might afford a measure of parallax through the periodical variations it might be expected to produce in the relative situation of the small attendant star, that Sir W. Herschel was induced (between the years 1779 and 1784) to form the first extensive catalogues of them, under the scrutiny of higher magnifying powers than had ever previously been applied to such purposes. In the pursuit of this object, the end to which it was instituted as a means was necessarily laid aside for a time, until the accumulation of more abundant materials should have afforded a choice of stars favourably circumstanced for systematic observation. Epochal measures however, of each star, were secured, and, on resuming the subject, his attention was altogether diverted from the original object of the inquiry by phænomena of a very unexpected character, which at once engrossed his whole attention. Instead of finding, as he expected, that annual fluctuation to and fro of one star of a double star with respect to the other,—that alternate annual increase and decrease of their distance and angle of position, which the parallax of the earth's annual motion would produce,—he observed,

nstances, a regular progressive change; in some cases bearing
their distance, — in others on their position, and advancing
one direction, so as clearly to indicate either a real motion of
emselves, or a general rectilinear motion of the sun and whole
n, producing a parallax of a higher order than would arise from
s orbitual motion, and which might be called systematic

Supposing the two stars, and also the sun, in motion independ-
ch other, it is clear that for the interval of several years, these
ust be regarded as rectilinear and uniform. Hence, a very
aintance with geometry will suffice to show that the *apparent*
one star of a double star, referred to the other as a centre, and
wn, as it were, on a plane in which that other shall be taken
or zero point, can be no other than a right line. This, at
be the case if the stars be independent of each other; but it
erwise if they have a physical connexion, such as, for instance,
nity and mutual gravitation would establish. In that case, they
ribe orbits round each other, and round their common centre
and therefore the apparent path of either, referred to the other

position of double stars, are adduced in the memoirs above mentioned; many of which are too decided, and too regularly progressive, to allow of their nature being misconceived. In particular, among the more conspicuous stars,—Castor, γ Virginis, ξ Ursæ, 70 Ophiuchi, σ and η Coronæ, ξ Bootis, η Cassiopeiæ, γ Leonis, ζ Herculis, δ Cygni, μ Bootis, ε 4 and ε 5 Lyræ, λ Ophiuchi, μ Draconis, and ζ Aquarii, are enumerated as among the most remarkable instances of the observed motion; and to some of them even periodic times of revolution are assigned; approximative only, of course, and rather to be regarded as rough guesses than as results of any exact calculation, for which the data were at the time quite inadequate. For instance, the revolution of Castor is set down at 334 years, that of γ Virginis at 708, and that of γ Leonis at 1200 years.

(842.) Subsequent observation has fully confirmed these results. Of all the stars above named, there is not one which is not found to be fully entitled to be regarded as binary; and, in fact, this list comprises nearly all the most considerable objects of that description which have yet been detected, though (as attention has been closely drawn to the subject, and observations have multiplied) it has, of late, received large accessions. Upwards of a hundred double stars, certainly known to possess this character, were enumerated by M. Mädler in 1841,[1] and more are emerging into notice with every fresh mass of observations which come before the public. They require excellent telescopes for their effective observation, being for the most part so close as to necessitate the use of very high magnifiers (such as would be considered extremely powerful microscopes if employed to examine objects within our reach), to perceive an interval between the individuals which compose them.

(843.) It may easily be supposed, that phænomena of this kind would not pass without attempts to connect them with dynamical theories. From their first discovery, they were naturally referred to the agency of some power, like that of gravitation, connecting the stars thus demonstrated to be in a state of circulation about each other; and the extension of the Newtonian law of gravitation to these remote systems was a step so obvious, and so well warranted by our experience of its all-sufficient agency in our own, as to have been expressly or tacitly made by every one who has given the subject any share of his attention. We owe, however, the first distinct system of calculation, by which the elliptic elements of the orbit of a binary star could be deduced from observations of its angle of position and distance at different epochs, to M. Savary, who showed[2] that the motions of one of the most remarkable among them (ξ Ursæ) were

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[1] Dorpat Observations, vol. ix. 1840 and 1841.

[2] Connoiss. des Temps, 1830.

within the limits allowable for error of observation, on the
of an elliptic orbit described in the short period of 58¼ years.
t process of computation conducted Professor Encke[1] to an
it for 70 Ophiuchi, described in a period of seventy-four years.
has especially signalized himself in this line of inquiry (see
veral orbits have also been calculated by Mr. Hind and Cap-
, and the author of these pages has himself attempted to con-
mite to these interesting investigations. The following may
s the chief results which have been hitherto obtained in this
stronomy: —

hem. 1832.

ents Nos. 1, 2, 3, 4 c, 5, 6 c, 7, 11 b, 12 a, are extracted from M. Măd-
view of the history of double stars in vol. ix. of the Dorpat Observations:
e Connoiss. des Temps, 1830: 4 b, 6 b, and 11 a, from vol. v. Trans.
Lond.: 6 a, from Berlin Ephemeris, 1832: No. 8, from Trans. Astron.
No. 9, 11 c, 12 b, and 13 from Notices of the Astronomical Society,
, and viii. p. 159, and No. 10 from the author's "Results of Astronomical
, &c. at the Cape of Good Hope," p. 297. The Σ prefixed to No. 7,
number of the star in M. Struve's Dorpat Catalogue (Catalogus Novus
uplicium, &c. Dorpat, 1827), which contains the places for 1826 of 3112

Star's Name.	Apparent semi-axis.	Excentricity.	Position of Node.	Perihelion from Node or Orbit.	Inclination.	Period in Years.	Perihelion Passage.	By whom computed.
1. Herculis	1″·189	0·44454	39° 26′	262° 4′	50° 53′	31·468	1829·50	Mädler.
2. η Coronæ B.	1·088	0·33760	24 18	261 21	71 8	43·246	1815·23	Ditto.
3. ζ Cancri	1·292	0·23486	1 28	266 0	63 17	58·910	1853·37	Ditto.
4. a. ξ Ursæ Majoris	3·857	0·41640	95 22	131 38	50 40	58·262	1817·25	Savary.
4. b. Ditto.	3·278	0·37770	97 47	134 22	56 6	60·720	1816·73	Herschel, junior.
4. c. Ditto.	2·417	0·41350	98 52	130 48	54 56	61·464	1816·44	Mädler.
4. d. Ditto.	2·439	0·43148	95 50	128 57	52 49	61·576	1816·86	Villarceaux.
5. ω Leonis	0·857	0·64338	135 11	185 27	46 33	82·533	1840·76	Ditto.
6. a. p. Ophiuchi	4·328	0·43007	147 12	125 22	46 25	73·862	1806·88	Encke.
6. b. Ditto	4·392	0·46670	137 2	145 46	48 5	80·340	1807·06	Herschel, junior.
6. c. Ditto	4·192	0·44380	126 55	142 53	64 51	92·870	1812·73	Mädler.
7. Σ 3062	1·255	0·44958	15 3	137 27	35 31	94·765	1837·41	Ditto.
8. ξ Bootis	12·560	0·59374	359 59	100 59	80 5	117·140	1779·88	Herschel, junior.
9. δ Cygni	1·811	0·60667	24 54	243 24	46 23	178·700	1862·87	Hind.
10. γ Virginis	3·580	0·87952	5 33	313 45	23 36	182·120	1836·43	Herschel, junior.
11. a. Castor	8·086	0·75820	58 6	97 29	70 3	252·660	1855·83	Ditto.
11. b. Ditto	7·008	0·79725	23 5	87 37	70 58	232·124	1913·90	Mädler.
11. c. Ditto	6·300	0·24050	11 24	356 22	43 14	632·270	1609·26	Hind.
12. a. σ Coronæ B	3·918	0·69978	25 7	64 38	29 29	608·450	1826·60	Mädler.
12. b. Ditto	5·194	0·72560	21 3	69 24	25 39	736·880	1826·48	Hind.
13. μ 2 Bootis	8·218	0·84010	117 21	103 17	46 57	649·720	1852·50	Ditto.
14. α Centauri	15·500	0·95000	86 7	291 22	47 56	77·000	1851·50	Jacob.
15. ζ Herculis	1·254	0·44821	34 21	104 55	43 43	36·357	1830·48	Villarceaux.

Of the stars in the above list, that which has been most assidu-
hed, and has offered phænomenon of the greatest interest, is
It is a star of the vulgar 3rd magnitude (3.08 = Photom.
d its component individuals are very nearly equal, and as it
n in some slight degree variable, since, according to the obser-
M. Struve, the one is alternately a little greater and a little
the other, and occasionally exactly equal to it. It has been
onsist of two stars since the beginning of the eighteenth cen-
distance being then between six and seven seconds, so that any
ood telescope would resolve it. When observed by Herschel
was 5″·66, and continued to decrease gradually and regularly
h, in 1836, the two stars had approached so closely as to appear
ound and single under the highest magnifying power which
pplied to most excellent instruments — the great refractor at
lone, with a magnifying power of 1000, continuing to indicate,
lge-shaped form of the disc of the star its composite nature.
ing the ratio of its length to its breadth and measuring the
Struve concludes that, at this epoch (1836·41), the distance
stars, centre from centre, might be stated at 0″·22. From that
ar again opened, and at present (1849) the individuals are more

of nearly nine-tenths of a complete circuit), both in angle and distance, with a degree of exactness *fully equal to that of observation itself.* No doubt can, therefore, remain as to the prevalence in this remote system of the Newtonian law of gravitation.

(845.) The observations of ξ Ursæ Majoris are equally well represented by M. Mädler's elements (4 c of our table,) thus fully justifying the assumption of the Newtonian law as that which regulates the motions of their binary systems. And even should it be the case, as M. Mädler appears to consider, that in one instance at least (that of p Ophiuchi,) deviations from elliptic motion, too considerable to arise from mere error of observation, exist (a position we are by no means prepared to grant,)[1] we should rather be disposed to look for the cause of such deviations in perturbations arising (as Bessel has suggested) from the large or central star itself being actually a close and hitherto unrecognized double star than in any defect of generality in the Newtonian law.

(846.) If the great length of the periods of some of these bodies be remarkable, the shortness of those of others is hardly less so. ζ Herculis has already completed two revolutions since the epoch of its first discovery, exhibiting in its course the extraordinary spectacle of a sidereal occultation, the small star having twice been completely hidden behind the large one. η Coronæ, ζ Cancri, and ξ Ursæ have each performed more than one entire circuit, and 70 Ophiuchi and γ Virginis have accomplished by far the larger portion of one in angular motion. If any doubt, therefore, could remain as to the reality of their orbitual motions, or any idea of explaining them by mere parallactic changes, or by any other hypothesis than the agency of centripetal force, these facts must suffice for their complete dissipation. We have the same evidence, indeed, of their rotations about each other, that we have of those of Uranus and Neptune about the sun; and the correspondence between their calculated and observed places in such very elongated ellipses, must be admitted to carry with it proof of the prevalence of the Newtonian law of gravity in their systems, of the very same nature and cogency as that of the calculated and observed places of comets round the central body of our own.

(847.) But it is not with the revolutions of bodies of a planetary or

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[1] p Ophiuchi belongs to the class of very unequal double stars, the magnitudes of the individuals being 4 and 7. Such stars present difficulties in the exact measurement of their angles of position which even yet continue to embarrass the observer, though, owing to later improvements in the art of executing such measurements, their influence is confined within much narrower limits than in the earlier history of the subject. In simply placing a fine single wire parallel to the line of junction of two such stars it is easily possible to commit an error of 3° or 4°. By placing them between two parallel thick wires such errors are in great measure obviated.

ature round a solar centre that we are now concerned; it is
f sun round sun — each, perhaps, at least in some binary sys-
the individuals are very remote and their period of revolution
ccompanied with its train of planets and *their* satellites, closely
rom our view by the splendour of their respective suns, and
o a space bearing hardly a greater proportion to the enormous
ich separates *them*, than the distances of the satellites of our
n their primaries bear to their distances from the sun itself.
inctly characterized subordination would be incompatible with
of their systems, and with the planetary nature of their orbits.
ely nestled under the protecting wing of their immediate supe-
eep of their other sun in its perihelion passage round their
carry them off, or whirl them into orbits utterly incompatible
onditions necessary for the existence of their inhabitants. It
onfessed, that we have here a strangely wide and novel field
tive excursions, and one which it is not easy to avoid luxu-

The discovery of the parallaxes of α Centauri and 61 Cygni,
are above enumerated among the "conspicuous" double stars
class (a distinction fully merited in the case of the former by

to be 0·353, the mass of our sun being 1. The sun, therefore, is neither vastly greater nor vastly less than the stars composing 61 Cygni.

(849.) The data in the case of α Centauri are more uncertain. Since the year 1822, the distance has been steadily and pretty rapidly decreasing at the rate of about half a second per annum, and that with very little change in the angle of position. Hence, it follows evidently that the plane of its orbit passes nearly through the earth, and (the distance about the middle of 1834 having been 17½″) it is very probable that either an occultation, like that observed in ζ Herculis, on a close appulse of the two stars, will take place about the year 1867. As the observations we possess afford no sufficient grounds for a satisfactory calculation of elliptic elements[1] we must be content to assume what, at all events, they fully justify, viz., that the major semiaxis must exceed 12″, and is very probably considerably greater. Now this with a parallax of 0″·913 would give for the real value of the semiaxis 13·15 radii of the earth's orbit, as a minimum. The real dimensions of their ellipse, therefore, cannot be so small as the orbit of Saturn; in all probability exceeds that of Uranus; and may possibly be much greater than either.

(850.) The parallel between these two double stars is a remarkable one. Owing no doubt to their *comparative* proximity to our system, their apparent proper motions are both unusually great, and for the same reason probably rather than owing to unusually large dimensions, their orbits appear to us under what, for binary double stars, we must call unusually large angles. Each consists, moreover, of stars, not very unequal in brightness, and in each both the stars are of a high yellow approaching to orange colour, the smaller individual, in each case, being also of a deeper tint. Whatever the diversity, therefore, which may obtain among other sidereal objects, these would appear to belong to the same family or genus.[2]

(851.) Many of the double stars exhibit the curious and beautiful phænomenon of contrasted or complementary colours.[3] In such instances,

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[1] Elements have been recently computed by Captain Jacob, for which see the table, p. 483.

[2] Similar combinations are very numerous. Many remarkable instances occur among the double stars catalogued by the author in the 2nd, 3rd, 4th, 6th and 9th volumes of Trans. Roy. Ast. Soc. and in the volume of Southern observations already cited. See Nos. 121, 375, 1066, 1907, 2030, 2146, 2214, 2772, 3853, 3395, 3998, 4000, 4055, 4196, 4210, 4615, 4649, 4765, 5003, 5012, of these catalogues. The fine binary star, B. A. C. No. 4923, has its constituents 15″ apart, the one 6*m*. yellow, the other, 7*m*. orange.

[3] "——— other suns, perhaps,
With their attendant moons, thou wilt descry,
Communicating male and female light,

tar is usually of a ruddy or orange hue, while the smaller one
e or green, probably in virtue of that general law of optics,
ides, that when the retina is under the influence of excitement
ht, coloured light; feebler lights, which seen alone would pro-
sation but of whiteness, shall for the time appear coloured with
nplementary to that of the brighter. Thus a yellow colour
ng in the light of the brighter star, that of the less bright one
e field of view will appear blue; while, if the tint of the
r verge to crimson, that of the other will exhibit a tendency to
even appear as a vivid green, under favourable circumstances.
contrast is beautifully exhibited by ι Cancri — the latter by γ
e', both fine double stars. If, however, the coloured star be
less bright of the two, it will not materially affect the other.
nstance, η Cassiopeiæ exhibits the beautiful combination of a
star, and a small one of a rich ruddy purple. It is by no
ever, intended to say, that in all such cases one of the colours
ffect of contrast, and it may be easier suggested in words, than
n imagination, what variety of illumination *two suns* — a red
, or a yellow and a blue one — must afford a planet circulating

(852.) Another very interesting subject of inquiry, in the physical history of the stars, is their proper motion. It was first noticed by Halley, that three principal stars, Sirius, Arcturus, and Aldebaran, are placed by Ptolemy, on the strength of observations made by Hipparchus, 130 years B. C., in latitudes respectively 20′, 22′, and 33′ more *northerly* than he actually found them in 1717.[1] Making due allowance for the diminution of obliquity of the ecliptic in the interval (1847 years) they ought to have stood, if really fixed, respectively 10′, 14′, and 0′ more *southerly*. As the circumstances of the statement exclude the supposition of error of transcription in the MSS., we are necessitated to admit a southward motion in latitude in these stars to the very considerable extent, respectively, of 37′, 42′, and 33′, and this is corroborated by an observation of Aldebaran at Athens, in the year A. D. 509, which star, on the 11th of March in that year, was seen immediately after its emergence from occultation by the moon, in such a position as it could not have had if the occultation were not nearly central. Now, from the knowledge we have of the lunar motions, this could not have been the case had Aldebaran at that time so much southern latitude as at present. *A priori*, it might be expected that apparent motions of some kind or other should be detected among so great a multitude of individuals scattered through space, and with nothing to keep them fixed. Their mutual attractions even, however inconceivably enfeebled by distance, and counteracted by opposing attractions from opposite quarters, must in the lapse of countless ages produce *some* movements—some change of internal arrangement—resulting from the difference of the opposing actions. And it is a fact, that such apparent motions are really proved to exist by the exact observations of modern astronomy. Thus, as we have seen, the two stars of 61 Cygni have remained constantly at the same, or very nearly the same, distance, of 15″, for at least fifty years past, although they have shifted their local situation in the heavens, in this interval of time, through no less than 4′ 23″, the annual proper motion of each star being 5″·3; by which quantity (exceeding a third of their interval) this system is every year carried bodily along in some unknown path, by a motion which, for many centuries, must be regarded as uniform and rectilinear. Among stars not double, and no way differing from the rest in any other obvious particular, ε Indi[2] and μ Cassiopeiæ are to be remarked as having the greatest proper motions of any yet ascertained, amounting respectively to 7″·74 and 3″·74 of annual displacement. And a great many others have

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[1] Phil. Trans. 1717, vol. xxx. fo. 736.

[2] D'Arrest. Astr. Nachr., No. 618.

ved to be thus constantly carried away from their places by
t not less unequivocal motions.[1]

Motions which require whole centuries to accumulate before
ce changes of arrangement, such as the naked eye can detect,
te sufficient to destroy that idea of mathematical fixity which
speculation, are yet too trifling, as far as practical applications
ce a change of language, and lead us to speak of the stars in
rlance as otherwise than fixed. Small as they are, however,
s, once assured of their reality, have not been wanting in at-
xplain and reduce them to general laws. No one, who reflects
ttention on the subject, will be inclined to deny the high proba-
certainty, that the *sun* as well as the stars must have a proper
some direction; and the inevitable consequence of such a mo-
participated by the rest, must be a slow *average* apparent ten-
ll the stars to the vanishing point of lines parallel to that
nd to the region which he is leaving, however greatly indi-
s might differ from such average by reason of their own pecu-
motion. This is the necessary effect of perspective; and it is
t it must be detected by observation, if we knew accurately the
roper motions of all the stars, and if we were sure that they

the foregoing, but differing from it about 27° in right ascension. Since that time methods of calculation have been improved and concinnated, our knowledge of the proper motions of the stars has been rendered more precise, and a greater number of cases of such motions have been recorded. The subject has been resumed by several eminent astronomers and mathematicians: viz. 1st, by M. Argelander, who, from the consideration of the proper motions of 21 stars exceeding 1″ per annum in arc, has placed the solar apex in R. A. 256° 25′, N. P. D. 51° 23′; from those of 50 stars between 0″·5 and 1″·0, in 255° 10′ 51° 26′; and from those of 819 stars having motions between 0″·1 and 0″·5 per annum, in 261° 11′ 59° 2′: 2ndly, by M. Luhndahl, whose calculations, founded on the proper motions of 147 stars, give 252° 53′, 75° 34′: and 3rdly, by M. Otto Struve, whose result 261° 22′, 62° 24′, emerges from a very elaborate discussion of the proper motions of 392 stars. All these places are for A. D. 1790.

(855.) The most probable mean of the results obtained by these three astronomers, is (for the same epoch) R. A.=259° 9′, N. P. D. 55° 23′. Their researches, however, extending only to stars visible in European observatories, it became a point of high interest to ascertain how far the stars of the southern hemisphere not so visible, treated independently on the same system of procedure, would corroborate or controvert their conclusion. The observations of Lacaille, at the Cape of Good Hope, in 1751 and 1752, compared with those of Mr. Johnson at St. Helena, in 1829–33, and of Henderson at the Cape in 1830 and 1831, have afforded the means of deciding this question. The task has very recently been executed in a masterly manner by Mr. Galloway, in a paper published in the Philosophical Transactions for 1847 (to which we may also refer the reader for a more particular account of the history of the subject than our limits allow us to give.) On comparing the records, Mr. Galloway finds eighty-one southern stars not employed in the previous investigations above referred to, whose proper motions in the intervals elapsed appear considerable enough to assure us that they have not originated in error of the earlier observations. Subjecting these to the same process of computation he concludes for the place of the solar apex, for 1790, as follows: viz. R. A. 260° 1′, N. P. D. 55° 37′, a result so nearly identical with that afforded by the northern hemisphere, as to afford a full conviction of its near approach to truth, and what may fairly be considered a demonstration of the physical cause assigned.

(856.) Of the mathematical conduct of this inquiry the nature of this work precludes our giving any account; but as the philosophical principle on which it is based has been misconceived, it is necessary to say a few

xplanation of it. Almost all the greatest discoveries in astron-
resulted from the consideration of what we have elsewhere
SIDUAL PHÆNOMENA[1], of a quantitative or numerical kind,
ay, of such portions of the numerical or quantitative results of
as remain outstanding and unaccounted for after subducting
ng for all that would result from the strict application of known
It was thus that the grand discovery of the precession of the
resulted as a residual phænomenon, from the imperfect explana-
return of the seasons by the return of the sun to the same
lace among the fixed stars. Thus, also, aberration and nutation
residual phænomena from that portion of the changes of the
laces of the fixed stars which was left unaccounted for by pre-
And thus again the *apparent* proper motions of the stars are
ed *residues* of their apparent movements outstanding and unac-
by strict calculation of the effects of precession, nutation, and
The nearest approach which human theories can make to
is to diminish this residue, this *caput mortuum* of observation,
be considered, as much as practicable, and, if possible, to reduce
ng, either by showing that something has been neglected in our

(857.) This is the process resorted to by all the geometers we have enumerated in the foregoing articles (arts. 854, 855). It gives not only the direction in space, but also the *velocity* of the solar motion, estimated on a scale conformable to that in which the velocity of the sidereal motions to be explained are given; *i. e.* in seconds of arc as subtended at the average distance of the stars concerned, by its annual motion in space. But here a consideration occurs which tends materially to complicate the problem, and to introduce into its solution an element depending on suppositions more or less arbitrary. The distance of the stars being, except in two or three instances, unknown, we are compelled either to restrict our inquiry to *these*, which are too few to ground any result on, or to make some supposition as to the relative distances of the several stars employed. In this we have nothing but general probability to guide us, and two courses only present themselves, either, 1st, To class the distances of the stars according to their magnitudes, or apparent brightnesses, and to institute separate and independent calculations for each class, including stars assumed to be equidistant, or nearly so: or, 2dly, To class them according to the observed amount of their apparent proper motions, on the presumption that those which appear to move fastest are really nearest to us. The former is the course pursued by M. Otto Struve, the latter by M. Argelander. With regard to this latter principle of classification, however, two considerations interfere with its applicability, viz. 1st that we see the *real* motion of the stars foreshortened by the effect of perspective; and 2dly, that that portion of the total *apparent* proper motion which arises from the real motion of the sun depends, not simply on the absolute distance of the star from the sun, but also on its angular apparent distance from the solar apex, being, *cæteris paribus*, as the sine of that angle. To execute such a classification correctly, therefore, we ought to know both these particulars for each star. The first is evidently out of our reach. We are therefore, for that very reason, compelled to regard it as casual, and to assume that on the average of a great number of stars it would be uninfluential on the result. But the second cannot be so summarily disposed of. By the aid of an approximate knowledge of the solar apex, it is true, approximate values may be found of the simply apparent portions of the proper motions, supposing all the stars equidistant, and these being subducted from the total observed motions, the residues might afford ground for the classification in question.[1] This, however, would be to a certain extent precarious system of procedure. On the , the classification by apparent brightness is open to no such since we are fully justified in assuming that, on a general e brighter stars are the nearer, and that the exceptions to this sual in that sense of the word which it always bears in such xpressing solely our ignorance of any ground for assuming a ny or other on either side of a determinate numerical rule. In ay's discussion of the southern stars the consideration of disived altogether, which is equivalent to an admission of complete on this point, as well as respecting the real directions and f the individual motions.

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[1] M. Argelander's classes, however, are constructed without reference to this consideration, on the sole basis of the total apparent amount proper motion, and are, therefore, *pro tanto*, questionable. It is the more satisfactory then to find so considerable an agreement among his partial results as actually obtains.

The velocity of the solar motion which results from M. Otto lculations is such as would carry it over an angular subtense if seen at right angles from the average distance of a star of gnitude. If we take, with M. Struve, senior, the parallax of as probably equal to 0″·209,[1] we shall at once be enabled to is annual motion with the radius of the earth's orbit, the result 3 of such units. The sun then advances through space (relaleast, among the stars,) carrying with it the whole planetary great distances from them respectively, in the manner of binary stars: [and it has been recently shown by M. Peters (Ast. Nachr. 748,) that, in the case of Sirius, such a circulation, performed in a period of 50·093 years in an ellipse whose excentricity is 0.7994, the perihelion passage taking place at the epoch A. D. 1701·431, would reconcile in a remarkable manner the observed anomalies, and reduce the residual motion to uniformity.]

(860.) The whole of the reasoning upon which the determination of the solar motion in space rests, is based upon the entire exclusion of any law either derived from observation or assumed in theory, affecting the amount and direction of the real motions both of the sun and stars. It supposes an absolute non-recognition, in those motions, of any general directive cause, such as, for example, a common circulation of all about a common centre. Any such limitation introduced into the conditions of the problem of the solar motion would alter *in toto* both its nature and the form of its solution. Suppose, for instance, that, conformably to the speculations of several astronomers, the whole system of the Milky Way, including our sun, and the stars, our more immediate neighbours, which constitute our sidereal firmament, should have a general movement of rotation in the plane of the galactic circle (any other would be exceedingly improbable, indeed hardly reconcilable with dynamical principles,) being held together in opposition to the centrifugal force thus generated by the mutual gravitation of its constituent stars. Except we at the same time admitted that the scale on which this movement proceeds is so enormous that all the stars whose proper motions we include in our calculations go together in a body, so far as that movement is concerned (as forming too small an integrant portion of the whole to differ sensibly in their relation to its central point;) we stand precluded from drawing any conclusion whatever, not only respecting the absolute motion of the sun, but respecting even its relative movement among those stars, until we have established some law, or at all events framed some hypothesis having the provisional force of a law, connecting the whole, or a part of the motion of each individual with its situation in space.

(861.) Speculations of this kind have not been wanting in astronomy, and recently an attempt has been made by M. Mädler to assign the local centre in space, round which the sun and stars revolve, which he places in the group of the Pleiades, a situation in itself improbable, lying as it does no less than 26° out of the plane of the galactic circle, out of which plane it is almost inconceivable that any *general* circulation can take place. In the present defective state of our knowledge respecting the proper motion of the smaller stars, especially in right ascension, (an element for

rt far less exactly ascertainable than the polar·distance, or at
has been hitherto far less accurately ascertained,) we cannot
all attempts of the kind as to a certain extent premature,
no means to be discouraged as forerunners of something more
The question, as a matter of fact, whether a rotation of the
ts own plane exist or not, might be at once resolved by the
bservation, both in R. A. and polar distance, of a considerable
stars of the Milky Way, judiciously selected for the purpose,
ing all magnitudes, down to the smallest distinctly identifiable,
e of being observed with normal accuracy: and we would re-
he inquiry to the special attention of the directors of permanent
es, provided with adequate instrumental means, in both hemi-
Thirty or forty years of observation, perseveringly directed to
n view, could not fail to settle the question.[1]

The solar motion through space, if real and not simply relative,
ise to uranographical corrections analogous to parallax and
The solar or systematic parallax is no other than that part
er motion of each star which is simply apparent, arising from
motion, and until the distances of the stars be known, must

ray in a period of 10,000 days. Then if both the sun and the centre of gravity of the binary system remained fixed in space, the relative apparent situation of the stars would be exactly restored to its former state after the lapse of this interval, and if the angle of position were 0° at first, after 10,000 days it would again be so. But now suppose that the centre of gravity of the star were in the act of receding in a direct line from the sun, with a velocity of one-tenth part of the radius of the earth's orbit *per diem*. Then at the expiration of 10,000 days it would be more remote from us by 1000 such radii, a space which light would require 57 days to traverse. Although really, therefore, the stars would have arrived at the position 0° at the exact expiration of 10,000 days, it would require 57 days more for the notice of that fact to reach our system. In other words, the period would appear to us to be 10,057 days, since we could only conclude the period to be completed, when to us, as observers, the original angle of position was again restored. A contrary motion would produce a contrary effect.

CHAPTER XVII.

CLUSTERS OF STARS AND NEBULÆ.

ERING GROUPS OF STARS. — GLOBULAR CLUSTERS. — THEIR
TY DYNAMICALLY POSSIBLE. — LIST OF THE MOST REMARK-
— CLASSIFICATION OF NEBULÆ AND CLUSTERS. — THEIR
UTION OVER THE HEAVENS. — IRREGULAR CLUSTERS. —
ABILITY OF NEBULÆ. — THEORY OF THE FORMATION OF
RS BY NEBULOUS SUBSIDENCE. — OF ELLIPTIC NEBULÆ.
OF ANDROMEDA. — ANNULAR AND PLANETARY NEBULÆ.
LE NEBULÆ. — NEBULOUS STARS. — CONNEXION OF NEBULÆ
DOUBLE STARS. — INSULATED NEBULÆ, OF FORMS NOT
IRREGULAR. — OF AMORPHOUS NEBULÆ. — THEIR LAW OF
UTION MARKS THEM AS OUTLIERS OF THE GALAXY —

insulated from the rest of the heavens. The constellation called Coma Berenices is another such group, more diffused, and consisting on the whole of larger stars.

(865.) In the constellation Cancer, there is a somewhat similar, but less definite, luminous spot, called Præsepe, or the bee-hive, which a very moderate telescope,—an ordinary night-glass for instance,—resolves entirely into stars. In the sword-handle of Perseus, also, is another such spot, crowded with stars, which requires rather a better telescope to resolve into individuals, separated from each other. These are called clusters of stars; and, whatever be their nature, it is certain that other laws of aggregation subsist in these spots, than those which have determined the scattering of stars over the general surface of the sky. This conclusion is still more strongly pressed upon us, when we come to bring very powerful telescopes to bear on these and similar spots. There are a great number of objects which have been mistaken for comets, and, in fact, have very much the appearance of comets without tails: small round, or oval nebulous specks, which telescopes of moderate power only show as such. Messier has given, in the *Connois. des Temps* for 1784, a list of the places of 103 objects of this sort; which all those who search for comets ought to be familiar with, to avoid being misled by their similarity of appearance. That they are not, however, comets, their fixity sufficiently proves; and when we come to examine them with instruments of great power,—such as reflectors of eighteen inches, two feet, or more in aperture,—any such idea is completely destroyed. They are then, for the most part, perceived to consist entirely of stars crowded together so as to occupy almost a definite outline, and to run up to a blaze of light in the centre, where their condensation is usually the greatest. (See *fig.* 1, pl. II., which represents (somewhat rudely) the thirteenth nebula of Messier's list (described by him as *nébuleuse sans étoiles*), as seen in a reflector of 18 inches aperture and 20 feet focal length.) Many of them, indeed, are of an exactly round figure, and convey the complete idea of a globular space filled full of stars, insulated in the heavens, and constituting in itself a family or society apart from the rest, and subject only to its own internal laws. It would be a vain task to attempt to count the stars in one of these *globular clusters*. They are not to be reckoned by hundreds; and on a rough calculation, grounded on the apparent intervals between them at the borders, and the angular diameter of the whole group, it would appear that many clusters of this description must contain at least five thousand stars, compacted and wedged together in a round space, whose angular diameter does not exceed eight or ten minutes; that is to say, in an area not more than a tenth part of that covered by the moon.

Perhaps it may be thought to savour of the gigantesque to look
individuals of such a group as suns like our own, and their mu-
nces as equal to those which separate our sun from the nearest
: yet, when we consider that their united lustre affects the eye
s impression of light than a star of the fourth magnitude, (for
t of these clusters is barely visible to the naked eye,) the idea
us compelled to form of their distance from us may prepare us
t any estimate of their dimensions. At all events, we can
k upon a group thus insulated, thus *in seipso totus, teres, atque*
as not forming a system of a peculiar and definite character.
nd figure clearly indicates the existence of some general bond
n the nature of an attractive force; and, in many of them, there
lent acceleration in the rate of condensation as we approach the
hich is not referable to a merely uniform distribution of equidis-
through a globular space, but marks an intrinsic *density* in their
ggregation, greater in the centre than at the surface of the mass.
icult to form any conception of the dynamical state of such a
On the one hand, without a rotatory motion and a centrifugal
s hardly possible not to regard them as in a state of progressive

an indefinite succession of ages. Supposing their motions, therefore, to be so adjusted at any one moment as that the orbits should not intersect each other, and so that the magnitude of each star, and the sphere of its more intense attraction, should bear but a small proportion to the distance separating the individuals, such a system, it is obvious, might subsist, and realize, in great measure, that abstract and ideal harmony, which Newton, in the 89th Proposition of the First Book of the *Principia*, has shown to characterize a law of force directly as the distance.[1]

(867.) The following are the places, for 1830, of the principal of these remarkable objects, as specimens of their class: —

R. A.			N. P. D.		R. A.			N. P. D.		R. A.			N. P. D.	
h.	m.	s.	°	′	h.	m.	s.	°	′	h.	m.	s.	°	′
0	16	25	163	2	15	9	56	87	16	17	26	51	143	34
9	8	33	154	10	15	34	56	127	13	17	28	42	93	8
12	47	41	159	57	16	6	55	112	33	11	26	4	114	2
13	4	30	70	55	16	23	2	102	40	18.	55	49	150	14
13	16	38	136	35	16	35	37	53	13	21	21	43	78	34
13	34	10	60	46	16	50	24	119	51	21	24	40	91	34

Of these, by far the most conspicuous and remarkable is ω *Centauri* the fifth of the list in order of Right Ascension. It is visible to the naked eye as a dim round cometic object about equal to a star 4·5 m., though probably if concentered in a single point, the impression on the eye would be much greater. Viewed in a powerful telescope it appears as a globe of fully 20′ in diameter, very gradually increasing in brightness to the centre, and composed of innumerable stars of the 13th and 15th magnitudes (the former probably being two or more of the latter closely juxtaposed). The 11th in order of the list (R. A. $16^h\ 35^m$) is also visible to the naked eye in *very* fine nights, between η and ζ Herculis, and is a superb object in a large telescope. Both were discovered by Halley, the former in 1677, and the latter in 1714.

(868.) It is to Sir William Herschel that we owe the most complete analysis of the great variety of those objects which are generally classed under the common head of Nebulæ, but which have been separated by him into — 1st. Clusters of stars, in which the stars are clearly distinguishable; and these, again, into globular and irregular clusters; 2d. Resolvable nebulæ, or such as excite a suspicion that they consist of stars, and which any increase of the optical power of the telescope may be expected to resolve into distinct stars; 3d. Nebulæ, properly so called, in which there is no appearance whatever of stars; which, again, have been subdivided into subordinate uses, according to their brightness and size;

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[1] See also *Quarterly Review*, No. 94, p. 540.

tary nebulæ; 5th. Stellar nebulæ; and, 6th. Nebulous stars.
power of his telescope disclosed the existence of an immense
these objects before unknown, and showed them to be distri-
the heavens, not by any means uniformly, but with a marked
to a certain district, extending over the northern pole of the
cle, and occupying the constellations Leo, Leo Minor, the body,
nd legs of Ursa Major, Canes Venatici, Coma Berenices, the
eg of Bootes, and the head, wing, and shoulders of Virgo. In
occupying about one-eighth of the whole surface of the sphere,
the entire nebulous contents of the heavens are congregated.
ner hand, they are very sparingly scattered over the constella-
, Taurus, the head and shoulders of Orion, Auriga, Perseus,
lalus, Draco, Hercules, the northern part of Serpentarius, the
ens, that of Aquila, and the whole of Lyra. The hours 3, 4,
17, 18, of right ascension in the northern hemisphere are sin-
r, and, on the other hand, the hours 10, 11, and 12 (but espe-
extraordinarily rich in these objects. In the southern hemi-
nuch greater uniformity of distribution prevails, and with
of two very remarkable centres of accumulation, called the
clouds (of which more presently), there is no very decided

7th magnitude downwards, eight of the more conspicuous of which are coloured with various shades of red, green, and blue, so as to give to the whole the appearance of a rich piece of jewellery.

(870.) Resolvable nebulæ can, of course, only be considered as clusters either too remote, or consisting of stars intrinsically too faint to affect us by their individual light, unless where two or three happen to be close enough to make a joint impression, and give the idea of a point brighter than the rest. They are almost universally round or oval — their loose appendages, and irregularities of form, being as it were extinguished by the distance, and the only general figure of the more condensed parts being discernible. It is under the appearance of objects of this character that all the greater globular clusters exhibit themselves in telescopes of insufficient optical power to show them well; and the conclusion is obvious, that those which the most powerful can barely render resolvable, and even those which, with such powers as are usually applied, show no sign of being composed of stars, would be completely resolved by a further increase of optical power. In fact, this probability has almost been converted into a certainty by the magnificent reflecting telescope constructed by Lord Rosse, of six feet in aperture, which has resolved or rendered resolvable multitudes of nebulæ which had resisted all inferior powers. The sublimity of the spectacle afforded by that instrument of some of the larger globular and other clusters enumerated in the list given in Art. 867, is declared by all who have witnessed it to be such as no words can express.

(871.) Although, therefore, nebulæ do exist, which even in this powerful telescope appear *as* nebulæ, without any sign of resolution, it may very reasonably be doubted whether there be really any essential physical distinction between nebulæ and clusters of stars, at least in the nature of the matter of which they consist, and whether the distinction between such nebulæ as are easily resolved, barely resolvable with excellent telescopes, and altogether irresolvable with the best, be any thing else than one of degree, arising merely from the excessive minuteness and multitude of the stars, of which the latter, as compared with the former, consist. The first impression which Halley, and other early discoverers of nebulous objects received from their peculiar aspect, so different from the keen, concentrated light of mere stars, was that of a phosphorescent vapour (like the matter of a comet's tail) or a gaseous and (so to speak) elementary form of luminous sidereal matter.[1] Admitting the existence of such a medium, dispersed in some cases irregularly through vast regions in space, in others confined to narrower and more definite limits, Sir W. Herschel was led to speculate on its gradual subsidence and condensation by the effect of its

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[1] Halley, Phil. Trans., xxix. p. 390.

y, into more or less regular spherical or spheroidal forms,
they must in that case be) towards the centre. Assuming that
ress of this subsidence, local centres of condensation, subordi-
general tendency, would not be wanting, he conceived that in
olid nuclei might arise, whose local gravitation still further
and so absorbing the nebulous matter, each in its immediate
ood, might ultimately become stars, and the whole nebulæ
on the state of a cluster of stars. Among the multitude of
ealed by his telescopes, every stage of this process might be
as displayed to our eyes, and in every modification of form to
eneral principle might be conceived to apply. The more or
ed state of a nebula towards its segregation into discrete stars
e stars themselves towards a denser state of aggregation round
nucleus, would thus be in some sort an indication of age.
there any variety of aspect which nebulæ offer, which stands at
adiction to this view. Even though we should feel ourselves
to reject the idea of a gaseous or vaporous "nebulous matter,"
e or none of its force. Subsidence, and the central aggrega-
uent on subsidence, may go on quite as well among a multi-

market-place, to which a large proportion of the inhabitants of a town frequently or occasionally resort, implies the permanent residence of each individual within its area. It is a fact that clusters thus centrally crowded do exist, and therefore the conditions of their existence must be dynamically possible, and in what has been said we may at least perceive some glimpses of the manner in which they are so. The actual intervals between the stars, even in the most crowded parts of a resolved nebula, to be seen at all by us, must be enormous. Ages, which to us may well appear indefinite, may easily be conceived to pass without a single instance of collision, in the nature of a catastrophe. Such may have gradually become rarer as the system has emerged from what must be considered its chaotic state, till at length, in the fulness of time, and under the pre-arranging guidance of that DESIGN which pervades universal nature, each individual may have taken up such a course as to annul the possibility of further destructive interference.

(873.) But to return from the regions of speculation to the description of facts. Next in regularity of form to the globular clusters, whose consideration has led us into this digression, are elliptic nebulæ, more or less elongated. And of these it may be generally remarked, as a fact undoubtedly connected in some very intimate manner with the dynamical conditions of their subsistence, that such nebulæ are, for the most part, beyond comparison more difficult of resolution than those of globular form. They are of all degrees of excentricity, from moderately oval forms to ellipses so elongated as to be almost linear, which are, no doubt, edge-views of very flat ellipsoids. In all of them the density increases towards the centre; and as a general law it may be remarked that, so far as we can judge from their telescopic appearance, their internal strata approach more nearly to the spherical form than their external. Their resolvability, too, is greater in the central parts, whether owing to a real superiority of size in the central stars or to the greater frequency of cases of close juxta-position of individuals, so that two or three united appear as one. In some the condensation is slight and gradual, in others great and sudden: so sudden, indeed, as to offer the appearance of a dull and blotted star, standing in the midst of a faint, nearly equable elliptic nebulosity, of which two remarkable specimens occur in R. A. $12^h\ 10^m\ 33^s$, N. P. D. 41° 46′, and in $13^h\ 27^m\ 28^s$, 119° 0′ (1830).

(874.) The largest and finest specimens of elliptic nebulæ which the heavens afford are that in the girdle of Andromeda (near the star ν of that constellation) and that discovered in 1783, by Miss Carolina Herschel, in R. A. $0^h\ 39^m\ 12^s$, N. P. D. 116° 13′. The nebula in Andromeda (Plate II. fig. 3.) is visible to the naked eye, and is continually mistaken

t by those unacquainted with the heavens. Simon Marius,
d it in 1612 (though it appears also to have been seen and
s oval, in 995), describes its appearance as that of a candle
ough horn, and the resemblance is not inapt. Its form, as seen
linary telescopes, is a pretty long oval, increasing by insensible
of brightness, at first very gradually, but at last more rapidly,
tral point, which, though very much brighter than the rest, is
ot a star, but nebula of the same general character with the
te of extreme condensation. Casual stars are scattered over
a reflector of 18 inches in diameter, there is nothing to excite
on of its consisting of stars. Examined with instruments of
fining power, however, the evidence of its resolvability into
be regarded as decisive. Mr. G. P. Bond, assistant at the
of Cambridge, U. S., describes and figures it as extending
in length, and upwards of a degree in breadth (so as to include
smaller adjacent nebulæ), of a form, generally speaking, oval,
considerably protuberant irregularity at its north following ex-
ry suddenly condensed at the nucleus almost to the semblance
and though not itself clearly resolved, yet thickly sown over
minute stars, so numerous as to allow of 200 being counted

(1830). The nebulæ in $12^h\ 27^m\ 3^s$, 63° 5′, and $12^h\ 31^m\ 11^s$, 100° 40′ present analogous features.

(875.) Annular nebulæ also exist, but are among the rarest objects in the heavens. The most conspicuous of this class is to be found almost exactly half way between β and γ Lyræ, and may be seen with a telescope of moderate power. It is small and particularly well defined, so as to have more the appearance of a flat oval solid ring than of a nebula. The axes of the ellipse are to each other in the proportion of about 4 to 5, and the opening occupies about half or rather more than half the diameter. The central vacuity is not quite dark, but is filled in with faint nebula, like a gauze stretched over a hoop. The powerful telescopes of Lord Rosse resolve this object into excessively minute stars, and show filaments of stars adhering to its edges.[1]

(876.) PLANETARY NEBULÆ are very extraordinary objects. They have, as their name imports, a near, in some instances, a perfect resemblance to planets, presenting discs round, or slightly oval, in some quite sharply terminated, in others a little hazy or softened at the borders. Their light is in some perfectly equable, in others mottled and of a very peculiar *texture*, as if curdled. They are comparatively rare objects, not above four or five and twenty having been hitherto observed, and of these nearly three-fourths are situated in the southern hemisphere. Being very interesting objects, we subjoin a list of the most remarkable.[2] Among these may be more particularly specified the sixth in order, situated in the Cross. Its light is about equal to that of a star of the 6·7 magnitude, its diameter about 12″, its disc circular or very slightly elliptic, and with a clear, sharp, well-defined outline, having exactly the appearance of a planet with the exception only of its colour, which is a fine and full blue verging somewhat upon green. And it is not a little remarkable that this phænomenon of a blue colour, which is so rare among *stars* (except when in the immediate proximity of yellow stars) occurs, though less strikingly, in three other objects of this class, viz. in No. 4, whose colour is sky-blue,

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[1] The places of the annular nebulæ, at present known (for 1830) are,

	R. A.	N. P. D.		R. A.	N. P. D.
1.	$17^h\ 10^m\ 39^s$	128° 18′	3.	$18^h\ 47^m\ 13^s$	57° 11′
2.	17 19 2	113	4.	20 9 33	59 57

[2] Places for 1830 of twelve of the most remarkable planetary nebulæ.

R. A.	N. P. D.	R. A.	N. P. D.	R. A.	N. P. D.
h. m. s.	° ′	h. m. s.	° ′	h. m. s.	° ′
1. 7 34 2	104 20	5. 11 4 49	34 4	9. 19 34 21	104 33
2. 9 16 39	147 35	6. 11 41 56	146 14	10. 19 40 19	39 54
3. 9 59 52	129 36	7. 15 5 18	135 1	11. 20 54 53	102 2
4. 10 16 36	107 47	8. 19 10 9	83 46	12. 23 17 44	48 24

s. 11 and 12, where the tint, though paler, is still evident.
9, and 12, are also exceedingly characteristic objects of this
. 3, 5, and 11 (the latter in the parallel of ν Aquarii, and
receding that star), are considerably elliptic, and (respectively)
30″ and 15″ in diameter. On the disc of No 3, and very
he centre of the ellipse, is a star 9^m, and the texture of its light,
ty, or as if formed of fine dust, clearly indicates its resolvability
The largest of these objects is No. 5, situated somewhat south
llel of β Ursæ Majoris and about 12^m following that star. Its
ameter is 2′ 40″, which, supposing it placed at a distance from
e than that of 61 Cygni, would imply a linear one seven times
n that of the orbit of Neptune. The light of this stupendous
rfectly equable (except just at the edge where it is slightly
nd of considerable brightness. Such an appearance would not
d by a globular space uniformly filled with stars or luminous
ch structure would necessarily give rise to an apparent increase
ss towards the centre in proportion to the thickness traversed
al ray. We might, therefore, be induced to conclude its real
to be either that of a hollow spherical shell or of a flat disc,
o us (by a highly improbable coincidence) in a plane precisely

assiduous application of the immense optical powers recently brought to bear on the heavens, will probably remove some portion of the mystery which at present hangs about these enigmatical objects.

(878.) Double nebulæ occasionally occur—and when such is the case, the constituents most commonly belong to the class of spherical nebulæ, and are in some instances undoubtedly globular clusters. All the varieties of double stars, in fact, as to distance, position, and relative brightness, have their counterparts in double nebulæ; besides which the varieties of form and gradation of light in the latter afford room for combinations peculiar to this class of objects. Though the conclusive evidence of observed relative motion be yet wanting, and though from the vast scale on which such systems are constructed, and the probable extreme slowness of the angular motion, it may continue for ages to be so, yet it is impossible, when we cast our eyes upon such objects, or on the figures which have been given of them,[1] to doubt their physical connexion. The argument drawn from the comparative rarity of the objects in proportion to the whole extent of the heavens, so cogent in the case of the double stars, is infinitely more so in that of the double nebulæ. Nothing more magnificent can be presented to our consideration, than such combinations. Their stupendous scale, the multitude of individuals they involve, the perfect symmetry and regularity which many of them present, the utter disregard of complication in thus heaping together system upon system, and construction upon construction, leave us lost in wonder and admiration at the evidence they afford of infinite power and unfathomable design.

(879.) Nebulæ of regular forms often stand in marked and symmetrical relation to stars, both single and double. Thus we are occasionally presented with the beautiful and striking phænomenon of a sharp and brilliant star concentrically surrounded by a perfectly circular disc or atmosphere of faint light, in some cases dying away insensibly on all sides, in others almost suddenly terminated. These are *Nebulous Stars.* Fine examples of this kind are the 45th and 69th nebulæ of Sir Wm. Herschel's fourth class[2] (R. A. $7^h\ 19^m\ 8^s$, N. P. D. 68° 45′, and $3^h\ 58^m\ 36^s$,

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less then could it do so if spread over a surface many million times exceeding in angular area the apparent disc of the central sun itself. (See Annuaire du Bureau des Longitudes, 1842, p. 409, 410, 411.) M. Arago *is expressly contending for reflected* light. If the envelope be self-luminous, his reasoning is perfectly sound.

[1] Phil. Trans., 1833. Plate vii.

[2] The classes here referred to are not the species described in Art. 868, but lists of nebulæ, eight in number, arranged according to brightness, size, density of clustering, &c., in one or other of which all nebulæ were originally classed by him. Class I. contains "Bright nebulæ;" II. "Faint do.;" III. "Very faint do.;" IV. "Planetary

n which stars of the 8th magnitude are surrounded by photo-
the kind above described respectively of 12″ and 25″ in dia-
mong stars of larger magnitudes, 55 Andromedæ and 8 Canum
m may be named as exhibiting the same phænomenon with
ney, but perhaps with less perfect regularity.

he connexion of nebulæ with double stars is in many instances
remarkable. Thus in R. A. $18^h\ 7^m\ 1^s$, N. P. D. 109° 56′,
lliptic nebula having its longer axis about 50″, in which, sym-
placed, and rather nearer the vertices than the foci of the ellipse,
ual individuals of a double star, each of the 10th magnitude.
r combination noticed by M. Struve (in R. A. $18^h\ 25^m$, N. P. D.
e stars are unequal and situated precisely at the two extremities
or axis. In R. A. $13^h\ 47^m\ 33^s$, N. P. D. 129° 9′, an oval
2′ in diameter has very near its centre a close double star, the
of which, slightly unequal, and about the 9·10 magnitude, are
than 2″ asunder. The nucleus of Messier's 64th nebula is
suspected" to be a close double star—and several other instances
ited.

Among the nebulæ which, though deviating more from symsurrounded by a ring at a considerable distance from the globe, very unequal in brightness in its different parts, and subdivided through about two-fifths of its circumference as if into two laminæ, one of which appears as if turned up towards the eye out of the plane of the rest. Near it (at about a radius of the ring distant) is a small bright round nebula. Viewed through the 6-feet reflector of Lord Rosse the aspect is much altered. The interior, or what appeared the upturned portion of the ring, assumes the aspect of a nebulous coil or convolution tending in a spiral form towards the centre, and a general tendency to a spiroid arrangement of the streaks of nebula connecting the ring and central mass which this power brings into view, becomes apparent, and forms a very striking feature. The outlying nebula is also perceived to be connected by a narrow, curved band of nebulous light with the ring, and the whole, if not clearly resolved into stars, has a "resolvable" character which evidently indicates its composition.[1]

(883.) We come now to a class of nebulæ of totally different character. They are of a very great extent, utterly devoid of all symmetry of form, — on the contrary, irregular and capricious in their shapes and convolutions to a most extraordinary degree, and no less so in the distribution of their light. No two of them can be said to present any similarity of figure or aspect, but they have one important character in common. They are all situated in, or very near, the borders of the Milky Way. The most remote from it is that in the sword handle of Orion, which being 20° from the galactic circle, and 15° from the visible border of the Via Lactea, might seem to form an exception, though not a striking one. But this very situation may be adduced as a corroboration of the general view which this principle of localization suggests. For the place in question is situated in the prolongation of that faint offset of the Milky Way which we traced (Art. 787.) from α and ι Persei towards Aldebaran and the Hyades, and in the zone of Great Stars noticed in Art. 785. as an appendage of, and probably bearing relation to that stratum.

(884.) From this it would appear to follow, almost as a matter of course, that they must be regarded as outlying, very distant, and as it were detached fragments of the great stratum of the Galaxy, and this view of the subject is strengthened when we find on mapping down their places that they may all be grouped in four great masses or nebulous regions, — that of Orion, of Argo, of Sagittarius, and of Cygnus. And thus, inductively, we may gather some information respecting the struc-

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[1] This description is from the recollection of a sketch exhibited by his Lordship at the British Association. Every astronomer must long for the publication of his own account of the wonders disclosed by this magnificent instrument.

orm of the Galaxy itself, which, could we view it as a whole,
ance such as that which separates us from these objects, would
bly present itself under an aspect quite as complicated and

The great nebula surrounding the stars marked θ 1 in the sword
Orion was discovered by Huyghens in 1656, and has been re-
gured and described by astronomers since that time. Its
varies greatly (as that of all nebulous objects does) with the
al power applied, so that it is difficult to recognize in represent-
le with inferior telescopes, even principal features, to say
subordinate details. Until this became well understood, it
ed to have changed very materially, both in form and extent,
interval elapsed since its first discovery. No doubt, however,
ns that these supposed changes have originated partly from the
e-mentioned, partly from the difficulty of correctly drawing,
more, engraving such objects, and partly from a want of suffi-
n the earlier delineators themselves in faithfully copying that
y really did see. Our figure (Plate IV., *fig.* 1,) is reduced
ger one made under very favourable circumstances, from draw-

achromatic at Cambridge, U. S., is evidently perceived to consist of clustering stars. There can therefore be little doubt as to the whole consisting of stars, too minute to be discerned individually even with these powerful aids, but which become visible as points of light when closely adjacent in the more crowded parts in the mode already more than once suggested.

(886.) The nebula is not confined to the limits of our figure. Northward of θ about 33′, and nearly on the same meridian are two stars marked C 1 and C 2 Orionis, involved in a bright and branching nebula of very singular form, and south of it is the star ι Orionis, which is also involved in strong nebula. Careful examination with powerful telescopes has traced out a continuity of nebulous light between the great nebula and both these objects, and there can be little doubt that the nebulous region extends northwards, as far as ε in the belt of Orion, which is involved in strong nebulosity, as well as several smaller stars in the immediate neighbourhood. Professor Bond has given a beautiful figure of the great nebula in Trans. American Acad. of Arts and Sciences, new series, vol. iii.

(887.) The remarkable variation in lustre of the bright star η in Argo, has been already mentioned. This star is situated in the most condensed region of a very extensive nebula or congeries of nebular masses, streaks and branches, a portion of which is represented in fig. 2, Plate IV. The whole nebula is spread over an area of fully a square degree in extent, of which that included in the figure occupies about one-fourth, that is to say, 28′ in polar distance, and 32′ of arc in R. A., the portion not included being, though fainter, even more capriciously contorted than that here depicted, in which it should be observed that the preceding side is towards the right hand, and the southern uppermost. Viewed with an 18-inch reflector, no part of this strange object shows any sign of resolution into stars, nor in the brightest and most condensed portion adjacent to the singular oval vacancy in the middle of the figure is there any of that curdled appearance, or that tendency to break up into bright knots with intervening darker portions which characterize the nebula of Orion, and indicate its resolvability. The whole is situated in a very rich and brilliant part of the Milky Way, so thickly strewed with stars (omitted in the figure), that in the area occupied by the nebula, not less than 1200 have been actually counted, and their places in R. A. and P. D. determined. Yet it is obvious that these have no connexion whatever with the nebula, being, in fact, only a simple continuation over it of the general ground of the galaxy, which on an average of two hours in Right Ascension in this period of its course, contains no less than 3138 stars to the

ree, all, however, distinct, and (except where the object in
situated) seen projected on a perfectly dark heaven, without
nce of intermixed nebulosity. The conclusion can hardly be
t in looking at it we see through, and beyond the Milky Way,
space, through a starless region, disconnecting it altogether
stem. "It is not easy for language to convey a full impres-
e beauty and sublimity of the spectacle which this nebula
enters the field of view of a telescope fixed in Right Ascen-
e diurnal motion, ushered in as it is by so glorious and innu-
rocession of stars, to which it forms a sort of climax," and in
he heavens otherwise full of interest. One other bright and
kably formed nebula of considerable magnitude precedes it
he same parallel, but without any traceable connexion between

The nebulous group of Sagittarius consists of several conspicuous
very extraordinary forms, by no means easy to give an idea of
scription. One of them (*h*, 1991[2]) is singularly trifid, con-
hree bright and irregularly formed nebulous masses, graduating
sibly externally, but coming up to a great intensity of light at
r edges, where they enclose and surround a sort of three-forked

nebula does not, as in the region of Orion, show any tendency to congregate about the stars.

(890.) The 19th nebula of Messier's Catalogue, though some degrees remote from the others, evidently belongs to this group. Its form is very remarkable, consisting of two loops like capital Greek Omegas, the one bright, the other exceedingly faint, connected at their bases by a broad and very bright band of nebula, insulated within which by a narrow comparatively obscure border, stands a bright, resolvable knot, or what is probably a cluster of exceedingly minute stars. A very faint round nebula stands in connexion with the upper or convex portion of the brighter loop.

(891.) The nebulous group of Cygnus consists of several large and irregular nebulæ, one of which passes through the double star *k* Cygni, as a long, crooked narrow streak, forking out in two or three places. The others,[1] observed in the first instance by Sir W. Herschel and by the author of this work as separate nebulæ, have been traced into connexion by Mr. Mason, and shown to form part of a curious and intricate nebulous system, consisting, 1st, of a long, narrow, curved, and forked streak, and 2dly, of a cellular effusion of great extent, in which the nebula occurs intermixed with, and adhering to stars around the borders of the cells, while their interior is free from nebula, and almost so from stars.

(892.) The Magellanic clouds, or the nubeculæ (major and minor,) as they are called in the celestial maps and charts, are, as their name imports, two nebulous or cloudy masses of light, conspicuously visible to the naked eye, in the southern hemisphere, in the appearance and brightness of their light not unlike portions of the Milky Way of the same apparent size. They are, generally speaking, round, or somewhat oval, and the larger, which deviates most from the circular form, exhibits the appearance of an axis of light, very ill defined, and by no means strongly distinguished from the general mass, which seems to open out at its extremities into somewhat oval sweeps, constituting the preceding and following portions of its circumference. A small patch, visibly brighter than the general light around, in its following part, indicates to the naked eye the situation of a very remarkable nebula (No. 30 Doradûs of Bode's catalogue,) of which more hereafter. The greater nubecula is situated between the meridians of $4^h\ 40^m$ and $6^h\ 0^m$ and the parallels of 156° and 162° of N.P.D., and occupies an area of about 42 square degrees. The lesser, between the meridians[2] $0^h\ 28^m$ and $1^h\ 15^m$ and the parallels of 162° and 165° N.P.D. covers about ten square degrees. Their degree of brightness may be judged of from the effect of strong moonlight, which totally obliterates the lesser, but not quite the greater.

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[1] R.A. $20^h\ 49^m\ 20^s$, N.P.D. 58° 27′.

[2] It is laid down nearly an hour wrong in all the celestial charts and globes.

When examined through powerful telescopes, the constitution
eculæ, and especially of the nubecula major, is found to be of
complexity. The general ground of both consists of large
patches of nebulosity in every stage of resolution, from light,
with 18 inches of reflecting aperture, up to perfectly separated
he Milky Way, and clustering groups sufficiently insulated and
to come under the designation of irregular, and in some cases
clusters. But besides those, there are also nebulæ in abun-
h regular and irregular; globular clusters in every state of
on; and objects of a nebulous character quite peculiar, and
no analogue in any other region of the heavens. Such is the
on of these objects, that in the area occupied by the nubecula
fewer than 278 nebulæ and clusters have been enumerated,
or 60 outliers, which (considering the general barrenness in
ts of the immediate neighbourhood) ought certainly to be
s its appendages, being about 6½ per square degree, which very
the average of any other, even the most crowded parts of the
eavens. In the nubecula minor, the concentration of such
ess, though still very striking, 37 having been observed within

10th to such as by their multitude and minuteness constitute irresolvable nebulosity, extending over tracts of many square degrees. Were there but one such object, it might be maintained without utter improbability that its apparent sphericity is only an effect of foreshortening, and that in reality a much greater proportional difference of distance between its nearer and more remote parts exists. But such an adjustment, improbable enough in one case, must be rejected as too much so for fair argument in two. It must, therefore, be taken as a demonstrated fact, that stars of the 7th or 8th magnitude and irresolvable nebula may co-exist within limits of distance not differing in proportion more than as 9 to 10, a conclusion which must inspire some degree of caution in admitting, *as certain*, many of the consequences which have been rather strongly dwelt upon in the foregoing pages.

(895.) Immediately preceding the centre of the nubecula minor, and undoubtedly belonging to the same group, occurs the superb globular cluster, No. 47, Toucani of Bode, very visible to the naked eye, and one of the finest objects of this kind in the heavens. It consists of a very condensed, spherical mass of stars, *of a pale rose-colour*, concentrically enclosed in a much less condensed globe of white ones, 15′ or 20′ in diameter. This is the first in order of the list of such clusters in Art. 867.

(896.) Within the nubecula major, as already mentioned, and faintly visible to the naked eye, is the singular nebula (marked as the star 30 Doradûs in Bode's Catalogue) noticed by Lacaille as resembling the nucleus of a small comet. It occupies about one-500th part of the whole area of the nubecula, and is so satisfactorily represented in plate V., fig. 1, as to render further description superfluous.

(897.) We shall conclude this chapter by the mention of two phænomena, which seem to indicate the existence of some slight degree of nebulosity about the sun itself, and even to place it in the list of nebulous stars. The first is that called the zodiacal light, which may be seen any very clear evening, soon after sunset, about the months of March, April, and May, or at the opposite seasons before sunrise, as a cone or lenticularly-shaped light, extending from the horizon obliquely upwards, and following generally the course of the ecliptic, or rather that of the sun's equator. The apparent angular distance of its vertex from the sun varies, according to circumstances, from 40° to 90°, and the breadth of its base perpendicular to its axis from 8° to 30°. It is extremely faint and ill defined, at least in this climate, though better seen in tropical regions, but cannot be mistaken for any atmospheric meteor or aurora borealis. It is manifestly in the nature of a lenticularly-formed envelope, surrounding

nd extending beyond the orbits of Mercury and Venus, and
haps quite, attaining that of the earth, since its vertex has been
90° from the sun's place in a great circle. It may be conjec-
e no other than the denser part of that medium, which, we
reason to believe, resists the motion of comets; loaded, per-
the actual materials of the tails of millions of those bodies, of
y have been stripped in their successive perihelion passages
). An *atmosphere* of the sun, in any proper sense of the word,
e, since the existence of a gaseous envelope propagating pres-
part to part; subject to mutual friction in its strata, and there-
g in the same or nearly the same time with the central body,
h dimensions and ellipticity, is utterly incompatible with dyna-
. If its particles have inertia, they must necessarily stand with
the sun in the relation of separate and independent minute
ch having its own orbit, plane of motion, and periodic time.
mass being almost nothing compared to that of the sun, mutual
ion is out of the question, though *collisions* among such as may
other's paths may operate in the course of indefinite ages to
sidence of at least some portion of it into the body of the sun

into fragments of a large fiery globe over a region of twenty or thirty square miles around the town of L'Aigle, in Normandy. The fact occurred at mid-day, and the circumstances were officially verified by a commission of the French government.[1] These, and innumerable other instances,[2] fully establish the general fact; and after vain attempts to account for it by volcanic projection, either from the earth or the moon, the planetary nature of these bodies seems at length to be almost generally admitted. The heat which they possess when fallen, the igneous phænomena which accompany them, their explosion on arriving within the denser regions of our atmosphere, &c., are all sufficiently accounted for on physical principles, by the condensation of the air before them in consequence of their enormous velocity, and by the relations of air in a highly attenuated state to heat.[3]

(899.) Besides stony and metallic masses, however, it is probable that bodies of very different natures, or at least states of aggregation, are thus circulating round the sun. Shooting stars, often followed by long trains of light, and those great fiery globes, of more rare, but not *very* uncommon occurrence, which are seen traversing the upper regions of our atmosphere, sometimes leaving trains behind them, remaining for many minutes, sometimes bursting with a loud explosion, sometimes becoming quietly extinct, may not unreasonably be presumed to be bodies extraneous to our planet, which only become visible when in the act of grazing the confines of our atmosphere. Among the last mentioned meteors are some which can hardly be supposed solid masses. The remarkable meteor of Aug. 18, 1783, traversed the whole of Europe, from Shetland to Rome, with a velocity of about 30 miles per second, at a height of 50 miles from the surface of the earth, with a light greatly surpassing that of the full moon, and a real diameter of fully half a mile. Yet with these vast dimensions, it changed its form visibly, and at length quietly separated into several distinct bodies, accompanying each other in parallel courses, and each followed by a tail or train.

(900.) There are circumstances in the history of shooting stars, which very strongly corroborate the idea of their extraneous or *cosmical* origin, and their circulation round the sun in definite orbits. On several occasions they have been observed to appear in unusual, and, indeed, astonish-

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[1] See M. Biot's report in Mém. de l'Institut. 1806.

[2] See a list of upwards of 200, published by Chladni, Annales du Bureau des Longitudes de France, 1825.

[3] Edinburgh Review, Jan. 1848, p. 195. It is very remarkable that no new chemical element has been detected in any of the numerous meteorolites which have been subjected to analysis.

rs, so as to convey the idea of a shower of rockets, or of snow-
ng, and brilliantly illuminating the whole heavens for hours
nd that not in one locality, but over whole continents and
even (in one instance) in both hemispheres. Now it is ex-
narkable that, whenever this great display has been exhibited
modern times), it has uniformly happened on the night be-
12th and 13th, or on that between the 13th and 14th of No-
Such cases occurred in 1799, 1823, 1832, 1833, and 1834.
back the records of similar phænomena, it has been ascertained,
that more often those identical nights, but sometimes those
y adjacent, have been, time out of mind, habitually signalized
hibitions. Another annually recurring epoch, in which, though
lliant, the display of meteors is more *certain* (for that of No-
ften interrupted for a great many years), is that of the 10th
on which night, and on the 9th and 11th, numerous, large,
shooting stars, with trains, are almost sure to be seen. Other
eriodic recurrence, less marked than the above, have also been
n extent established.

It is impossible to attribute such a recurrence of identical dates to its east and west points may be at the time of observation, the paths of the meteors all appear to diverge from that star. On the 9th—11th of August the geometrical fact is the same, the apex only differing; B Camelopardali being for that epoch the point of divergence. As we need not suppose the meteoric ring coincident in its plane with the ecliptic, and as for a *ring* of meteors we may substitute an elliptic annulus of any reasonable excentricity, so that both the velocity and direction of each meteor may differ to any extent from the earth's, there is nothing in the great and obvious difference *in latitude* of these apices at all militating against the conclusion.

(903.) If the meteors be uniformly distributed in such a ring or elliptic annulus, the earth's encounter with them in every revolution will be certain, if it occur once. But if the ring be broken, if it be a succession of groups revolving in an ellipse in a period *not* identical with that of the earth, years may pass without a rencontre; and when such happen, they may differ to any extent in their intensity of character, according as richer or poorer groups have been encountered.

(904.) No other plausible explanation of these highly characteristic features (the annual periodicity, and divergence from a common apex, *always alike for each respective epoch*) have been even attempted, and accordingly the opinion is generally gaining ground among astronomers that shooting stars belong to their department of science, and great interest is excited in their observation and the further development of their laws. The most connected and systematic series of observations of them, having for their object to trace out their *relative* paths with respect to the earth, are those of Benzenberg and Brandes, who, by noting the instants and apparent places of appearance and extinction, as well as the precise apparent paths among the stars, of individual meteors, from the extremities of a measured base line nearly 50,000 feet in length, were led to conclude that their heights at the instant of their appearance and disappearance vary from 16 miles to 140, and their relative velocities from 18 to 36 miles per second, velocities so great as clearly to indicate an independent planetary circulation round the sun. [A very remarkable meteor or bolide, which appeared on the 19th August, 1847, was observed at Dieppe and at Paris, with sufficient precision to admit of calculation of the elements of its orbit in absolute space. This calculation has been performed by M. Petit, director of the observatory of Toulouse, and the following *hyperbolic* elements of its orbit *round the sun* are stated by him (Astr. Nachr. 701) as its result; viz., Semimajor axis = — 0·3240083; excentricity = 3·95130; perihelion distance = 0·95626; inclination to plane of the earth's equator, 18° 20′ 18″; ascending node on the same

34′ 48″; motion direct. According to this calculation, the
d have occupied no less than 37340 years in travelling from the
the nearest fixed star supposed to have a parallax of 1″.]

It is by no means inconceivable that the earth approaching to
fer but little from it in direction and velocity, may have at-
ny of them to it as permanent satellites, and of these there *may*
large, and of such texture and solidity, as to shine by reflected
become visible (such, at least, as are very near the earth) for a
ent, suffering extinction by plunging into the earth's shadow;
ords undergoing total eclipse. Sir John Lubbock is of opinion
s the case, and has given geometrical formulæ for calculating
nces from observations of this nature.[1] The observations of M.
d lead us to believe in the existence of at least one such body,
ound the earth, as a satellite, in about 3 hours 20 minutes, and
t a distance equal to 2·513 radii of the earth from its centre,
iles above its surface.[2]

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[1] Phil. Mag. Lond. Ed. Dub. 1848, p. 80.

[2] Comptes Rendus, Oct. 12, 1846, and Aug. 9, 1847.

PART IV.

OF THE ACCOUNT OF TIME.

CHAPTER XVIII.

NATURAL UNITS OF TIME.—RELATION OF THE SIDEREAL TO THE SOLAR DAY AFFECTED BY PRECESSION.—INCOMMENSURABILITY OF THE DAY AND YEAR.—ITS INCONVENIENCE.—HOW OBVIATED.—THE JULIAN CALENDAR.—IRREGULARITIES AT ITS FIRST INTRODUCTION.—REFORMED BY AUGUSTUS.—GREGORIAN REFORMATION.—SOLAR AND LUNAR CYCLES.—INDICTION.—JULIAN PERIOD.—TABLE OF CHRONOLOGICAL ERAS.—RULES FOR CALCULATING THE DAYS ELAPSED BETWEEN GIVEN DATES.—EQUINOCTIAL TIME.

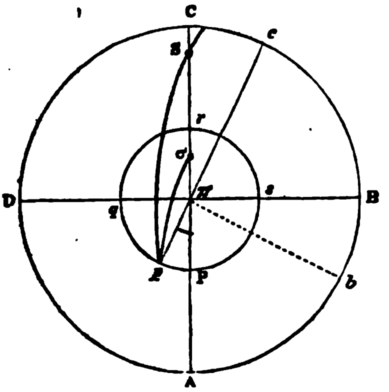
(906.) TIME, like distance, may be measured by comparison with standards of any length, and all that is requisite for ascertaining correctly the length of any interval, is to be able to apply the standard to the interval throughout its whole extent, without overlapping on the one hand, or leaving unmeasured vacancies on the other; to determine, without the possible error of a unit, the number of integer standards which the interval admits of being interposed between its beginning and end; and to estimate precisely the fraction, over and above an integer, which remains when all the possible integers are subtracted.

(907.) But though all standard units of time are equally possible, theoretically speaking, yet all are not, practically, equally convenient. The solar day is a natural interval which the wants and occupations of man in every state of society force upon him, and compel him to adopt as his fundamental unit of time. Its length as estimated from the departure of the sun from a given meridian, and its next return to the same, is subject, it is true, to an annual fluctuation in excess and defect of its mean value, amounting at its maximum to full half a minute. But except for astronomical purposes, this is too small a change to interfere in the slightest degree with its use, or to attract any attention, and the tacit substitu-

mean for its true (or variable) value may be considered as
en made from the earliest ages, by the ignorance of mankind
uch fluctuation existed.

The time occupied by one complete rotation of the earth on its
e mean[1] sidereal day, may be shown, on dynamical principles, to
to no variation from any external cause, and although its dura-
d be shortened by contraction in the dimensions of the globe
h as might arise from the gradual escape of its internal heat,
quent refrigeration and shrinking of the whole mass, yet theory,
hand, has rendered it almost certain that this cause cannot have
y perceptible amount of change during the history of the human
, on the other, the comparison of ancient and modern observa-
ds every corroboration to this conclusion. From such compari-
lace has concluded that the sidereal day has not changed by so
ne hundredth of a second since the time of Hipparchus. The
real day therefore possesses in perfection the essential quality of
unit, that of complete *invariability*. The same is true of the
real year, if estimated upon an average sufficiently large to com-
e minute fluctuations arising from the periodical variations of the

Fig. 112.



ecliptic A B C D. Let S be a star anywhere situated on the ecliptic, or between it and the small circle P *q* *r*. Then if the pole P were at rest, a meridian of the earth setting out from P S C, and revolving in the direction C D, will come again to the star after the exact lapse of one sidereal day, or one rotation of the earth on its axis. But P is not at rest. After the lapse of one such day it will have come into the situation (suppose) *p*, the vernal equinox B having retreated to *b*, and the colure P C having taken up the new position *p* *c*. Now a conical movement impressed on the axis of rotation of a globe already rotating is equivalent to a rotation impressed on the whole globe round the axis of the cone, in addition to that which the globe has and retains round its own independent axis of revolution. Such a new rotation, in transferring P to *p*, being performed round an axis passing through π, will not alter the situation of that point of the globe which has π in its zenith. Hence it follows that *p* π *c* passing through π will be the position taken up by the meridian P π C after the lapse of an exact sidereal day. But this does not pass through S, but falls short of it by the hour-angle π *p* S, which is yet to be described before the meridian comes up to the star. The meridian, then, has lost so much on, or lagged so much behind, the star in the lapse of that interval. The same is true whatever be the arc P *p*. After the lapse of any number of days, the pole being transferred to *p*, the spherical angle π *p* S will measure the total hour-angle which the meridian has lost on the star. Now where S lies any where between C and *r*, this angle continually increases (though not uniformly), attaining 180° when *p* comes to *r*, and still (as will appear by following out the movement beyond *r*) increasing thence till it attains 360° when *p* has completed its circuit.

a whole revolution of the equinoxes, the meridian will have lost revolution upon the star, or in 9448300 sidereal days, will have d the star only 9448299 times: in other words, the length of the ured by the mean of the successive arrivals of any star outside cle $P\,p\,q\,r$ on one and the same meridian is to the absolute time n of the earth on its axis as 9448300 : 9448299, or as 1·00000011

It is otherwise of a star situated *within* this circle, as at σ. For tar the angle $\pi\,p\,\sigma$, expressing the lagging of the meridian, in- o a maximum for some situation of p between q and r, and again to o at r; after which it takes an opposite direction, and lian begins to get in advance of the star, and continues to get more so, till p has attained some point between s and P, where ce is a maximum, and thence decreases again to o when p has l its circuit. For any star so situated, then, the mean of all the stimated through a whole period of the equinoxes is an absolute ay, as if precession had no existence.

If we compare the sun with a star situated in the ecliptic, the year is the mean of all the intervals of its arrival at that star which is the ratio of the solar to the sidereal day actually in use among astronomers.

(912.) It would be well for chronology if mankind would, or could have contented themselves with this one invariable, natural, and convenient standard in their reckoning of time. The ancient Egyptians did so, and by their adoption of an historical and official year of 365 days have afforded the only example of a practical chronology, free from all obscurity or complication. But the return of the seasons, on which depend all the more important arrangements and business of cultivated life, is not conformable to such a multiple of the diurnal unit. Their return is regulated by the *tropical year*, or the interval between two successive arrivals of the sun at the vernal equinox, which, as we have seen (Art. 383), differs from the sidereal year by reason of the motion of the equinoctial points. Now this motion is not absolutely uniform, because the ecliptic, upon which it is estimated, is gradually, though very slowly, changing its situation in space under the disturbing influence of the planets (Art. 640.) And thus arises a variation in the *tropical* year, which is dependent on the place of the equinox (Art. 383.) The *tropical* year is actually about $4{\cdot}21^{s}$ shorter than it was in the time of Hipparchus. This absence of the most essential requisite for a standard, viz. invariability, renders it necessary, since we cannot help employing the tropical year in our reckoning of time, to adopt an arbitrary or artificial value for it, so near the truth, as not to admit of the accumulation of its error for several centuries producing any practical mischief, and thus satisfying the ordinary wants of civil life; while, for scientific purposes, the tropical year, so adopted, is considered only as the representative of a certain number of integer days and a fraction—the day being, in effect, the only standard employed. The case is nearly analogous to the reckoning of value by guineas and shillings, an artificial relation of the two coins being fixed by law, near to, but scarcely ever exactly coincident with, the natural one, determined by the relative market price of gold and silver, of which either the one or the other — whichever is really the most invariable, or the most in use with other nations, — may be assumed as the true theoretical standard of value.

(913.) The other inconvenience of the tropical year as a greater unit is its incommensurability with the lesser. In our measure of space all our subdivisions are into aliquot parts: a yard is three feet, a mile eight furlongs, &c. But a year is no *exact* number of days, nor an integer number with any exact fraction, as one third or one fourth, over and above; but the surplus is an *incommensurable* fraction, composed of hours, minutes, seconds, &c., which produces the same kind of inconvenience in

oning of time that it would do in that of money, if we had
of the value of twenty-one shillings, with odd pence and far-
a fraction of a farthing over. For this, however, there is no
to keep a strict register of the surplus fractions; and, when
nt to a whole day, cast them over into the integer account.
To do this in the simplest and most convenient manner is the
well-adjusted calendar. In the Gregorian calendar, which we
accomplished with as much simplicity and neatness as the
, by carrying a little farther than is done above, the principle
ned or artificial year, and adopting *two* such years, both con-
n exact integer number of days, viz. one of 365 and the other
d laying down a simple and easily remembered rule for the
ich these years shall succeed each other in the civil reckoning
that during the lapse of at least some thousands of years the
integer artificial, or Gregorian, years elapsed shall not differ
ame number of real tropical years by a whole day. By this
, the equinoxes and solstices will always fall on days similarly
d bearing the same name in each Gregorian year; and the
l for ever correspond to the same months, instead of running

fraction, is 365·24224, so the error in the Gregorian rule on 10000 of the present tropical years, is 2·6, or $2^d\ 14^h\ 24^m$; that is to say, less than a day in 3000 years; which is more than sufficient for all human purposes, those of the astronomer excepted, who is in no danger of being led into error from this cause. Even this error is avoided by extending the wording of the Gregorian rule one step farther than its contrivers probably thought it worth while to go, and declaring that years divisible by 4000 should consist of 365 days. This would take off two integer days from the above calculated number, and 2·5 from a larger average; making the sum of days in 100000 Gregorian years, 36524225, which differs only by a single day from 100000 real tropical years, such as they exist at present.

(916.) In the historical dating of events there is no year A. D. 0. The year immediately previous to A. D. 1, is always called B. C. 1. This must always be borne in mind in reckoning chronological and astronomical intervals. The sum of the nominal years B. C. and A. D. must be diminished by 1. Thus, from Jan. 1, B. C. 4713, to Jan. 1, 1582, the years elapsed are not 6295, but 6294.

(917.) As any distance along a high road might, though in a rather inconvenient and roundabout way, be expressed without introducing error by setting up a series of milestones, at intervals of unequal lengths, so that every fourth mile, for instance, should be a yard longer than the rest, or according to any other fixed rule; taking care only to mark the stones so as to leave room for no mistake, and to advertise all travellers of the difference of lengths and their order of succession; so may any interval of time be expressed correctly by stating in what Gregorian years it begins and ends, *and whereabouts in each*. For this statement coupled with the declaratory rule, enables us to say how many integer years are to be reckoned at 365, and how many at 366 days. The latter years are called bissextiles, or leap-years, and the surplus days thus thrown into the reckoning are called *intercalary or leap days.*

(918.) If the Gregorian rule, as above stated, had always and in all countries been known and followed, nothing would be easier than to reckon the number of days elapsed between the present time, and any historical recorded event. But this is not the case; and the history of the calendar, with reference to chronology, or to the calculation of ancient observations, may be compared to that of a clock, going regularly when left to itself, but sometimes forgotten to be wound up; and when wound, sometimes set forward, sometimes backward, either to serve particular purposes and private interests, or to rectify blunders in setting. Such, at least, appears to have been the case with the Roman calendar, in which our own originates, from the time of Numa to that of Julius Cæsar, when the lunar

onths, or 355 days, was augmented at pleasure to correspond
, by which the seasons are determined, by the arbitrary inter-
the priests, and the usurpations of the decemvirs and other
, till the confusion became inextricable. To Julius Cæsar,
Sosigenes, an eminent Alexandrian astronomer and mathema-
we the neat contrivance of the two years of 365 and 366 days,
ertion of one bissextile after three common years. This im-
nge took place in the 45th year before Christ, which he ordered
ce on the 1st of January, *being the day of the new moon im-*
following the winter solstice of the year before. We may judge
into which the reckoning of time had fallen, by the fact, that
the new system it was necessary to enact that the previous
c., should consist of 445 days, a circumstance which obtained
pithet of "the year of confusion."

Had Cæsar lived to carry out into practical effect, as Chief
own reformation, an inconvenience would have been avoided,
he very outset threw the whole matter into confusion. The
is edict establishing the Julian system have not been handed
, but it is probable that they contained some expression equi-

of securing the intercalation as a matter of precedent, made his initial year 45 B. C. a leap year. Whenever, therefore, in the relation of any event, either in ancient history or in modern, previous to the change of style, the time is specified in our modern nomenclature, it is always to be understood as having been identified with the assigned date by threading the mazes (often very tangled and obscure ones) of special and national chronology, and referring the day of its occurrence to its place in the Julian system *so interpreted.*

(920.) Different nations in different ages of the world have of course reckoned their time in different ways, and from different epochs, and it is therefore a matter of great convenience that astronomers and chronologists (as they have agreed on the uniform adoption of the Julian system of years and months) should also agree on an epoch antecedent to them all, to which, as to a fixed point in time, the whole list of chronological eras can be differentially referred. Such an epoch is the noon of the 1st of January, B. C. 4713, which is called the epoch of the Julian period, a cycle of 7980 Julian years, to understand the origin of which we must explain that of three subordinate cycles, from whose combination it takes its rise, by the multiplication together of the numbers of years severally contained in them, viz.:—the Solar and Lunar cycles, and that of the indictions.

(921.) The Solar cycle consists of 28 Julian years, after the lapse of which the same days of the week on the Julian system would always return to the same days of each month throughout the year. For four such years consisting of 1461 days, which is not a multiple of 7, it is evident that the least number of years which will fulfil this condition must be seven times that interval, or 28 years. The place in this cycle for any year A. D., as 1849, is found by adding 9 to the year, and dividing by 28. The remainder is the number sought, 0 being counted as 28.

(922.) The Lunar cycle consists of 19 years or 235 lunations, which differ from 19 Julian years of 365¼ days only by about an hour and a half, so that, supposing the new moon to happen on the first of January, in the first year of the cycle, it will happen on that day (or within a very short time of its beginning or ending) again after a lapse of 19 years, and almost certainly on that day, and within an hour and a half of the same hour of the day, after the lapse of four such cycles, or 76 years; and all the new moons in the interval will run on the same days of the month as in the preceding cycle. This period of 19 years is sometimes called the Metonic cycle, from its discoverer Meton, an Athenian mathematician, a

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ring the rectification of their accumulated error to the fourth, by inserting *there* the intercalary day. For the correction of Roman dates during the fifty-two years between the Julian and Augustan reformations, see Ideler, "Handbuch der Mathematischen und Technischen Chronologie," which we take for our guide throughout this chapter.

uly appreciated by his countrymen, as ensuring the correspon-
een the lunar and solar years, the former of which was followed
eks. Public honours were decreed to him for this discovery,
nce very expressive of the annoyance which a lunar year of
nflicts on a civilized people, to whom a regular and simple
one of the first necessities of life. The cycle of 76 years, a
ovement on the Metonic cycle, was first proposed by Callippus,
efore called the Callippic cycle. To find the place of a given
lunar cycle, (or as it is called the Golden Number,) add 1 to
r of the year A. D., and divide by 19, the remainder (or 19 if
isible) is the Golden Number.

The cycle of the indictions is a period of 15 years used in the
w and in the fiscal organization of the Roman empire, under
and his successors, and thence introduced into legal dates, as
Number, serving to determine Easter, was in to ecclesiastical
find the place of a year in the indiction cycle, add 3 and divide
he remainder (or 15 if 0 remain) is the number of the indic-

f we multiply together the numbers 28, 19, and 15, we get

(925.) Given the year of the Julian period, those of the subordinate cycles are easily determined as above. Conversely, given the years of the solar and lunar cycles, and of the indiction, to determine the year of the Julian period proceed as follows: — Multiply the number of the year in the solar cycle by 4845, in the lunar by 4200, and in the Cycle of the Indictions by 6916, divide the sum of the products by 7980, and the remainder is the year of the Julian period sought.

(926.) The following table contains these intervals for some of the more important historical eras: —

Intervals in Days between the Commencement of the Julian Period, and that of some other remarkable chronological and astronomical Eras.

Names by which the Era is usually cited.	First day current of the era.	Chronolog'l designation of the year.	Current year of the Julian Period.	Interval days.
Julian Epochs.	*Julian Dates.*			
Julian period	Jan. 1.	B. C. 4713	1	0
Creation of the world (Usher)	(Jan. 1.)	4004	710	258,963
Era of the Deluge (Aboulhassan Kushiar)	Feb. 18.	3102	1612	588,466
Ditto Vulgar Computation	(Jan. 1.)	2348	2366	863,817
Era of Abraham (Sir H. Nicholas)	Oct. 1.	2015	2699	985,718
Destruction of Troy, (ditto)	July 12.	1184	3530	1,289,160
Dedication of Solomon's Temple	(May 1.)	1015	3699	1,350,815
Olympiads (mean epoch in general use)	July 1.	776	3938	1,438,171
Building of Rome (Varronian epoch, U. C.)	April 22.	753	3961	1,446,502
Era of Nabonassar	Feb. 26.	747	3967	1,448,638
Metonic cycle (Astronomical epoch)	July 15.	432	4282	1,563,831
Callippic cycle, do. (Biot)	June 28.	330	4384	1,599,608
Philippic era, or era of Philip Aridæus	Nov. 12.	324	4390	1,603,398
Era of the Seleucidæ	Oct. 1.	312	4402	1,607,739
Cæsarean era of Antioch	Sept. 1.	49	4665	1,703,770
Julian reformation of the Calendar	Jan. 1.	45	4669	1,704,987
Spanish era	Jan. 1.	38	4676	1,707,544
Actian era in Rome	Jan. 1.	30	4684	1,710,466
Actian era of Alexandria	Aug. 29.	30	4684	1,710,706
Vulgar or Dionysian era	Jan. 1.	A. D. 1	4714	1,721,424
Era of Diocletian	Aug. 29.	284	4997	1,825,030
Hejira (astronomical epoch, new moon)	July 15.	622	5335	1,948,439
Era of Yezdegird	June 16.	632	5345	1,952,063
Gelalæan era (Sir H. Nicholas)	March 14.	1079	5792	2,115,285
Last day of Old Style (Catholic nations)	Oct. 4.	1582	6295	2,299,160
Last day of Old Style in England)	Sept. 2.	1752	6465	2,361,221
Gregorian Epochs.	*Gregorian Dates.*			
New Style in Catholic nations	Oct. 15.	1582	6295	2,299,161
Ditto in England	Sept. 14.	1752	6465	2,361,222
Commencement of the 19th century, epoch of Bode's catalogue of stars	Jan. 1.	1801	6514	2,378,862
Epoch of the catalogue of stars of the R. Astronomical Society	Jan. 1.	1830	6543	2,389,454
Epoch of the catalogue of the British Association	Jan. 1.	1850	6563	2,396,759

civil epochs of the Metonic cycle, and the Hejira, are each one day later
onomical, the latter being the epochs of the absolute *new moons*, the former
earliest possible visibility of the lunar crescent in a tropical sky. M. Biot
at the solstice and new moon not only coincided on the day here set down
ncement of the Callippic cycle, but that, by a happy coincidence, a *bare*
isted of seeing the crescent moon at Athens *within that day, reckoned from*
idnight.

he determination of the exact interval between any two given
atter of such importance, and, unless methodically performed,
iable to error, that the following rules will not be found out of
the first place it must be remarked, generally, that a date,
a day or year, always expresses the day or year *current* and not
d that the designation of a year by A. D. or B. C. is to be re-
he *name* of that year, and *not as a mere number uninterrupt-*
ating the place of the year in the scale of time. Thus, in
an. 5, B. C. 1, Jan. 5 does not mean that 5 days of January in
question have elapsed, but that 4 have elapsed, and the 5th is
nd the B. C. 1, indicates that *the first day* of the year so
first year current before Christ,) preceded the first day of the
by one year. The scale of A. D. and B. C. is not continuous,

Table 1. Multiples of 1461, the days in a Julian *Quadriennium*.

1	1461	4	5844	7	10227
2	2922	5	7305	8	11688
3	4383	6	8766	9	13149

Table 2. Days in Residual years.

0	0
1	366
2	731
3	1096

Table 3. Days elapsed from Jan. 1 to the 1st of each Month.

	In a common year.	In a leap year.		In a common year.	In a leap year.
Jan. 1	0	0	July 1	181	182
Feb. 1	31	31	Aug. 1	212	213
March 1	59	60	Sept. 1	243	244
April 1	90	91	Oct. 1	273	274
May 1	120	121	Nov. 1	304	305
June 1	151	152	Dec. 1	334	335

Example.—What is the current day of the Julian period corresponding to the last day of Old Style in England, on Sept. 2, A. D. 1752?

1752
4713
6465 year current.
1
4)6464 years elapsed.
Q = 1616
R = 0

1000	1,461,000
600	876,600
10	14.610
6	8,766
R = 0	0
Jan. 1 to Sept. 1,	244
Sept. 1 to Sept. 2,	1
	2,361,221 days elapsed.
Current day the	2,361,222d.

(930.) To find the same for any given date, New Style. Proceed as above, considering the date as a Julian date, and disregarding the change of style. Then, from the resulting days, subtract as follows:—

For any date of New Style, antecedent to March 1, A. D. 1700........ 10 days.
After Feb. 28, 1700, and before March 1, A. D. 1800.................. 11 days.
" 1800 " " 1900.................. 12 days.
" 1900 " " 2100.................. 13 days, &c.

(931.) *To find the interval between any two dates, whether of Old or New Style, or one of one, and one of the other.* Find the day current of the Julian period corresponding to each date, and their difference is the interval required. If the dates contain hours, minutes, and seconds, they must be annexed to their respective days current, and the subtraction performed as usual.

(932.) The Julian rule made every fourth year, without exception, a bissextile. This is, in fact, an over-correction; it supposes the length of the tropical year to be $365\frac{1}{4}^d$, which is too great, and thereby induces an ays in 908 years, as will easily appear on trial. Accordingly, the year 1414, it began to be perceived that the equinoxes ally creeping away from the 21st of March and September, ought to have always fallen had the Julian year been exact, ing (as it appeared) too early. The necessity of a fresh and orm in the calendar was from that time continually urged, and dmitted. The change (which took place under the popedom XIII.) consisted in the omission of ten[1] nominal days after October, 1582, (so that the next day was called the 15th, and ,) and the promulgation of the rule already explained for future The change was adopted immediately in all Catholic coun- nore slowly in Protestant. In England, "the change of style," alled, took place after the 2d of September, 1752, eleven no- being then struck out; so that, the last day of Old Style being first of New Style (the next day) was called the 14th, instead The same legislative enactment which established the Gre- in England in 1752, shortened the preceding year, 1751, by er. Previous to that time, the year was held to begin with March, and the year A. D. 1751 did so accordingly; but that woven with the account given by an ancient author of some historical event, as to indicate precisely the interval of time between the eclipse and the event, and at the same time completely to identify the eclipse, that date is recovered and fixed for ever.[1]

(934.) The days thus parcelled out into years, the next step to a perfect knowledge of time is to secure the identification of each day, by imposing on it a name universally known and employed. Since, however, the days of a whole year are too numerous to admit of loading the memory with distinct names for each, all nations have felt the necessity of breaking them down into parcels of a more moderate extent; giving names to each of these parcels, and particularizing the days in each by numbers, or by some special indication. The lunar month has been resorted to in many instances; and some nations have, in fact, preferred a lunar to a solar chronology altogether, as the Turks and Jews continue to do to this day, making the year consist of 12 lunar months, or 354 days. Our own division into twelve unequal months is entirely arbitrary, and often productive of confusion, owing to the equivoque between the lunar and calendar month.[2] The intercalary day naturally attaches itself to February as the shortest.

(935.) Astronomical time reckons from the noon of the current day, civil from the preceding midnight, so that the two dates coincide only during the earlier half of the astronomical, and the later of the civil day. This is an inconvenience which might be remedied by shifting the astronomical epoch to coincidence with the civil. There is, however, another inconvenience, and a very serious one, to which both are liable, inherent in the nature of the day itself, which is a local phænomenon, and commences at different instants of absolute time, under different meridians, whether we reckon from noon, midnight, sunrise, or sunset. In consequence all astronomical observations require, in addition to their date, to render them comparable with each other, the longitude of the place of observation from some meridian, commonly respected by all astronomers. For geographical longitudes, the Isle of Ferroe has been chosen by some as a common meridian, indifferent (and on that very account offensive) to all nations. Were astronomers to follow such an example, they would probably fix upon Alexandria, as that to which Ptolemy's observations and computations were reduced, and as claiming on that account the respect of all while offending the national egotism of none. But even this will not meet the whole difficulty. It will still remain doubtful, on a meridian 180° remote from that of Alexandria, what day is intended by any given date. Do what we will, when it is the Monday, the 1st of January, 1849, in one part of the world, it will be Sunday, the 31st of December, 1848, in another, so long as time is reckoned by local hours. This equivoque, and the necessity of specifying the geographical locality as an element of the date, can only be got over by a reckoning of time which refers itself to some event, real or imaginary, common to all the globe. Such an event is the passage of the sun through the vernal equinox, or rather the passage of an imaginary sun, supposed to move with perfect equality, through a vernal equinox supposed free from the inequalities of nutation, and receding upon the ecliptic with *perfect* uniformity. The actual equinox is variable, not only by the effect of nutation, but by that of the inequality of precession, resulting from the change in the plane of the ecliptic due to planetary perturbation. Both variations are, however, periodical; the one in the short period of 19 years, the other in a period of enormous length, hitherto uncalculated, and whose maximum of fluctuation is also unknown. This would appear, at first sight, to render impracticable the attempt to obtain from the sun's motion any rigorously uniform measure of time. A little consideration, however, will satisfy us that such is not the case. The solar tables, by which the apparent place of the sun in the heavens is represented with almost absolute precision from the earliest ages to the present time, are constructed upon the supposition that a certain angle, which is called "the sun's mean longitude," (and which is, in effect, the sum of the mean sidereal motion of the sun, *plus* the mean sidereal motion of the equinox in the opposite direction, *as near as it can be obtained* from the accumulated observations of twenty-five centuries,) increases with rigorous uniformity as time advances. The conversion of this mean longitude into time at the rate of 360° to the mean tropical year, (such as the tables assume it,) will therefore give us both the unit of time, and the uniform measure of its lapse which we seek. It will also furnish us with an epoch, not indeed marked by any real event, but not on that account the less positively fixed, being connected, through the medium of the tables, with every single observation of the sun on which they have been constructed and with which compared.

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[1] See the remarkable calculations of Mr. Baily relative to the celebrated solar eclipse which put an end to the battle between the kings of Media and Lydia, B. C. 610, Sept. 30. Phil. Trans. ci. 220.

[2] "A month in law is a lunar month or twenty-eight days, (! !) unless otherwise expressed."—*Blackstone*, ii, chap. 9. "A lease for twelve months is only for forty-eight weeks."—*Ibid.*

(936.) Such is the simplest abstract conception of equinoctial time. It is the mean longitude of the sun of *some one approved set of solar* tables, converted into time at the rate of 360° to the tropical year. Its unit is the mean tropical year which those tables assume *and no other*, and its epoch is the mean vernal equinox of these tables for the current year, or the instant when the mean longitude of the tables is rigorously 0, according to the assumed mean motion of the sun and equinox, the assumed epoch of mean longitude, and the assumed equinoctial point on which the tables have been computed, *and no other*. To give complete effect to this idea, it only remains to specify the particular tables fixed upon for the purpose, which ought to be of great and admitted excellence, since, once decided on, the very essence of the conception is that *no subsequent alteration in any respect should be made, even when the continual progress of astronomical science shall have shown any one or all of the elements concerned to be in some minute degree erroneous* (as necessarily they must,) *and shall have even ascertained the corrections they require* (to be themselves again corrected, when another step in refinement shall have been made.)

(937.) Delambre's solar tables (in 1828) when this mode of reckoning time was first introduced, appeared entitled to this distinction. According to these tables, the sun's mean longitude was 0°, or the mean vernal equinox occurred, in the year 1828, on the 22d of March at $1^h\ 2^m\ 59^s{\cdot}05$ mean time at Greenwich, and therefore at $1^h\ 12^m\ 20^s{\cdot}55$ mean time at Paris, or $1^h\ 56^m\ 34^s{\cdot}55$ mean time at Berlin, at which instant, therefore, the equinoctial time was $0^d\ 0^h\ 0^m\ 0^s{\cdot}00$, being the commencement of the 1828th year current of equinoctial time, if we choose to date from the mean tabular equinox, nearest to the vulgar era, or of the 6541st year of the Julian period, if we prefer that of the first year of that period.

(938.) Equinoctial time then dates from the mean vernal equinox of Delambre's solar tables, and its unit is the mean tropical year of these tables ($365^d{\cdot}242264$.) Hence, having the fractional part of a day expressing the difference between the mean local time at any place (suppose Greenwich) on any one day between two consecutive mean vernal equinoxes, that difference will be the same for every other day in the same interval. Thus, between the mean equinoxes of 1828 and 1829, the difference between equinoctial and Greenwich time is $0^d{\cdot}956261$ or $0^d\ 22^h\ 57^m\ 0^s{\cdot}95$, which expresses the equinoctial day, hour, minute, and second, corresponding to mean noon at Greenwich on March 23, 1828, and for the noons of the 24th, 25th, &c., we have only to substitute 1d, 2d, &c., for 0^d, retaining the same decimals of a day, or the same hours, minutes, &c., up to and including March 22, 1829. Between Greenwich noon of the 22d and 23d of March, 1829, the 1828th equinoctial year terminates, and the 1829th commences. This happens at $0^d{\cdot}286003$, or at $6^h\ 51^m\ 50^s{\cdot}66$ Greenwich mean time, after which hour, and until the

the Greenwich hour added to equinoctial time 364^{d}·956261 t to more than 365·242264, a complete year, which has there- subtracted to get the equinoctial date in the next year, cor- to the Greenwich time. For example, at $12^{h}\ 0^{m}\ 0^{s}$ Greenwich or 0^{d}·500000, the equinoctial time will be 364·956261+ =365·456261, which being greater than 365·242264, shows quinoctial year current has changed, and the latter number racted, we get 0^{d}·213977 for the equinoctial time of the r current corresponding to March 22, 12^{h} Greenwich mean

Having, therefore, the fractional part of a day for any one year the equinoctial hour, &c., at the mean noon of any given place, cceeding years will be had by subtracting 0^{d}·242264, and its from such fractional part (increased if necessary by unity,) and ng years by adding them. Thus, having found 0·198525 for al part for 1827, we find for the fractional parts for succeeding 1853 as follows[1]:—

·956261	1835	·260413	1842	·564565	1848	·110981
·713997	1836	·018149	1843	·322301	1849	·868717
·471733	1837	·775885	1844	·080037	1850	·626453

APPENDIX.

I. LISTS OF NORTHERN AND SOUTHERN STARS, WITH THEIR APPROXIMATE MAGNITUDES, ON THE VULGAR AND PHOTOMETRIC SCALE.

1. NORTHERN STARS.

STAR.	MAG.		STAR.	MAG.		STAR.	MAG.		STAR.	MAG.	
	Vulg.	Phot.		Vulg.	Phot.		Vulg.	Phot.		Vulg.	Phot.
Arcturus	0·77	1·18	γ Cassiopeiæ	2·52	2·93	η Draconis	3·02	3·43	ε Aurigæ (Var.)	3·37	3·78
Capella	1·0 :	1·4 :	α Andromedæ	2·54	2·95	β Draconis	3·06	3·47	γ Lyncis	3·39	3·80
Lyra	1·0 :	1·4 :	α Cassiopeiæ	2·57	2·98	β Arietis	3·09	3·50	ζ Draconis	3·40	3·81
Procyon	1·0 :	1·4 :	γ Geminorum	2·59	3·00	γ Pegasi	3·11	3·52	π Herculis	3·41	3·82
α Orionis	1·0 :	1·43	Algol (Var.)	2·62	3·03	ε Virginis?	3·14	3·55	β Canis Min.?	3·41	3·82
Aldebaran	1·1 :	1·5 :	ε Pegasi	2·62	3·03	θ Aurigæ	3·17	3·58	ζ Tauri	3·42	3·83
α Aquilæ	1·28	1·69	γ Draconis	2·62	3·03	β Herculis	3·18	3·59	δ Draconis	3·42	3·83
Pollux	1·6 :	2·0 :	β Leonis	2·63	3·04	α Canum Ven.	3·22	3·63	μ Geminorum	3·42	3·83
Regulus	1·6 :	2·0 :	α Ophiuchi	2·63	3·04	β Ophiuchi	3·23	3·64	γ Boötis	3·43	3·84
α Cygni	1·90	2·31	β Cassiopeiæ	2·63	3·04	δ Cygni	3·24	3·65	ε Geminorum	3·43	3·84
Castor	1·94	2·35	γ Cygni	2·63	3·04	ε Persei	3·26	3·67	δ Herculis	3·44	3·85
ε Ursæ (Var.)	1·95	2·36	α Pegasi	2·65	3·06	η Tauri?	3·26	3·67	δ Geminorum	3·44	3·85
α Ursæ (Var.)	1·96	2·36	β Pegasi	2·65	3·06	ζ Persei	3·27	3·68	q Orionis	3·45	3·86
α Persei	2·07	2·48	α Coronæ	2·69	3·10	ζ Herculis	3·28	3·69	β Cephei	3·45	3·86
η Ursæ (Var.)	2·18	2·59	γ Ursæ	2·71	3·12	ι Aurigæ	3·29	3·70	θ Ursæ	3·45	3·86
γ Orionis	2·18	2·59	β Ursæ	2·77	3·18	γ Ursæ Minor	3·30	3·71	ι Ursæ	3·46	3·87
β Tauri	2·28	2·69	ε Boötis	2·80	3·21	η Pegasi	3·31	3·72	η Aurigæ	3·46	3·87
Polaris	2·28	2·69	ε Cygni	2·88	3·29	ζ Aquilæ	3·32	3·73	γ Lyræ	3·47	3·88
γ Leonis	2·34	2·75	α Cephei	2·90	3·31	β Cygni	3·33	3·74	η Geminorum	3·48	3·89
α Arietis	2·40	2·81	α Serpentis	2·92	3·33	γ Persei	3·34	3·75	γ Cephei	3·48	3·89
ζ Ursæ	2·43	2·84	δ Leonis	2·94	3·35	μ Ursæ	3·35	3·76	κ Ursæ	3·49	3·90
β Andromedæ	2·45	2·86	γ Aquilæ	2·98	3·39	β Triang. Bor.	3·35	3·76	ε Cassiopeiæ	3·49	3·90
β Aurigæ	2·48	2·89	δ Cassiopeiæ	2·99	3·40	δ Persei	3·36	3·77	θ Aquilæ	3·50	3·91
γ Andromedæ	2·50	2·91	η Boötis	3·01	3·42	ψ Ursæ	3·36	3·77			

STAR.	MAG.		STAR.	MAG.	
	Vulg.	Phot.		Vulg.	Phot.
corpii	2·91	3·32	ϑ Corvi	3·22	3·63
ntauri	2·91	3·32	β Eridani	3·26	3·67
rgûs.............	2·94	3·35	θ Argûs...............	3·26	3·67
rvi	2·95	3·36	β Hydri...............	3·27	3·68
corpii............	2·96	3·37	ε Corvi	3·28	3·69
entauri	2·96	3·37	β Aræ.................	3·31	3·72
phiuchi..........	2·97	3·38	α Toucani...........	3·32	3·73
quarii............	2·97	3·38	β Capricorni	3·32	3·73
rgûs..............	2·98	3·39	ρ Argûs...............	3·32	3·73
entauri	2·99	3·40	π Scorpii............	3·35	3·76
eporis	3·00	3·41	β Leporis	3·35	3·76
phiuchi..........	3·00	3·41	γ Lupi................	3·36	3·77
gittarii	3·01	3·42	υ Scorpii	3·37	3·78
phiuchi..........	3·05	3·46	ι Orionis............	3·37	3·78
ibræ..............	3·07	3·48	α Aræ.................	3·40	3·81
irginis	3·08	3·49	π Sagittarii..........	3·40	3·81
rgûs	3·08	3·49	α Muscæ.............	3·43	3·84
gittarii..........	3·11	3·52	α Hydri ?............	3·44	3·85
ibræ..............	3·12	3·53	τ Scorpii	3·44	3·85
gittarii	3·13	3·54	ζ Hydræ	3·45	3·86
pi	3·14	3·55	γ Hydræ	3·46	3·87
olumbæ..........	3·15	3·56	β Trianguli..........	3·46	3·87
entauri..........	3·20	3·61	σ Scorpii.............	3·50	3·91
pricorni	3·20	3·61	τ Argûs..............	3·50	3·91

II. SYNOPTIC TABLE OF THE ELEMENTS OF THE PLANETARY SYSTEM.

Name of Body.	Mean Distance from the Sun or Semi-axis.	Mean Sidereal Period in Mean Solar Days.	Eccentricity in parts of the Semi-axis.	Inclination of Orbit to the Ecliptic.			Longitude of the Ascending Node.		
				°	′	″	°	′	″
Sun									
Mercury	0·3870981	87·9692580	0·2055149	7	0	9·1	45	57	30·9
Venus.........	0·7233316	224·7007869	0·0068607	3	23	28·5	74	54	12·9
Earth	1·0000000	365·2563612	0·0167836						
Mars	1·5236923	686·9796458	0·0933070	1	51	6·2	48	0	3·5
Flora..........	2·2016870	1193·249	0·1565570	5	53	4·8	110	18	12·0
Victoria......	2·3348770	1303·148	0·2179220	8	22	15·3	235	28	25·3
Vesta	2·3610810	1325·147	0·0895694	7	8	29·7	103	23	31·6
Iris............	2·3850809	1345·403	0·2322236	5	28	14·0	259	42	47·6
Metis..........	2·3868070	1346·940	0·1228221	5	35	54·7	68	28	57·7
Hebe..........	2·4261120	1380·184	0·2005582	14	47	1·8	138	30	32·5
Parthenope .	2·4465170	1397·719	0·0963995	4	36	44·1	125	1	32·1
Astræa........	2·5770470	1511·095	0·1880586	5	19	22·7	141	25	14·6
Egeria	2·5783650	1512·217	0·1115620	15	55	44·9	43	37	7·1
Irene..........	2·5849050	1517·975	0·1678530	9	5	47·0	86	51	28·7
Juno	2·6708370	1594·296	0·2548847	13	3	22·1	170	54	45·6
Ceres..........	2·7680510	1682·125	0·0766523	10	37	4·4	80	48	46·6
Pallas	2·7728580	1686·510	0·2398150	34	37	33·1	172	43	59·7
Hygeia	3·1208900	2020·761	0·0892461	3	47	14·2	287	52	25·2
Jupiter	5·2027760	4332·5848212	0·0481621	1	18	51·3	98	26	18·9
Saturn	9·5387861	10759·2198174	0·0561505	2	29	35·7	111	56	37·4
Uranus	19·1823900	30686·8208296	0·0466794	0	46	28·4	72	59	35·3
Neptune	30·0368000	60126·7100000	0·0087195	1	46	59·0	130	5	11·0

Elements in M. T. nwich (G). in (B).	Mass (denominator of fraction, the Sun being 1.)	Diameter in Miles.	Density.	Time of Rotation on Axis.	
				h.	m.
....	1	882000	0·25	607	48
an. 1. 0^h G.	4865751	3140	1·12	24	5
Do.	401839	7800	0·92	23	21
Do.	389551	7926	1·00	24	0
Do.	2680337	4100	0·95	24	37
an. 1. 0^h B.					
ct. 0. 0^h Par.					
an. 9. 0^h B.	....	250?			
lar. 31. 0^h B.					
une 4. 0^h B.					
uly 10. 0^h B.					
lay 11. 0^h G.					
an. 1. 0^h B.					
ov. 2. 0^h B.					
uly 1. 10^h B.					
pr. 8. 0^h B.	....	79?	......	27	0?
ept. 25. 0^h B.	....	163?			
ug. 23. 0^h B.					
une 6. 0^h G.					
an. 1. 0^h G.	1047·871	87000	0·24	9	56
Do.	3501·600	79160	0·14	10	29
Do.	24905	34500	0·24	9	30?
an. 1. 0^h G.	18780	41500	0·14		

Note.—The elements of the orbits of Mercury, Venus, the Earth, Mars, Jupiter, Saturn, and Uranus, are those given by the late F. Baily, Esq., in his "Astronomical Tables and Formulæ," and are the same with those which form the basis of Delambre's tables, embodying the formulæ of Laplace. The elements of Uranus and Neptune can only be regarded as provisional; those of the former requiring considerable corrections, necessitated by the discovery of Neptune, but which, not being yet finally ascertained, by reason of the uncertainty still attending on the mass and elements of the latter planet, it was thought better to leave the old elements untouched than to give an imperfect rectification of them. The masses of the planets are those most recently adopted by Encke (Ast. Nachr. No. 443), on mature consideration of all the authorities, that of Neptune excepted, which is Prof. Pierce's determination from Bond's and Lassell's observation of the satellite discovered by the latter. The densities are Hansen's (A. N. 443.)

The elements of Vesta, Juno, Ceres, and Pallas, are the osculating elements for 1850, computed by Encke (A. N. 636.) [Those of Flora are from the computations of Brunnow (A. N. 645); of Victoria, Villarceaux (A. N. 741); of Iris, Schubert (A. N. 730); of Metis, Wolfers (A. N. 764); of Hebe, Luther (A. N. 721); of Parthenope, Galen (A. N. 757); of Astræa, D'Arrest (A. N. 626); of Egeria, D'Arrest (A. N. 749); of Irene, Vogel and Rümker (A. N. 765); and of Hygeia, Santini (A. N. 702.)

Of these last-mentioned small planets, Hygeia, Parthenope, and Egeria were discovered by Dr. Gasparis, at Naples, on April 12, 1849, May 11 and Nov. 2, 1850, respectively; Iris, Flora, Victoria, and Irene, by Mr. Hind, on Aug. 13 and Oct. 18, 1847, Sept. 13, 1850, and May 19, 1851, respectively. The elements of the recently-discovered small planets may undergo material corrections from further observation. Irene has a blue colour and a faint nebulous envelope. The orbits of Astræa and Hygeia approach at one point (their common node) within 0·006 of the radius of the earth's orbit. It will not be long before the planets themselves come within that proximity to each other (A. N. 752.) Victoria and Astræa are subject to variations of brightness, which indicate rotations on their axes, and dark spots (A. N. 760.) D'Arrest (A. N. 752) remarks that a relation subsists between the excentricities of the orbits of the small planets, and the inclinations of the planes in which they lie to the sun's equator, the more excentric orbits being the more inclined. While these sheets pass through the press, another, yet unnamed, is announced by M. de Gasparis.]

III.

SYNOPTIC TABLE OF THE ELEMENTS OF THE ORBITS OF THE SATELLITES, SO FAR AS THEY ARE KNOWN.

1. THE MOON.

Mean distance from earth	59^{r}·96435000
Mean sidereal revolution	27^{d}·321661418
Mean synodical ditto	29^{d}·530588715
Excentricity of orbit	0·054844200
Mean revolution of nodes	6793^{d}·391080
Mean revolution of apogee	3232^{d}·575343
Mean longitude of node at epoch	13° 53′ 17″·7
Mean longitude of perigee at do.	266 10 7 ·5
Mean inclination of orbit	5 8 47 ·9
Mean longitude of moon at epoch	118 17 8 ·3
Mass, that of earth being 1,	0·011399
Diameter in miles	2153
Density, that of the earth being 1,	0·5657

[1] The distances are expressed in equatorial radii of the primaries. The epoch is Jan. 1, 1801, unless otherwise expressed. The periods, &c. are expressed in mean solar days.

Plane to Jupiter's Equator.	of Nodes on the fixed Plane.	being 1,000,000,000.
° ′ ″	Years.	
0 0 6	....	17328
0 1 5	29·9142	23235
0 5 2	141·7390	88497
0 24 4	531·0000	42659

he 3d and 4th small, but variable, in consequence of their mutual

Saturn.

Ele- s.	Mean Longitude at the Epoch.	Excentricity.	Perisaturnium.
	° ′ ″		° ′ ″
·0	256 58 48		
·0	67 41 36		
	313 43 48	0·04 ?	54 ?
	327 40 48	0·02 ?	42 ?
	353 44 0	0·02 ?	95 ?
·0	137 21 24	0·029314	256 38 11
	... ...		
·0	269 37 48		

node with the ecliptic. The first seven satellites move in, or very
y intermediate between the planes of the ring and of the planet's
ngitude (on the ecliptic).
me night (Sept. 19, 1848), by Mr. Lassell, of Liverpool, and Prof.
han conjecture. Messrs. Kater, Encke, and Lassell agree in repre-
s the broad black divisions which ordinary telescopes show.

4. SATELLITES OF URANUS.

Sat.	Sidereal Revolution.				Mean distance.	Epochs of Passage and most ascending Node of Orbits. G. T.	Nodes and Inclinations.
	d	h	m	s			
1	4	?					The orbits are inclined at an angle of about 78° 58′ to the ecliptic, in a plane whose ascending node is in long. 165° 30′ (Equinox of 1798). Their motion is *retrograde*. The orbits are nearly circular.
2	8	16	56	31·3	17·0	1787. Feb. 16th. 0^h 10^m	
3	10	23	?		19·8 ?		
4	13	11	7	12·6	22·8	1787. Jan. 7th. 0^h 28^m	
5	38	2	?		45·5 ?		
6	107	12	?		91·0		

5. SATELLITES OF NEPTUNE.

One only has been *certainly* observed,—its approximate period being 5^d 20^h 50^m 45^s,—distance about 12 radii of the planet.

AT THEIR LAST APPEARANCE.

Faye's.	De Vico's.	Brorsen's.
1843. Oct. 17.	1844. Sept. 2.	1846. Feb. 25.
$3^h\ 42^m\ 16^s$	$11^h\ 36^m\ 53^s$	$9^h\ 13^m\ 35^s$
19° 34′ 19″	342° 31′ 15″	116° 28′ 34″
209 29 19	63 49 31	102 39 36
11 22 31	2 54 45	30 55 7
3·81179	3·09946	3·15021
0·555962	0·617256	0·793629
2718^d·26	1993^d·09	2042·24
Direct.	Direct.	Direct.

d Ω that of the ascending node for the epoch of the perihelion;
period in days.

s up to June, 1847, by *all* their several computors, in Prof. Encke's
die Bahn eines Cometen zu berechnen." The list is compiled by
of these orbits we collect the following, as a more correct statement
10° inclination, 3 out of 15; under 20°, 9 out of 29. Retrograde
In such orbits, of all inclinations from 0 to 90°, 11 out of 37.
the enlarged field of comparison.

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THE END.

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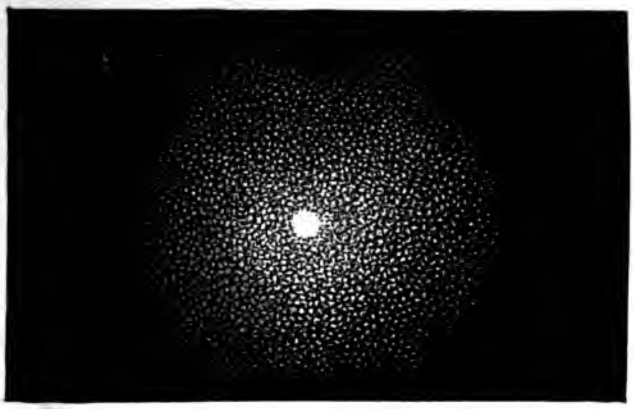
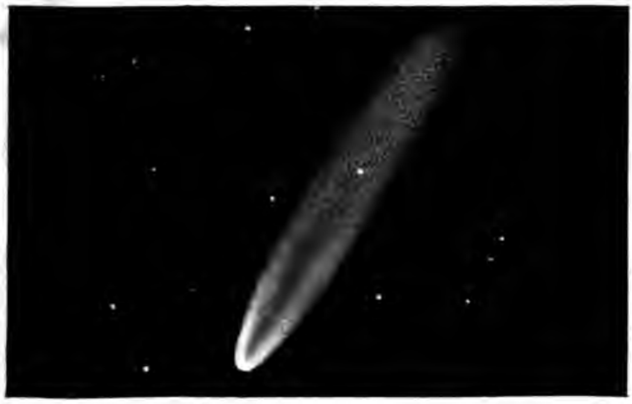


Fig. I



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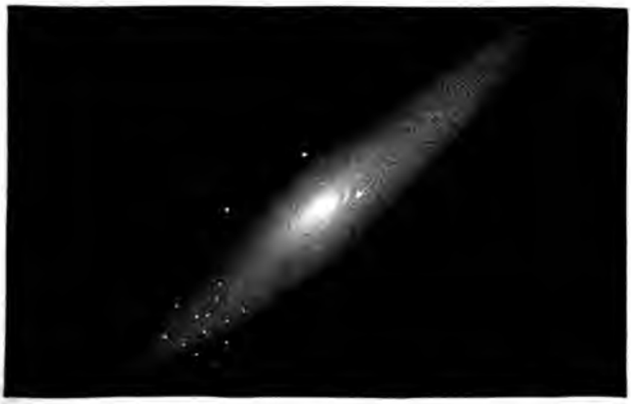
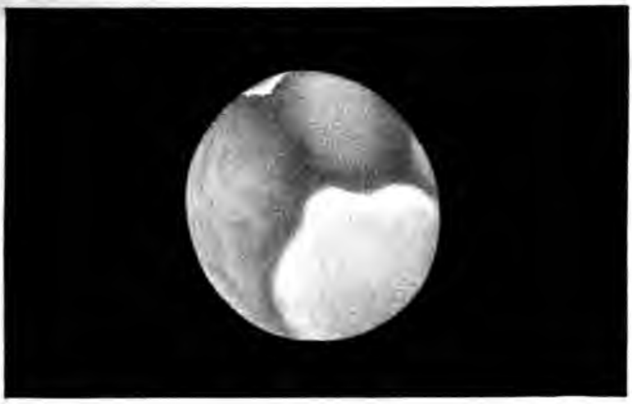


Fig. 3

T. Sinclair's lith.

Plate III



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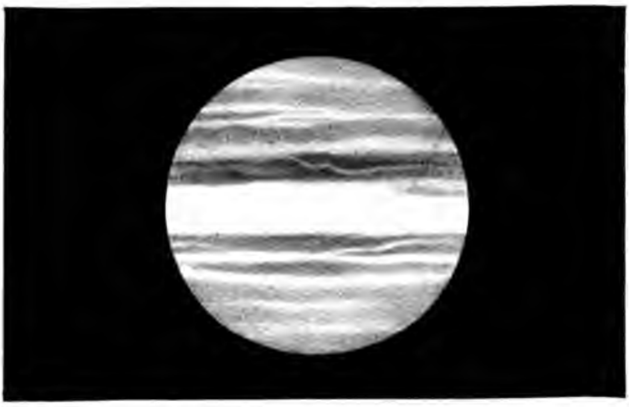

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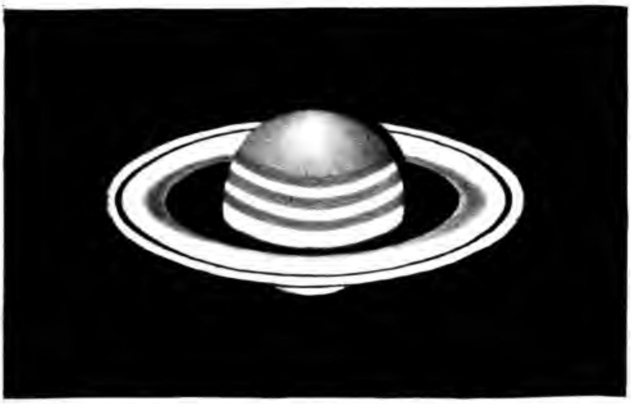
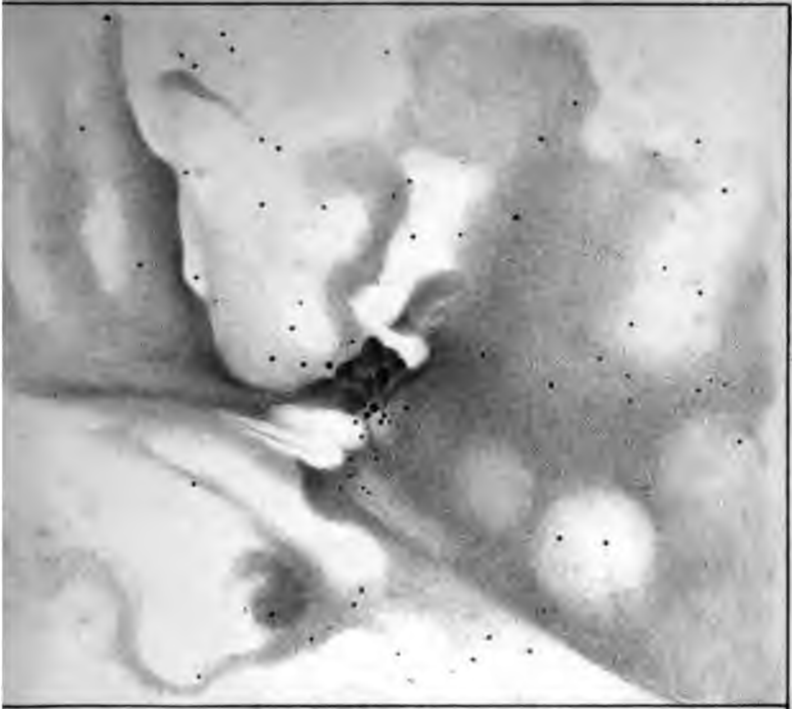


Fig 3

T. Jenkinson lith.

Plate IV.



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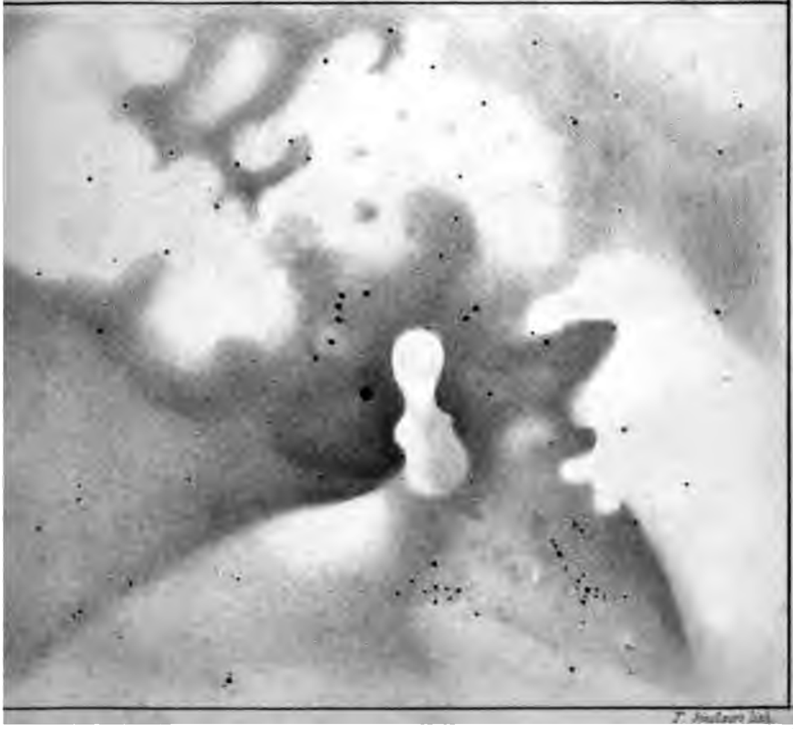

Fig.

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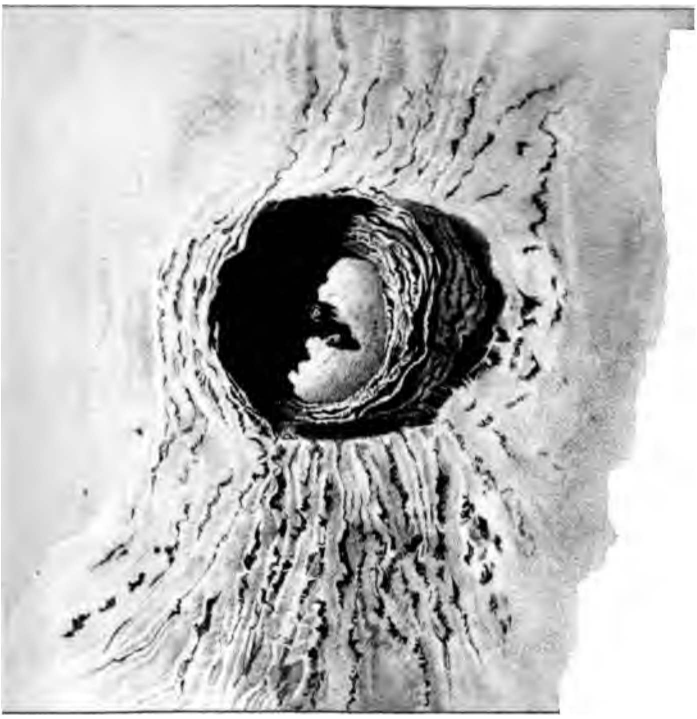

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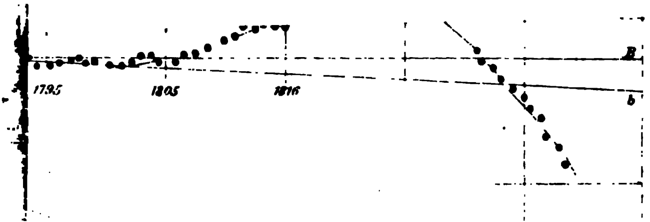

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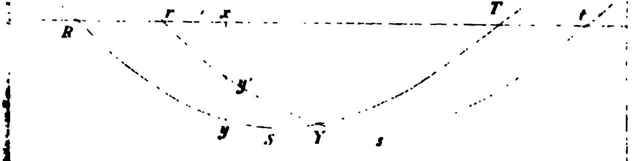

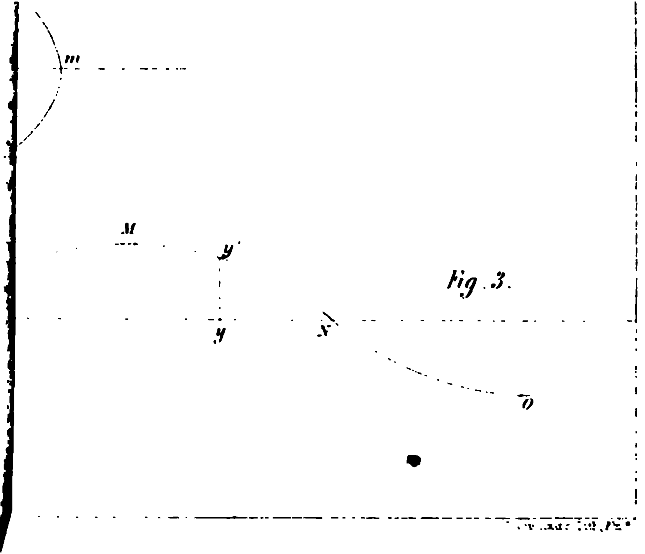

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